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Simulating the anisotropic clustering of luminous red galaxies with subhaloes: a direct confrontation with observation and cosmological implications

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ABSTRACT

We model the apparent clustering anisotropy of luminous red galaxies (LRGs) in the Sloan Digital Sky Survey using subhaloes identified in cosmological N -body simulations. We first conduct a Markov chain Monte Carlo analysis on the parameters characterizing subhaloes hosting LRGs assuming a specific Λ cold dark matter cosmology on which we run the simulations. We show that simple models with central and satellite subhaloes can explain the observed multipole moments of the power spectrum up to hexadecapole on large scales ($k \lesssim 0.3 h \text{Mpc}^{-1}$). A satellite fraction of 20–30 per cent is favoured weakly depending on the detail of the model. The fraction is shown to be robust when we adopt a more refined model based on the halo occupation number from the literature. We then vary cosmological parameters controlling the anisotropy in redshift space effectively by deforming the simulation box (the Alcock–Paczynski effect) and changing the amplitude of the velocities (the redshift-space distortions). We demonstrate that we can constrain the geometry of the Universe, the structure growth rate and the parameters characterizing LRGs simultaneously. This is a step towards cosmological analysis with realistic bias description beyond empirical bias functions with nuisance parameters.

Key words: methods: numerical – cosmology: theory – large-scale structure of Universe.

1 INTRODUCTION

Understanding the nature of a population of galaxies is the key to derive unbiased cosmological constraints using their spatial clustering pattern. This issue of galaxy bias is one of the biggest obstacles in the modern cosmology (Kaiser 1984). In other words, selecting a controlled sample of galaxies whose environmental properties are well understood is an appropriate way for this purpose. Luminous red galaxies (LRGs) collected by the Sloan Digital Sky Survey (SDSS; York et al. 2000) are believed to be such a galaxy sample (Eisenstein et al. 2001): a population mostly composed of central galaxies associated with massive haloes. This understanding can be inferred from a number of observational facts such as a low number density, a large fraction of single-LRG systems (Reid & Spergel 2009) and a high bias factor at large scale relative to the clustering amplitude of the underlying matter distribution (Eisenstein et al. 2005). All these observed features make them useful for cosmological applications, and indeed a lot of important cosmological implications have been derived using this sample including the first

clear detection of baryon acoustic oscillations (BAOs) by Eisenstein et al. (2005).

One of the most popular analytical approaches to model the clustering of galaxies is so-called the halo occupation distribution (HOD) approach based on the halo model (e.g. Ma & Fry 2000; Peacock & Smith 2000; Seljak 2000; Scoccimarro et al. 2001; Berlind & Weinberg 2002). One assumes that all the observed galaxies live in haloes, and one can calculate the clustering properties once the probability of having N galaxies in a halo with mass M_{host} is given. One usually assumes a simple functional form for the (mean) halo occupation number as a function of the halo mass, and determines the model parameters by fitting to some observed properties. In particular, the HOD parameters have been investigated for LRGs based on the spatial clustering on relatively small scales (Kulkarni et al. 2007; White et al. 2007; Blake, Collister & Lahav 2008; Wake et al. 2008; Padmanabhan et al. 2009; Zheng et al. 2009), galaxy–galaxy lensing (Mandelbaum et al. 2006; Hikage et al. 2013) or a direct measurement of the number distribution of LRGs forming groups (Ho et al. 2009; Reid & Spergel 2009). LRGs may also be able to be modelled with simulated subhaloes employing abundance matching schemes (e.g. Conroy, Wechsler & Kravtsov 2006; Masaki et al. 2013).

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On another front, there have been a lot of attempts to model the statistics of galaxy spatial clustering at the scale of BAOs based on perturbation theory (PT; see Bernardeau et al. 2002 for a thorough review of the standard PT and also Jeong & Komatsu 2006; Eisenstein, Seo & White 2007; Nishimichi et al. 2007; Jeong & Komatsu 2009; Padmanabhan & White 2009; Sherwin & Zaldarriaga 2012; McCullagh et al. 2013 for its application to BAOs) beyond linear theory in the light of ongoing/near future galaxy redshift surveys. Since these survey projects aim at a precise determination of cosmological distances, the required accuracy in the theoretical modelling is highly demanding. In such a situation, a number of ‘renormalized’ PT techniques have been developed to have a better convergence of perturbative series expansions on statistical quantities at the scale of interest (Crocce & Scoccimarro 2006a,b, 2008; McDonald 2007; Matarrese & Pietroni 2007, 2008; Valageas 2007; Bernardeau, Crocce & Scoccimarro 2008; Matsubara 2008a,b; Pietroni 2008; Taruya & Hiramatsu 2008; Hiramatsu & Taruya 2009; Anselmi, Matarrese & Pietroni 2011; Okumura, Taruya & Matsubara 2011; Sugiyama & Futamase 2012a,b; Bernardeau, Van de Rijt & Vernizzi 2013; Sugiyama & Spergel 2013; Valageas, Nishimichi & Taruya 2013). It is also worth mentioning the importance of cosmological N -body simulations in developing these analytical models. Since all the perturbative schemes involve some approximation, ansatz or truncation, and moreover, there are fundamental limitations of perturbative approaches such as the breakdown of the single-stream approximation after shell crossing (Valageas 2011, 2013), one has to confirm the validity of one’s scheme and determine with care the applicable range in wavenumber and redshift by testing with fully non-linear predictions based on simulations (Crocce & Scoccimarro 2008; Carlson, White & Padmanabhan 2009; Nishimichi et al. 2009; Taruya et al. 2009; Sato & Matsubara 2011; Valageas & Nishimichi 2011). Recently, some codes to evaluate the non-linear power spectrum of the cosmic density field very rapidly based on new perturbative approaches have been made publicly available (Crocce, Scoccimarro & Bernardeau 2012; Taruya et al. 2012). These tools are practically useful in confronting the analytical predictions with observational data.

Despite all the recent progress, galaxy bias is still difficult to implement into these new techniques in a consistent manner without losing the non-perturbative properties in those theories (some recent attempts along this line can be found in e.g. McDonald 2006; Matsubara 2008b; Nishizawa, Takada & Nishimichi 2013). Also, the effect of one-halo term (or, equivalently, satellite galaxies) could be significant on the power spectrum especially in redshift space as recently suggested by Hikage & Yamamoto (2013), which is beyond the scope of PT calculations. Thus it is not straightforward to analyse galaxy clustering with renormalized PT techniques and extract cosmological information robustly, even if the environmental properties of the observed galaxy sample, such as those of LRGs mentioned above, are very well understood.

The purpose of this study is to see whether the current state-of-the-art N -body simulations can explain the clustering of galaxies (LRGs, more specifically) in redshift space on large scales where most of the cosmological information exists. Also, we explore the possibility of extracting cosmological information by confronting simulations with observation instead of using an analytical model. Studies along this line are not straightforward for various reasons. One is from the fact that we simulate *one realization* of the cosmological random field in finite volume drawn from an assumed cosmological model while we observe *another realization* under the correct cosmology. Another is the high computational cost to cover the multidimensional cosmological parameter space with sim-

ulations. Since this parameter space has typically as many as six dimensions in the standard Λ cold dark matter (Λ CDM) model (and even more when one wishes to test some non-standard models), this is not realistic with high-resolution simulations with sufficient volume. Studies on the rescaling of the simulation outputs to different cosmological models can be found in e.g. Tormen & Bertschinger (1996), Cole (1997), Angulo & White (2010), Mead & Peacock (2013), and its applications to extract cosmological information or to infer the cosmological model dependence from galaxy properties are performed in Simha & Cole (2013) and Guo et al. (2013).

We partly overcome these difficulties in this study by employing the following methodology. First, by taking an ensemble average over different random realizations of simulations whose total volume is much larger than the observed volume, we obtain a well converged prediction that can directly be compared with observation as one does with analytical models. We next introduce three parameters, which affect the apparent *anisotropy* of the clustering and are closely related to some cosmological parameters, and float them to see their impact on the power spectrum without re-running new simulations starting from new initial conditions.

Our new parameters are responsible for the growth rate of the cosmic perturbations, the Hubble parameter and the angular diameter distance at the effective redshift of the observed galaxies. The first one can be observed through the redshift-space distortions (RSDs) caused by peculiar velocities of galaxies (Kaiser 1987). Our first parameter, that scales the amplitude of velocities in simulation outputs, amplifies or suppresses the magnitude of RSDs. Since the significance of the anisotropy induced by RSDs is a good indicator of the growth rate of the cosmic structure, $f\sigma_8$, where $f \equiv d \ln D_+ / d \ln a$ with D_+ being the linear growth factor, we can test the underlying gravity theory by measuring it (e.g. Percival et al. 2004; Guzzo et al. 2008; Linder 2008; Yamamoto, Sato & Hütsi 2008; Song & Percival 2009; Blake et al. 2011a; Beutler et al. 2012; Reid et al. 2012; Samushia et al. 2013). The other two parameters deform the simulation box and induce apparent anisotropy to the clustering of the mock galaxies. By doing this, we simulate the Alcock–Paczynski (AP) effect (Alcock & Paczynski 1979), which should be there in the observed clustering if the cosmological model assumed in the conversion of redshifts to the three-dimensional positions has a mismatch with the true one. One can determine the Hubble parameter, H , and the angular diameter distance, D_A , through this effect combined with the characteristic scale of BAOs (e.g. Ballinger, Peacock & Heavens 1996; Matsubara & Suto 1996; Hu & Haiman 2003; Okumura et al. 2008; Padmanabhan & White 2008; Shoji, Jeong & Komatsu 2009; Blake et al. 2011b, 2012; Taruya, Saito & Nishimichi 2011; Chuang & Wang 2012, 2013a,b; Reid et al. 2012; Kazin et al. 2013; Sánchez et al. 2013; Xu et al. 2013; Anderson et al. 2014). Although more involved approaches such as one in Angulo & White (2010) may rescale the simulations more accurately, our simple method is computationally very easy to implement and can be safely applied as far as one focuses on the anisotropy of the clustering.

An important question that we would like to ask here is whether we can distinguish these distortions with uncertainties in the modelling of galaxies, in particular velocities of galaxies. Adopting different prescriptions for the LRG–subhalo connection, we show how much the resultant cosmological constraints are affected. In the accompanying paper (Oka et al. 2014), we present a similar analysis using the same observed data set but with an analytical model for the non-linear galaxy power spectrum in redshift space. In the model, the galaxy bias and the Fingers of God suppression of the power spectrum (Jackson 1972) are modelled by rather simple functional

forms with free parameters without much physical justification at the quantitative level. On the other hand, we expect that our model based on simulated subhaloes physically describes these effects in a fully non-linear manner as long as the gravitational interaction is concerned. Given these differences, it would also be interesting to discuss the consistency between the cosmological parameters derived with two fundamentally different prescriptions for LRGs.

This paper is organized as follows. We first briefly explain the observed power spectrum in Section 2. We then discuss our methods to model LRGs using simulated subhaloes in Section 3. Our main results are presented in Section 4. We compare our cosmological constraints with those in Oka et al. (2014) to check the consistency of the two analyses in that section. A comparison with a model based on the observed HOD is also discussed. We summarize the findings of this study and add some discussion on the possible future generalization of our method towards a fully consistent cosmological parameter estimation in Section 5.

2 DATA

In this section, we briefly explain the galaxy sample analysed in this study and the measurement of the power spectrum done in Yamamoto et al. (2010). The LRGs used in the measurement is drawn from the SDSS Data Release 7 (Abazajian et al. 2009). 100 157 LRGs with spectra in the redshift range of $0.16 < z < 0.47$ are selected from the Northern Galactic Cap covering $\sim 7150 \text{ deg}^2$. The anisotropic power spectrum is expanded into multipole moments:

$$P(k, \mu_k) = \sum_{\ell=\text{even}} \mathcal{P}_\ell(\mu_k) P_\ell(k), \quad (1)$$

where μ_k is the directional cosine of the wave vector with respect to the line of sight and $\mathcal{P}_\ell(\mu_k)$ denotes the Legendre polynomial. The multipole moments of the power spectrum, $P_\ell(k)$, are measured with the procedure developed in Yamamoto et al. (2006) at the effective redshift of $z = 0.3$. We analyse the moments up to hexadecapole ($\ell = 4$) on large scales ($k \leq 0.305 h \text{ Mpc}^{-1}$). The estimation of the statistical error on the measured spectrum is also described in Yamamoto et al. (2006), in which they employ a method based on Feldman, Kaiser & Peacock (1994) assuming Gaussianity of the density field. Although different multipole moments have non-zero cross-covariance even when the underlying density field obeys the Gaussian statistics, we ignore it in this study for simplicity. At leading order, this contribution is proportional to $\beta = f/b$, where b is the linear bias factor, and is less important for highly biased tracers (Taruya et al. 2011). Since the LRGs used in this analysis have bias as large as ~ 2 , ignoring the cross-covariance would not affect the final result significantly.

Although we basically analyse the power spectrum measured from the same data set as in our accompanying paper, Oka et al. (2014), we employ a different assumption in the underlying cosmology when redshifts of galaxies are converted to distances. In this study, we adopt a flat Λ CDM cosmology with $\Omega_m = 0.28$ and $h = 0.7$ that is exactly the same as in Yamamoto et al. (2010), while Oka et al. (2014) re-measure the power spectrum with $\Omega_m = 0.32$ and $h = 0.67$ motivated by the recent *Planck* result (Planck Collaboration XVI 2013). Since the cosmological model used in the simulations is closer to the former, we simply adopt the original measurement by Yamamoto et al. (2010). See Table 1 for cosmological models discussed in this study. In Section 4, we discuss the consistency between the derived cosmological parameters and those assumed here (and in the N -body simulations) and propose a

possible iterative scheme to perform a fully consistent cosmological analysis in Section 5.

3 MODELLING LRGs WITH SIMULATIONS

Here we describe our model of mock LRGs constructed from cosmological N -body simulations. After showing the detail of the simulations and subhaloes in Section 3.1, we explain how we connect them and mock LRGs in Section 3.2. We then summarize the method to measure the model power spectrum and to fit to the observed data in Section 3.3. Our method to simulate the cosmological dependence of the apparent anisotropy is described in Section 3.4.

3.1 Simulations and subhalo identification

The cosmological N -body simulations used in this study are performed in Nishimichi & Taruya (2011). Employing 1280^3 collisionless particles in periodic cubes with the side length of $1144.72 h^{-1} \text{ Mpc}$, we simulate the gravitational growth of structure with a publicly available tree-PM code, GADGET2 (Springel 2005). The initial conditions are set by a code developed in Nishimichi et al. (2009) and parallelized in Valageas & Nishimichi (2011) based on the second-order Lagrangian perturbation theory (2LPT; e.g. Scoccimarro 1998; Crocce, Pueblas & Scoccimarro 2006). We assume a flat Λ CDM universe with the parameters derived by the 5-year observation by *Wilkinson Microwave Anisotropy Probe* (WMAP) satellite (Komatsu et al. 2009; WMAP5+BAO+SNALL in the reference, and see also Table 1) to compute the linear power spectrum using CAMB (Lewis, Challinor & Lasenby 2000). 15 independent random realizations are simulated and snapshots at $z = 0.35$ are stored.

