



HAL
open science

Secular resonant dressed orbital diffusion - I. Method and WKB limit for tepid discs

Jean-Baptiste Fouvy, Christophe Pichon, Simon Prunet

► **To cite this version:**

Jean-Baptiste Fouvy, Christophe Pichon, Simon Prunet. Secular resonant dressed orbital diffusion - I. Method and WKB limit for tepid discs. Monthly Notices of the Royal Astronomical Society, Oxford University Press (OUP): Policy P - Oxford Open Option A, 2015, 449, pp.1967-1981. 10.1093/mnras/stv359 . insu-03644951

HAL Id: insu-03644951

<https://hal-insu.archives-ouvertes.fr/insu-03644951>

Submitted on 26 Apr 2022

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



Secular resonant dressed orbital diffusion – I. Method and WKB limit for tepid discs

Jean-Baptiste Fouvry,^{1,2★} Christophe Pichon^{1,2,3} and Simon Prunet^{1,2,4}

¹ *Institut d'Astrophysique de Paris, CNRS (UMR7095), 98 bis Boulevard Arago, 75014 Paris, France*

² *UPMC Univ. Paris 06, UMR7095, 98 bis Boulevard Arago, 75014 Paris, France*

³ *Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK*

⁴ *CFHT Corporation, 65-1238 Mamalahoa Hwy, Kamuela, Hawaii 96743, USA*

Accepted 2015 February 17. Received 2015 January 29; in original form 2014 December 1

ABSTRACT

The equation describing the secular diffusion of a self-gravitating collisionless system induced by an exterior perturbation is derived while assuming that the time-scale corresponding to secular evolution is much larger than that corresponding to the natural frequencies of the system. Its two-dimensional formulation for a tepid galactic disc is also derived using the epicyclic approximation. Its Wentzel-Kramers-Brillouin (WKB) limit is found while assuming that only tightly wound transient spirals are sustained by the disc. It yields a simple quadrature for the diffusion coefficients which provides a straightforward understanding of the loci of maximal diffusion within the disc.

Key words: diffusion – gravitation – galaxies: evolution – galaxies: kinematics and dynamics – galaxies: spiral.

1 INTRODUCTION

Understanding the secular dynamical evolution of galaxies over cosmic time has been a long standing subject of interest. Indeed, self-gravitating collisionless systems such as galaxies may, over cosmic times, change their kinematical structure as they respond secularly to their evolving environment, in a manner which depends both on their internal orbital structure, but also on how this structure resonates with its environment or with itself. It is therefore critical to distinguish in the physical properties of galaxies the contributions from the cosmic environment (nurture) and its induced perturbations from the ones coming from the intrinsic properties of the galaxies (nature). For thermodynamically improbable cold systems such as galactic discs, their gravitational susceptibility should also play a specific role which must be taken into account when studying their long-term evolution.

To tackle this question, one can rely on numerical N -body simulations of higher resolutions to take into account non-linear physical processes (e.g. Dubois et al. 2014) or perform idealized well-crafted numerical experiments (Sellwood & Athanassoula 1986; Earn & Sellwood 1995; Sellwood 2012). With such statistical investigations, one can assess the importance and the role of the orbital structure of a galactic disc to drive its secular evolution. Angle-action variables (Goldstein 1950; Born 1960; Binney & Tremaine 2008) and the matrix method (Kalnajs 1976) also allow us to take into account the self-gravitating amplification of such collisionless systems. From this analytical framework, one should therefore be able to derive flexible qualitative and quantitative equations describ-

ing the secular dynamics of discs, without relying on the implementation of the corresponding demanding numerical models.

This topic of secular evolution has been addressed via the dressed Fokker–Planck equation, where the source of secular evolution for a self-gravitating system is taken to be potential fluctuations from an external bath, e.g. corresponding to the cosmic environment. Binney & Lacey (1988) computed the first- and second-order diffusion coefficients describing the orbits deviation induced by fluctuations in the gravitational potential. Weinberg (1993) showed the importance of self-gravity on the non-local and collective relaxation of stellar systems. Weinberg (2001a,b) considered the dressed gravitational amplification of Poisson shot noise in stellar systems and the impact of the properties of the noise processes. Ma & Bertschinger (2004) used a quasi-linear approach to investigate dark matter diffusion induced by cosmological fluctuations. Pichon & Aubert (2006) sketched a time-decoupling approach to solve the collisionless Boltzmann equation and applied it to the statistical study of dynamical flows through dark matter haloes. Chavanis (2012a) considered the evolution of homogeneous collisionless systems forced by an external perturbation. Nardini et al. (2012) also considered the evolution of such long-range interacting systems when driven by external stochastic forces.

