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▶ **To cite this version:**

Wencheng D. Shao, Xi Zhang, João Mendonça, Thérèse Encrenaz. Local-time Dependence of Chemical Species in the Venusian Mesosphere. The Planetary Science Journal, 2022, 3, 10.3847/PSJ/ac3bd3 . insu-03713342

HAL Id: insu-03713342

https://insu.hal.science/insu-03713342

Submitted on 4 Jul 2022

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The Planetary Science Journal, 3:3 (17pp), 2022 January https://doi.org/10.3847/PSJ/ac3bd3
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Local-time Dependence of Chemical Species in the Venusian Mesosphere

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Received 2021 February 7; revised 2021 November 19; accepted 2021 November 19; published 2022 January 12

Abstract

Observed chemical species in the Venusian mesosphere show local-time variabilities. SO_2 at the cloud top exhibits two local maxima over local time, H_2O at the cloud top is uniformly distributed, and CO in the upper atmosphere shows a statistical difference between the two terminators. In this study, we investigated these local-time variabilities using a three-dimensional (3D) general circulation model (GCM) in combination with a two-dimensional (2D) chemical transport model (CTM). Our simulation results agree with the observed local-time patterns of SO_2, H_2O, and CO. The two-maximum pattern of SO_2 at the cloud top is caused by the superposition of the semidiurnal thermal tide and the retrograde superrotating zonal (RSZ) flow. SO_2 above 85 km shows a large day–night difference resulting from both photochemistry and the subsolar-to-antisolar (SS-AS) circulation. The transition from the RSZ flows to SS-AS circulation can explain the CO difference between two terminators and the displacement of the CO local-time maximum with respect to the antisolar point. H_2O is long-lived and exhibits very uniform distribution over space. We also present the local-time variations of HCl, ClO, OCS, and SO simulated by our model and compare to the sparse observations of these species. This study highlights the importance of multidimensional CTMs for understanding the interaction between chemistry and dynamics in the Venusian mesosphere.

Unified Astronomy Thesaurus concepts: Venus (1763); Planetary atmospheres (1244); Atmospheric variability (2119); Atmospheric composition (2120); Atmospheric circulation (112); Theoretical models (2107)

1. Motivation

The Venusian mesosphere (∼70–100 km) is characterized by complex photochemistry (e.g., Yung & DeMore 1982; Mills 1998; Zhang et al. 2012; Krasnopolsky 2012). Species like SO_2 and CO—fundamental components of the photochemical network—have shown significant spatial and temporal variabilities. For example, Venus Express detected that the SO_2 mixing ratio at 70–80 km varies by orders of magnitude over time and space (Vandaele et al. 2017a, 2017b). Ground-based observations by TEXES/IRTF (Texas Echelon Cross Echelle Spectrograph/Infrared Telescope Facility) show that SO_2 around the cloud top exhibits plumes and patchy features over the Venus disk (Encrenaz et al. 2012, 2013, 2016, 2019, 2020). Venus Express also observed that CO has strong short-term variabilities up to one order of magnitude (Vandaele et al. 2016).

Some chemical species in the Venusian mesosphere show strong local-time variabilities. Encrenaz et al. (2020) extracted the local-time dependence of SO_2 at ∼64 km from TEXES/IRTF and found that the SO_2 mixing ratio generally exhibits two local maxima around the morning and evening terminators, respectively. SPICAV (SPectroscopy for the Investigation of the Characteristics of the Atmosphere of Venus) on board Venus Express also observed a similar SO_2 local-time pattern at the cloud top on the dayside (Vandaele et al. 2017a; Encrenaz et al. 2019; Marcq et al. 2020). Sandor et al. (2010) used microwave spectra to obtain day–night differences of SO_2 and SO at 70–100 km. Despite the scarcity of the data, SO_2 appears more abundant at night than during the day, while SO is likely to have a reversed day–night difference. Belyaev et al. (2017) observed that midnight SO_2 abundance appears 3–4 times higher than at the terminators around 95 km through SPICAV occultations. For CO, Clancy & Muhleman (1985) observed a day–night difference from microwave measurements. They found that the CO bulge (i.e., the local maximum of CO mixing ratio) shifts from midnight to the morning as altitude decreases from above ∼95 km to 80–90 km. Clancy et al. (2003) showed similar CO patterns in subsequent microwave observations. Vandaele et al. (2016) summarized the CO data observed by SOIR (Solar Occultation in the InfraRed) on board Venus Express and found a statistical difference between the morning and evening terminators, and the difference also depends on altitude. Compared to SO_2 and CO, H_2O seems to vary insignificantly with local time. Encrenaz et al. (2012, 2013, 2016, 2019, 2020), using TEXES, observed that the H_2O mixing ratio at ∼64 km, obtained from the HDO spectra, distributes uniformly over the Venus disk. Chamberlain et al. (2020) showed that the H_2O profiles above 80 km observed by SOIR do not exhibit dependence on terminators. Sandor & Clancy (2012, 2017), using JCMT (James Clerk Maxwell Telescope), observed that the HCl mixing ratio above 85 km exhibits no evident day–night difference. Krasnopolsky (2010), using the CSHELL spectrograph at NASA IRTF, observed that the morning OCS can be more abundant than the afternoon OCS.

The origin of these local-time variabilities has not been thoroughly investigated but likely relates to atmospheric chemistry and dynamics. Intense photochemistry occurs in the Venusian mesosphere (e.g., Zhang et al. 2012), in which the dependence of solar irradiance on local time affects the local distribution of chemical species. On Venus, the cloud region (∼47–70 km) is characterized by a retrograde superrotating

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zonal (RSZ) flow (e.g., Sánchez-Lavega et al. 2008; Lebonnois et al. 2010; Mendonça & Read 2016; Mendonça & Buchhave 2020). In the thermosphere (>110 km), strong evidence shows a subsolar-to-antisolar (SS-AS) circulation pattern (e.g., Bougher et al. 2006). The upper mesosphere (90–110 km) might be a region where SS-AS circulation is superimposed on the RSZ flow (e.g., Lellouch et al. 1994). Besides, thermal tides excited by the solar heating also strongly perturb the temperature and winds in the mesosphere (e.g., Taylor et al. 1980; Limaye 2007; Fukuya et al. 2021). These dynamical flow patterns transport chemical species and modulate their local-time variabilities.

A few theoretical studies have investigated the local-time variabilities of chemical species. Jessup et al. (2015) studied spatial variations of SO_2 and SO observed by the Space Telescope Imaging Spectrograph on the Hubble Space Telescope. They showed through one-dimensional (1D) photochemical models that solar zenith angle could significantly affect the SO_2 variability. Gilli et al. (2017) presented CO and O density profiles in the upper atmosphere at different local times using a three-dimensional (3D) general circulation model (GCM) coupling chemistry and dynamics (Stolzenbach et al. 2015; Stolzenbach 2016). Their results indicate the importance of the SS-AS circulation in the CO and O distributions. Navarro et al. (2021) and Gilli et al. (2021) used an improved GCM to study CO's spatial variabilities in the upper atmosphere. Their simulated CO pattern shows a CO bulge shift toward the morning by 2–3 hr in the mesosphere, attributed to a weak westward retrograde wind. However, a dedicated study of the local-time variability of SO_2 is still lacking, and mechanisms controlling CO's local-time distributions need further investigation.

As a preliminary step toward fully understanding the spatial and temporal variabilities of chemical species in the Venusian atmosphere, in this study we investigate the local-time dependence of multiple chemical species including SO_2 and CO using a 3D dynamical model in combination with a two-dimensional (2D; longitude–pressure) chemical model with the state-of-the-art full photochemical network on Venus. Our simulated local-time distributions of species like SO_2, CO, and H_2O show agreement with observations. We explore underlying mechanisms determining species' local-time distributions and find that the relative importance of dynamics and chemistry depends on altitude and species. Our study indicates that the local-time distributions of SO_2 and CO can constrain important dynamical patterns in the Venusian atmosphere.

This paper is structured as follows. Section 2 and Appendix A provide technical details of our models. In Section 3, we present simulations from our nominal case and study the local-time dependence of SO_2, CO, and other species. In Section 4, we do sensitivity tests and discuss the influences of our models' parameters and resolution on our results. Finally, we conclude our results and discuss future work in Section 5.

2. Methodology

We use a 3D GCM in combination with a 2D chemical transport model (CTM) to study the local-time variabilities of chemical species in the Venusian mesosphere. We adopt this combination method because fully coupling chemistry with dynamics in the GCM or utilizing a 3D CTM with a full chemical network is computationally expensive. A 2D (longitude–pressure) chemical model is sufficient to study the chemical species' local-time variabilities that we focused on in this work. For example, combining a 3D GCM with a 2D CTM has been used to study chemical species in Earth's atmosphere (e.g., Smyshlyaev et al. 1998).

We adopt the OASIS GCM, a novel and flexible 3D planetary model (Mendonça & Buchhave 2020). OASIS is a dedicated model that incorporates multiple self-consistent modules. For our Venus dynamical simulations, we use the nonhydrostatic dynamical core coupled with physics modules that represent a basalt soil/surface, convective adjustments, and the radiative processes from the gas and clouds (a nongray scheme with multiple scattering). The simulated atmosphere extends from the surface to 100 km, with a horizontal resolution of 2° and a vertical resolution of ∼2 km. The model was integrated for 25,000 Earth days (∼214 Venus solar days; one Venus solar day is ∼117 Earth days) with a time step of 50 s. The model and bulk planet parameters (e.g., specific heat, gravity, and mean radius) are the same as the ones used in Mendonça & Buchhave (2020, see their Table 2). One of the main weaknesses of current Venus GCMs is the poor representation of the circulation in the deep atmosphere, which is also poorly constrained by observational data (refer to Mendonça & Read 2016 for more details). To represent a deep circulation in our 3D simulations closer to the observations, we applied a Newtonian relaxation method to force the zonal winds in the deep atmosphere toward the observed values. The forcing acts only at 44 km altitude, which is below the cloud region and the region explored in this study. The equilibrium winds were constructed assuming the atmosphere at 44 km rotating as a solid body with a maximum velocity of 50 m s^{-1} at the equator based on the estimated observed values from Kerzhanovich & Limaye (1985). For the Newtonian relaxation timescale, we have assumed a value of 2000 Earth days (∼17 Venus solar days), which is close to the radiative timescale at 44 km (Pollack & Young 1975). At 44 km, the temperature difference between the dayside and nightside of the planet is small (less than 10 K) because the radiative timescale is much longer than the dynamical timescale. The value chosen for the relaxation timescale ensures a good model performance and low impact in the wave activity in the lower atmosphere. Our converged simulations were further integrated to 5000 Earth days (∼43 Venus solar days) to produce the temperature and wind fields for the CTM input.

We do not directly couple the GCM and CTM in the sense that the simulated gas distributions in the CTM are not used as the GCM input. As described in Mendonça & Buchhave (2020), the 3D GCM itself only uses simple representations of the clouds and chemistry. The cloud structure remains constant with time, and three different cloud particle size modes (Knollenberg & Hunten 1980; Crisp 1986) are used. The GCM considers four main chemical species in the atmosphere: 96.5% of CO_2 in mole, ∼3.5% of N_2, 50 ppm of H_2O, and 100 ppm of SO_2. Their volume mixing ratios are assumed to be well mixed in the GCM and meant not to be exactly equal to the values observed but to capture the main bulk conditions of the Venusian atmosphere.

The 2D CTM is generalized from the 1D state-of-the-art Caltech/JPL kinetics model (Yung & DeMore 1982; Mills 1998; Mills & Allen 2007; Zhang et al. 2010, 2012; Bierson & Zhang 2020; Shao et al. 2020). This model resolves complex chemistry for carbon, oxygen, hydrogen, nitrogen,

Table 1
Boundary Conditions for Several Important Species in the 2D CTM

Species	Lower Boundary Condition	Upper Boundary Condition
SO_2	$f = 1.0$ ppm	$\phi = 0$
H_2O	$f = 1.0$ ppm	$\phi = 0$
CO	$f = 45$ ppm	$\phi = 0$
NO	$f = 5.5$ ppb	$\phi = 0$
HCl	$f = 0.4$ ppm	$\phi = 0$
CO_2	f=0.965	$\phi = 0$
OCS	$f = 1.0$ ppm	$\phi = 0$
S_8	$v = v_m$	$\phi = -6.0 \times 10^7$ cm^{-2} s^{-1}

Note. f means the fixed volume mixing ratio, ϕ means the diffusive boundary flux, and v is the deposition velocity. Values here are referred to those in Zhang et al. (2012) and Bierson & Bierson & Zhang (2020). Species not specified here all have $\phi = 0$ at the upper boundary and the maximum deposition velocity v_m (see Zhang et al. 2012) at the lower boundary (58 km).

sulfur, and chlorine species. This model includes 52 chemical species and over 400 reactions (refer to Zhang et al. 2012). We generalized this 1D model to 2D and included the advection for each chemical tracer in the longitude–(log-)pressure coordinate plane. See Appendix A for the derivation of the meridionally mean continuity equation. The meridionally mean advection is constructed from the output of the 3D GCM. We implemented the flux-limiting Prather scheme (Prather 1986; Shia et al. 1989, 1990) to calculate the advection of species in the 2D continuity equation. This scheme has several advantages, including the conservation of chemical species, maintenance of positive concentration, and stability for large time steps. The chemical model incorporating this scheme has been applied to Earth's atmosphere to study variabilities of chemical species like ozone (Jiang et al. 2004). We have also implemented a parallel computing technique using Message Passing Interface (MPI) in our 2D CTM to improve simulation efficiency.