Unfortunately, the spatial distribution of dark matter particles are available only for 11 realizations out of 15 due to a problem in our hard disc. Using the remaining 11 realizations, we identify subhaloes using an independent implementation of SUBFIND algorithm (Springel et al. 2001). In the code, we first find friends-of-friends (FoF; e.g. Davis et al. 1985) groups and then search for gravitationally bound particles inside each FoF group. In this paper, we conventionally refer to the most massive subhalo in a FoF group as a *central*, while the rest of the subhaloes are called as *satellites*. We keep all the subhaloes with mass larger than $5.5 \times 10^{11} h^{-1} M_\odot$ ($10 N$ -body particles) hosted by haloes larger than $1.8 \times 10^{12} h^{-1} M_\odot$. These subhaloes are used to reproduce the anisotropic clustering of observed LRGs in what follows.

3.2 Connecting subhaloes to LRGs

Parameters describing the properties of galaxies such as the HOD are often discussed using clustering measures on relatively small scales or one-point statistics. They are discussed usually based on a specific cosmological model. On the other hand, the large-scale clustering, often aiming at extraction of cosmological information, is expected to be insensitive to the detail of the nature of galaxies. For example, a scale-independent linear bias model might be fine, though not fully validated, at the large-scale limit. Since we here discuss the clustering in linear to weakly non-linear regime, we wish to avoid introducing many parameters and/or employing a specific functional form for bias such as those for the HOD. The determination of the detailed model parameters must be difficult unless we extend the analysis to sufficiently smaller scales. We employ rather simple models based on LRG-(sub)halo connection and discuss the validity with the goodness-of-the-fit to the observed data.

Table 1. Summary of the cosmological parameters of flat Λ CDM model. The first row shows the model assumed for the distance–redshift relation in the measurement of the power spectrum. The parameters in the second row are adopted in the N -body simulations. The third to the sixth rows are the maximum likelihood parameters from CMB observations without combining with other experiments, and they are plotted in Figs 11 and 13 to see the consistency with our analysis.

	Ω_m	h	σ_8	r_s^a	H^b	D_A^c	$f\sigma_8^d$	Reference
Dist.– z rel.	0.28	0.7	–	–	80.88	922.81	–	Yamamoto et al. (2010)
Sim.	0.279	0.701	0.817	104.02	80.88	921.69	0.462	Nishimichi & Taruya (2011)
WMAP5	0.249	0.725	0.787	108.74	82.57	897.51	0.420	Komatsu et al. (2009)
WMAP7	0.271	0.703	0.801	105.08	80.88	920.91	0.448	Komatsu et al. (2011)
WMAP9	0.282	0.697	0.820	103.46	80.59	926.46	0.461	Hinshaw et al. (2013)
Planck	0.318	0.671	0.834	98.87	78.84	954.68	0.489	Planck Collaboration XVI (2013)

^aThe sound horizon scale at the baryon drag epoch in h^{-1} Mpc.

^bThe Hubble parameter in $\text{km s}^{-1} \text{Mpc}^{-1}$ evaluated at $z = 0.3$.

^cThe angular diameter distance in Mpc at $z = 0.3$.

^dThe growth rate parameter, f , multiplied by σ_8 . The value is computed at the redshift of the simulation output, $z = 0.35$, for the model adopted in the simulations (second row), while this parameter is evaluated at $z = 0.3$, the effective redshift of the survey, for the other models.

In this study, we examine several different descriptions of mock LRGs using subhaloes identified above. In each of the five models below, we have two options: the model ‘a’ assigns the centre-of-mass position and velocity of a subhalo to a mock LRG, while in the model ‘b’ we regard the most bound N -body particle in a subhalo as a mock LRG. While the mock LRGs in these two models have almost the same position, the velocity can significantly be different. The modelling of velocities of mock LRGs is crucially important in analysing the anisotropic clustering because it directly affects the signal of RSDs, and the two models investigated here serve as the two extreme cases; the motion of LRGs perfectly coincides with that of their host subhaloes in model ‘a’, while the model ‘b’ takes account of the relative motion of mock LRGs with respect to their hosts. Naively, we expect that the velocity of a LRG aligns to that of the host subhalo. However, assigning the centre-of-mass velocities to mock LRGs (i.e. model ‘a’) is already a big assumption. We thus supplementally infer the results of model ‘b’ to see the impact of an unexpected velocity component, if exists, especially on the cosmological constraints. In what follows we employ notations such as ‘Model 3b’ to label the total of 10 models.

The five models we test in this study are as follows.

Model 1: centrals only. The first model assumes one-to-one correspondence between a mock LRG and a FoF group in the simulations. Mock LRGs populate only in the most massive subhaloes (i.e. centrals) in the host FoF groups in this model. The only model parameter is the minimum host halo mass, $M_{\text{host, min}}$. This model is motivated by the observational fact that about 95 per cent of LRGs do not have close companions (Reid & Spergel 2009). Though this model might be too naive, it is still interesting to see how well we can reproduce the observed anisotropy of LRGs with the central population alone.

Model 2: centrals + satellites. A natural extension of Model 1 is to add satellites on top of centrals. In addition to the parameter $M_{\text{host, min}}$ in Model 1, this model has the second parameter, $M_{\text{sub, min}}$, the minimum mass of subhaloes needed to host mock LRGs. The former mainly controls the environment of mock LRGs and thus the strength of bias, while the latter effectively determines the fraction of satellites. Model 2 is the baseline model in this study.

Model 3: different criteria for centrals and satellites (minimum mass). The next model we consider treats centrals and satellites differently. Since we expect that these two species of subhaloes have different velocity structure, we expect that the fraction of satellites might be important to explain the observed anisotropic power spectrum in redshift space. In order to test the robustness of the satellite fraction derived with Model 2, we introduce different minimum subhalo masses for the two species: $M_{\text{cen, min}}$ for centrals and $M_{\text{sat, min}}$ for satellites. This model has three parameters in total ($M_{\text{host, min}}$, $M_{\text{cen, min}}$ and $M_{\text{sat, min}}$).

Model 4: different criteria for centrals or satellites (random sampling). As in Model 3, the fourth model also assigns mock LRGs to centrals and satellites in a different manner. This time, we assign mock LRGs only to subhaloes randomly chosen out of those that satisfy the condition set by $M_{\text{host, min}}$ and $M_{\text{sub, min}}$. Although we can in principle conduct a random sampling to both of centrals and satellites with different probabilities, we do so only either population at one time and keep all the subhaloes of the other population to reduce the shot noise. This is because the resultant power spectrum should be unchanged when we conduct a random sampling equally to centrals and satellites, and the only effect is an amplified shot noise, because this operation is a Poisson sampling. After some tests, we find that we cannot find a better fit when we randomly discard satellite subhaloes while assigning mock LRGs to all the centrals. We thus conduct random sampling only to the central population. The probability of a central subhalo to host a mock LRG, p_{cen} , is the third parameter of this model.

Model 5: centrals + satellites, cosmology varied. Finally, we fit the observed power spectrum varying the cosmological parameters. We adopt a model similar to the baseline model, Model 2. In addition to the two parameters in Model 2, we simultaneously vary three parameters that will be introduced shortly in Section 3.4. We investigate with this model the robustness of the result obtained with Model 2 for the parameters that characterize the properties of LRGs. At the same time, we discuss the prospects to simultaneously determine the cosmological parameters together with the model parameters in Model 2, and the possible systematic bias in the derived cosmological parameters when we misunderstand the nature of galaxies.

3.3 Determining the model parameters with the Markov chain Monte Carlo method

For a given mock LRG catalogue specified with the model parameters described above, we measure the multipole moments of the power spectrum, evaluate the goodness of fit to the observed data and finally put constraints on the parameters as described in this section.

We first explain the measurement of the multiple moments of the power spectrum. We adopt the distant observer approximation and displace the mock LRGs as

$$\mathbf{s} = \mathbf{x} - f u_z \hat{\mathbf{z}}, \quad (2)$$

where \mathbf{x} and \mathbf{s} , respectively, denote the position of a mock LRG in real and redshift space, $u_z = -v_z/(aHf)$ is the z component of the normalized velocity and $\hat{\mathbf{z}}$ is the unit vector along the line of sight. We then construct a density field in redshift space on a regular lattice with 1024^3 grid points using the nearest grid point (NGP) interpolation scheme (Hockney & Eastwood 1981). We have 11 density fields, $\delta^{(n)}$, from 11 random realizations given the model parameters, where $n = 1, \dots, 11$. After applying the fast Fourier transformation (FFT) to $\delta^{(n)}$, we correct for the window effect arising from the NGP interpolation by dividing $\delta_k^{(n)}$ by an appropriate window kernel (e.g. Jing 2005). We record the square of the density contrast multiplied by an appropriate weight

$$\hat{P}_\ell^{(n)}(\mathbf{k}) \equiv \frac{2\ell + 1}{2} \mathcal{P}_\ell(\mu_k) \left| \delta_k^{(n)} \right|^2 \quad (3)$$

for $\ell = 0, 2$ and 4 . Note that we subtract the shot noise contribution, the inverse of the number density of the mock LRGs, from the monopole moment. We store all the modes up to $k < 0.31 h \text{Mpc}^{-1}$, and the total number of available modes from the 11 realizations is about $(3.77 \times 10^5) \times 11 \simeq 4.15 \times 10^6$ in Models 1, 2, 3 and 4 where we do not consider the AP distortion, and it changes with the model parameters in Model 5.

Once a set of \hat{P}_ℓ is measured, we compare it with the observed multipole moments. We fit \hat{P}_ℓ with the cubic B-spline function (note that this procedure is not an interpolation but a fit) to have smooth predictions of the multipole moments of the power spectrum and evaluate them at the exact wavenumbers where observational data are available. The breakpoints to construct the B-spline function are chosen so as not to smooth the multipoles too much to erase the feature of BAOs (see Appendix A). This procedure greatly reduces the artificial effect arising from the discrete sampling of modes in Fourier space along the μ_k direction, which adds a significant noisy pattern depending on the box size of the simulation especially to the higher multipoles (i.e., $\ell = 2$ and 4 ; see Nishimichi & Taruya 2011; Taruya et al. 2012). See Appendix A for more detail.

After the multipole moments of the power spectrum are obtained, we evaluate the goodness of fit defined as

$$\chi^2 \equiv \sum_{\ell=0,2,4} \sum_{k < k_{\text{max}}} \left[\frac{P_{\ell, \text{mock}}(k) - P_{\ell, \text{obs}}(k)}{\Delta P_{\ell, \text{obs}}(k)} \right]^2, \quad (4)$$

where $P_{\ell, \text{mock}}$ and $P_{\ell, \text{obs}}$, respectively, denote the simulated and observed power spectra, and $\Delta P_{\ell, \text{obs}}$ is the statistical error (i.e. the standard deviation) on the observed spectrum. We simply neglect the statistical error on the model power spectrum because it is small (typically 10 per cent of $\Delta P_{\ell, \text{obs}}$; see Appendix A).

We search for the parameter set that gives the smallest χ^2 employing a Markov chain Monte Carlo (MCMC) method. Since we have to perform the FFT to the 11 density fields and subsequently the B-spline fitting to a lot of data points ($\sim \mathcal{O}(10^6)$) at each chain

of the MCMC, the computational cost is quite expensive. It takes about 1 min with a processor (16 cores) in the Cray XC30 cluster to evaluate the power spectrum for a given model parameter, and we need up to 100 000 chains to have a converged result depending on the model. We can do this analysis in about 1 week using 10 processors parallelly.

3.4 Simulating the redshift-space distortions

If one wishes to constrain cosmological parameters, one may need to prepare theoretical predictions for the whole parameter space of interest. This is computationally unfeasible with N -body simulations and, as a matter of fact, we have a converged prediction of the power spectrum only for one cosmological model. However, we can still simulate the anisotropy of the clustering expected for different cosmologies using that simulation data set for one particular cosmological model.

We start with the positions and velocities of mock LRGs given a set of model parameters described in Section 3.2. We then introduce one parameter that controls the amplitude of RSDs. We denote it by α_v , and we map the positions of mock LRGs from real to redshift space as

$$\mathbf{s} = \mathbf{x} - \alpha_v f u_z \hat{\mathbf{z}}, \quad (5)$$

instead of equation (2). The distribution of mock LRGs is still given in a cube up to here. We then deform this periodic cube according to two parameters, α_{\parallel} and α_{\perp} . Namely, we stretch the simulation box by a factor of α_{\parallel} (α_{\perp}) along the line-of-sight (plane-of-sky) direction under the distant observer approximation. We then measure the multipoles of the power spectrum with the same procedure as before, but from 11 rectangular cuboids instead of 11 cubes.

The parameter α_v carries information about the growth rate as explained in what follows. Adopting a theoretical template based on general relativity (GR), one usually treats the parameter combination $f(z)\sigma_8(z)$ as an independent free parameter that may differ from the value for GR. This combination controls the amplitude of the linear velocity field, and can be wavenumber dependent in some modified gravity scenarios such as the $f(R)$ gravity (Hu & Sawicki 2007; Starobinsky 2007). In practice, a constant $f\sigma_8$ is more tractable and thus usually adopted in the literature. What we do with the parameter α_v is the *non-linear* analogy of this procedure. By selecting a value of α_v different from unity, we change the non-linear velocity field from what is expected for GR. On sufficiently large scales where non-linear contamination is negligible, this parameter can be identified with the following combination:

$$\alpha_v = \frac{(f\sigma_8)(z=0.3)}{(f\sigma_8)_{\text{fid}}(z=0.35)}. \quad (6)$$

Note that we take into account in the above a small difference of the redshift between the survey ($z=0.3$) and the simulation ($z=0.35$). Here and hereafter, variables with a subscript ‘fid’ refer to those assumed in the simulations or used in the redshift–distance relation when the power spectrum is measured.

The AP distortion induced to the simulations with our procedure is based on the mapping of wave vectors under the distant observer approximation (Ballinger et al. 1996; Matsubara & Suto 1996):

$$P_{\text{obs}}(\mathbf{k}) = [\alpha_{\parallel} \alpha_{\perp}^2] P_{\text{true}}(\mathbf{q}), \quad (7)$$

$$q_{\parallel} = \alpha_{\parallel} k_{\parallel}, \quad q_{\perp} = \alpha_{\perp} k_{\perp},$$

where \mathbf{q} is the wave vector in the unknown true coordinate system with q_{\parallel} and q_{\perp} , respectively, being its components parallel and

Table 2. 10 best-fitting models and their goodness of fit. Masses are given in unit of $10^{13} h^{-1} M_{\odot}$, while the number density n_g is in $10^{-4} h^3 \text{Mpc}^{-3}$. The first column shows the name of the model. The best-fitting model parameters are shown in the second to the sixth column. The mean number density n_g , the fraction of satellite LRGs f_{sat} and the fraction f_{sat}^* when a different definition of satellites are adopted (see text) are listed in the seventh to the eleventh column. We finally show the chi-squared statistics and the figure number in which the multipoles are shown.