Using an argument based on time-scale decoupling inspired from Pichon & Aubert (2006), we present here a careful and detailed derivation of the secular resonant dressed orbital diffusion equation for a general collisionless self-gravitating system undergoing external perturbations¹. We then develop this formalism for the secular

¹ We also recover a missing $1/2$ factor absent from this previous work, due to an error in temporal integration bounds.

* E-mail: fouvry@iap.fr

evolution of an infinitely thin galactic disc. In order to circumvent the complex direct or analytical calculation of the modes of a galactic disc carried only for a small number of disc models (Zang 1976; Kalnajs 1977; Goodman 1988; Weinberg 1991; Vauterin & Dejonghe 1996; Pichon & Cannon 1997; Evans & Read 1998; Jalali & Hunter 2005), we then rely on the Wentzel-Kramers-Brillouin (WKB) approximation (Liouville 1837; Toomre 1964; Kalnajs 1965; Lin & Shu 1966) to obtain a tractable algebraic expression for both the gravitational susceptibility of the system and the associated diffusion coefficients. Within the realm of this approximation, which should apply to cold enough discs, the diffusion coefficient reduce to simple quadratures.

The paper is organized as follows. Section 2 presents a derivation of the general dressed Fokker–Planck equation for a perturbed self-gravitating collisionless system. Appendix A provides a complementary derivation based on Hamilton’s equation, extending Binney & Lacey (1988) to self-gravitating systems. Section 3 focuses on razor thin axisymmetric galactic discs within the WKB approximation. Some of the underlying calculations are postponed to Appendices B and C. Finally, Section 4 wraps up.

2 SECULAR DIFFUSION EQUATION

The secular diffusion equation aims at describing the long-term aperiodic evolution of a self-gravitating collisionless system, perturbed by exterior potential fluctuations. A typical application for this formalism is the study of a galactic disc undergoing (cosmic) perturbations from its surrounding dark matter halo or the secular diffusion of accretion streams within the Galactic halo. We will suppose that the background gravitational potential of the system is stationary and integrable, so that we may always remap the usual phase-space coordinates (x, v) to the angle-action coordinates (θ, \mathbf{J}) . This is a strong assumption, as one could imagine situations where the secular evolution breaks symmetry warranting integrability. The angles θ are 2π -periodic, whereas the actions \mathbf{J} are conserved for a few dynamical times and secularly drift with cosmic time.

2.1 Evolution equations

We consider a stationary Hamiltonian $H_0(\mathbf{J})$, associated with a stationary background gravitational potential ψ_0 . We also consider a quasi-stationary distribution function $F_0(\mathbf{J}, t)$, which, at fixed secular time, only depends on the actions thanks to Jeans theorem (Jeans 1915). Finally, we suppose that an exterior source is perturbing this stationary system, so that we can expand the distribution function and the Hamiltonian of the system as

$$\begin{cases} F(\mathbf{J}, \theta, t) = F_0(\mathbf{J}, t) + f(\mathbf{J}, \theta, t), \\ H(\mathbf{J}, \theta, t) = H_0(\mathbf{J}) + \psi^e(\mathbf{J}, \theta, t) + \psi^s(\mathbf{J}, \theta, t), \end{cases} \quad (1)$$

where f is the perturbation of the distribution function, ψ^e is the perturbing *exterior* potential generated by the exterior source, and ψ^s is the *self-response* from the system induced by its self-gravity. This decomposition now involves two main temporal scales. The shortest scale is the fluctuation time-scale, during which $F_0(\mathbf{J})$ may be considered constant. The longest time-scale corresponds to the secular evolution time-scale. The perturbations are supposed to be small so that $f \ll F_0$ and $\psi^e, \psi^s \ll \psi_0$. The evolution of the collisionless system is then driven by Boltzmann collisionless equation which reads

$$\frac{dF}{dt} = \frac{\partial F}{\partial t} + \{H, F\} = 0, \quad (2)$$

where $\{H, F\}$ is the Poisson bracket. Injecting the decomposition from equation (1) in Boltzmann’s equation (2), we obtain