In our 2D CTM, photon density reaching the top of the atmosphere (TOA) is set as equal to $\pi/4$ times the equatorial value on Venus, so as to represent the meridionally mean value considering the latitudinal dependence of solar zenith angle. The solar zenith angle in our 2D CTM varies with longitude, and the solar zenith angle at each longitude also changes with time. Our 2D CTM has a vertical resolution of ~ 2 km and a horizontal resolution of 12°. The altitude range is ~58–100 km. The time step is set as 10 minutes. In Section 4, we show that increasing horizontal resolution does not change the simulated local-time variabilities of chemical species.

The 3D distributions of temperature and wind patterns from GCM simulations in the last ~ 4 Venus solar days are first averaged meridionally. To match the spatial and temporal grids in our 2D CTM, we then smooth and interpolate the GCM data to obtain temperature and wind fields in 1 hr resolution. Finally, we average the fields temporally to obtain the diurnally varying one-Venus-day (~117 Earth days) fields and repeatedly input them into the 2D CTM. Table 1 lists boundary conditions for several important species. For other species, zero flux at the upper boundary and maximum deposition velocity at the lower boundary (58 km) are applied. We apply the same lower and upper boundary conditions to all longitudes. The unknown sulfur reservoir in the upper atmosphere (e.g., Zhang et al. 2010, 2012; Vandaele et al. 2017b) is represented by a downward S_8 flux at the upper boundary in our model, as used in Bierson & Zhang (2020). The specified flux at the upper boundary (e.g., Table 1) is separate from the advective flux and is used to provide extra sources outside the domain (e.g., the S_8 flux). We calculate the advective flux above (below) the upper (lower) boundary by setting a ghost box with species' mixing ratios the same as those at the boundary. In the zonal direction, a periodic boundary condition (i.e., species abundances at 0° and 360° are equal) is adopted.

In this study, we treat the subgrid diffusivity parameters K_{xx}, K_{xz}, K_{zx}, and K_{zz} in the meridionally mean continuity equation (see Appendix A) as free parameters. For simplicity, we assume zero K_{xz} and K_{zx}. The meridionally mean zonal wind is usually larger than the eddy wind. For example, the meridionally mean zonal wind is ~ 100 m s^{-1} at ~60 km, while the eddy wind is ~10–20 m s^{-1} at ~60 km at the equatorial region in the GCM output. If we assume that the sub-grid eddy length scale is 10–100 km (the horizontal grid size is about 200 km around the equator in the GCM), the horizontal diffusivity K_{xx} is about 10^9–10^{10} cm^2 s^{-1}. Here we use $K_{xx} = 10^9$ cm^2 s^{-1} to represent the horizontal transport by eddies. The K_{zz} vertical profile in our CTM is the same as the 1D K_{zz} profile in Zhang et al. (2012) and is applied to all longitudes. In Section 4, we will explore the sensitivity of our results to these parameters.

3. Local-time Dependence of Chemical Species

In this section, we discuss the local-time dependence of SO_2, CO, H_2O, HCl, ClO, OCS, and SO. These chemical species' distributions are averaged over the last 5 Venus solar days from our simulations. To better understand the interaction between atmospheric dynamics and chemistry and the influence on local-time variability, we also estimate the chemical loss timescale t_{Chemical} and horizontal transport timescale $t_{\mathrm{Transport}}$. The chemical loss timescale for each species is equal to the number density of the species divided by its total chemical loss rate. For SO_2, CO, and SO, the chemical loss in the fast cycles (see Appendix B) is excluded. The horizontal transport timescale is estimated using the planetary radius divided by the zonal wind speed (e.g., Zhang & Showman 2018).

The 3D GCM results show that the diurnal cycle excites various harmonics of the thermal tides in the Venus atmosphere (Figures 1(a)–(b)). The thermal tides are stationary with respect to the subsolar point. Observational (e.g., Pechmann & Ingersoll 1984; Zasova et al. 2007) and theoretical (e.g., Lebonnois et al. 2010; Mendonça & Read 2016) studies suggest that the semidiurnal component has the largest amplitude of the thermal tide harmonics in the upper cloud region. Our GCM successfully simulates this semidiurnal component in the upper cloud (58–70 km). In the vertical wind pattern shown in Figure 1(b), the semidiurnal tide at 58–70 km induces an upwelling branch in the afternoon. A similar branch also appears at 0:00–6:00 but with a weaker amplitude. At 18:00–0:00 and 6:00–12:00, the semidiurnal tide in the upper cloud induces downwelling motions. The evening downwelling is stronger than the morning downwelling. Positive temperature anomalies are found around midnight and noon as a result of perturbations by the semidiurnal tide.

Above 85 km, the diurnal thermal tide dominates the wind and temperature local variations in our simulations (Figures 1(a)–(b)). As altitude increases, the phase of the semidiurnal tide shifts eastward. Above 85 km, two upwelling branches of the semidiurnal tide merge into a dayside upwelling branch, while the morning downwelling of the semidiurnal tide extends and becomes the nightside downwelling branch above 85 km. In this

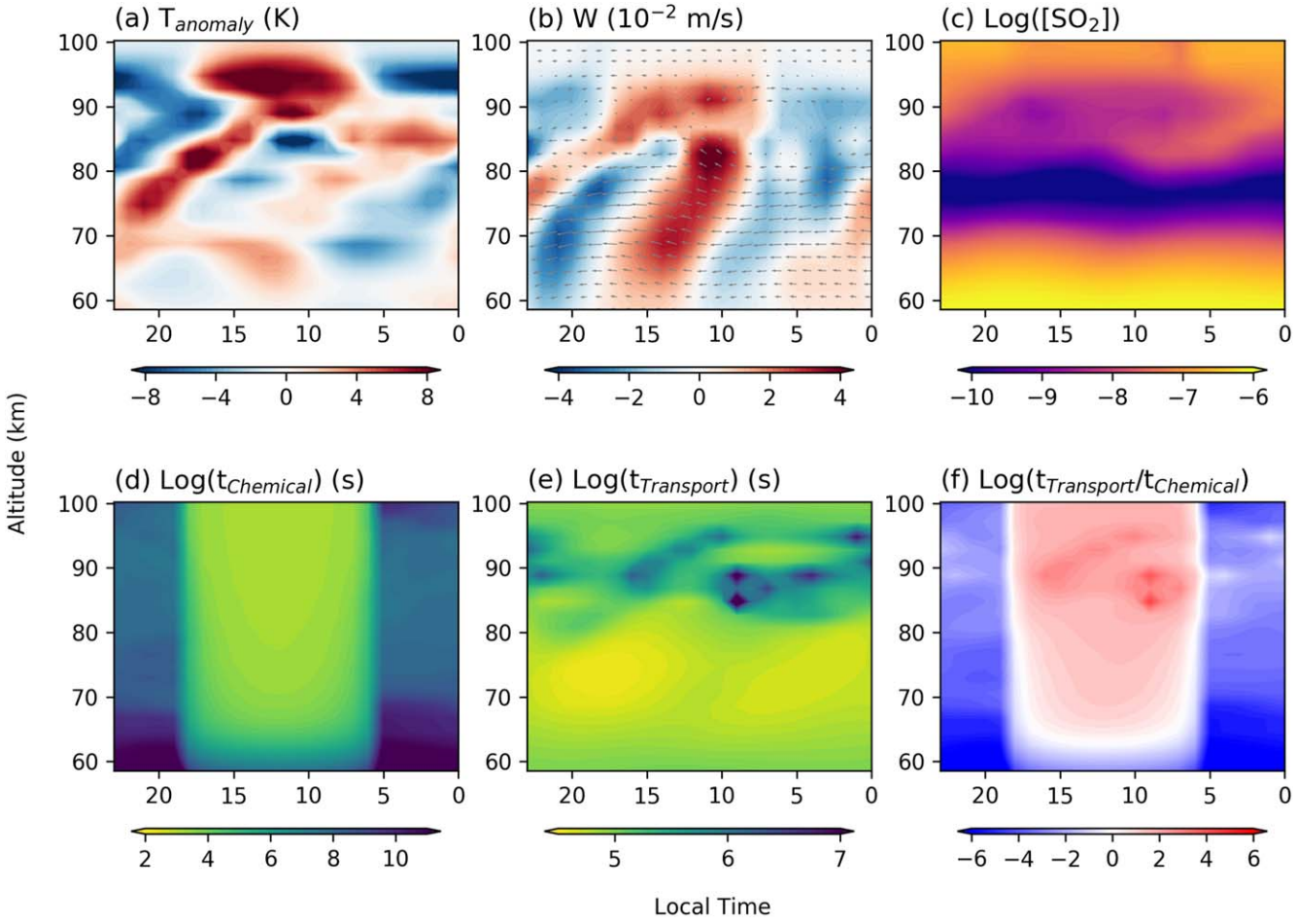


Figure 1. Local-time dependence of (a) temperature anomalies T_{anomaly}, (b) vertical velocity W, (c) SO_2 mixing ratio, (d) SO_2 chemical loss timescale t_{Chemical}, (e) transport timescale $t_{\text{Transport}}$, and (f) ratio of $t_{\text{Transport}}$ to t_{Chemical}. The wind field (m s^{-1}) is superposed on panel (b). Temperature and wind fields are from the OASIS simulations (Mendonça & Buchhave 2020), and the SO_2 mixing ratios are from the 2D CTM. Temperature anomaly is the deviation from an average temperature profile shown in Figure C1. Note that both the 2D CTM and OASIS use (log-)pressure coordinates. Height at the vertical axis in this plot represents the isobaric level and is derived from pressure by using the VIRA model (see Table 1 of Mendonça & Read 2016). The local time 06:00 is the morning terminator, 18:00 is the evening terminator, 12:00 is noon, and 00:00 is midnight. Earlier local time means eastward shift on Venus. The ratio of $t_{\text{Transport}}$ to t_{Chemical} indicates the main driven mechanism for the species distribution: the ratio smaller than unity (blue region in panel (f)) implies a mainly transport-driven regime; the ratio larger than unity (red region in panel (f)) implies a mainly photochemistry-driven regime; the ratio around unity (white region in panel (f)) implies the transition between the two regimes.

Table 2
Observations Used in This Paper

Observation	Altitude (km)	Species	Mixing Ratio Range	Reference
TEXES	~64	SO_2	150–400 ppb	Encrenaz et al. (2020)
JCMT	70–100		0–90 ppb	Sandor et al. (2010)
SPICAV	95–100		50–200 ppb	Belyaev et al. (2017)
Microwave	80–100	CO	30–1000 ppm	Clancy & Muhleman (1985)
JCMT	75–100		50–1000 ppm	Clancy et al. (2003)
SOIR	85–130		10^{-4}–10^{-1}	Vandaele et al. (2016)
TEXES	~64	H_2O	~1 ppm	Encrenaz et al. (2020)
JCMT	70–100	HCl	0–450 ppb	Sandor & Clancy (2012, 2017)
SOIR	70–105		30–800 ppb	Mahieux et al. (2015)
JCMT	70–100	ClO	1.5–3.7 ppb	Sandor & Clancy (2018)
CSHELL	~65	OCS	0.3–9 ppb	Krasnopolsky (2010)
JCMT	70–100	SO	0–30 ppb	Sandor et al. (2010)
SPICAV	85–105		10–1000 ppb	Belyaev et al. (2012)