Model	$M_{\text{host, min}}$	$M_{\text{sub, min}}$	$M_{\text{sat, min}}$	$M_{\text{cen, min}}$	p_{cen}	n_g	f_{sat}	$\pm(1\sigma)$	$\pm(2\sigma)$	f_{sat}^*	$\chi_{0, \text{min}}^2$	$\chi_{2, \text{min}}^2$	$\chi_{4, \text{min}}^2$	χ_{red}^2	$P_{\ell}(k)$
1a	2.50	–	–	–	–	1.28	0			0	109.2	916.8	27.1	11.8	Fig. 1
1b	2.70	–	–	–	–	1.16	0			0	119.0	656.5	34.6	9.10	
2a	1.29	0.211	–	–	–	4.07	0.300	+0.017 –0.012	+0.032 –0.024	0.098	48.0	28.4	27.8	1.18	Fig. 2
2b	1.45	0.299	–	–	–	3.30	0.248	+0.015 –0.010	+0.029 –0.024	0.093	31.2	28.7	30.7	1.03	
3a	1.29	–	0.216	1.27	–	3.76	0.316	+0.009 –0.019	+0.023 –0.033	0.099	39.9	28.6	26.3	1.09	Fig. 4
3b	1.47	–	0.316	1.46	–	2.97	0.259	+0.009 –0.019	+0.024 –0.033	0.094	26.7	29.2	29.1	0.976	
4a	0.981	0.886	–	–	0.149	0.740	0.260	+0.020 –0.006	+0.036 –0.019	0.073	26.8	26.8	21.6	0.865	Fig. 6
4b	1.16	1.13	–	–	0.167	0.622	0.213	+0.019 –0.008	+0.038 –0.020	0.060	21.7	27.7	22.9	0.830	
5a	1.81	0.427	–	–	–	2.85	0.214	+0.017 –0.030	+0.046 –0.050	0.089	21.6	31.1	26.4	0.930	Fig. 8
5b	1.87	0.465	–	–	–	2.67	0.201	+0.018 –0.027	+0.046 –0.049	0.088	20.5	29.5	30.8	0.950	

perpendicular to the line of sight, while the observed coordinate is \mathbf{k} with k_{\parallel} , k_{\perp} defined analogously. The pre-factor in the square bracket represents the change in the volume element by the coordinate transformation. This mapping formula derived for analytical models of the power spectrum is automatically realized by measuring the power spectrum from the deformed simulation boxes.

With a help of the BAO feature clearly visible in the observed spectrum, we can conduct a geometrical test. The parameter α_{\parallel} is related to the comoving distance along the line of sight, and thus reflects the ratio of the Hubble parameter between the true unknown cosmology and the one assumed when we convert the redshifts into comoving distances. In addition, the parameter α_{\parallel} depends on the acoustic horizon scale at the baryon drag epoch (Eisenstein & Hu 1998), since the true acoustic scale might be different from the one realized in the simulations. We denote the acoustic scale at the drag epoch by r_s , and compute it with the CAMB code. Indeed, this quantity computed for our fiducial cosmological model is larger than that for the *Planck* cosmology by about 5 per cent when measured in units of $h^{-1} \text{Mpc}$ (this difference mainly comes from the difference in h ; see Table 1). Taking this difference into account, we have

$$\alpha_{\parallel} = \frac{H(z)r_s}{H_{\text{fid}}(z)r_{s, \text{fid}}} \quad (8)$$

at the effective redshift of the measured power spectrum, $z = 0.3$. Similarly, the angular diameter distance, D_A , can be constrained through the parameter α_{\perp} :

$$\alpha_{\perp} = \frac{D_{A, \text{fid}}(z)r_s}{D_A(z)r_{s, \text{fid}}} \quad (9)$$

at $z = 0.3$.

We simultaneously vary α_v , α_{\parallel} and α_{\perp} as well as $M_{\text{host, min}}$ and $M_{\text{sub, min}}$ to find the best-fitting parameter set for Model 5. We discuss the robustness of the constraints on the parameters for the mock LRGs when the cosmological assumptions are relaxed by comparing the results of Models 2 and 5. Also, we show the derived cosmological constraints and compare them with those in the literature to demonstrate the prospect of analysing observational data with theoretical predictions from simulations instead of analytical models.

4 RESULTS

Now we are in position to show the results of the MCMC analysis explained so far. We first discuss the importance of the satellite population to model the anisotropic clustering of LRGs by showing the results of the fit with Models 1 and 2 in Section 4.1. We then compare the satellite fraction derived with Model 2 and that with Models 3 and 4 where centrals and satellites host mock LRGs with different criteria in Section 4.2. We further discuss the robustness of the results against cosmological uncertainties in Section 4.3. Some cosmological implications are given in Section 4.4. We finally compare the multiplicity function of our best-fitting models with observation in Section 4.5. The best-fitting model parameters as well as the goodness of fit are summarized in Table 2.

4.1 Importance of satellites

We start this section by showing the result of Model 1. The best-fitting multipoles of the power spectrum are shown in Fig. 1 together with the observation. The results of Models 1a and 1b are, respectively, plotted in solid and dashed lines. It is clear from the figure that we cannot simultaneously fit the three multipoles with the central population alone. There exists a mismatch between the observed and the model quadrupole at $k \gtrsim 0.1 h \text{Mpc}^{-1}$. The difference between the two models (Models 1a and 1b) is visible, but is much smaller than the discrepancy between the observation and the two models. Also, we can observe a difference in the broad-band shape of the monopole moment. Since we have only one parameter in Model 1 (i.e. the minimum mass of the host haloes), it might be difficult to simultaneously adjust the amplitude and the shape of the spectrum.

To be more quantitative, we show the minimum value of χ^2 in Table 2. The values for the three multipoles, $\chi_{\ell, \text{min}}^2$ ($\ell = 0, 2$ and 4) and their sum divided by the degree of freedom (the reduced chi-squared, χ_{red}^2) are listed in the table. Note that each multipole moment has 30 data points up to $k_{\text{max}} = 0.305 h \text{Mpc}^{-1}$. Clearly, both Models 1a and 1b give a poor fit to the observed data with Model 1b being slightly better ($\chi_{\text{red}}^2 = 11.8$ and 9.10 , respectively), and most of the discrepancy comes from the quadrupole moment.

We next discuss the result of Model 2 in which we assign mock LRGs to both centrals and satellites. We can see in Fig. 2 a substantial improvement of fit over Model 1 plotted in Fig. 1. The

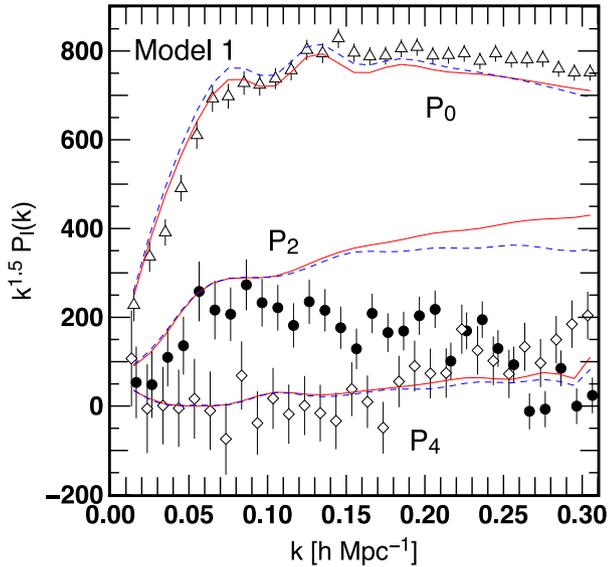


Figure 1. Best-fitting power spectra against observation assuming one-to-one correspondence between LRGs and haloes (Model 1). Symbols show the observed multipoles of the power spectrum (open triangles: monopole; filled circles: quadrupole; open diamonds: hexadecapole). The solid (dashed) curves depict the result of Model 1a (b). Note that symbols for the quadrupole (hexadecapole) moment are horizontally offset by $+0.0015$ (-0.0015) to avoid overlap.

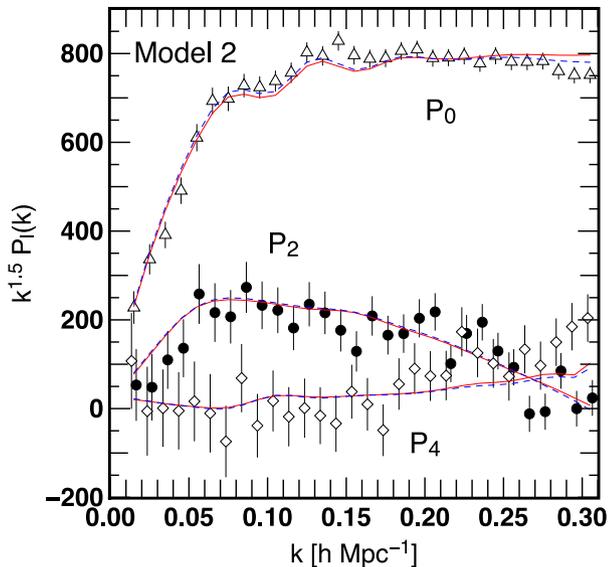


Figure 2. Same as in Fig. 1, but for Model 2.

values of χ^2_{red} are 1.18 and 1.03 for Models 2a and 2b, respectively, implying that the modelling is sufficient to explain the observed multipole moments up to $k \simeq 0.3 h \text{ Mpc}^{-1}$. The constraints on the model parameters are shown in the left-hand panel of Fig. 3 with the best-fitting values indicated by the plus (cross) symbols for Model 2a (2b). The two parameters, $M_{\text{host, min}}$ and $M_{\text{sub, min}}$, show a positive correlation in both models, and the allowed regions for the two models are significantly different.

We also plot in the right-hand panel the constraint on the fraction of LRGs hosted by satellite subhaloes, f_{sat} . A satellite fraction of $f_{\text{sat}} \sim 0.3$ is favoured by the observed power spectrum for Model 2a, while a smaller value is derived for Model 2b ($f_{\text{sat}} \sim 0.25$;

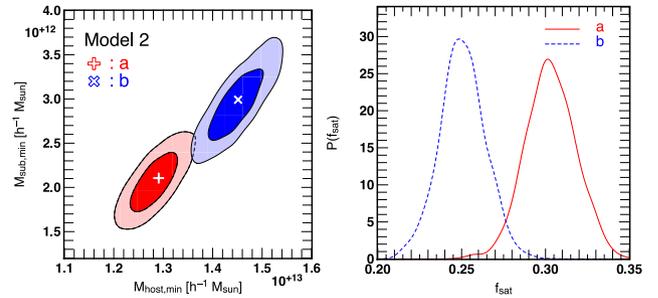


Figure 3. Constraints on the parameters of Model 2. We plot 1σ and 2σ allowed regions for parameters characterizing the mock LRGs (left), and the posterior distribution of the fraction of satellites (right). In the left-hand panel, Model 2a is shown by red contours while Model 2b is plotted with blue. We depict by solid (dashed) line the result of Model 2a (2b) in the right-hand panel.

dashed). The difference between Models 2a and 2b mainly comes from the different velocity dispersion of mock LRGs. We have seen that Model 1 cannot explain the suppress of the quadrupole at $k \gtrsim 0.1 h \text{ Mpc}^{-1}$. By adding a significant fraction of satellite LRGs, Model 2 can explain the quadrupole thanks to a larger velocity dispersion of the mock LRGs in satellite subhaloes than those in centrals. Furthermore, since the mock LRGs in Model 2b have a larger velocity dispersion than in Model 2a, the former needs a smaller fraction of satellite LRGs to match the observed power spectrum. The result of this section highlights the importance of satellite LRGs to understand the observed anisotropic clustering. Since the χ^2 values are very similar in these two models, we conclude that it is difficult to constrain the additional velocity contribution in Model 2b from the large-scale analysis alone. This component is degenerate with the fraction of satellites, and thus we can put a constraint only on the total velocity dispersion of centrals and satellites, which determines the magnitude of the Fingers of God effect.

Our constraint on the fraction of satellite LRGs might be counterintuitive given the small fraction of multiple LRG systems suggested from observations (~ 5 per cent). Indeed, the lower bound on this parameter is given as $f_{\text{sat}} > 0.27$ (0.22) at 95 per cent confidence level (c.l.) in Model 2a (2b) from our analysis, while the value is constrained to $6.36^{+0.38}_{-0.39}$ per cent (95 per cent c.l.) by Reid & Spergel (2009) based on the observed multiplicity function of LRGs. As discussed in more detail below, this apparent discrepancy in the satellite fraction partly comes from different definitions of satellites. Also, our result suggests that the central subhaloes in our terminology do not always host a LRG. We reserve a thorough discussion on the compatibility of our mock LRGs until Section 4.5, and keep testing the robustness of our constraints against different assumptions in the model of LRGs in what follows.

4.2 Extension of the simplest model

Now the important question to ask here is whether the seemingly large satellite fraction obtained in the previous section using Model 2 is real or just a consequence of incorrect modelling of LRGs. We here adopt Models 3 and 4 and allow additional parameters to float. Since these parameters control the selection criteria of satellites and centrals individually, we can check the robustness of the satellite fraction derived in the previous section.

The result of Model 3 is plotted in Figs 4 and 5 for the best-fitting power spectrum and the constraints on the model parameters,

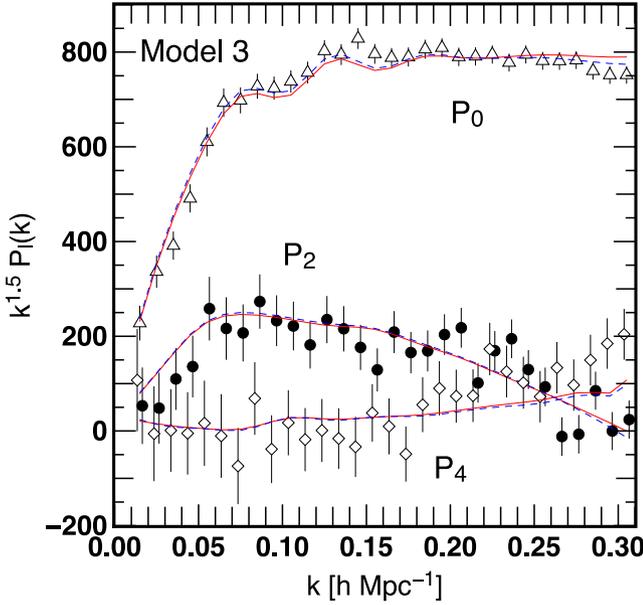


Figure 4. Same as in Fig. 1, but for Model 3.

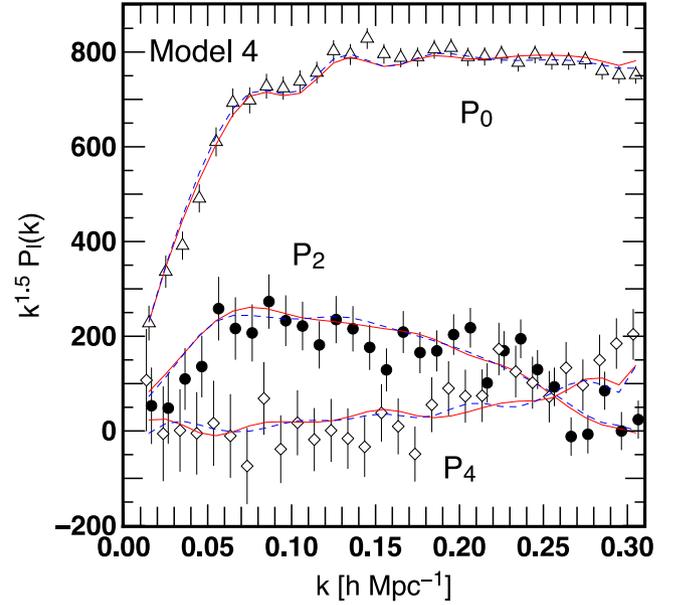


Figure 6. Same as in Fig. 1, but for Model 4.