$$\begin{aligned} 0 = & \frac{\partial F_0}{\partial t} + \frac{\partial f}{\partial t} - \left[\frac{\partial \psi^e}{\partial \theta} + \frac{\partial \psi^s}{\partial \theta} \right] \cdot \frac{\partial F_0}{\partial \mathbf{J}} - \left[\frac{\partial \psi^e}{\partial \theta} + \frac{\partial \psi^s}{\partial \theta} \right] \cdot \frac{\partial f}{\partial \mathbf{J}} \\ & + \boldsymbol{\Omega} \cdot \frac{\partial f}{\partial \theta} + \left[\frac{\partial \psi^e}{\partial \mathbf{J}} + \frac{\partial \psi^s}{\partial \mathbf{J}} \right] \cdot \frac{\partial f}{\partial \theta}, \end{aligned} \quad (3)$$

where we have defined the frequencies of the motion on the action-torii as

$$\dot{\theta} = \boldsymbol{\Omega} = \frac{\partial H_0}{\partial \mathbf{J}}. \quad (4)$$

In order to derive the corresponding secular equation, we perform an angle-average on θ of equation (3). All terms involving a single derivation $\partial/\partial\theta$ are equal to 0, since the angles θ are 2π -periodic. Moreover, we have $\int_{\theta} d\theta f = 0$, because all the variations independent of θ are included in $F_0(\mathbf{J}, t)$. As F_0 is independent of θ , we obtain

$$\begin{aligned} \frac{\partial F_0}{\partial t} = & \frac{1}{(2\pi)^d} \int d\theta \left[\frac{\partial \psi^e}{\partial \theta} + \frac{\partial \psi^s}{\partial \theta} \right] \cdot \frac{\partial f}{\partial \mathbf{J}} \\ & - \frac{1}{(2\pi)^d} \int d\theta \left[\frac{\partial \psi^e}{\partial \mathbf{J}} + \frac{\partial \psi^s}{\partial \mathbf{J}} \right] \cdot \frac{\partial f}{\partial \theta}, \end{aligned} \quad (5)$$

where d is the dimension of the physical space. Using Schwartz theorem, this secular diffusion equation can be written under the shorter form

$$\frac{\partial F_0}{\partial t} = \frac{1}{(2\pi)^d} \frac{\partial}{\partial \mathbf{J}} \cdot \left[\int d\theta f \frac{\partial[\psi^e + \psi^s]}{\partial \theta} \right]. \quad (6)$$

Equation (6), written as the divergence of a flux, emphasizes the fact that during orbital diffusion the total number of stars is strictly conserved. Recalling that $f \ll F_0$ and $\psi^e, \psi^s \ll \psi_0$, the secular evolution equation (6) shows that $\partial F_0/\partial t$ is in fact a second-order term.

Correspondingly, keeping only first-order terms in equation (3), we obtain the second diffusion equation for the short time-scale, which reads

$$\frac{\partial f}{\partial t} + \boldsymbol{\Omega} \cdot \frac{\partial f}{\partial \theta} - \frac{\partial F_0}{\partial \mathbf{J}} \cdot \frac{\partial[\psi^e + \psi^s]}{\partial \theta} = 0. \quad (7)$$

This equation describes the evolution of the perturbation distribution function f on the fast fluctuating time-scale. On such time-scales, $\partial F_0/\partial \mathbf{J}$ will be considered as independent of t . The next step is to study the fast fluctuating equation (7), whose solutions will allow us to estimate the diffusion coefficients for the secular evolution given by equation (6) and describe the diffusion of the quasi-stationary distribution function F_0 in action-space.

2.2 Fourier expansion

One of the many assets of the angle-action variables is that the angles θ are 2π -periodic allowing us to perform discrete Fourier expansions with respect to these variables. We define the Fourier transform in angles of a function $X(\theta, \mathbf{J})$ as

$$\begin{cases} X(\theta, \mathbf{J}) = \sum_{m \in \mathbb{Z}^d} X_m(\mathbf{J}) e^{im \cdot \theta}, \\ X_m(\mathbf{J}) = \frac{1}{(2\pi)^d} \int d\theta X(\theta, \mathbf{J}) e^{-im \cdot \theta}. \end{cases} \quad (8)$$