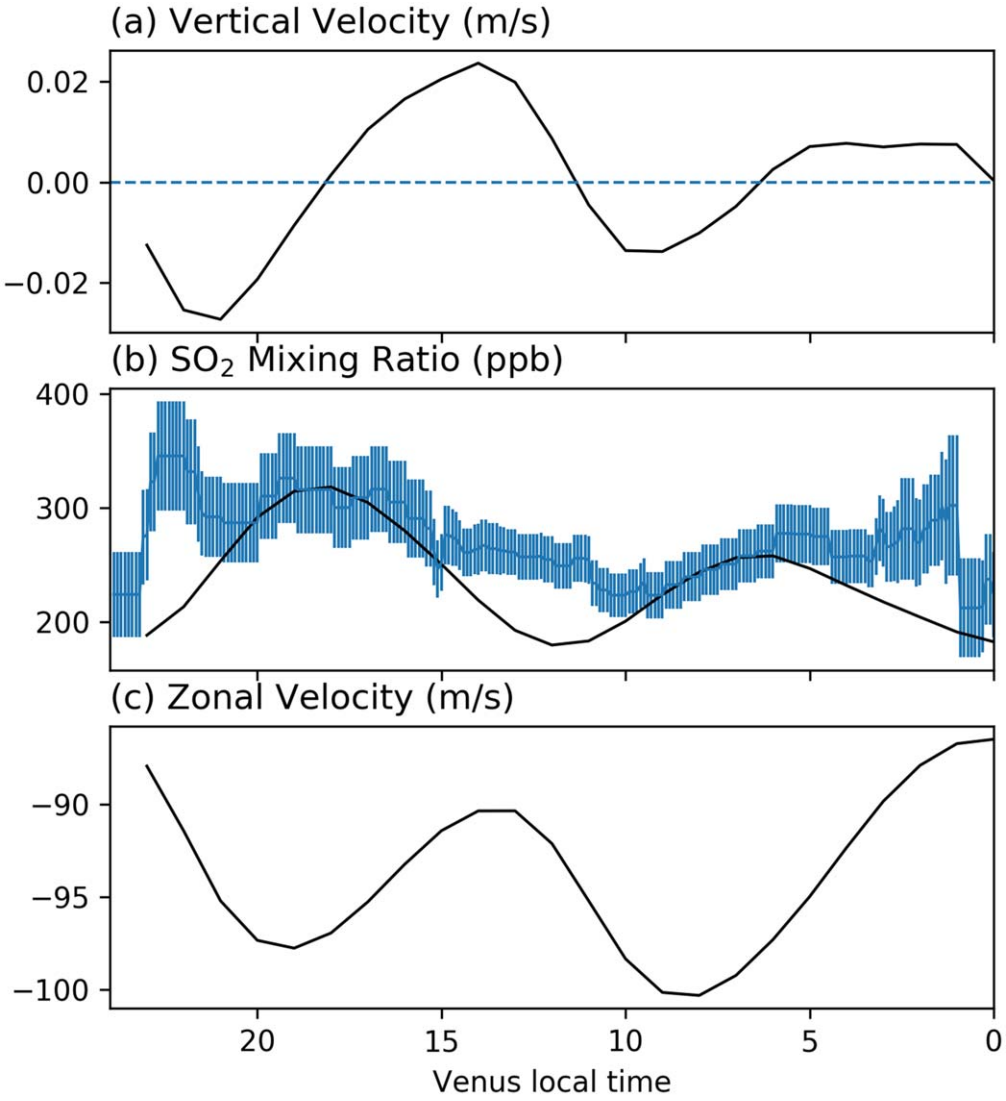


Figure 2. Local-time distributions of (a) vertical velocity, (b) SO_2 mixing ratio, and (c) zonal velocity around 64 km. Observational data (error bars) in panel (b) are from TEXES/IRTF (Encrenaz et al. 2020). Positive vertical velocity is upward, and negative zonal velocity is westward.

altitude region, the vertical wind field is mainly composed of the wavenumber-one diurnal component. Temperature distribution is also affected by the diurnal thermal tide above 85 km, with positive anomaly on the dayside. At 85–100 km, the SS-AS circulation is imposed on the RSZ flows, marking a transition region from RSZ flow to SS-AS circulation where the wind pattern is important for chemical tracer exchange between the lower and higher altitudes, as well as that between the dayside and the nightside. The chemical tracers at the lower altitude are first transported upward by the upwelling branch on the dayside and then delivered to the nightside by horizontal day–night flows of the SS-AS circulation. On the nightside, the chemicals are recycled back to the lower region by the downwelling branch of the SS-AS circulation and transported to the dayside by the RSZ flows.

3.1. SO_2

Our simulation shows that the semidiurnal thermal tide is essential to explain the SO_2 local-time pattern in the upper cloud region (Figure 1). Below 80 km, the SO_2 mixing ratio decreases as altitude increases. Vertical mixing occurs when the SO_2-rich air is transported upward from a lower altitude and the SO_2-poor air is transported downward from a higher altitude. In the upper cloud region (58–70 km), the two upwelling branches of the semidiurnal tide produce two local maxima in the SO_2 local-time distribution. The two maxima are shifted westward by the RSZ flow and located around two terminators. This local-time pattern is more clearly seen in Figure 2. Our simulations successfully reproduce the SO_2 observations by TEXES at ∼64 km (Encrenaz et al. 2020). Note that SPICAV also observed a similar SO_2 local-time distribution at ∼70 km on the dayside (Vandaele et al. 2017a; Encrenaz et al. 2019; Marcq et al. 2020). The TEXES data exhibit more complicated features, like a peak around 22 hr and another around 2 hr (Figure 2). The causes of these peaks are not well understood and might be associated with small-scale dynamics.

In the region above the clouds, photochemistry drives the SO_2 behavior on the dayside (Figure 1(f)). Above 85 km, SO_2 day–night difference becomes evident; SO_2 is less abundant on the dayside than on the nightside. Both photochemistry and dynamics drive this day–night difference. On the dayside, photolysis destroys SO_2; on the nightside, the descending branch of the SS-AS circulation brings SO_2-rich air downward

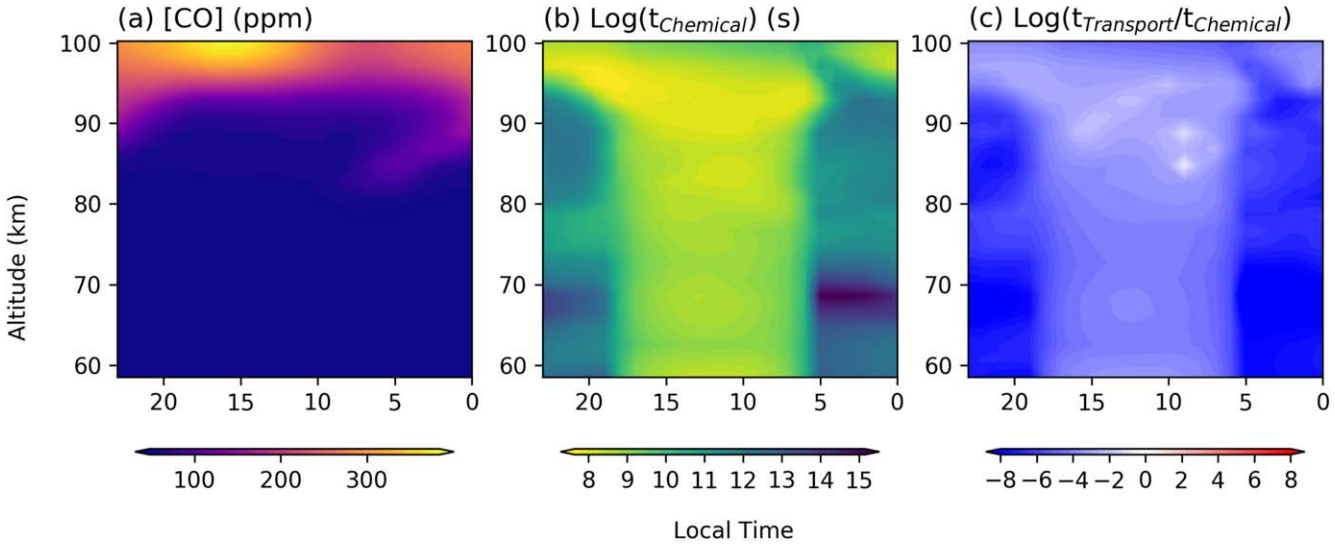


Figure 3. Local-time dependence of (a) CO mixing ratio, (b) CO chemical loss timescale t_{Chemical}, and (c) ratio of $t_{\mathrm{Transport}}$ to t_{Chemical}.

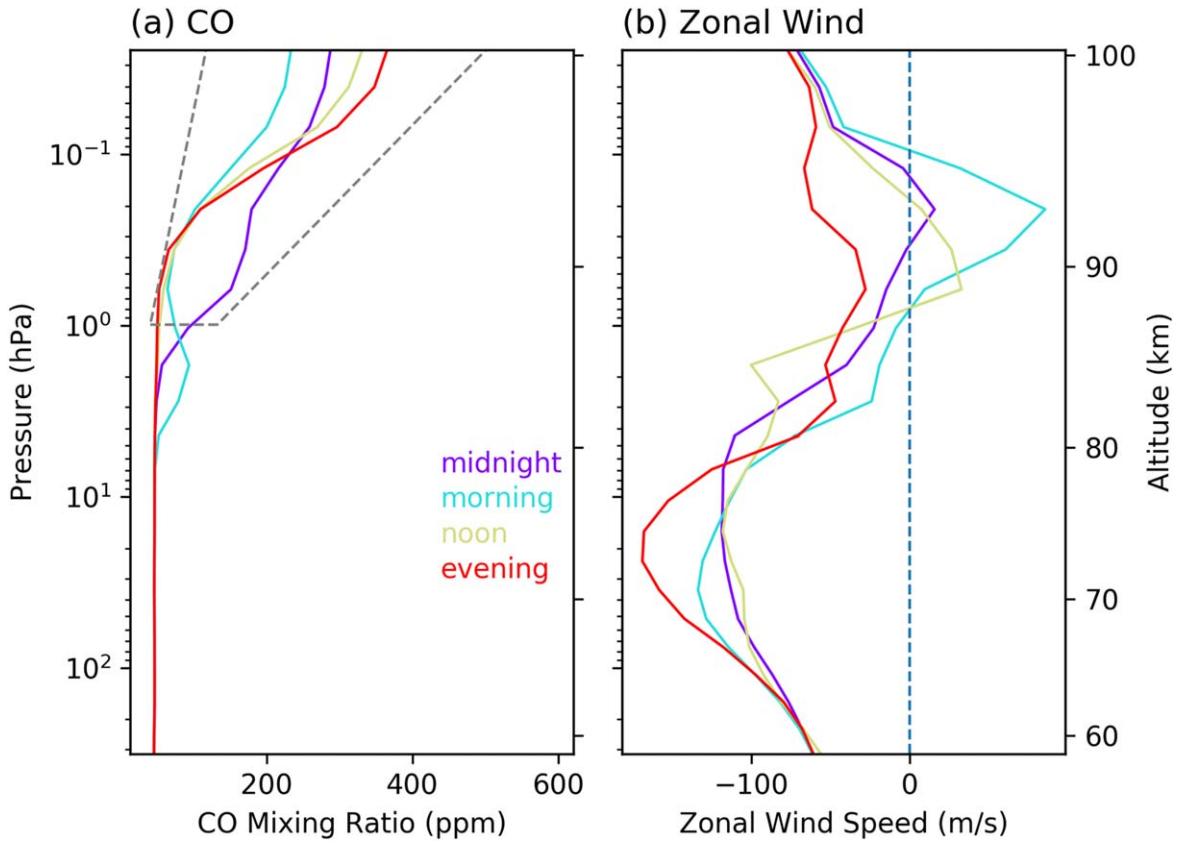


Figure 4. Vertical profiles of (a) CO mixing ratio and (b) zonal wind at different local times. Altitude derived from the VIRA model is shown on the right axis. The gray dashed line in panel (a) encloses a region corresponding to a rough range of the observations by Vandaele et al. (2016). The blue dashed line in panel (b) is zero zonal wind.

because SO_2 mixing ratio generally increases as altitude increases above 85 km owing to the assumed S_8 downward flux in our model. Sandor et al. (2010) implied an SO_2 day–night difference from microwave measurements at 70–100 km, despite the scarcity of their data. Belyaev et al. (2017) observed 150–200 ppb SO_2 at midnight versus 50 ppb SO_2 at terminators around 95 km from SPICAV occultations. Our simulations roughly agree with the observed SO_2 pattern around 95 km (Figure 10(a)).