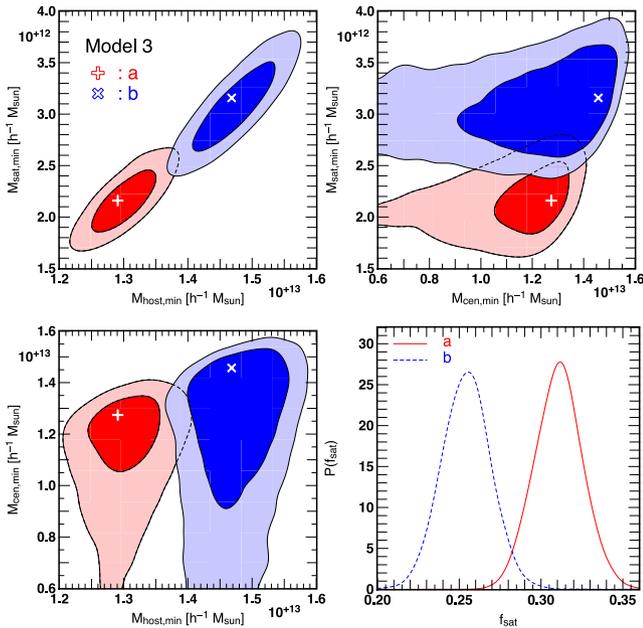


Figure 5. Same as in Fig. 3, but for Model 3.

respectively. It is difficult to judge by eye, but the fit is slightly improved over Model 2 as the values of χ_{red}^2 suggest (see Table 2). The constraints in $M_{\text{host, min}}-M_{\text{sat, min}}$ plane shown in the top left-hand panel of Fig. 5 are very similar to those in the left-hand panel of Fig. 3 for $M_{\text{host, min}}$ and $M_{\text{sub, min}}$. This suggests that the satellite LRGs in Model 3 are almost the same as in Model 2. By contrast, we cannot put a stringent constraint on the other parameter, $M_{\text{cen, min}}$ (see top right- and bottom left-hand panel of Fig. 5). This is because given $M_{\text{host, min}}$, there are only small number of centrals having mass much smaller than $M_{\text{host, min}}$. The best-fitting parameter set, indicated by the plus and cross symbols, suggests a large $M_{\text{cen, min}}$ compared with $M_{\text{sat, min}}$. Thus the slightly improved fit compared to Model 2 comes from a reduced number of the central LRGs. As a result, the satellite fraction plotted in the bottom right-hand

panel of Fig. 5 shows a slightly larger value than in Model 2 (right-hand panel of Fig. 3), but the difference is much smaller than the statistical uncertainty on that parameter.

Next, the best-fitting power spectrum of Model 4 is shown in Fig. 6. Judging from the χ_{red}^2 listed in Table 2, this model gives a better fit than Models 2 and 3 ($\chi_{\text{red}}^2 = 0.865$ and 0.830 for Models 4a and 4b, respectively). Although the values themselves should be used with caution because we ignored the off-diagonal components of the covariance matrix, these values suggest a slight overfit to the observational data. This might be explained as follows. In Model 4, central subhaloes are randomly selected with a probability p_{cen} to host LRGs. Because of this random process, the resultant multipole moments can be different from one time to another even one employs exactly the same model parameters. In other words, the increased shot noise in the model spectrum, which is not taken into account in the MCMC analysis, can sometimes mimic the observed noise pattern by chance and reduce the χ^2 statistics. Indeed, the best-fitting multipoles are less smooth than in previous figures because of this effect.

The constraints on the model parameters are shown in Fig. 7. Interestingly, the constraints in the $M_{\text{host, min}}-M_{\text{sub, min}}$ plane shown in the top left-hand panel are significantly different from that in Model 2 (Fig. 3): the minimum host halo mass is smaller, and the minimum subhalo mass is larger in Model 4. Because of the smaller host halo mass, the number of centrals that pass the mass criterion increases, but it is then reduced by random sampling. The number of satellites is smaller because of a larger $M_{\text{sub, min}}$, and as a result, the mean number density of the final mock LRGs are much smaller than in Model 2 (see Table 2). These differences might suggest two possible different nature of LRGs inferred with Model 2 and with Model 4. Nevertheless, looking at the bottom right-hand panel of Fig. 7, the fraction of satellites is still high and is in the range of 20–30 per cent. This is broadly consistent with the result of Model 2 shown in the right-hand panel of Fig. 3. Thus the anisotropic clustering on large scales serves as a good probe of the fraction of satellites robustly against different assumptions in the model.

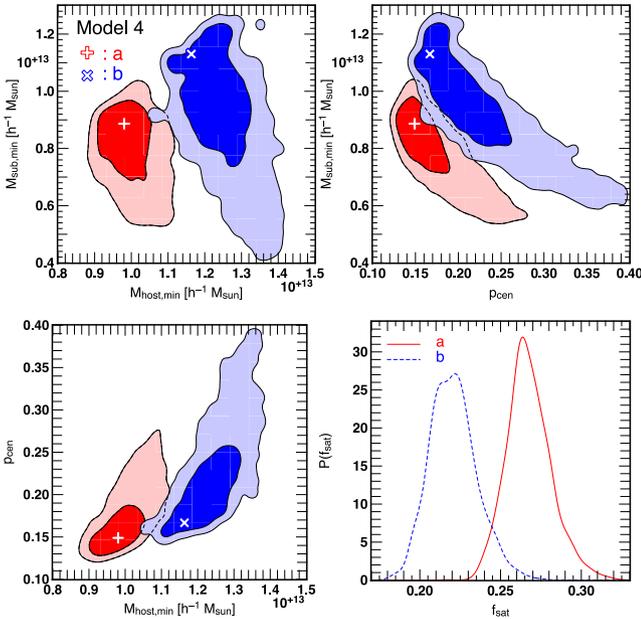


Figure 7. Same as in Fig. 3, but for Model 4.

4.3 Robustness of the parameters against cosmological uncertainties

So far we have seen that the fraction of satellites is an important ingredient to understand the anisotropic clustering of LRGs in redshift space. The discussion so far is, however, done keeping the underlying cosmological model fixed to the one adopted in the simulations. We relax the cosmological assumptions in this section, and discuss the possibility of a simultaneous determination with cosmological parameters. Here we focus on the constraints on the parameters that describe the connection between LRGs and subhaloes, and leave the cosmological implications to the next section.

We first show in Fig. 8 the best-fitting power spectrum of Models 5a and 5b. Compared to the results of Model 2, the goodness of fit is improved ($\chi^2_{\text{red}} = 1.18\text{--}0.93$ for Model a and $1.03\text{--}0.95$ for Model

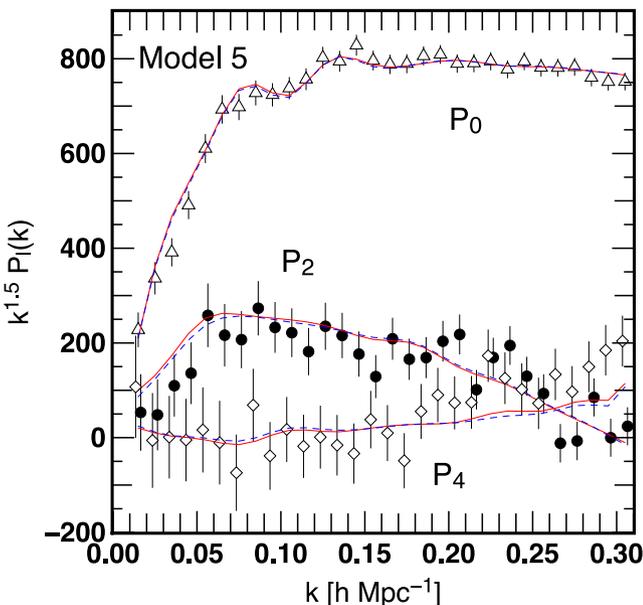


Figure 8. Same as in Fig. 1, but for Model 5.

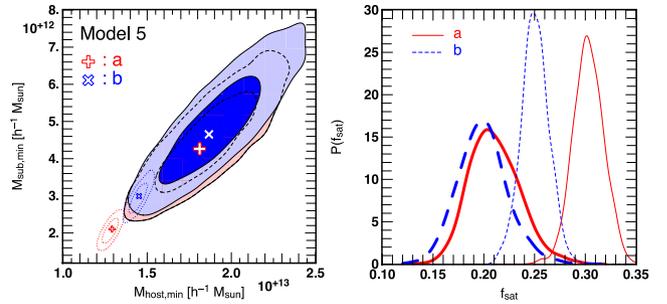


Figure 9. Constraints on the parameters describing the properties of LRGs when cosmological parameters are varied (Model 5). We plot the 1σ and 2σ allowed regions for $M_{\text{host},\text{min}}$ and $M_{\text{sub},\text{min}}$ (coloured contours in the left-hand panel) and the fraction of satellites (thick lines in the right-hand panel). We also plot the results when cosmological parameters are fixed (Model 2) by dotted contour lines in the left-hand panel and thin lines in the right-hand panel.

b). Looking at each of the three multipoles, we can see that most of the improvement comes from the monopole moment (see Table 2). We discuss in more detail on this improvement in the subsequent subsection.

The constraints on the parameters, $M_{\text{host},\text{min}}$ and $M_{\text{sub},\text{min}}$ are shown in the left-hand panel of Fig. 9. We also show the result of Model 2 in dotted contour lines. Compared to Model 2, the results are greatly altered in two ways: first, the best-fitting minimum masses are shifted towards larger values and secondly the statistical uncertainties on them are increased by a factor of about 5.

We plot in the right-hand panel of Fig. 9 the satellite fraction derived with Model 5 (thick lines) and with Model 2 (thin lines). Compared with the result in the left-hand panel, the statistical uncertainty on the fraction of satellites does not increase so dramatically when cosmological assumptions are relaxed, although a smaller fraction is favoured for Model 5 (~ 20 per cent). This basically confirms that the fraction of satellites still play a dominant role to explain the anisotropy of the apparent clustering pattern of LRGs even when we introduce additional sources of anisotropy.

Another interesting difference compared to Model 2 is that the difference between Models 5a and 5b is much smaller than that between Models 2a and 2b. As we will discuss shortly, the distortion parameters introduced in Model 5 explain the observed anisotropy of the clustering of LRGs partly, and the difference between Models 5a and 5b is somewhat absorbed by these new parameters. Our final estimate of the 1σ (2σ) allowed region of the satellite fraction marginalized over cosmological uncertainties is $f_{\text{sat}} = 0.214^{+0.018}_{-0.030}$ ($0.201^{+0.018}_{-0.028}$) for Model 5a and $f_{\text{sat}} = 0.201^{+0.018}_{-0.028}$ ($0.201^{+0.047}_{-0.050}$) for Model 5b.

4.4 Prospects to derive cosmological parameters using N -body simulations

Now we turn to the cosmological constraints marginalized over the parameters that describe the nature of LRGs. We first show the constraints on our parameters that induce distortions to the clustering of mock LRGs in Fig. 10. We plot the 67 and 95 per cent confidence regions of α_v and two parameter combinations, $\alpha \equiv (\alpha_{\parallel}\alpha_{\perp}^2)^{1/3}$ and $\epsilon \equiv \alpha_{\parallel}/\alpha_{\perp}$. The parameter α is responsible for the mismatch between the true and assumed distance scale, and is expected to be determined by the BAO feature in the monopole moment. On the other hand, the other parameter, ϵ , determines the

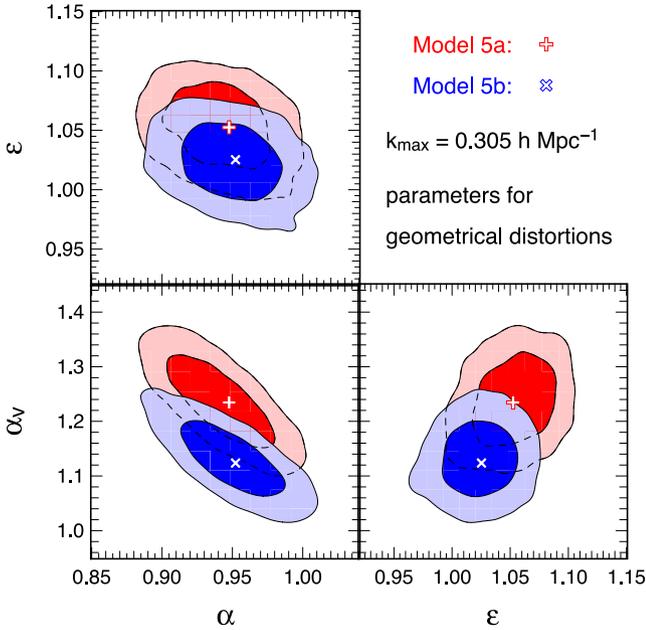


Figure 10. Constraints on the distortion parameters marginalized over $M_{\text{host, min}}$ and $M_{\text{sub, min}}$. We plot the 1σ and 2σ allowed regions of the parameters, $\alpha = (\alpha_{\parallel}\alpha_{\perp}^2)^{1/3}$, $\epsilon = \alpha_{\parallel}/\alpha_{\perp}$ and α_v . We include wavenumbers up to $k_{\text{max}} = 0.305 h \text{ Mpc}^{-1}$ into the analysis.

significance of the AP effect that induces a geometrical distortion, and thus is sensitive to the higher multipoles.

Fig. 10 suggests that the different modelling of LRGs using subhaloes (i.e. Models 5a and 5b) leads to different values of the distortion parameters favoured by the observation. We can find that the parameter α is the least affected among the three by the difference in the prescription for mock LRGs. This is presumably owing to the robustness of the distance determination using the feature of BAOs against different bias models. In both models, $\alpha < 1$ is strongly suggested, and the best-fitting parameter is around 0.95. As we discussed in Section 3.4, the sound horizon scale, r_s , in our fiducial cosmology is larger than that recently suggested. Our constraint on α is consistent with this expectation, and supports recent observational results. A value of α smaller than unity helps to reduce $\chi_{0, \text{min}}^2$ from that in Model 2, and we can actually observe an improved fit to the BAO wiggles in the monopole moment in Fig. 8. Note also that $\chi_{0, \text{min}}^2$ is the smallest in Model 5b (and almost the same in Model 5a) compared to more complicated descriptions for the LRG–subhalo correspondence in Model 3 or 4.

The situation seems a bit different for the other two parameters. They are more sensitive to the detail of the model and the confidence regions for the two models, Models 5a and 5b, have some non-negligible offset in the bottom right-hand panel of Fig. 10. In both cases, $\alpha_v > 1$ and $\epsilon > 1$ are suggested. In particular, Model 5a needs a large velocity boost factor, $\alpha_v \sim 1.2$. Since the mock LRGs in this model have a smaller velocity dispersion than in Model 5b, a larger value of α_v is required to explain the observed power spectrum. This situation is quite similar to the results of Models 2a and 2b: we have shown that we need a larger satellite fraction for Model 2a than Model 2b (see Fig. 3). Instead of the fraction of satellite, the parameter α_v mainly absorbs the difference between the two models in this case.