Thanks to this transformation, the evolution equation (7) takes the form

$$\frac{\partial f_m}{\partial t} + i\mathbf{m} \cdot \boldsymbol{\Omega} f_m - i\mathbf{m} \cdot \frac{\partial F_0}{\partial \mathbf{J}} [\psi_m^e + \psi_m^s] = 0. \quad (9)$$

At this stage, we introduce the assumption of time-scale decoupling and *push* the secular time to infinity. As a consequence, in the upcoming calculations, we will suppose that $\partial F_0/\partial \mathbf{J} = \text{cst}$. Forgetting transient terms and bringing the initial time to $-\infty$, to focus only on the forced regime, the equation (9) can be solved explicitly, leading to

$$f_m(\mathbf{J}, t) = \int_{-\infty}^t d\tau e^{-im\cdot\Omega(t-\tau)} i\mathbf{m} \cdot \frac{\partial F_0}{\partial \mathbf{J}} [\psi_m^e + \psi_m^s](\mathbf{J}, \tau). \quad (10)$$

We define the temporal Fourier transform as

$$\begin{cases} \hat{f}(\omega) = \int_{-\infty}^{+\infty} dt f(t) e^{i\omega t}, \\ f(t) = \frac{1}{2\pi} \int_{-\infty}^{+\infty} d\omega \hat{f}(\omega) e^{-i\omega t}. \end{cases} \quad (11)$$

Taking the Fourier transform of equation (9) at the frequency ω , one can write

$$\hat{f}_m(\mathbf{J}, \omega) = -\frac{\mathbf{m} \cdot \partial F_0/\partial \mathbf{J}}{\omega - \mathbf{m} \cdot \boldsymbol{\Omega}} [\hat{\psi}_m^e(\mathbf{J}, \omega) + \hat{\psi}_m^s(\mathbf{J}, \omega)]. \quad (12)$$

2.3 Matrix method

An important property of this self-gravitating system is that the perturbed distribution function f is consistent with the self-gravitating potential ψ^s and its associated density ρ^s , so that we have

$$\rho^s(\mathbf{x}, t) = \int d\mathbf{v} f(\mathbf{x}, \mathbf{v}, t) = \sum_m \int d\mathbf{v} f_m(\mathbf{J}, t) e^{im\cdot\theta}. \quad (13)$$

In order to simplify further equation (13), we follow Kalnajs matrix method (Kalnajs 1976) and introduce a complete biorthonormal basis of potentials and densities $\psi^{(p)}(\mathbf{x})$ and $\rho^{(p)}(\mathbf{x})$, such that

$$\begin{cases} \nabla^2 \psi^{(p)} = 4\pi G \rho^{(p)}, \\ \int d\mathbf{x} [\psi^{(p)}(\mathbf{x})]^* \rho^{(q)}(\mathbf{x}) = -\delta_p^q. \end{cases} \quad (14)$$

On such a basis, we can write the exterior and the self-potentials as

$$\begin{cases} \psi^s(\mathbf{x}, t) = \sum_{p \in \mathbb{N}} a_p(t) \psi^{(p)}(\mathbf{x}), \\ \psi^e(\mathbf{x}, t) = \sum_{p \in \mathbb{N}} b_p(t) \psi^{(p)}(\mathbf{x}). \end{cases} \quad (15)$$

The linearity of Poisson's equation ensures that this decomposition translates into $\rho^s(\mathbf{x}, t) = \sum_p a_p(t) \rho^{(p)}(\mathbf{x})$. Using the biorthogonality of the basis, we multiply equation (13) by $[\psi^{(p)}(\mathbf{x})]^*$ and integrate over all positions to obtain

$$a_p(t) = -\sum_m \int d\mathbf{x} d\mathbf{v} f_m(\mathbf{J}, t) e^{im\cdot\theta} [\psi^{(p)}(\mathbf{x})]^*. \quad (16)$$

2.4 Response matrix and self-consistency

As the transformation $(\mathbf{x}, \mathbf{v}) \mapsto (\boldsymbol{\theta}, \mathbf{J})$ is canonical, we have $d\mathbf{x} d\mathbf{v} = d\boldsymbol{\theta} d\mathbf{J}$. The integration on $\boldsymbol{\theta}$ in equation (16) is straightfor-

ward since only $[\psi^{(p)}(\mathbf{x})]^* e^{im\cdot\theta}$ depends on it, so that it becomes

$$a_p(t) = -(2\pi)^d \sum_m \int d\mathbf{J} f_m(\mathbf{J}, t) [\psi_m^{(p)}(\mathbf{J})]^*. \quad (17)$$