3.2. CO

CO is a long-lived chemical species whose local-time distribution is determined mainly by dynamics (Figure 3). Figures 3(b) and (c) show that CO has a long chemical loss timescale above the clouds. Below 85 km, CO is well mixed and almost exhibits no diurnal variations. Above 85 km, CO shows a day–night difference similar to SO_2 (Figure 3(a)). This difference is caused by the SS-AS circulation. The CO mixing

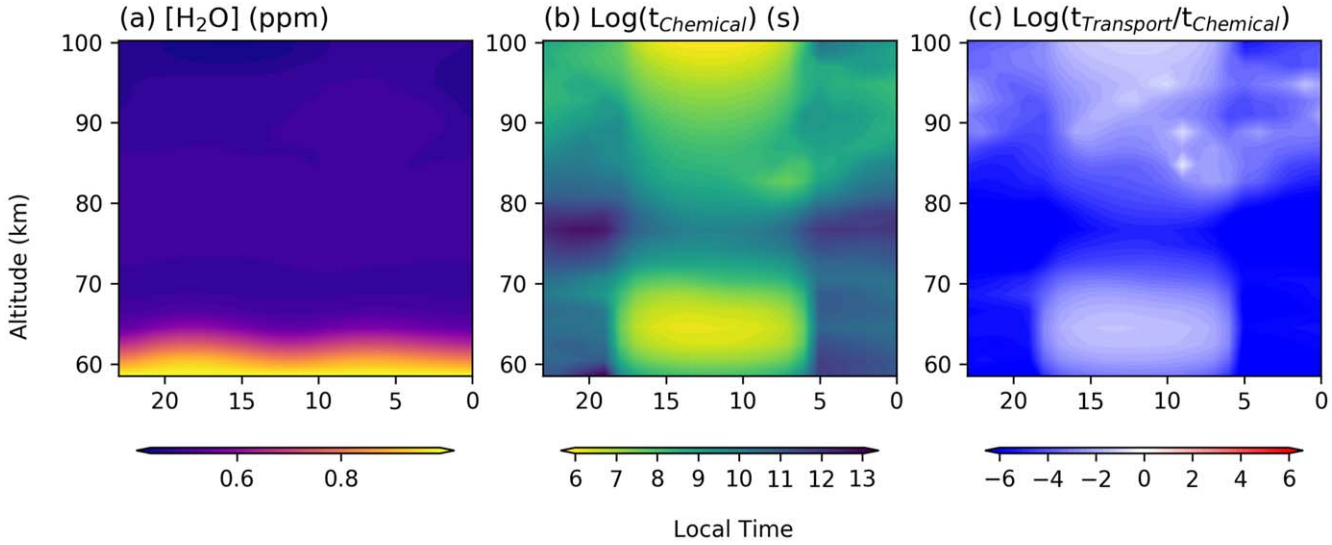


Figure 5. Same as Figure 3, but for H_2O.

ratio generally increases as altitude increases above 85 km because it is mainly produced by CO_2 photolysis in the upper atmosphere and transported downward. Above 85 km, the SS-AS circulation reduces CO on the dayside by mixing the CO-poor air upward from below. Although photochemistry produces more CO on the dayside, the SS-AS circulation regulates and dominates the CO local-time distribution by transporting the CO from the dayside to the nightside. As a result, CO is accumulated on the nightside and appears more abundant than on the dayside.

The CO maximum is located around midnight at ∼95 km and shifts to dawn at ∼85 km (also see Figures C3(c)–(d)). This pattern has been observed by microwave instruments (Clancy & Muhleman 1985; Clancy et al. 2003). The CO maximum is shifted westward owing to zonal winds in the transition region, where SS-AS circulation transits to RSZ flow as altitude decreases. A similar CO pattern is also seen in the 3D GCM of Navarro et al. (2021) and Gilli et al. (2021). In their simulations, the CO maximum shifts westward toward the morning at 85–100 km, caused by a westward flow imposed on the SS-AS circulation. Our results overall agree with their results.

SOIR has observed a statistical difference of the CO mixing ratio profiles between two terminators via solar occultation (Vandaele et al. 2016). In Figure 6 of Vandaele et al. (2016), the CO mixing ratio below 95 km is larger at the morning terminator than at the evening terminator, while CO above 105 km shows a reversed pattern. Our simulation reproduces such a pattern, but the reversal of the terminator difference occurs at a lower altitude (∼90 km; see Figure 4(a)) than in the SOIR observations (∼90–110 km).

This reversed CO terminator difference originates from transition of atmospheric flows. At 80–90 km in our model, thermal tides transport CO-rich air downward on the nightside. The RSZ flows shift the CO-rich air toward the morning terminator. CO-poor air is pumped up by the upwelling branch of the tides on the dayside and is shifted toward the evening terminator. This process results in a larger CO mixing ratio in the morning than in the evening. Above 90 km, the SS-AS circulation transports CO produced on the dayside toward both terminators. Theoretically, if the dynamical pattern is symmetric about noon, there should be no difference between the two terminators. However, the circulation from our 3D GCM simulations is asymmetric at these altitudes. For example, at 90–95 km, zonal flows at the two terminators have different amplitudes with opposite directions (Figure 4(b)). This asymmetry could cause the terminator difference of the CO mixing ratio above 90 km. The wind pattern from the GCM in Gilli et al. (2021) is also asymmetric above 110 km owing to perturbations of gravity waves. The CO observations by SOIR do not show a large difference between terminators until above 120 km (Vandaele et al. 2016). This may imply that only above 120 km does the asymmetric wind pattern become significant enough to affect CO local-time patterns.

The reversal altitude of the CO terminator difference might be closely related to the transition from RSZ flow to SS-AS circulation on Venus. That our simulated reversal level is lower than in the SOIR observations might imply that the transition from RSZ flow to SS-AS circulation occurs at a lower level in our GCM simulations than that in the real Venusian atmosphere. Because the transition level could also vary with time and space, future observations of CO distributions are useful to constrain the flow pattern transition in the upper atmosphere of Venus.

3.3. H_2O, HCl, ClO, OCS, and SO

H_2O distributes almost uniformly over local time and altitude in our simulations (Figure 5(a)). This is because H_2O is a long-lived species (Figures 5(b)–(c)). Due to thermal tides, H_2O exhibits small local-time variations in the upper cloud region (58–70 km), and the amplitudes of these variations are generally less than 30%. The uniform distribution of H_2O over local time is consistent with observations by both SPICAV (e.g., Fedorova et al. 2008) and TEXES (e.g., Encrenaz et al. 2020). SOIR observations also show no significant difference of H_2O between morning and evening terminators in the upper mesosphere (Chamberlain et al. 2020).

HCl, like H_2O, has a long chemical lifetime and distributes uniformly over space (Figure 6). Its vertical profile in our simulations, like in previous 1D models (e.g., Yung & DeMore 1982), shows a weak decrease from the cloud top to above 90 km. This simulated profile disagrees with JCMT

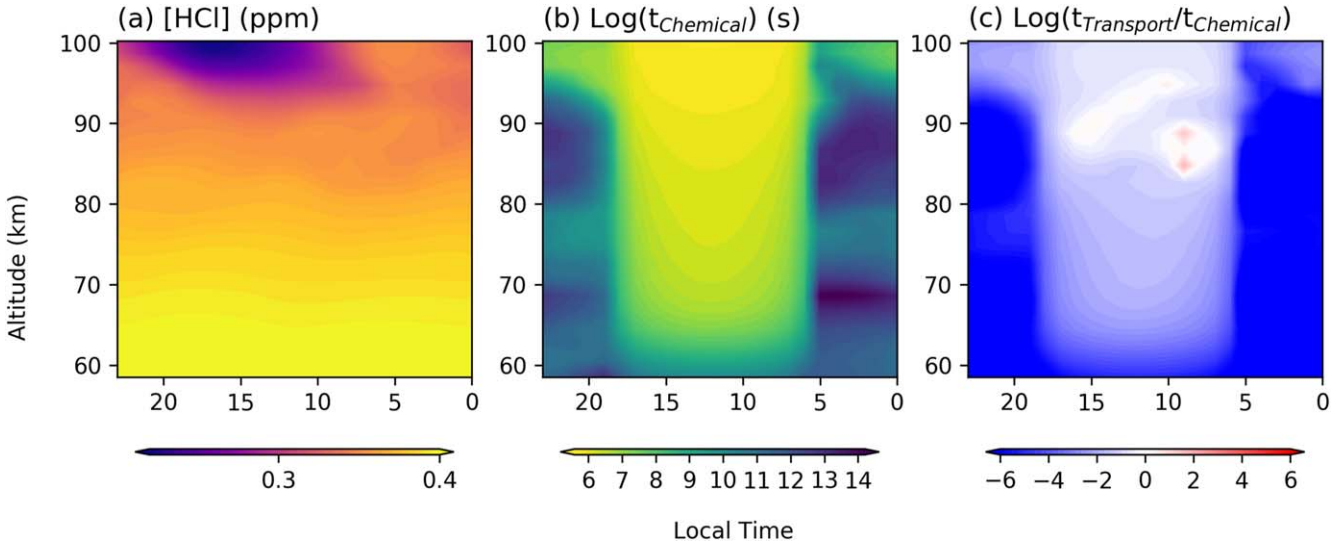


Figure 6. Same as Figure 3, but for HCl.

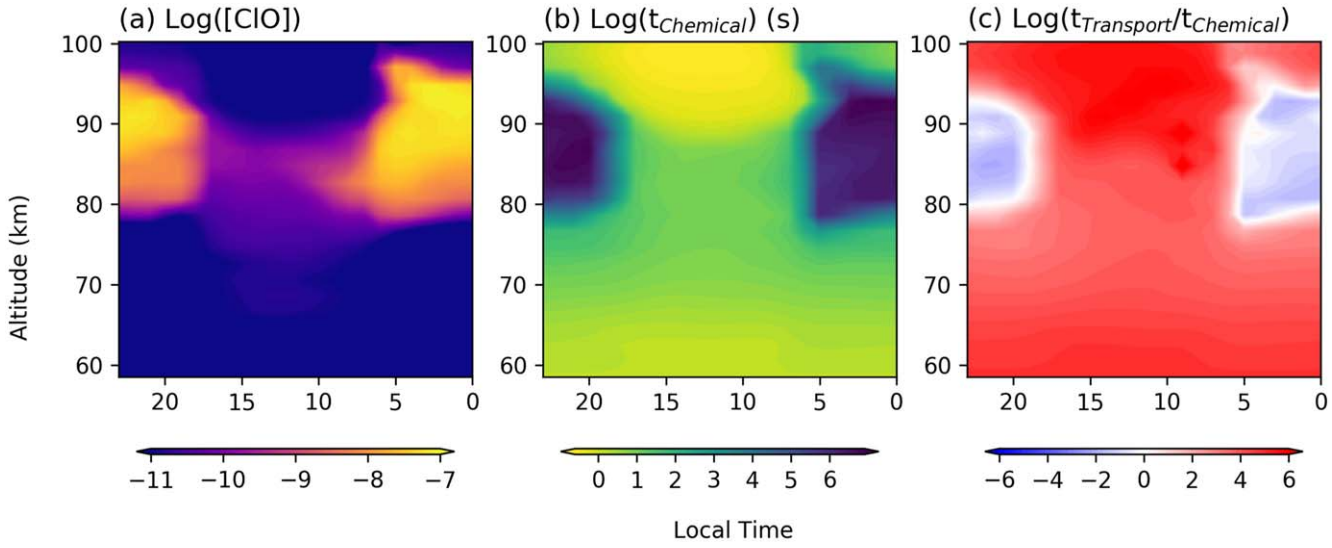


Figure 7. Same as Figure 3, but for ClO.

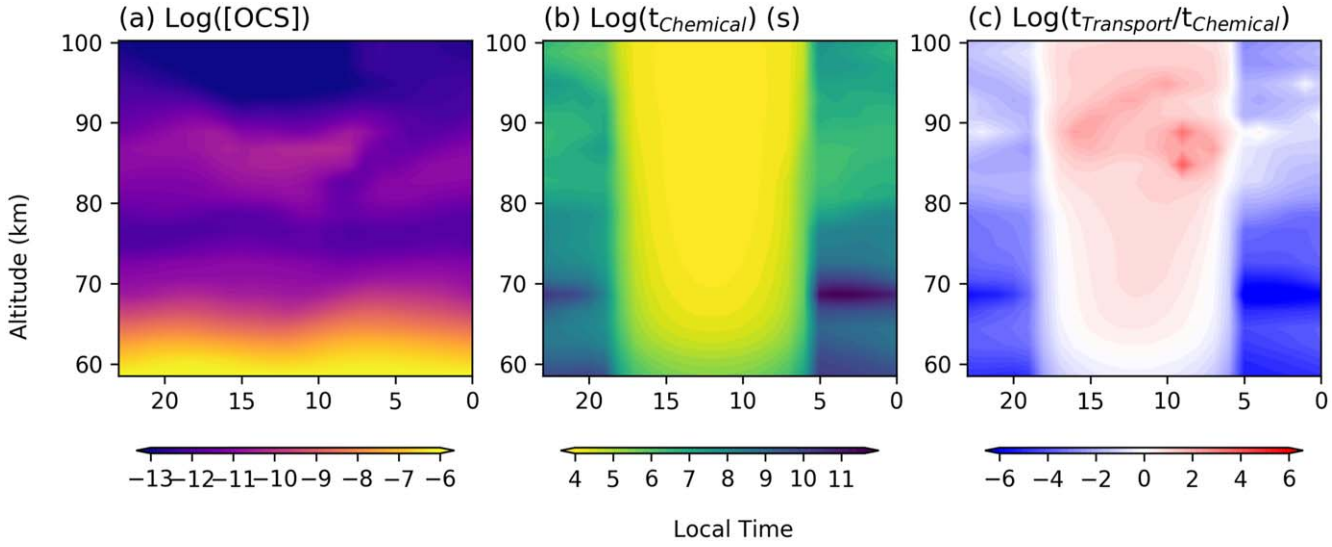


Figure 8. Same as Figure 3, but for OCS.