Since both α_v and ϵ are related to the amplitude of the apparent anisotropy in redshift space, it is natural to see that they are degenerate with each other. While a value of α_v larger than unity *squashes*

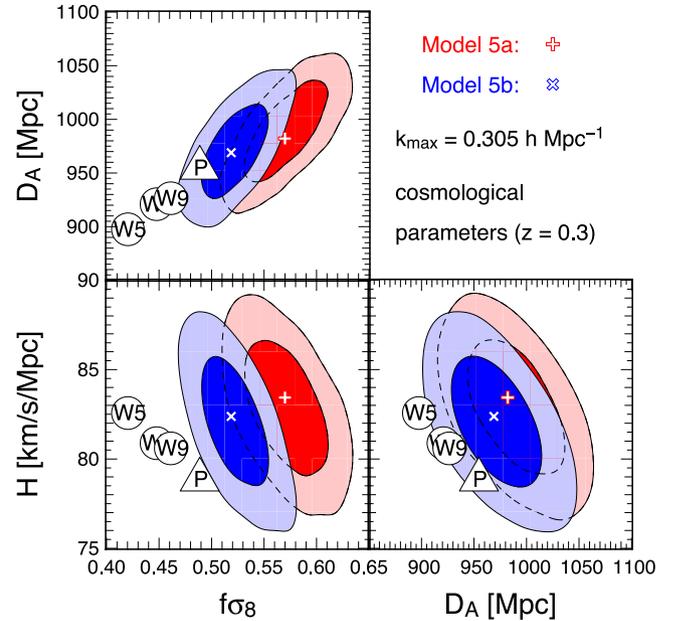


Figure 11. Cosmological constraints. We assume the value of r_s from the *Planck* observation when we convert the constraints on α_{\parallel} and α_{\perp} to H and D_A . We also show by symbols the results of CMB experiments (W5, W7, W9 and P respectively refer to *WMAP5*, *WMAP7*, *WMAP9* and *Planck*; best-fitting flat Λ CDM to the CMB data alone).

the apparent clustering pattern on large scales and *elongates* it on small scales along the line-of-sight direction (the Kaiser and the Fingers of God effect, respectively), a large ϵ always *elongates* the apparent clustering independent of the distance scale. As we will show later, the large velocity boost factor, α_v , mainly comes from the power spectrum on small scales (i.e. a prominent Fingers of God suppression in the observed spectrum). Then, a large α_v determined on small scales might result in an overprediction of the Kaiser distortion on large scales. A large ϵ is chosen such that it partly cancels the strong Kaiser effect. This way, the difference in the modelling of velocities in the two models propagates to the derived distortion parameter, ϵ .

It is worth noting that both models statistically exclude the fiducial parameter set, $\alpha_v = \alpha = \epsilon = 1$. This suggests that the cosmological model assumed in the simulations and in the redshift–distance relation might be different from the underlying true cosmology: the BAO scale in the simulation is too long, mock LRGs need larger velocities and a geometrical distortion is induced in the observed clustering to elongate the structure along the line of sight. Though it might be too early to falsify the cosmological model from this analysis alone, our results are encouraging. We can tell the difference between our model based on simulations and the clustering of LRGs in the real Universe with the current accuracy of the observational data. As we will outline in Section 5, it might be very interesting to repeat the analysis with some additional cosmological simulations and a new measurement of the power spectrum assuming a cosmological model suggested by the current analysis.

We then derive the constraints on the parameters, $f\sigma_8$, H and D_A using the relations, equations (6), (8) and (9). In doing so, we assume the value of r_s derived from the *Planck* observation (see Table 1). The results are shown in Fig. 11. For reference, we also mark the best-fitting Λ CDM models from some cosmic microwave background (CMB) observations by symbols (W5, W7 and W9 for *WMAP*, and P for *Planck* in the figure, and see Table 1

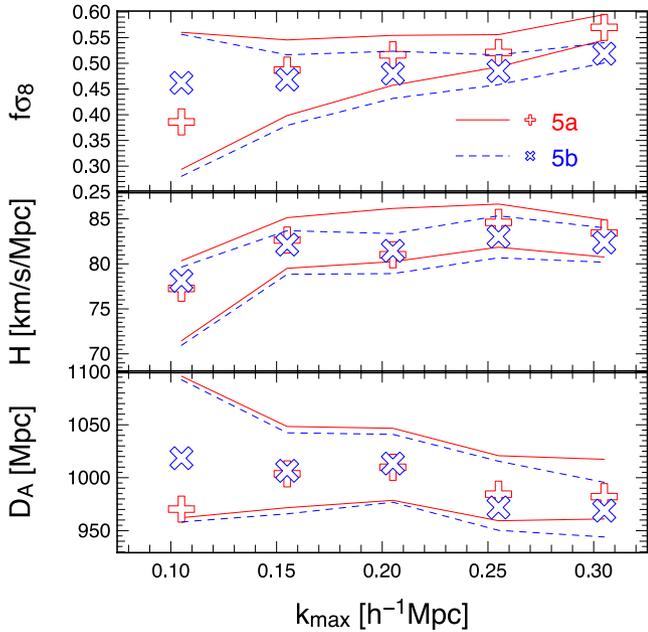


Figure 12. Robustness of the results against the maximum wavenumber. The 67 per cent confidence intervals of the derived cosmological parameters are plotted as a function of the maximum wavenumber (bands in solid lines for Model 5a, dashed lines for Model 5b). The best-fitting parameters are shown as plus symbols for Model 5a and as cross symbols for Model 5b.

for detail). Our constraints in the H - D_A plane (bottom right) are broadly consistent with the CMB observations (interestingly, more consistent to more recent results). On the other hand, the derived growth rate parameter, $f\sigma_8$, is larger than the predictions of the cosmological models favoured by CMB observations. We should interpret this with extreme caution. The discrepancy between the Models 5a and 5b in the parameter combination, $f\sigma_8$, is as large as 10 per cent, and thus our estimate might have a systematic error of this size. As we will see shortly, the constraint on $f\sigma_8$ from this analysis is likely to be contaminated significantly by non-linearity of the velocity field, and the derived combination might not be a linear growth rate but its non-linear counterpart.

For a deeper understanding of the situation, we change the wavenumber range to be taken into account in the analysis, and discuss the sensitivity of the results against the maximum wavenumber, k_{\max} . The resulting one-dimensional marginalized constraints on the three cosmological parameters are shown in Fig. 12. We show the 1σ confidence intervals in the bands (solid for Model 5a, dashed for Model 5b). When we adopt a value of k_{\max} as small as $\sim 0.1 h \text{ Mpc}^{-1}$, the intervals for the two models almost coincide with each other. This suggests that the difference in the velocity structure on small scales is not relevant for the anisotropy on very large scales. The discrepancy between the two models grows with k_{\max} , and finally the two 1σ intervals become exclusive for $f\sigma_8$ when $k_{\max} = 0.305 h \text{ Mpc}^{-1}$, which is our fiducial value.

The parameter $f\sigma_8$ appears to be monotonically increasing with k_{\max} , and Model 5a shows a more sensitive response to k_{\max} than Model 5b. As we have already discussed, our estimate of $f\sigma_8$ through α_v is affected by non-linearity of the velocity field, and indeed the result indicates a scale-dependent velocity bias between our mock LRGs and the observation. Larger velocities are required on smaller scales. Our estimator of $f\sigma_8$ is expected to have no dependence on k_{\max} when the simulations are perfect (i.e. $\alpha_v = 1$ independent of k_{\max}). Since the amplitude of the velocity perturbations is ex-

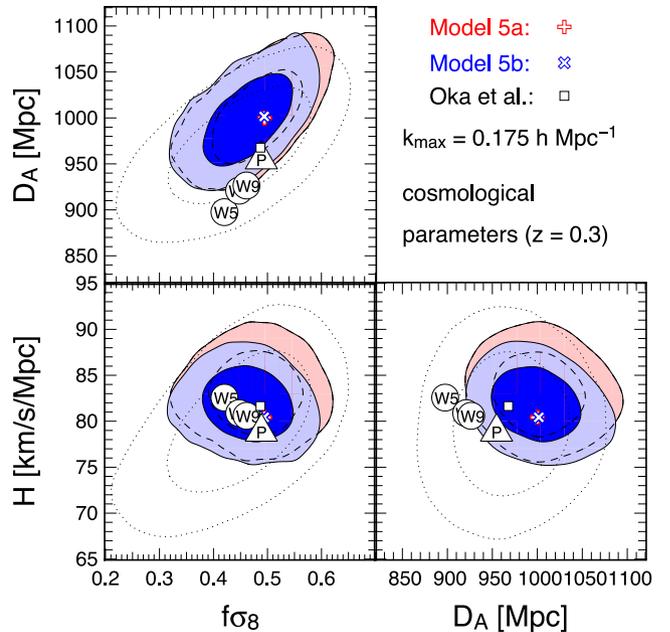


Figure 13. Comparison between our cosmological constraints (coloured contours) and those obtained by fitting the same data but with an analytical model described in the accompanying paper (Oka et al. 2014; dotted contours). We change the maximum wavenumber to $k_{\max} = 0.175 h \text{ Mpc}^{-1}$ so as to match with that in Oka et al. (2014).

pect to be linear at the large-scale limit and our two models give almost the same answer at $k_{\max} = 0.105 h \text{ Mpc}^{-1}$, the result at $k_{\max} = 0.105 h \text{ Mpc}^{-1}$ can be considered as a weak but reliable constraint on the linear growth rate. We will comment on a possible solution for the non-linear contaminations to the estimation of $f\sigma_8$ later in Section 5, and leave further investigations of the reliability of the constraint to future studies.

The other two parameters, H and D_A , show a weaker dependence on k_{\max} . Although the situation seems to be better for these parameters than $f\sigma_8$, one should still take the results with caution. The incompleteness of our modelling of the velocities can propagate to these parameters. The parameters, H and D_A , are determined through α and ϵ . We show that α is robust against systematics in Fig. 10. Nevertheless, a misestimation of $f\sigma_8$ propagates to ϵ , that results in a systematic error on H and D_A . It is of interest to further test the constraints on these parameters with new simulations that can reproduce the observed RSDs without a large velocity boost α_v , and a study along this line is now ongoing.

We have discussed so far the constraints on the cosmological parameters using N -body simulations and their possible systematics. Now we compare our results with those obtained with an independent method. We plot in Fig. 13 the constraints on the three cosmological parameters from our analysis (coloured contours) and those by the accompanying paper (Oka et al. 2014) in which we adopt an analytical model for the parameter estimation (contours in dashed lines with the best-fitting values depicted by cubes). In this figure, we adopt $k_{\max} = 0.175 h \text{ Mpc}^{-1}$, which is the same as in Oka et al. (2014).

Overall, the contours obtained by Oka et al. (2014) enclose the allowed regions obtained by our analysis. Especially, the best-fitting parameters obtained with different models are consistent with each other at about 1σ level. This is encouraging since the two analyses rely on totally different prescriptions for the non-linear growth of structure, the RSDs and the galaxy bias. The weaker constraints by

Oka et al. (2014) probably come from the fact that the analytical model has free parameters to control the galaxy bias as well as the velocity dispersion in the analytical model. These nuisance parameters are effectively fixed in our analysis once the connection between LRGs and subhaloes is assumed.

The result suggests that it is potentially possible to tighten the constraints by adopting a reasonable prescription for galaxies using subhaloes identified in simulations. Alternatively, giving a prior information to the velocity dispersion based on simulations rather than letting it float as a free parameter might be another way to obtain tighter constraints. The best-fitting parameters derived from CMB observations are in good agreement with the contours obtained by Oka et al. (2014). If we believe the contours obtained in this study, we can exclude some of the CMB results. This suggests that the anisotropic clustering of galaxies have a strong statistical power if one successfully models the non-linear growth of the velocity field (velocity dispersion, in this particular case) and the galaxy bias.

4.5 Multiplicity function

So far we have focused on the clustering of LRGs in redshift space on relatively large scales and discussed how well we can explain the observational data with mock LRGs in cosmological simulations. It is of course important to construct a mock galaxy catalogue that has a mean number density consistent with observations. Indeed mock galaxy catalogue is usually constructed based on HOD or abundance matching techniques, in which *one-point* statistics are directly drawn from observation or constrained from clustering information on *small scales*. In these methods, mock galaxies are modelled so that the observed mean number density is recovered automatically. It is interesting to check the compatibility of our mock LRGs with observed one-point statistics. In this section, we compare our best-fitting models with the multiplicity function estimated by Reid & Spergel (2009).

Before making a comparison, we note some differences in the LRG samples used in our study and in Reid & Spergel (2009). As summarized in Section 2, our sample is based on the DR7 and we use LRGs only on the north cap in the redshift range of $0.16 < z < 0.47$. By contrast, the sample analysed in Reid & Spergel (2009) is based on DR4+ (Adelman-McCarthy et al. 2007) and they include both caps but restrict the redshift range to $0.16 < z < 0.36$. They also supplementary add the imaging sample to correct for fibre collisions, incomplete sky and complex angular masks. It is thus not straightforward to directly compare our results with theirs, but nevertheless the comparison is still meaningful to test our alternative mock making scheme and infer the nature of LRGs.

First of all, the mean number density of mock LRGs in some of our best-fitting models are significantly larger than the observation. The sample in Reid & Spergel (2009) has $n_g \simeq 9.7 \times 10^{-5} h^3 \text{Mpc}^{-3}$ (Zehavi et al. 2005), and this is two to four times smaller than the results of Models 2, 3 and 5 (see Table 2). On the other hand, Model 4 has a slightly smaller number density than that from observation.

This discrepancy in the mean number density implies that the mass of the haloes as well as that of subhaloes do not uniquely determine whether a LRG can live there or not: our model simply assumes that *all* the subhaloes above a mass threshold have LRGs, and this assumption is so simple that lead the number density larger than the observed. A simple way towards a more sophisticated modelling of LRGs is to introduce a second physical parameter that determines the habitability of a LRG. An example is the presence of a massive halo at an earlier epoch (i.e. $z = 2$ or 3) adopted in the abundance matching technique (Masaki et al. 2013). This approach

was shown to recover the HOD and the angular clustering very well in addition to the number density, which is automatically adjusted by the matching by construction.

Instead of pursuing along this direction, we adopt a simple statistical approach to adjust the mean density in this study. We conduct random sampling and reduce the number of mock LRGs assuming that the observed LRG catalogue is a Poisson sample of our mock LRGs. As we have already noted, this procedure does not affect the expectation value of the power spectrum only resulting in a larger shot noise and thus the successful fitting results in the previous sections are not diminished. Also, the fraction of satellites derived by our analysis are not affected by random sampling.

We might be able to justify the adjustment of the number density with random sampling as follows. Galaxy bias is a *stochastic* process by nature and is difficult to connect simulated subhaloes to observed LRGs *deterministically* (Dekel & Lahav 1999). In other words, there might exist some unknown parameters to determine whether a subhalo hosts a LRG or not (Taruya & Suto 2000). Modelling LRGs from the first principle is well beyond the scope of this study based on *N*-body simulations that solves a purely gravitational system, and baryonic physics must be essential in order to fully predict (i.e. deterministically simulate) the clustering of LRGs or galaxies in general (e.g. Somerville et al. 2001; Yoshikawa et al. 2001). We give up a deterministic prediction of LRGs but instead effectively take into account the stochasticity of galaxy bias by random sampling. In what follows, we adjust the spatial density to $n_g = 10^{-4} h^3 \text{Mpc}^{-3}$ for Models 2, 3 and 5, while we keep all the mock LRGs in Model 4.

We now compare the multiplicity function with that measured by Reid & Spergel (2009). We denote the number of LRGs in a host halo by N_{LRG} . In our case, the host haloes are simply the FoF groups while Reid & Spergel (2009) identify groups with the count-in-cylinder method. Then the multiplicity function, $P(N_{\text{LRG}})$, is defined as the fraction of LRG systems with N_{LRG} members. Note that we do not count FoF groups without a mock LRG in the denominator to match the definition of the observed multiplicity function. We plot in Fig. 14 $P(N_{\text{LRG}})$ against N_{LRG} . The overall decreasing trend of $P(N_{\text{LRG}})$ with N_{LRG} seen in the observed data (histograms) is well

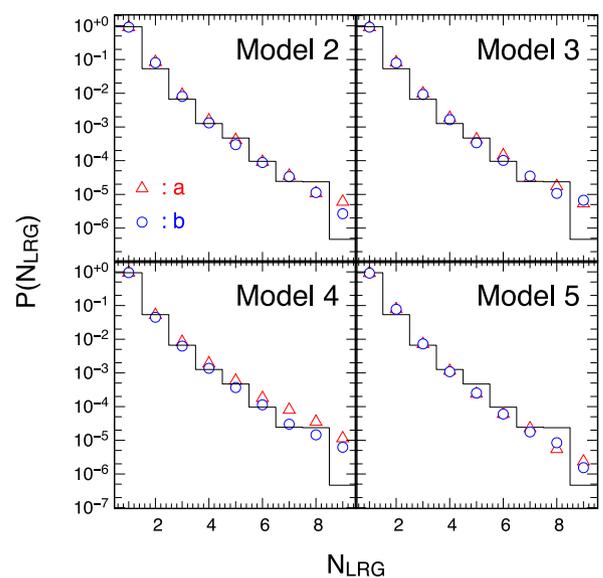


Figure 14. Multiplicity function of LRGs from our best-fitting models (symbols) against observation (histogram). The triangles (circles) show the result of Model a (b).

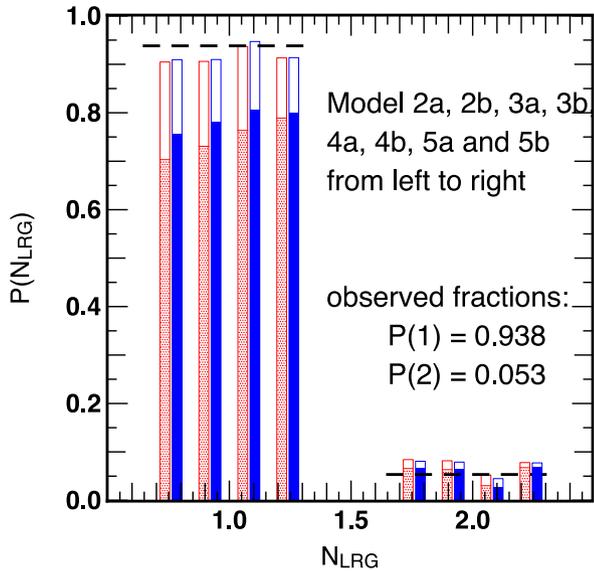


Figure 15. Zoom of Fig. 14 for $N_{\text{LRG}} = 1$ and 2. The filled part of the bars show the contribution for systems with a central LRG and the open part is responsible for the systems without a central LRG.

explained by all the models (symbols; triangles for Model a and circles for Model b). This coincidence is notable since we construct the catalogues guided only by the clustering pattern in redshift space on large scales. Also, remember that the allowed regions of the parameters characterizing the condition of subhaloes to host LRGs are significantly different in different models. These results suggest that the multiplicity function is an important ingredient that determines the multipole moments of the power spectrum in redshift space on large scales.

We plot in Fig. 15 the first two bins, $N_{\text{LRG}} = 1$ and 2 in linear scale. This time the models are shown in bars (Model 2a, 2b, . . . , 5b from left to right), while the observational data are shown by horizontal dashed lines. The filled part of the bars shows mock LRG systems that contain a central LRG, while the open part indicates those consisting of only satellites. The observational data as well as our models suggest that most of the LRGs are distributed in a single- or double-LRG system, and these two bins mostly determine the clustering properties. The fraction of single-LRG systems ranges from 90 to 95 per cent with our eight models, and this number is in good agreement with the observed fraction, 93.8 per cent. Models 4a and 4b give the most consistent fractions to the observation, and might be a better description of LRGs than other models. Indeed, these models have the smallest χ_{red}^2 to the multipole moments (see Table 2).

An important difference in the interpretation of the observed multiplicity function in our study and in Reid & Spergel (2009) is the definition of centrals and satellites. Our definition of central LRGs are those hosted by subhaloes that have the largest mass in the same host haloes, and the rest of LRGs are satellites. Consequently, the most massive subhaloes sometimes do not host a LRG in our model and a LRG system can be exclusively composed of satellite LRGs. On the other hand, in Reid & Spergel (2009), only LRGs in multiple LRG systems are considered as satellites and all the single-LRG systems are regarded as central LRGs. The satellite fraction as large as 20–30 per cent derived in our analysis mainly comes from this difference. If we instead employ the same definition of satellites as in Reid & Spergel (2009), the fraction becomes much smaller. We list in Table 2 the fraction of satellites, f_{sat}^* , derived with this

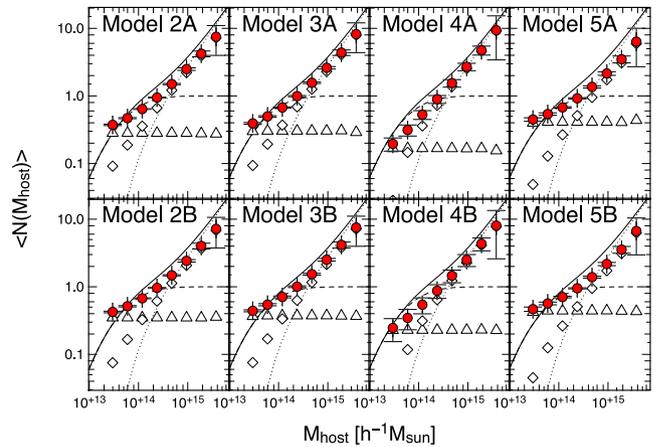


Figure 16. Mean halo occupation number as a function of the halo mass derived from our eight models. Triangles and diamonds show the central and satellite contributions, respectively, while the sum is plotted as circles with error bars. Also shown by lines is the model described by equation (10) with the parameters determined by Reid & Spergel (2009). Solid, dashed and dotted lines show the total, central and satellite HOD, respectively.

definition. The values are now 6–10 per cent, and are roughly the same as the observation. Thus the seemingly large satellite fraction of our mock LRGs does not conflict with the result of Reid & Spergel (2009), in which they derive a satellite fraction of $6.36^{+0.38}_{-0.39}$.

One can say that our definition is based on the kinematics of LRGs: what matters to explain the feature of RSDs in observation is how many LRGs have velocities relative to their host haloes. One important lesson we have learned through this analysis is that a non-negligible fraction of LRGs without a close companion (i.e. single-LRG systems) are not located at the centre of the host haloes. The situation is consistent with the result of a cross-correlation analysis of single-LRG systems with galaxy–galaxy lensing signal (and also cross-correlation with the photometric galaxies) performed by Hikage et al. (2013). Their result suggests that about 24 per cent of these LRGs have offset from the true gravitational centre of the systems (see fig. 10 in the reference).

4.6 Halo occupation distribution

The HOD is a useful approach to connect galaxies to dark matter haloes frequently adopted in the literature. This function is closely related to the multiplicity function presented in the previous subsection. Indeed, Reid & Spergel (2009) give a tight constraint on the HOD by fitting the observed multiplicity function. We here compare the HOD of our models with that in the literature and discuss the compatibility of a model based on the observed HOD to explain the anisotropic clustering of LRGs.

Fig. 16 shows the HODs measured from our mock catalogues with the best-fitting parameters. Centrals and satellites are, respectively, plotted by triangles and diamonds, while the sum of the two populations is shown as circles with error bar showing the scatter among different random realizations. Also shown by lines are the functions,

$$N_{\text{tot}} = N_{\text{cen}}(1 + N_{\text{sat}}), \quad (10)$$

$$N_{\text{cen}} = \frac{1}{2} \left[1 + \text{erf} \left(\frac{\log_{10} M_{\text{host}} - \log_{10} M_{\text{min}}}{\sigma} \right) \right], \quad (11)$$

$$N_{\text{sat}} = \left(\frac{M_{\text{host}} - M_{\text{cut}}}{M_1} \right)^\alpha, \quad (12)$$

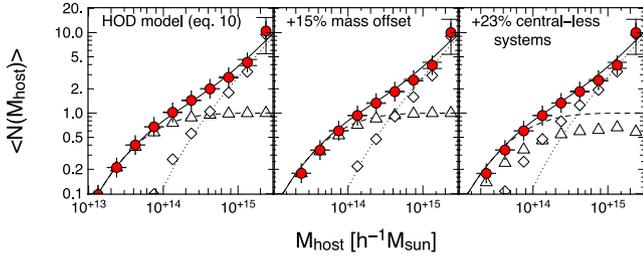


Figure 17. Mean halo occupation number as a function of the halo mass for models based on the HOD by Reid & Spergel (2009) (lines). We follow the central HOD in equation (11) and then assign satellites to subhaloes more massive than $10^{13} h^{-1} M_{\odot}$ whose host haloes have a central LRG to obtain almost identical HOD in the left-hand panel. We scale the masses by 15 per cent to adjust the mean number density in the middle panel. We finally replace centrals with satellites randomly following the procedure described in the text in the right-hand panel.

with the parameters, $M_{\min} = 5.796 \times 10^{13} h^{-1} M_{\odot}$, $M_{\text{cut}} = 3.6 \times 10^{13} h^{-1} M_{\odot}$, $M_1 = 3.564 \times 10^{14} h^{-1} M_{\odot}$, $\sigma = 0.7$ and $\alpha = 1.035$ determined by Reid & Spergel (2009) based on the observation. The dashed and dotted lines are the central and satellite contribution and the sum is depicted by the solid line.

Overall, the total HOD is very similar among our eight models, and its slope is close to that by Reid & Spergel (2009). There exists, however, an offset in the overall amplitude. This is likely to be due to the difference in the halo mass function between ours and that in Reid & Spergel (2009): the halo finder [FoF and spherical overdensity (SO)], output redshifts ($z = 0.35$ and 0.2) and a different cosmological model (WMAP 5- and 3-year) can give a noticeable change in the halo mass function. With a different mass function, the HOD that gives the correct mean number density should be different. Another difference in the HOD can be seen at the low-mass end. It is natural for our models to have a rather sharp cut because we impose a minimum halo mass to host a LRG. We need a more elaborate model to have a smoother HOD.

When focusing on each of the central and satellite contributions, the difference among the models are clearer. The central contribution is smaller for models in which we randomly discard more central subhaloes (i.e. Model 4), while the observed HOD is unity when the host halo mass is large. As already discussed in the previous subsection, an important point is that it is the HOD of the sum of the two populations that is directly constrained from observation, and the contributions from each population strongly depends on the definition of central and satellite LRGs. We thus do not consider the different HOD for each population problematic.

Given the similar but different HODs, we next examine whether the HOD derived from observation can explain the anisotropic clustering on large scales at the same time. We do this by employing again the subhaloes found in the simulations, but enforcing that the HOD is the same as that in Reid & Spergel (2009). We here consider the centre-of-mass positions and velocities for simplicity. We first assign LRGs to central subhaloes with the probability of equation (11) depending on the mass of the host haloes. We then consider haloes with a central LRG and assign satellite LRGs to other member subhaloes. We empirically find that we can recover the HOD for satellites given by equation (12) almost perfectly when we assign satellite LRGs to subhaloes with mass larger than $10^{13} h^{-1} M_{\odot}$. The resultant HOD is shown in the left-hand panel of Fig. 17. Despite almost the same HOD, the mean number density of these mock LRGs is about 20 per cent larger than the observed one. This is

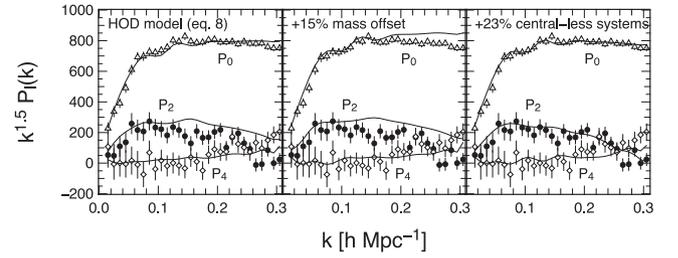


Figure 18. Multipole moments of the power spectrum measured from the mock LRGs constructed based on the observed HOD. The three panels correspond to those in Fig. 17.

again because of the difference in the halo mass function. We adjust the mean number density by scaling the masses, M_{\min} , M_{cut} and M_1 in the HOD function by 15 per cent (middle panel).

We compare the multipole moments of the power spectrum in redshift space measured from these new mock LRGs in the left-hand and middle panels of Fig. 18 before and after the mass scaling, respectively. We can see in the left-hand panel that the amplitude and the slope of the monopole and quadrupole moments are somewhat different from observed moments. Indeed the χ^2 statistics for these mocks is 222.1 for 90 data points, which is significantly larger than what we find for our other mock catalogues discussed so far. After the mass scaling, the amplitude of the monopole moment is now closer to the observation on large scales ($k \lesssim 0.15 h \text{ Mpc}^{-1}$) at the cost of larger gap on smaller scales as well as on the quadrupole moment (middle panel). The value of χ^2 gets larger, 405.1, after shifting the masses.

The middle panel of Fig. 18 suggests that the observed multipole moments are affected more strongly by the Fingers of God suppression than the mock LRGs constructed based on the observed HOD. Motivated by the high satellite fraction derived from the main analysis, we consider the following simple model. Starting from the mock catalogues with the HOD shown in the middle panel of Fig. 17, we randomly select a central LRG and reassign it to the most massive satellite subhalo without a LRG in the same host halo. Note that the total number of LRGs within a host halo is conserved throughout this process. We repeat this procedure to 23 per cent of the haloes with a central LRG to finally have a satellite fraction of 30 per cent without changing the total HOD (see right-hand panel of Fig. 17). We now have some LRG systems, which have only satellites. Note that such systems do not exist in the other two mock catalogues discussed in this subsection, nor in the model described by equation (10) by construction. We finally show the power spectrum measured from the mock LRGs with increased satellite fraction in the right-hand panel of Fig. 18. We now have an excellent agreement with observation with the χ^2 statistics being 82.0. This is comparable to the other models discussed in previous subsections.

We conclude that the large satellite fraction is always necessary to explain the multipole moments of the power spectrum independent of the detail of the model. This subsection also demonstrates that one needs a statistical procedure such as random sampling to reduce the total number or random swapping of centrals to satellites to explain both the mean number density (and also the HOD) and the clustering in redshift space, as long as we rely only on the mass of haloes and of subhaloes. Although more involved models with further physical inputs such as the formation epoch of the subhaloes (Masaki et al. 2013) would be an interesting next step, we leave it to a future investigation given the already successful fit to the observation.