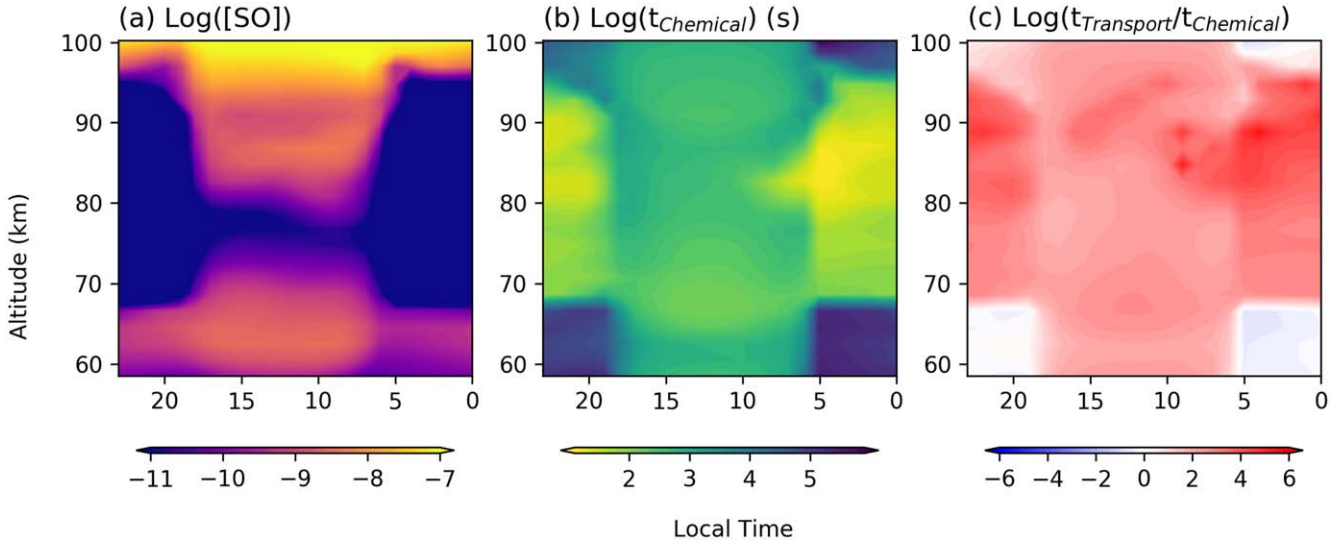


Figure 9. Same as Figure 3, but for SO.

observations (Sandor & Clancy 2012, 2017), which show a large decrease as altitude increases. Our model seems to support the conclusion of Sandor & Clancy (2017) that the large decrease of HCl mixing ratio observed by JCMT does not originate from the SS-AS circulation. However, note that SOIR observed that HCl mixing ratio increases as altitude increases (Mahieux et al. 2015), which disagrees with the JCMT observations and also our model (and previous models). The SOIR observation suggests a chlorine source at high altitude, but no chemical hypothesis could support this source. Future observations are needed to further investigate the discrepancy among models and observations.

ClO is a short-lived species except at 80–95 km on the nightside (Figure 7). In the entire mesosphere, ClO mixing ratio is rather small, mostly <1 ppb. But at 80–95 km on the nightside, where ClO chemical lifetime is longer, ClO can reach a few tens of ppb (Figure C6). Our simulated nightside ClO is roughly consistent with the results in a 1D nightside model from Krasnopolsky (2013), but the abundance is much larger than the 1D diurnal-mean photochemical model results from Zhang et al. (2012) and Krasnopolsky (2012). Sandor & Clancy (2018) observed the nighttime ClO using JCMT and retrieved a few ppb of ClO above 85 km, which is an order of magnitude smaller than our simulated ClO mixing ratio on the nightside. Because the observed HCl from JCMT is also smaller than our simulated HCl in the upper atmosphere, we hypothesize that there might be some unidentified sinks for ClO and HCl.

The vertical profile of OCS mixing ratio shows a small peak at 80–90 km. This peak is due to the downward S_8 flux from the top boundary in order to explain the SO_2 inversion. Part of the S_8 also converts to OCS to form a peak at 80–90 km. OCS is a short-lived chemical species on the dayside above the clouds and a long-lived species on the nightside (Figure 8). Its distribution thus is largely affected by photochemistry on the dayside in the upper atmosphere and by dynamics on the nightside. OCS around 95 km exhibits a smaller mixing ratio on the dayside than on the nightside and a reversed local-time pattern around 85 km (Figure C7), as a result of competition between photochemistry and dynamics. However, since OCS mixing ratios at these altitudes do not exceed 1 ppb, these local-time variations are not easily observed. At ~65 km in the upper cloud, OCS mixing ratio can exceed 1 ppb, and the local-time difference of the OCS mixing ratio can reach ~10 ppb. This may be an observable pattern in the future. OCS also exhibits a two-maxima local-time pattern at ~65 km, similar to SO_2. But the larger maximum of OCS locates around the morning terminator, while that of SO2 is around the evening terminator. Krasnopolsky (2010) observed a few ppb of OCS near 65 km using the CSHELL spectrograph at NASA IRTF and indicated a pattern in which the morning OCS is more abundant than the afternoon OCS, supporting our simulated OCS local-time pattern here (Figure 10(c)). The OCS decrease from morning to afternoon should be related to that around 65 km; the OCS behavior is driven by both photochemistry and dynamics (Figure 8(c)), unlike SO_2, which is more driven by dynamics.

SO exhibits a complex spatial pattern (Figure 9). Since SO is a short-lived species and mainly produced by SO_2 photolysis, SO is more abundant on the dayside than on the nightside. But in the upper cloud region on the nightside, SO has a longer chemical lifetime than the transport timescale by the RSZ flow (Figure 9(c)), leading to a smaller day–night contrast than that at 70–95 km. The day–night difference of SO in the upper mesosphere is consistent with the JCMT observation by Sandor et al. (2010) (see Figure 10(b)). However, the SO mixing ratio around 95 km is lower in our model than the SPICAV observations (Belyaev et al. 2012). The SO mixing ratio shows a very strong local-time dependence (Figure 9(a)). Therefore, only observing the terminator SO is insufficient to understand the SO behavior. To better understand the sulfur cycle in the upper mesosphere of Venus, observations covering multiple local times on both the dayside and nightside are required.

4. Sensitivity Test

We conduct sensitivity tests to explore the effects of the horizontal diffusion coefficient K_{xx}, the vertical diffusion coefficient K_{zz}, and the horizontal resolution on our results. For simplicity, we still assume the diffusion coefficients K_{xz} and K_{zx} as zero.

Our sensitivity tests show that K_{zz} augment does not affect the overall local-time patterns of all species discussed above (Figures C2–C8). For example, the two-maximum pattern of SO_2 at ~64 km is still well produced in the cases with a larger K_{zz} (Figure C2(a)). The major effect of increasing K_{zz} is to

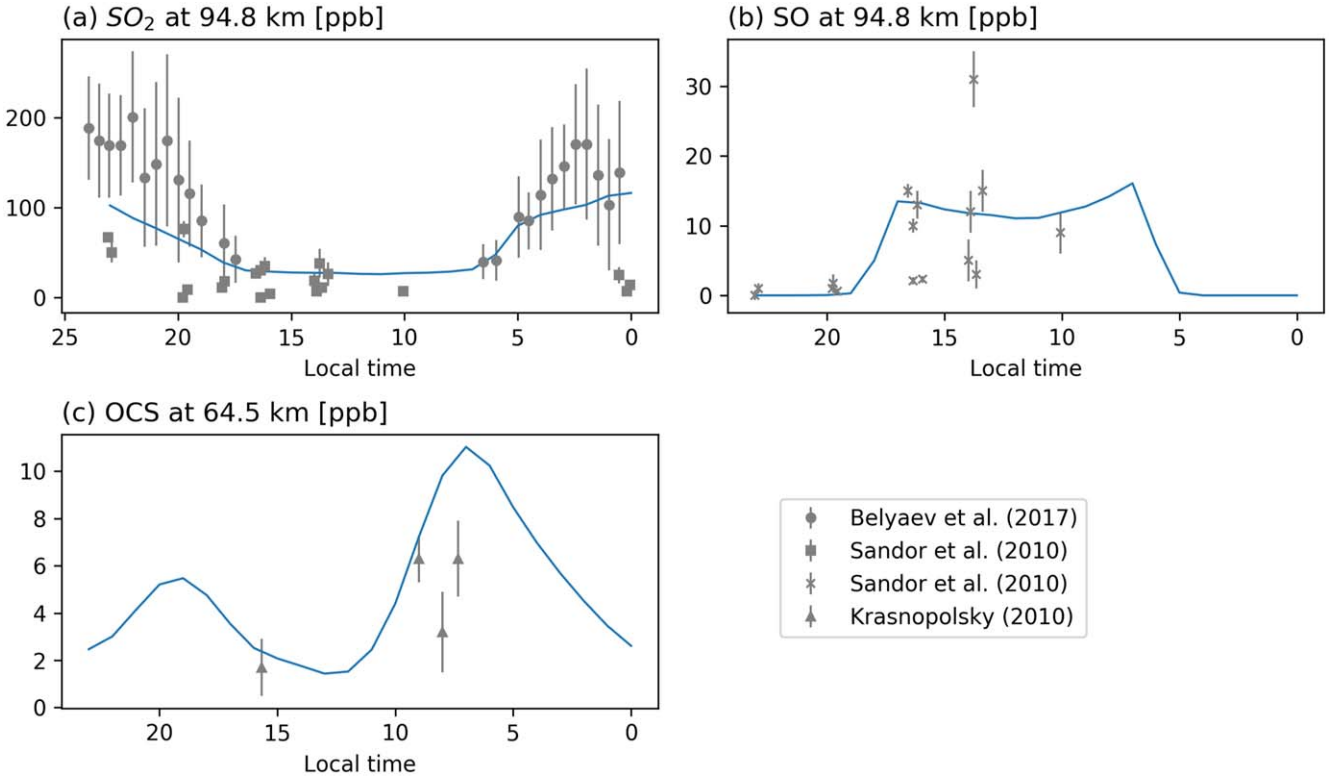


Figure 10. Local-time variations of volume mixing ratios of SO_2, SO, and OCS from both our model and observations (Figure 11(b) of Belyaev et al. 2017; Sandor et al. 2010; Krasnopolsky 2010). Note that the observation altitude is not necessarily exactly the value shown in this plot (refer to Table 2 to see the observation altitude range).

increase the SO_2 mixing ratio below 80 km. The increase below 80 km is due to more diffusion from the lower sulfur reservoir (at ~58 km in our model). As a result, the mixing ratio of SO—a photochemical product of SO_2—below 80 km also increases. OCS is also sensitive to K_{zz} value. As K_{zz} increases, the amplitude of the OCS local-time variation at ~65 km increases despite the qualitative pattern being unchanged (Figure C7(a)). This implies that the OCS local-time pattern at ~65 km can be a good indicator of the strength of atmospheric vertical mixing.

Changing K_{xx} from $10^9\,cm^2\,s^{-1}$ (the value in the nominal case) by a factor of 10 does not affect the local-time patterns of all species discussed above (Figures C2–C8). It exerts almost no effect on the mixing ratios of the species. This is because horizontal transport by eddies only contributes a small proportion to chemical transport compared to the meridionally mean zonal wind; the horizontal diffusion timescale is $\sim 10^7$–10^9 s (estimated by L^2/K_{xx}, where L is planetary radius), compared to the advection timescale of $\sim 10^4$ s in the upper cloud (Figure 1(e)). Our test also shows that increasing the horizontal resolution from 12° to 6° does not change the local-time patterns of the species discussed in this work (Figures C2–C8).

5. Conclusion and Discussions

In this paper, we investigated the local-time dependence of chemical species in the Venusian mesosphere. We used a 3D GCM and a 2D CTM to simulate species' local-time distributions and investigate the underlying mechanisms. Our models reproduce the observed local-time patterns of many chemical species such as SO_2 and CO. Dynamics and photochemistry play different roles in controlling the local-time patterns for different chemical species in the Venusian atmosphere.

As observed by TEXES, the local-time pattern of the SO_2 at ~64 km features two local maxima around terminators (Encrenaz et al. 2020). Using our model, we found that this feature is caused by the superposition of the semidiurnal thermal tide and the RSZ flow in the upper cloud. The two upwelling branches of the semidiurnal tide produce two local SO_2 maxima, and the superrotating wind advects the maxima toward terminators. SO_2 above 85 km has a large day–night difference with more SO_2 on the nightside, due to both chemistry and dynamics; SO_2 on the dayside is destroyed by photolysis, while SO_2 on the nightside is enriched by downwelling motions. This day–night difference of SO_2 in our model agrees with SPICAV occultation observations.

Circulation patterns control the CO local-time pattern over photochemical processes in the upper mesosphere. Above 80 km, CO increases as altitude increases. The upwelling of SS-AS circulation transports the CO-poor air on the dayside, while the downwelling does the opposite on the nightside. This circulation pattern decreases CO on the dayside and increases CO on the nightside. The CO local-time maximum shifts westward from midnight to the morning as altitude decreases in the upper mesosphere. This shift is consistent with microwave observations and is due to the transition from the SS-AS circulation to the RSZ flow. Below 80 km, the CO mixing ratio is nearly constant over space owing to its long chemical loss timescale.