5 SUMMARY AND DISCUSSIONS

The clustering statistics of galaxies on large scales provides us with a wealth of cosmological information. A simple parametric approach is usually adopted to absorb the uncertainty of galaxy bias. This simple treatment is still expected to work well as long as one restricts the analysis to sufficiently large scales where the nature of galaxies does not affect the statistics significantly. We consider in this study a bias model for LRGs based on subhaloes identified in a series of cosmological N -body simulations to pursue more realistic bias model for analyses on large scales. We determine the parameters that characterize the condition of subhaloes to host a LRG by directly fitting the first three non-zero multipole moments of the power spectrum in redshift space on scales $k \lesssim 0.3 h \text{ Mpc}^{-1}$. We find that models employing both central and satellite subhaloes successfully reproduce the clustering pattern within the statistical error, while a model with the central population alone clearly fails to do so.

An important ingredient in the model is the fraction of the satellite LRGs, and this parameter is constrained to 20–30 per cent slightly depending on the detail of the model. Though the large-scale clustering alone cannot constrain the detail of the model such as the HOD, the fraction of satellites is shown to be robust against the assumptions in the model. Indeed, we cannot explain the multipole moments unless we increase the satellite fraction by hand to 30 per cent even when we adopt the observed HOD of LRGs.

Since our model for LRGs is rather simple relying only on the mass of the host haloes and of subhaloes, it is difficult to explain the clustering and number density simultaneously without a statistical procedures such as random sampling or the random conversion of LRG hosts from centrals into satellites. We definitely need a more involved model or eventually need proper account of baryonic physics to have a prediction without random processes. We expect that these complications are more responsible for the clustering properties on small scales where the one-halo contribution is dominant. Although the range of wavenumber that we consider here is naively expected to be in a regime where the impact of non-linear physics is mild, it would be important to test the compatibility of these models against observation based on a proper statistical argument.

Although we use the simulation output obtained for one particular cosmological model, we induce additional anisotropic clustering signal by deforming the simulation box and changing the magnitude of RSDs. By doing this, we demonstrate that we can determine cosmological and LRG parameters simultaneously. The constraints on the distances (H and D_A) are found to be more robust against different prescriptions for LRGs thanks to the clear signal of BAOs in the observed spectrum, while non-linear corrections can give a systematic correction to the measurement of the growth rate parameter $f\sigma_8$, when the modelling of the velocities of LRGs is imperfect. We here propose a possible extension of the current analysis and a solution to the non-linear contamination to the measurement of the growth rate.

Our constraints on the parameters α_v , α_{\parallel} and α_{\perp} suggest that the cosmological model adopted in our simulations is slightly disfavoured by the observation. The most significant deviation from the fiducial value is seen in the parameter α_v especially on small scales. A simple step to go beyond the present analysis is to resimulate the clustering of LRGs with the best-fitting cosmological model obtained here and repeat the same analysis again. Also, we suggest to repeat the measurement of the power spectrum from the observation assuming a new distance–redshift relation. When the cosmological model adopted in the simulations as well as in

the distance–redshift relation is close enough to the true underlying model, these deformation parameters should be consistent with unity (i.e. no deformation, no velocity bias). We can iterate the analysis until this condition is satisfied. Although it is computationally too expensive to simulate the clustering over a cosmological parameter space, several times of iterations are well within reach of the current power of supercomputers. Since we wish to constrain the growth rate parameter in the *linear* regime, it might be helpful to introduce one more parameter in our model which controls the motion of mock LRGs relative to the host haloes. The growth rate parameter less affected by non-linearity may be derived after marginalizing over this new parameter in return for a looser constraint in the fraction of satellites.

Before presenting the final estimate of the cosmological parameters with these modifications to the analysis, several issues must be carefully treated. One of our approximate treatments is in the estimate of the statistical error. We ignore here the off-diagonal components of the covariance matrix of the observational data and also the statistical error in the template power spectrum measured from the simulations in finite volume. We discuss the current level of accuracy of the error estimation in Appendix A. Another is the effect of fibre collisions which is ignored in measuring the power spectrum analysed in this study. Although this effect is smaller than the current statistical accuracy on large scales (see Appendix B), it should be properly included in the construction of mock LRGs when we extend the analysis to smaller scales. Finally, the result of the analysis may depend on the subhalo finder (Pujol et al. 2013). We naively expect that the dependence is most prominent in the parameters that describe the nature of galaxies (e.g. the minimum halo mass to host a galaxy). However, the constraints on the cosmological parameters should also be verified carefully especially when the statistical error becomes very small with future ambitious survey projects.

We believe that this study is an important first step towards cosmological analyses using simulations as a theoretical template. The use of simulations allows us to take into account the non-linear and non-local nature of galaxy bias at full order. Also, our analysis demonstrates that we can break the degeneracy between the uncertainty in the nature of galaxies and the cosmological model. An important ingredient in this analysis is the fraction of satellite LRGs, and we show that we can still put a meaningful constraint on this parameter when we relax the assumption of the cosmological model. This is encouraging to us to continue along this line and apply our methodology to some future galaxy surveys targeting galaxies whose environmental properties are not understood very well.

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REFERENCES

- Abazajian K. N. et al., 2009, *ApJS*, 182, 543
- Adelman-McCarthy J. K. et al., 2007, *ApJS*, 172, 634
- Alcock C., Paczynski B., 1979, *Nature*, 281, 358
- Anderson L. et al., 2014, *MNRAS*, 439, 83
- Angulo R. E., White S. D. M., 2010, *MNRAS*, 405, 143
- Anselmi S., Matarrese S., Pietroni M., 2011, *J. Cosmol. Astropart. Phys.*, 6, 15
- Ballinger W. E., Peacock J. A., Heavens A. F., 1996, *MNRAS*, 282, 877
- Berlind A. A., Weinberg D. H., 2002, *ApJ*, 575, 587
- Bernardeau F., Colombi S., Gaztañaga E., Scoccimarro R., 2002, *Phys. Rep.*, 367, 1
- Bernardeau F., Crocce M., Scoccimarro R., 2008, *Phys. Rev. D*, 78, 103521
- Bernardeau F., Van de Rijt N., Vernizzi F., 2013, *Phys. Rev. D*, 87, 043530
- Beutler F. et al., 2012, *MNRAS*, 423, 3430
- Blake C., Collister A., Lahav O., 2008, *MNRAS*, 385, 1257
- Blake C. et al., 2011a, *MNRAS*, 415, 2876
- Blake C. et al., 2011b, *MNRAS*, 418, 1707
- Blake C. et al., 2012, *MNRAS*, 425, 405
- Carlson J., White M., Padmanabhan N., 2009, *Phys. Rev. D*, 80, 043531
- Chuang C.-H., Wang Y., 2012, *MNRAS*, 426, 226
- Chuang C.-H., Wang Y., 2013a, *MNRAS*, 431, 2634
- Chuang C.-H., Wang Y., 2013b, *MNRAS*, 435, 255
- Cole S., 1997, *MNRAS*, 286, 38
- Conroy C., Wechsler R. H., Kravtsov A. V., 2006, *ApJ*, 647, 201
- Crocce M., Scoccimarro R., 2006a, *Phys. Rev. D*, 73, 063519
- Crocce M., Scoccimarro R., 2006b, *Phys. Rev. D*, 73, 063520
- Crocce M., Scoccimarro R., 2008, *Phys. Rev. D*, 77, 023533
- Crocce M., Pueblas S., Scoccimarro R., 2006, *MNRAS*, 373, 369
- Crocce M., Scoccimarro R., Bernardeau F., 2012, *MNRAS*, 427, 2537
- Davis M., Efstathiou G., Frenk C. S., White S. D. M., 1985, *ApJ*, 292, 371
- Dekel A., Lahav O., 1999, *ApJ*, 520, 24
- Eisenstein D. J., Hu W., 1998, *ApJ*, 496, 605
- Eisenstein D. J. et al., 2001, *AJ*, 122, 2267
- Eisenstein D. J. et al., 2005, *ApJ*, 633, 560
- Eisenstein D. J., Seo H.-J., White M., 2007, *ApJ*, 664, 660
- Feldman H. A., Kaiser N., Peacock J. A., 1994, *ApJ*, 426, 23
- Guo H., Zehavi I., Zheng Z., 2012, *ApJ*, 756, 127
- Guo Q., White S., Angulo R. E., Henriques B., Lemson G., Boylan-Kolchin M., Thomas P., Short C., 2013, *MNRAS*, 428, 1351
- Guzzo L. et al., 2008, *Nature*, 451, 541
- Hikage C., Yamamoto K., 2013, *J. Cosmol. Astropart. Phys.*, 8, 19
- Hikage C., Mandelbaum R., Takada M., Spergel D. N., 2013, *MNRAS*, 435, 2345
- Hinshaw G. et al., 2013, *ApJS*, 208, 19
- Hiramatsu T., Taruya A., 2009, *Phys. Rev. D*, 79, 103526
- Ho S., Lin Y.-T., Spergel D., Hirata C. M., 2009, *ApJ*, 697, 1358
- Hockney R. W., Eastwood J. W., 1981, *Computer Simulation Using Particles*. McGraw-Hill, New York
- Hu W., Haiman Z., 2003, *Phys. Rev. D*, 68, 063004
- Hu W., Sawicki I., 2007, *Phys. Rev. D*, 76, 064004
- Jackson J. C., 1972, *MNRAS*, 156, 1P
- Jeong D., Komatsu E., 2006, *ApJ*, 651, 619
- Jeong D., Komatsu E., 2009, *ApJ*, 691, 569
- Jing Y. P., 2005, *ApJ*, 620, 559
- Kaiser N., 1984, *ApJ*, 284, L9
- Kaiser N., 1987, *MNRAS*, 227, 1
- Kazin E. A. et al., 2013, *MNRAS*, 435, 64
- Komatsu E. et al., 2009, *ApJS*, 180, 330
- Komatsu E. et al., 2011, *ApJS*, 192, 18
- Kulkarni G. V., Nichol R. C., Sheth R. K., Seo H.-J., Eisenstein D. J., Gray A., 2007, *MNRAS*, 378, 1196
- Lewis A., Challinor A., Lasenby A., 2000, *ApJ*, 538, 473
- Linder E. V., 2008, *Astropart. Phys.*, 29, 336
- Ma C.-P., Fry J. N., 2000, *ApJ*, 543, 503
- McCullagh N., Neyrinck M. C., Szapudi I., Szalay A. S., 2013, *ApJ*, 763, L14
- McDonald P., 2006, *Phys. Rev. D*, 74, 103512
- McDonald P., 2007, *Phys. Rev. D*, 75, 043514
- Mandelbaum R., Seljak U., Kauffmann G., Hirata C. M., Brinkmann J., 2006, *MNRAS*, 368, 715
- Masaki S., Hikage C., Takada M., Spergel D. N., Sugiyama N., 2013, *MNRAS*, 433, 3506
- Matarrese S., Pietroni M., 2007, *J. Cosmol. Astropart. Phys.*, 6, 26
- Matarrese S., Pietroni M., 2008, *Modern Phys. Lett. A*, 23, 25
- Matsubara T., 2008a, *Phys. Rev. D*, 77, 063530
- Matsubara T., 2008b, *Phys. Rev. D*, 78, 083519
- Matsubara T., Suto Y., 1996, *ApJ*, 470, L1
- Mead A., Peacock J., 2014, *MNRAS*, 440, 1233
- Nishimichi T., Taruya A., 2011, *Phys. Rev. D*, 84, 043526
- Nishimichi T. et al., 2007, *PASJ*, 59, 1049
- Nishimichi T. et al., 2009, *PASJ*, 61, 321
- Nishizawa A. J., Takada M., Nishimichi T., 2013, *MNRAS*, 433, 209
- Oka A., Saito S., Nishimichi T., Taruya A., Yamamoto K., 2014, *MNRAS*, 439, 2515
- Okamura T., Taruya A., Matsubara T., 2011, *J. Cosmol. Astropart. Phys.*, 8, 12
- Okumura T., Matsubara T., Eisenstein D. J., Kayo I., Hikage C., Szalay A. S., Schneider D. P., 2008, *ApJ*, 676, 889
- Padmanabhan N., White M., 2008, *Phys. Rev. D*, 77, 123540
- Padmanabhan N., White M., 2009, *Phys. Rev. D*, 80, 063508
- Padmanabhan N., White M., Norberg P., Porciani C., 2009, *MNRAS*, 397, 1862
- Peacock J. A., Smith R. E., 2000, *MNRAS*, 318, 1144
- Percival W. J. et al., 2004, *MNRAS*, 353, 1201
- Pietroni M., 2008, *J. Cosmol. Astropart. Phys.*, 10, 36
- Planck Collaboration XVI, 2013, preprint ([arXiv:1303.5076](https://arxiv.org/abs/1303.5076))
- Pujol A. et al., 2014, *MNRAS*, 438, 3205
- Reid B. A., Spergel D. N., 2009, *ApJ*, 698, 143
- Reid B. A. et al., 2012, *MNRAS*, 426, 2719
- Samushia L. et al., 2013, *MNRAS*, 429, 1514
- Sánchez A. G. et al., 2013, *MNRAS*, 433, 1202
- Sato M., Matsubara T., 2011, *Phys. Rev. D*, 84, 043501
- Scoccimarro R., 1998, *MNRAS*, 299, 1097
- Scoccimarro R., Sheth R. K., Hui L., Jain B., 2001, *ApJ*, 546, 20
- Seljak U., 2000, *MNRAS*, 318, 203
- Sherwin B. D., Zaldarriaga M., 2012, *Phys. Rev. D*, 85, 103523
- Shoji M., Jeong D., Komatsu E., 2009, *ApJ*, 693, 1404
- Simha V., Cole S., 2013, *MNRAS*, 436, 1142
- Somerville R. S., Lemson G., Sigad Y., Dekel A., Kauffmann G., White S. D. M., 2001, *MNRAS*, 320, 289
- Song Y.-S., Percival W. J., 2009, *J. Cosmol. Astropart. Phys.*, 10, 4
- Springel V., 2005, *MNRAS*, 364, 1105
- Springel V., White S. D. M., Tormen G., Kauffmann G., 2001, *MNRAS*, 328, 726
- Starobinsky A. A., 2007, *Soviet J. Exp. Theor. Phys. Lett.*, 86, 157
- Sugiyama N. S., Futamase T., 2012a, *ApJ*, 760, 114
- Sugiyama N. S., Futamase T., 2012b, preprint ([arXiv:1210.7499](https://arxiv.org/abs/1210.7499))
- Sugiyama N. S., Spergel D. N., 2014, *J. Cosmol. Astropart. Phys.*, 2, 42
- Takahashi R. et al., 2009, *ApJ*, 700, 479
- Taruya A., Hiramatsu T., 2008, *ApJ*, 674, 617
- Taruya A., Suto Y., 2000, *ApJ*, 542, 559
- Taruya A., Nishimichi T., Saito S., Hiramatsu T., 2009, *Phys. Rev. D*, 80, 123503
- Taruya A., Saito S., Nishimichi T., 2011, *Phys. Rev. D*, 83, 103527
- Taruya A., Bernardeau F., Nishimichi T., Codis S., 2012, *Phys. Rev. D*, 86, 103528
- Taruya A., Nishimichi T., Bernardeau F., 2013, *Phys. Rev. D*, 87, 083509
- Tormen G., Bertschinger E., 1996, *ApJ*, 472, 14
- Valageas P., 2007, *A&A*, 465, 725
- Valageas P., 2011, *A&A*, 526, A67
- Valageas P., 2013, *Phys. Rev. D*, 88, 083524
- Valageas P., Nishimichi T., 2011, *A&A*, 527, A87
- Valageas P., Nishimichi T., Taruya A., 2013, *Phys. Rev. D*, 87, 083522
- Wake D. A. et al., 2008, *MNRAS*, 387, 1045

- White M., Zheng Z., Brown M. J. I., Dey A., Jannuzi B. T., 2007, *ApJ*, 655, L69
- Xu X., Cuesta A. J., Padmanabhan N., Eisenstein D. J., McBride C. K., 2013, *MNRAS*, 431, 2834
- Yamamoto K., Nakamichi M., Kamino A., Bassett B. A., Nishioka H., 2006, *PASJ*, 58, 93
- Yamamoto K., Sato T., Hütsi G., 2008, *Progress Theor. Phys.*, 120, 609
- Yamamoto K., Nakamura G., Hütsi G., Narikawa T., Sato T., 2010, *Phys. Rev. D*, 81, 103517
- York D. G. et al., 2000, *AJ*, 120, 1579
- Yoshikawa K., Taruya A., Jing Y. P., Suto Y., 2001, *ApJ*, 558, 520
- Zehavi I. et al., 2005, *ApJ*, 621, 22
- Zheng Z., Zehavi I., Eisenstein D. J., Weinberg D. H., Jing Y. P., 2009, *ApJ*, 707, 554

APPENDIX A: CALCULATION OF THE MODEL MULTIPOLES AND THEIR STATISTICAL ERROR

In this appendix, we show the accuracy of the multipole moments of the power spectrum measured from N -body simulations and discuss the validity of our approximate treatment of the covariance matrix.