Our models also explain the CO terminator difference observed by SOIR. CO at the morning terminator is more abundant than that at the evening terminator at lower

altitudes, while this pattern is reversed at higher altitudes. The difference at lower altitudes is due to thermal tides combined with the RSZ flows. The difference at higher altitudes might relate to the zonally asymmetric circulation. The reversal level simulated by our models is lower than the SOIR observations. This could indicate that the transition level from RSZ flow to SS-AS circulation in our GCM is lower than that in the Venusian atmosphere. The CO local-time variability could thus be used to constrain the atmospheric circulation of Venus.

H_2O and HCl are long-lived like CO and distribute almost uniformly over both local time and altitude. The uniform distribution of H_2O is qualitatively consistent with the TEXES observations. HCl vertical profiles simulated by our models disagree with JCMT observations and support that SS-AS circulation is unlikely to produce the large decrease of HCl in the upper mesosphere. ClO shows a maximum at 80–95 km on the nightside. OCS is observable in the upper cloud and also exhibits a two-maxima local-time pattern in the upper cloud. SO is a short-lived species whose mixing ratio is larger on the dayside than on the nightside.

The disagreement of RSZ-to-SS-AS transition level between the model and the SOIR data needs further investigation. This transition occurs where the semidiurnal tides dissipate in the upper mesosphere. The thermal tidal waves transport retrograde angular momentum downward to the superrotation region and decelerate the atmosphere above (Mendonça & Read 2016). These waves dissipate/break in the upper layers by radiative damping. Improving the representation of gas absorbers in the upper atmospheric region of the 3D simulations and moving the top of the model domain to higher altitudes might help reduce the disagreement in the RSZ-to-SS-AS transition altitude between the data and the model. The latter will mitigate the inaccuracies due to the top rigid model boundary, which may impact the atmospheric flow in the transition region. Also, moving the top boundary to a higher altitude will diminish the impact of the sponge layer scheme in the model's uppermost layers in the GCM.

Furthermore, in the future, new observations from the Venus missions (DAVINC+, VERITAS, and EnVision) will reveal more spatial and temporal variabilities of chemical species on Venus. To understand these variabilities, the 3D GCM + 3D CTM approach could be a better way than our current approach despite a more expensive computational cost. A future 3D GCM + 3D CTM model set will show how 3D circulations (including meridional circulations) and photochemistry together control species' variabilities in the middle atmosphere of Venus.

This work is supported by NSF grant AST1740921 to X.Z. W.D.S. is supported by the China Scholarship Council Fellowship. We thank Carver J. Bierson for the discussion of the 2D CTM. We also acknowledge use of the lux supercomputer at UC Santa Cruz, funded by NSF MRI grant AST 1828315. We thank two anonymous reviewers for constructive comments on the manuscript.

Appendix A
Meridionally Mean Chemical Transport Equation

Our CTM uses the log-pressure–longitude coordinate to solve the continuity equation. From Chapters 9 and 10 in Andrews et al. (1987), the continuity equation for volume mixing ratio χ of a minor species is

$$(\rho_0\chi)_t + \frac{(\rho_0\chi u)_\lambda + (\rho_0\chi v\cos\phi)_\phi}{a\cos\phi} + (\rho_0\chi w)_z = \rho_0 S. \tag{A1}$$

Here $\rho_0 = \rho_s \exp\{-z/H\}$ is the reference background density and ρ_s and H are the density at a reference level (bottom boundary) and a characteristic scale height that does not vary with height, respectively. t is time. λ, ϕ, and z are longitude, latitude, and height in the log-pressure coordinate, respectively, and u, v, and w are velocities in three directions. a is the planetary radius. $\rho_0 S$ represents chemical production and loss rates. Subscript represents the partial derivative with respect to each coordinate.

Multiplying Equation (A1) by $\cos\phi\, d\phi$, integrating it from one pole to the other pole over the meridional direction, and dividing it by $\int_{-\pi/2}^{\pi/2}\cos\phi d\phi$, we get

$$(\rho_0\overline{\chi})_t + \frac{(\rho_0\overline{\chi u^*})_\lambda}{a} + (\rho_0\overline{\chi w})_z = \rho_0\overline{S}, \tag{A2}$$

where the overbar represents the average of any quantity x over latitude

$$\overline{x} = \frac{\int_{-\pi/2}^{\pi/2} x\cos\phi d\phi}{\int_{-\pi/2}^{\pi/2}\cos\phi d\phi}, \tag{A3}$$

and $u^* = u/\cos\phi$.

Doing the same operation to the continuity equation for the background atmosphere

$$\frac{(\rho_0 u)_\lambda + (\rho_0 v\cos\phi)_\phi}{a\cos\phi} + (\rho_0 w)_z = 0, \tag{A4}$$

we can get a similar expression

$$\frac{(\rho_0\overline{u^*})_\lambda}{a} + (\rho_0\overline{w})_z = 0. \tag{A5}$$

Combining Equations (A2) and (A5) and using

$$\overline{x_1 x_2} = \overline{x_1}\,\overline{x_2} + \overline{x_1' x_2'}, \tag{A6}$$

where x_1 and x_2 are any two quantities and $x_i' = x_i - \overline{x_i}$ $(i = 1, 2)$ is the deviation from the mean, we get

$$(\rho_0\overline{\chi})_t + \frac{\rho_0\overline{u^*}(\overline{\chi})_\lambda}{a} + \rho_0\overline{w}(\overline{\chi})_z = \rho_0\overline{S} - \left\{\frac{(\rho_0\overline{\chi' u^{*'}})_\lambda}{a} + (\rho_0\overline{\chi' w'})_z\right\}. \tag{A7}$$

We can parameterize the deviation term in the curly brackets as diffusion:

$$(\rho_0\overline{\chi})_t + \frac{\rho_0\overline{u^*}(\overline{\chi})_\lambda}{a} + \rho_0\overline{w}(\overline{\chi})_z = \rho_0\overline{S} - \left\{\left(\frac{\rho_0}{a^2}K_{xx}\overline{\chi}_\lambda + \frac{\rho_0}{a}K_{xz}\overline{\chi}_z\right)_\lambda + \left(\rho_0 K_{zz}\overline{\chi}_z + \frac{\rho_0}{a}K_{zx}\overline{\chi}_\lambda\right)_z\right\}. \tag{A8}$$

This is the tracer continuity equation in the log-pressure–longitude plane, derived based on the traditionally defined longitude–latitude coordinate.

Appendix B
Fast Chemical Cycles

When calculating the chemical lifetimes of SO_2, SO, and CO, we exclude some fast chemical cycles. We list these cycles here for a reference.

There are two fast cycles in the SO_2-related network. One involves species Cl_2:

$$\left.\begin{array}{c} 2Cl + 2SO_2 + 2M \rightarrow 2ClSO_2 + 2M \\ 2ClSO_2 \rightarrow Cl_2 + 2SO_2 \\ Cl_2 \rightarrow 2Cl \end{array}\right\}. \quad \text{(B1)}$$

The other involves species SO:

$$\left.\begin{array}{c} Cl + SO_2 + M \rightarrow ClSO_2 + M \\ SO + ClSO_2 \rightarrow OSCl + SO_2 \\ OSCl + M \rightarrow Cl + SO + M \end{array}\right\}. \quad \text{(B2)}$$

For SO, besides the cycle (B2), there is another fast cycle involving the SO dimer:

$$\left.\begin{array}{c} 2SO + M \rightarrow (SO)_2 + M \\ (SO)_2 + M \rightarrow 2SO + M \end{array}\right\}. \quad \text{(B3)}$$

For CO, there are two fast cycles:

$$\left.\begin{array}{c} Cl + CO + CO_2 \rightarrow ClCO + CO_2 \\ ClCO + CO_2 \rightarrow Cl + CO + CO_2 \end{array}\right\} \quad \text{(B4)}$$

and

$$\left.\begin{array}{c} Cl + CO + N_2 \rightarrow ClCO + N_2 \\ ClCO + N_2 \rightarrow Cl + CO + N_2 \end{array}\right\}. \quad \text{(B5)}$$

In Figures 1, 3, and 9, we have shown the chemical lifetime calculations for SO_2, CO, and SO with these fast cycles excluded. Figure B1 shows the chemical lifetime calculations when these cycles are included. Comparing Figure B1 to Figures 1, 3, and 9, we found that including these cycles when calculating the chemical lifetime would give unreasonable results. For example, CO appears short-lived in the upper cloud on the dayside (panel (e) of Figure B1).

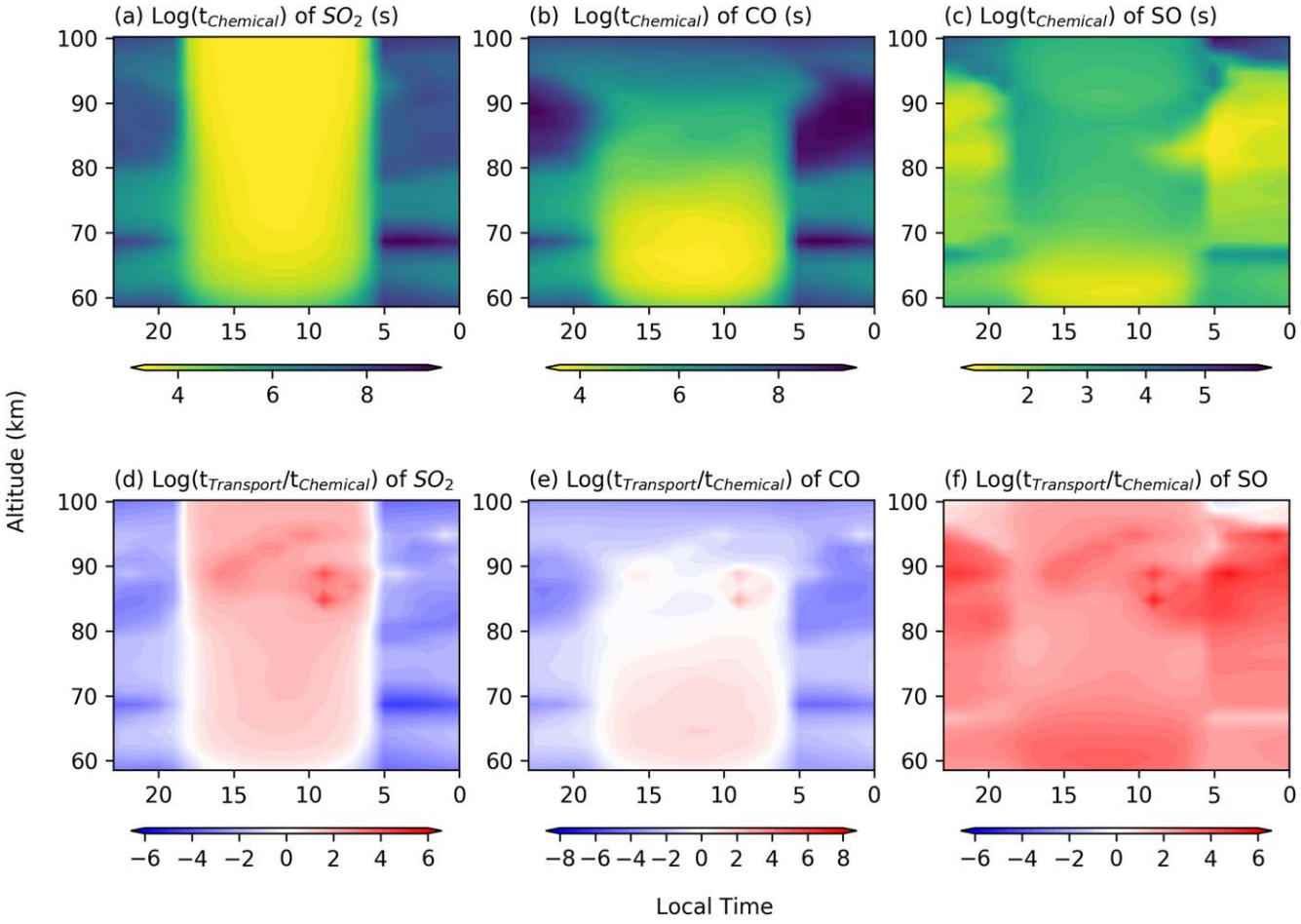


Figure B1. Chemical lifetimes of (a) SO_2, (b) CO, and (c) SO if the fast cycles in Appendix B are included. Panels (d)–(f) are the ratios of $t_{\rm Transport}$ to $t_{\rm Chemical}$ for the three species when the fast cycles are included.

Appendix C
Supporting Data

This appendix provides supporting figures. Figure C1 shows the average temperature profile simulated by our GCM. Figures C2–C8 show the sensitivity of the species concerned in this paper on diffusivity and horizontal resolution.

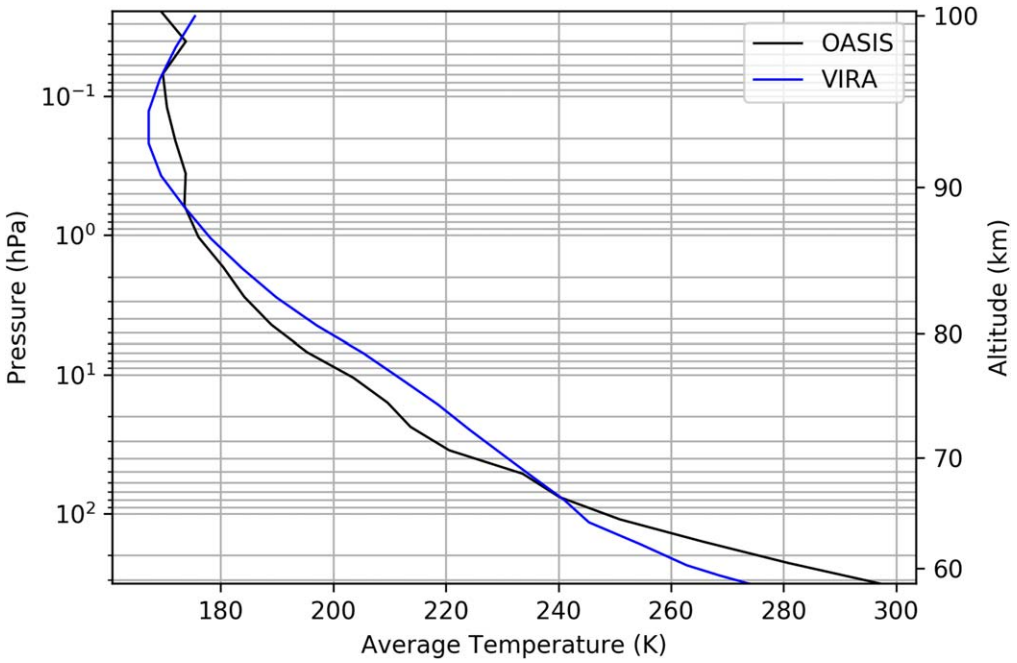


Figure C1. Average temperature profile simulated by our GCM (black) and VIRA temperature profile (blue). The temperature anomaly in Figure 1 is the deviation from this GCM's average temperature profile.

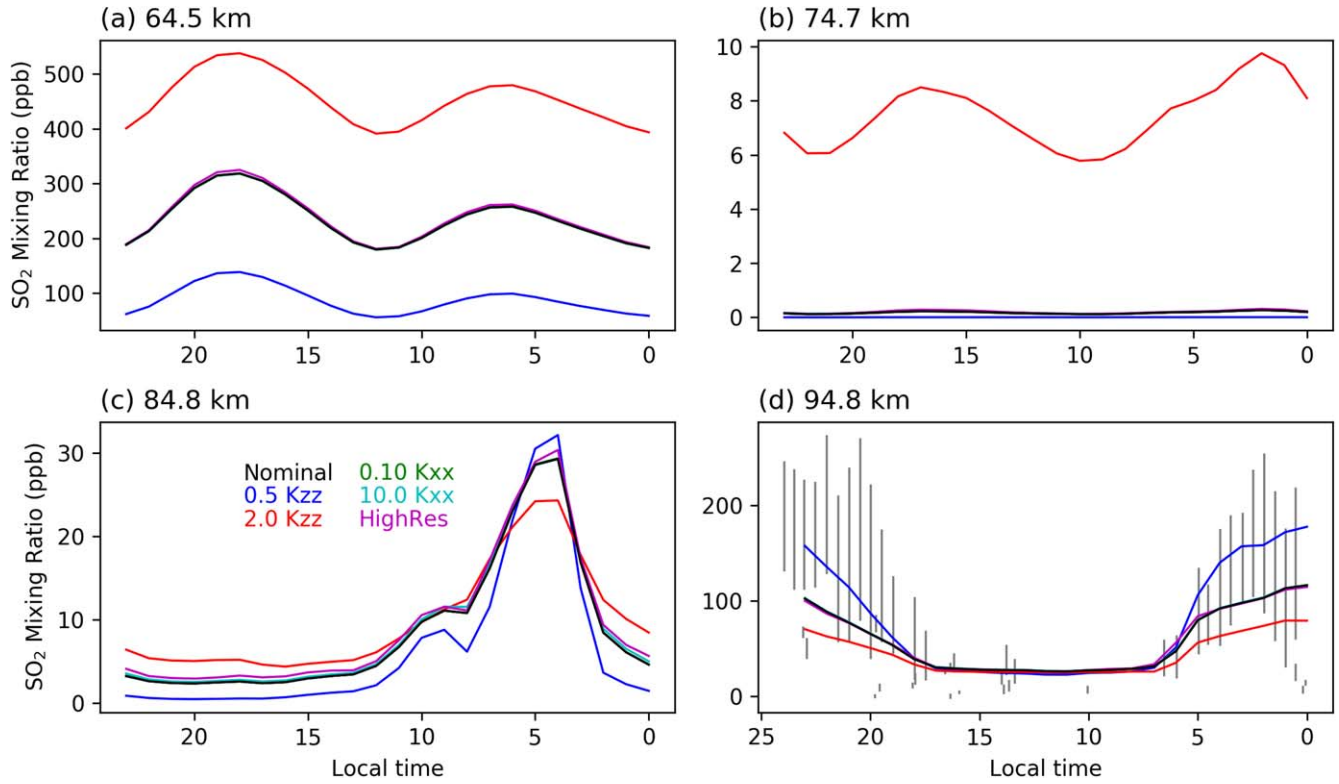


Figure C2. Local-time distributions of SO_2 mixing ratio at different altitudes for different cases: (a) our nominal case (black); (b) cases with K_{zz} enlarged by 1.5 (blue) and 2.0 (red); (c) cases with K_{xx} changed by a factor of 0.1 (green) and 10.0 (cyan); (d) cases with a higher (double) horizontal resolution (magenta). Note that green, cyan, and black lines are almost overlapping with each other. In panel (d), error bars show observations at 95–100 km from SPICAV/VEx by Belyaev et al. (2017) and observations at 70–100 km from JCMT by Sandor et al. (2010).

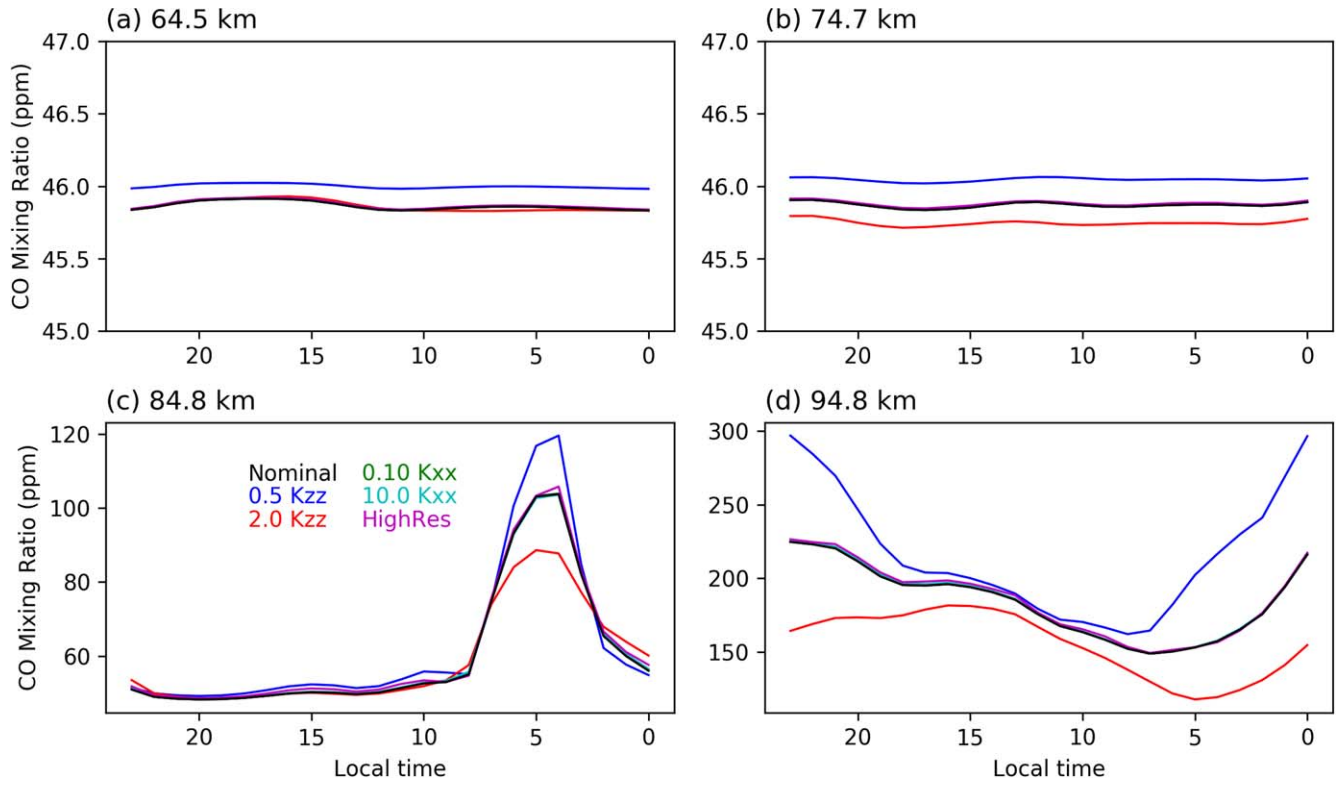


Figure C3. Same as Figure C2, but for CO.

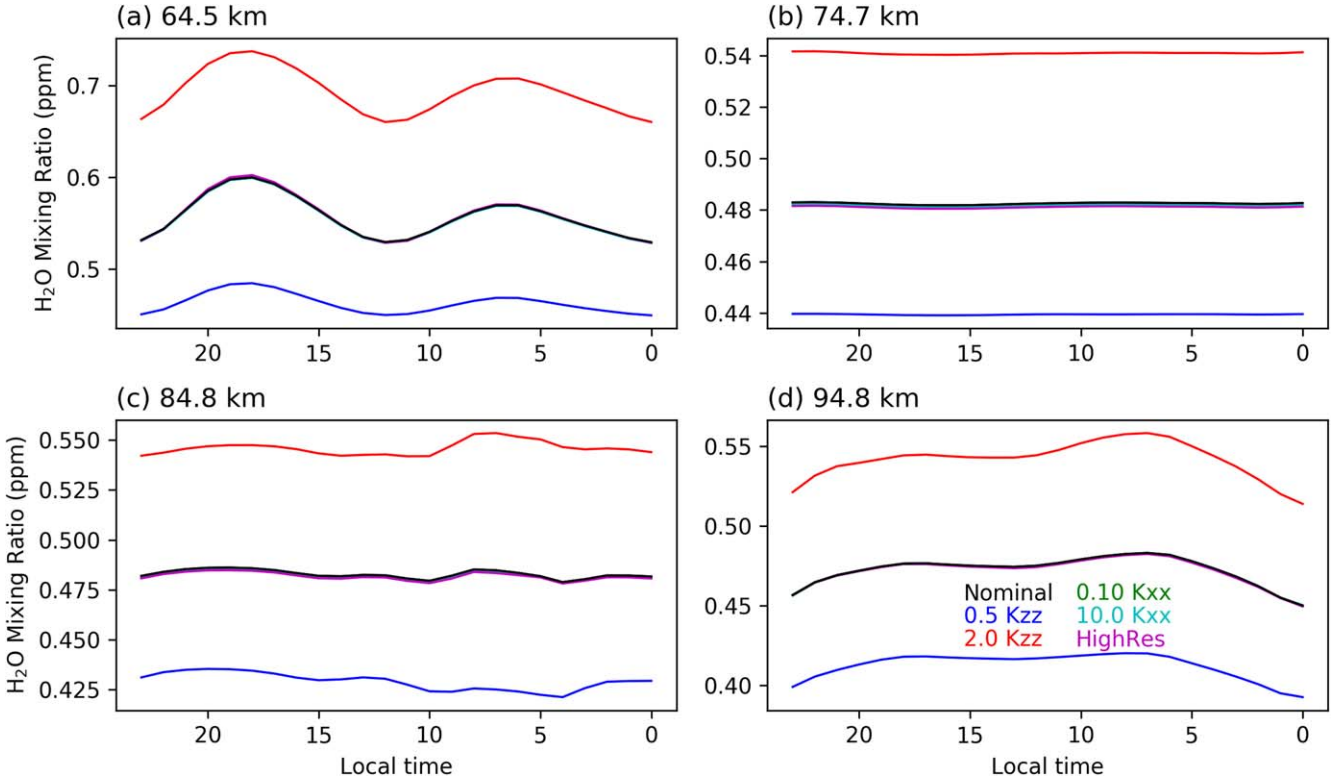


Figure C4. Same as Figure C2, but for H_2O.