First, we discuss the estimator of the multipoles used in their measurement from N -body simulations. One of a simple estimator adopts a binning scheme and assigns each wave vector to a wavenumber bin:

$$\hat{P}_{\ell, \text{binned}}^{(n)}(k_i) = \frac{1}{N_i} \frac{2\ell + 1}{2} \sum_{k \in i\text{th bin}} \mathcal{P}_{\ell}(\mu_k) \left| \delta_k^{(n)} \right|^2, \quad (\text{A1})$$

where N_i stands for the number of modes available in the i th wavenumber bin. This estimator, however, suffers from an artificial effect caused by the discreteness of wave vectors in finite volume. On large scales, the number of modes N_i is usually small, and therefore we cannot sample modes efficiently along the μ_k direction. Note that this effect cannot be mitigated even when we increase the number of independent random realizations, since we sample exactly the same wave vectors from all the simulations unless we change the box size for different realizations. Higher multipole moments ($\ell = 2$ and 4 in our case) are more sensitive to this effect since the kernel $\mathcal{P}_{\ell}(\mu_k)$ strongly depends on the argument, μ_k . This effect is discussed e.g. in Nishimichi & Taruya (2011), Taruya, Nishimichi & Bernardeau (2013) and its impact on the hexadecapole is shown to be significant even when the simulation box size is as large as $\sim 2 h^{-1} \text{ Gpc}$.

In this study, we thus implement a different estimator of the multipole moments. As outlined in Section 3.3, we apply the B-spline fitting to the *unbinned* weighted square of the density contrast (i.e. equation 3). This is demonstrated in the top panel of Fig. A1. We explain the detail of this procedure here. We choose 13 breakpoints equally spaced from 0 to $0.31 h \text{ Mpc}^{-1}$ to construct the basis function (vertical dashed lines). This is chosen so that the resultant curve is smooth enough but the wiggles of BAOs are not degraded significantly. We employ a cubic B-spline function and fit the data points shown by points.

With this procedure, the aforementioned artificial noise in the multipoles is greatly reduced. We show the model power spectrum based on the B-spline fitting described above and compare it with the binned power spectrum adopting the best-fitting parameters of Model 5a. The error bars on the binned power spectrum indicate the 1σ statistical error of the data points expected from the scatter among different realizations. Higher multipole moments, especially the hexadecapole moment, show a notable noisy pattern larger than the typical size of error bars. On the other hand, the multipoles based

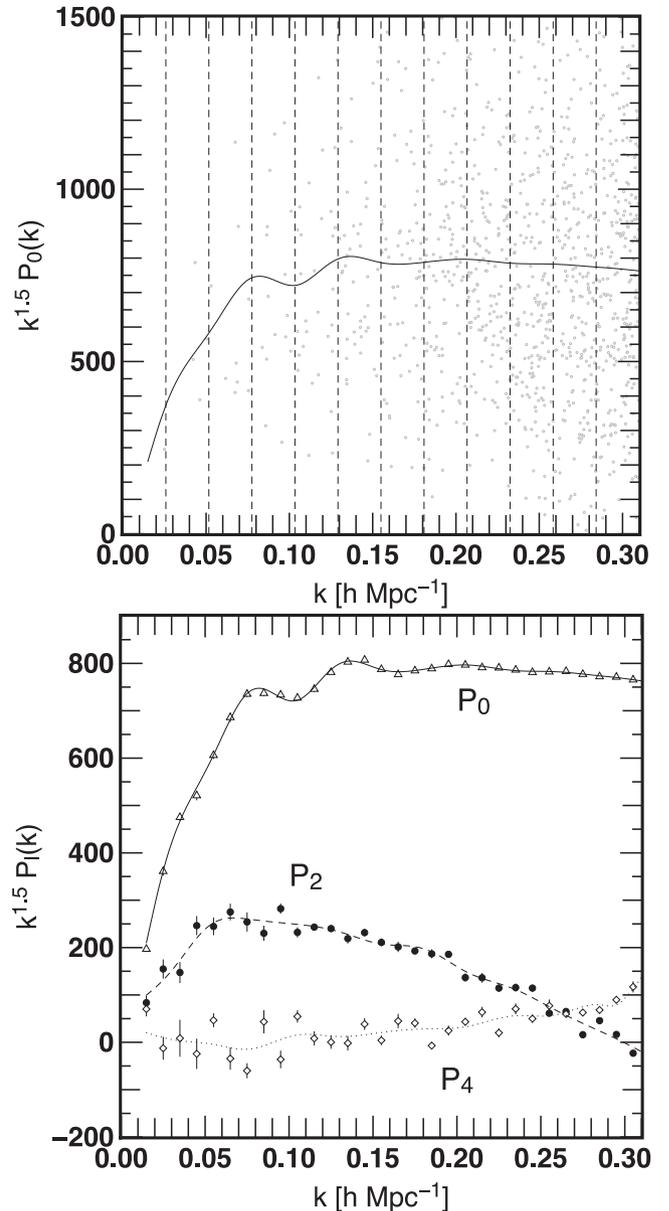


Figure A1. Measurement of the power spectrum based on the cubic B-spline fitting. We plot in the top panel a random sample of 1000 data points, $|\delta_k^{(n)}(k)|$, out of 3 880 650 fitted to compute the monopole moment (points) and the result of the fit (solid line). The vertical dashed lines show the positions of breakpoints used to construct the basis function. The bottom panel shows a comparison of multipole moments of the power spectrum calculated from different estimators. Lines show the multipoles based on the cubic B-spline fitting while the results of the binned estimator are depicted by symbols with error bars.

on the B-spline fitting are much smoother. We can still find clearly by eye the baryon acoustic signature on the monopole moment even after the B-spline fitting.

We next discuss the statistical accuracy of the model power spectrum. One of an approximate treatment in this study is that we take account of the statistical error only on the observed power spectrum but neglect that on the model power spectrum estimated from finite simulation volume. Even though the total volume of 11 realizations ($16 h^{-3} \text{ Gpc}^3$, and can be slightly different when we consider the AP distortion) is larger than the observed volume ($1.39 h^{-3} \text{ Gpc}^3$ in

our fiducial cosmological model), the accuracy of the model power spectrum should carefully be checked.

We plot in Fig. A2 the statistical error on the simulated multipole moments of the power spectrum divided by that on the observational data. In the plot, we again adopt the best-fitting parameters of Model 5a and estimate the statistical error on the model power spectrum by

$$[\Delta P_{\ell, \text{sim}}(k)]^2 = \frac{1}{N_r - 1} \sum_n \left[\hat{P}_\ell^{(n)}(k) - \bar{P}_\ell(k) \right]^2, \quad (\text{A2})$$

where $N_r = 11$ is the number of independent realizations and $\hat{P}_\ell^{(n)}$ and $\bar{P}_\ell(k)$ denote the ℓ th multipole moment of the power spectrum measured from the n th realization and that averaged over N_r realizations, respectively. We plot in Fig. A2 $\Delta P_{\ell, \text{sim}}(k)$ divided by the statistical error on the observed multipoles.

Overall, the ratio is 10–30 per cent depending on the wavenumber and is almost independent of multipoles. We can roughly estimate it from the fact that the error bar scale as volume $^{-1/2}$, and this gives a 29 per cent smaller error on the model multipoles. This is consistent with the result of Fig. A2 on large scales ($k \sim 0.03 h \text{ Mpc}^{-1}$). On smaller scales since our mock LRGs have a larger number density than the observed LRGs, the mock multipoles are less affected by shot noise. That is why the ratio decreases with wavenumber.

On the other hand, the estimate of the covariance matrix of the observed multipoles adopted in this study is expected to be similarly accurate. For example, the cross-covariance between the different multipoles, which we ignore in this study, can be evaluated by linear theory, and the correlation coefficient for the monopole–quadrupole cross-covariance is ~ 20 per cent for $\beta = f/b \sim 0.28$ (see fig. 4 of Taruya et al. 2011). Also, the non-Gaussian component of the covariance matrix is $\lesssim 30$ per cent in redshift space on the scale of our interest according to Takahashi et al. (2009) in which they perform 5000 cosmological N -body simulation to estimate the covariance matrix (see figs 2 and 3 in the reference). We thus expect that the final error ellipses of the model parameters is accurate at this level (10–30 per cent).

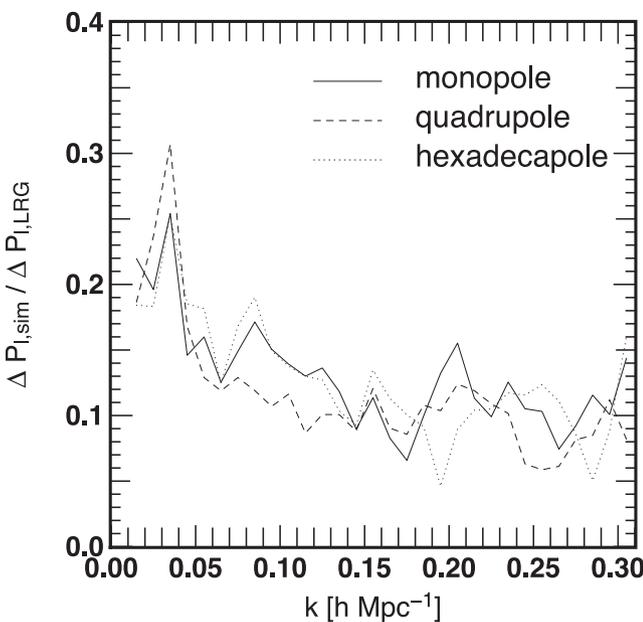


Figure A2. Statistical error on the model power spectrum compared with the measured one.

APPENDIX B: IMPACT OF FIBRE COLLISIONS

We discuss the possible impact of the fibre collisions on the multipole moments of the power spectrum in this appendix. The LRG sample analysed in Yamamoto et al. (2010) suffers from spectroscopic fibre collisions, and some targeted LRG candidates do not have spectra when they have a close companion within 55 arcsec. Although both of close pairs have redshifts thanks to multiple observations in some region of the sky, the redshifts about 7 per cent of the targeted objects are not measured. In the measurement of the power spectrum, Yamamoto et al. (2010) do not make a correction for fibre collisions for simplicity. However, since fibre collisions occur pairs of galaxies that are very close in the plane-of-sky direction, they can be a source of apparent anisotropy of the clustering of galaxies. This effect is studied in Guo, Zehavi & Zheng (2012) for the multipole moment of the two-point correlation function. Although they find that it is significant only on small scales ($\lesssim 10 h^{-1} \text{ Mpc}$), its impact on the Fourier counterpart is not trivial. Here we give a rough estimate of this effect using the mock LRGs.

We use mock LRGs in our simulation and simulate this effect by searching for fibre collision pairs and randomly discarding one of two galaxies subject to a fibre collision. There are small number of mock LRGs that have more than one companions that are closer than 55 arcmin, and we discard these mock LRGs preferentially to reduce the total number of discarded mock LRGs. We keep the distant observer approximation and convert the separation of pairs perpendicular to the line of sight to the observed angle assuming the angular diameter distance at the effective redshift of observation, $z = 0.3$. We do not take into account of multiple observations of plates to assess the maximum impact of fibre collisions. Strictly speaking, we should redo the analysis by including this effect into the MCMC analysis to be more consistent. However, we simply take the best-fitting mock catalogue explained in the main text (Model 5a) and then implement the effect of fibre collisions since a MCMC analysis with this effect is time consuming.

We show the result in Fig. B1. Notice that we plot a wider wavenumber range than in the main text to see the impact of fi-

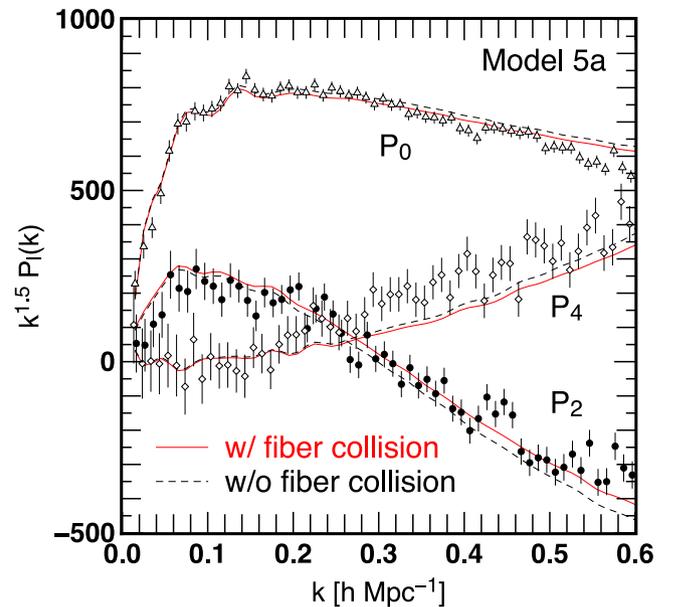


Figure B1. Impact of fibre collisions on the power spectrum in redshift space. We plot the multipole moments measured from our best-fitting mock LRGs (Model 5a) with (solid) and without (dashed) fibre collisions.

bre collisions down to smaller scales, but the fitting range is same as before (i.e. $k_{\text{max}} = 0.305 h \text{ Mpc}^{-1}$). The figure shows the multipole moments with or without fibre collisions (solid and dashed lines, respectively). We also plot the observational data by symbols with error bars. Fibre collisions suppress the monopole and hexadecapole, and amplify the quadrupole moment. The effect is more prominent on smaller scales. However, at $k \lesssim 0.3 h \text{ Mpc}^{-1}$, the difference between the solid and the dashed lines is typically much smaller than the statistical error on the observational data. Thus we expect that fibre collisions do not affect the parameter constraints

significantly as long as we focus on relatively large scales. Since our fitting procedure using N -body data is in principle applicable to non-linear scales unlike perturbation theory based models that break down on certain scales where non-linearity is very strong, a proper inclusion of fibre collisions in the analysis might be important to enlarge the wavenumber range from which we can extract cosmological information robustly.

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