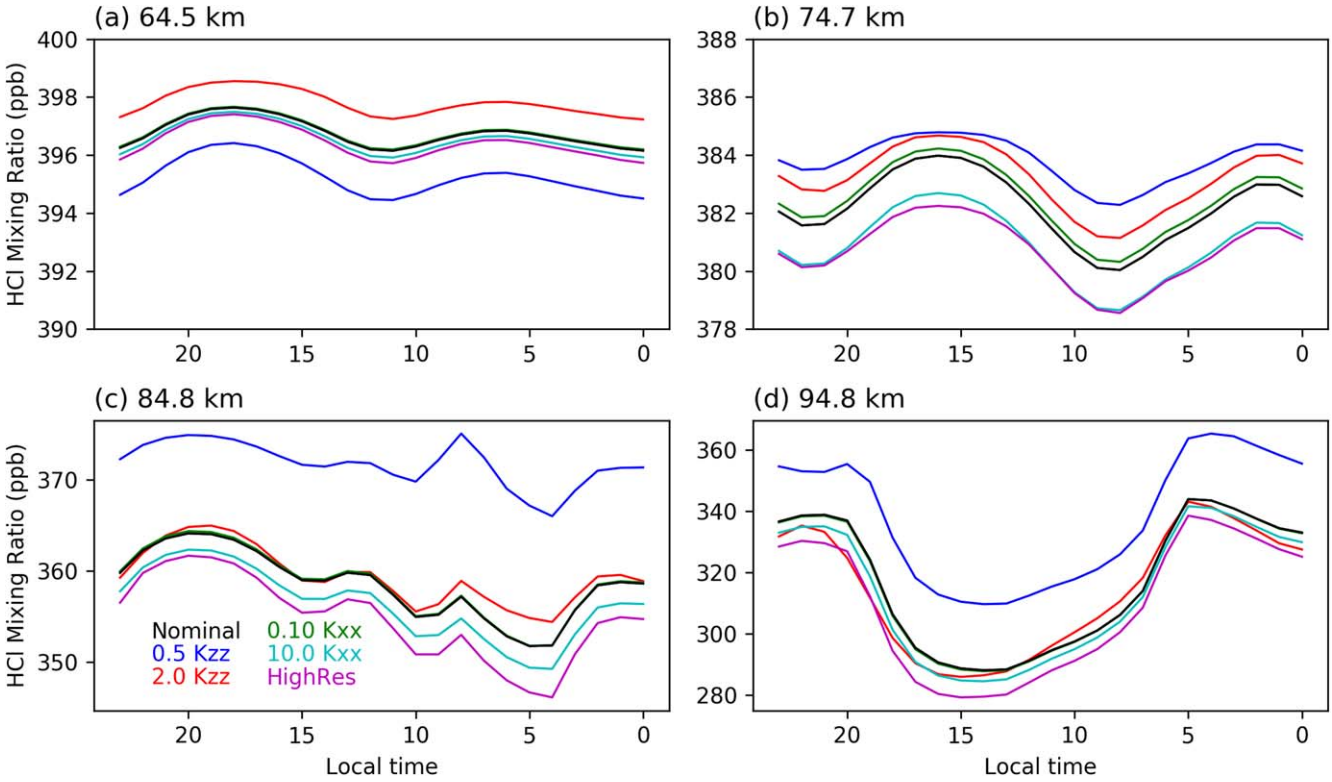


Figure C5. Same as Figure C2, but for HCl.

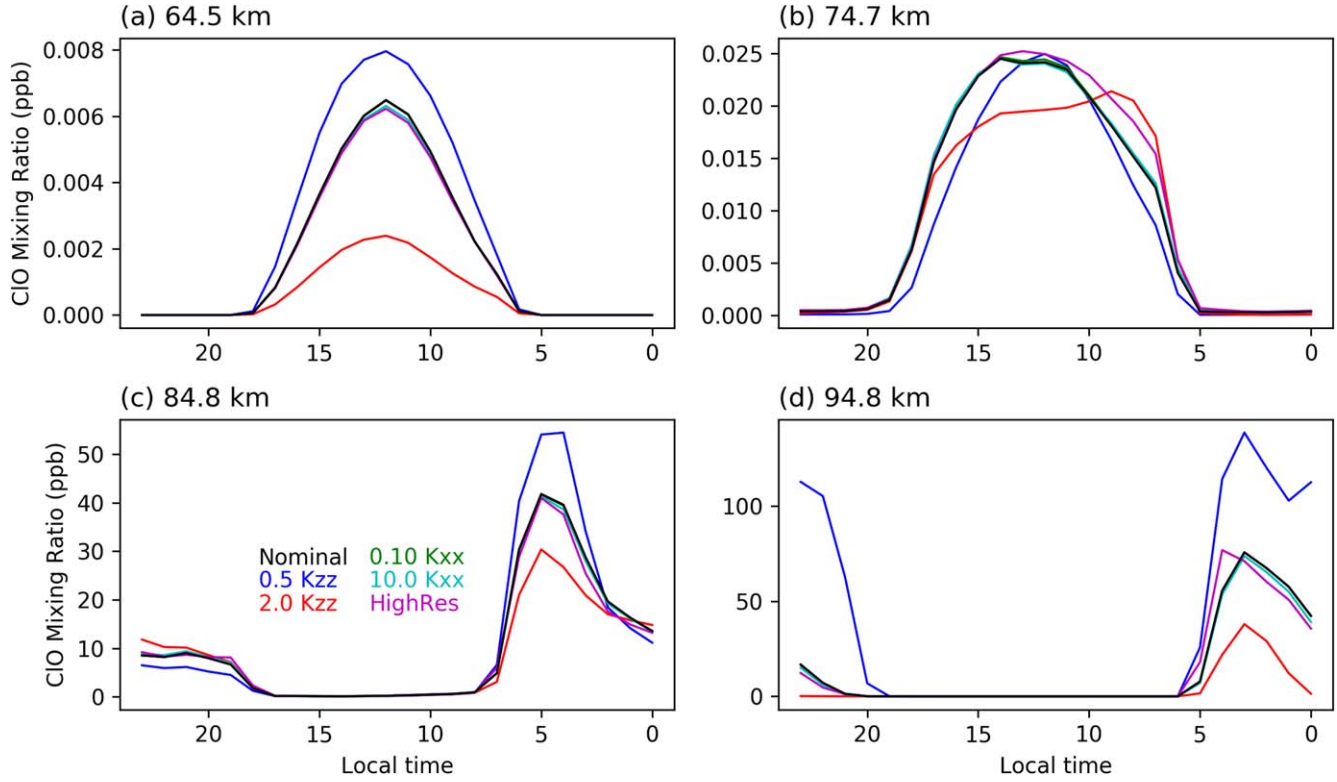


Figure C6. Same as Figure C2, but for ClO.

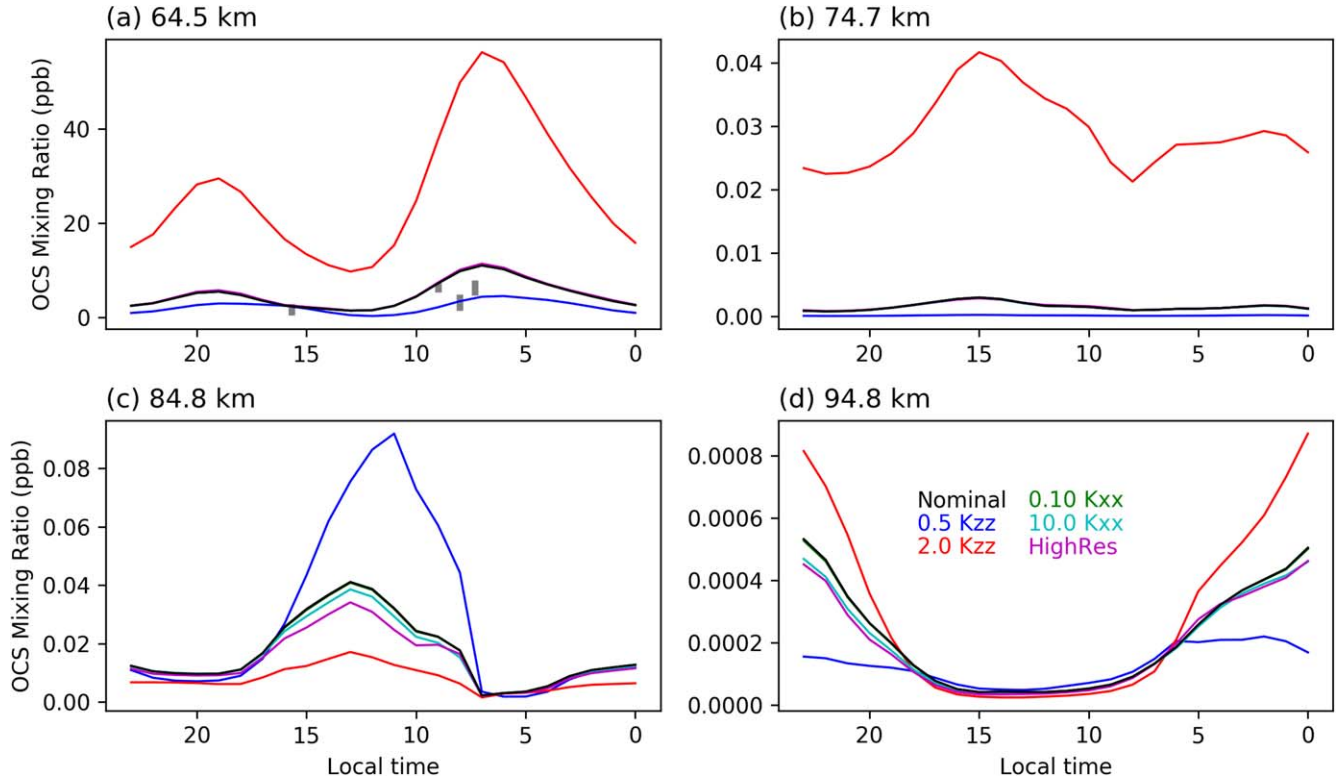


Figure C7. Same as Figure C2, but for OCS. In panel (a), gray bars show a few observation points near 65 km from CSHELL/IRTF by Krasnopolsky (2010).

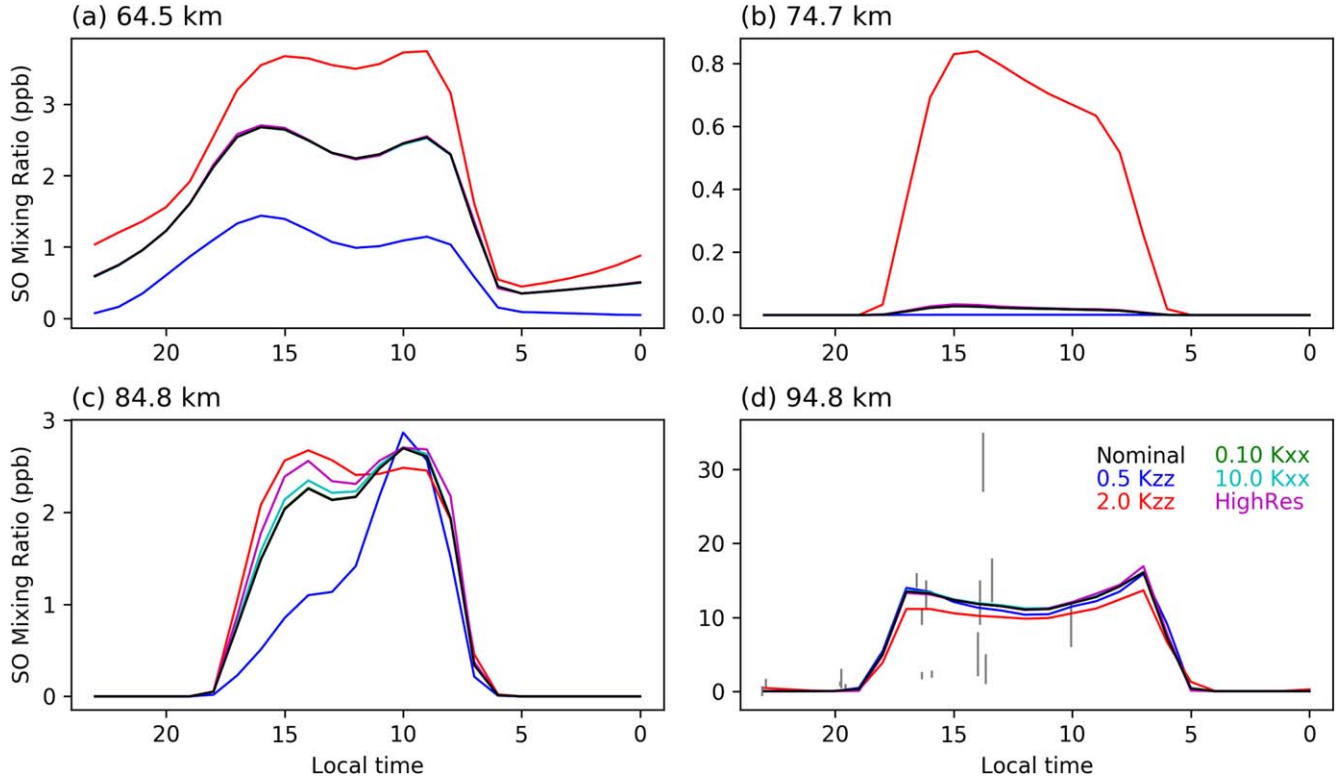


Figure C8. Same as Figure C2, but for SO. In panel (d), error bars show observations at 70–100 km from JCMT by Sandor et al. (2010).

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