Using the expression (12) and taking the temporal Fourier transform of equation (17) at the frequency ω , one obtains

$$\begin{aligned} \hat{a}_p(\omega) &= (2\pi)^d \sum_q [\hat{a}_q(\omega) + \hat{b}_q(\omega)] \\ &\times \sum_m \int d\mathbf{J} \frac{\mathbf{m} \cdot \partial F_0/\partial \mathbf{J}}{\omega - \mathbf{m} \cdot \boldsymbol{\Omega}} [\psi_m^{(p)}(\mathbf{J})]^* \psi_m^{(q)}(\mathbf{J}). \end{aligned} \quad (18)$$

We define the response matrix of the system $\hat{\mathbf{M}}$ as

$$\hat{\mathbf{M}}_{pq}(\omega) = (2\pi)^d \sum_m \int d\mathbf{J} \frac{\mathbf{m} \cdot \partial F_0/\partial \mathbf{J}}{\omega - \mathbf{m} \cdot \boldsymbol{\Omega}} [\psi_m^{(p)}(\mathbf{J})]^* \psi_m^{(q)}(\mathbf{J}), \quad (19)$$

where one must note that the response matrix depends only on the initial equilibrium state of the disc, since $\partial F_0/\partial \mathbf{J}$ evolves only on the secular scale, the perturbing and self-gravitating potentials are absent, and the basis elements $\psi^{(p)}$ are chosen once for all. Finally, in order to shorten the notations, the amplitudes of the self- and exterior potentials are defined as $\mathbf{a}(t) = (a_1(t), \dots, a_p(t), \dots)$ and $\mathbf{b}(t) = (b_1(t), \dots, b_p(t), \dots)$. Thanks to these notations, one can simplify equation (18), and rewrite it under the form

$$\hat{\mathbf{a}}(\omega) + \hat{\mathbf{b}}(\omega) = [\mathbf{I} - \hat{\mathbf{M}}(\omega)]^{-1} \cdot \hat{\mathbf{b}}(\omega). \quad (20)$$

One should note that the matrix $[\mathbf{I} - \hat{\mathbf{M}}]$ is invertible only if the self-gravitating system is linearly stable so that all the eigenvalues of $\hat{\mathbf{M}}$ are assumed to be strictly smaller than 1 for all values of ω .

2.5 Diffusion coefficients

The amplification relation (20) corresponds to the short time-scale (dynamical) response of the system, driven by the evolution equation (7). We will now describe the impact of these solutions on the long time-scale diffusion equation given by equation (6). Starting from equation (6), one has to evaluate an expression of the form

$$\begin{aligned} &\frac{1}{(2\pi)^d} \int d\boldsymbol{\theta} f(\mathbf{J}, \boldsymbol{\theta}, t) \frac{\partial [\psi^e + \psi^s]}{\partial \boldsymbol{\theta}} \\ &= \frac{1}{(2\pi)^d} \sum_{m_1, m_2} \int d\boldsymbol{\theta} f_{m_1} i\mathbf{m}_2 [\psi_{m_2}^e + \psi_{m_2}^s] e^{i(m_1+m_2)\cdot\boldsymbol{\theta}}. \end{aligned} \quad (21)$$

Here, only terms with $\mathbf{m}_1 = -\mathbf{m}_2$ are different from 0. Using equation (10) and the fact that $\psi_{-m} = \psi_m^*$, we can finally rewrite the diffusion equation (6) under the form

$$\frac{\partial F_0}{\partial t} = \sum_m \mathbf{m} \cdot \frac{\partial}{\partial \mathbf{J}} \left[D_m(\mathbf{J}) \mathbf{m} \cdot \frac{\partial F_0}{\partial \mathbf{J}} \right], \quad (22)$$

where the anisotropic diffusion coefficients $D_m(\mathbf{J})$ are given by

$$\begin{aligned} D_m(\mathbf{J}, t) &= [\psi_m^{e*}(\mathbf{J}, t) + \psi_m^{s*}(\mathbf{J}, t)] \\ &\times \int_{-\infty}^t d\tau e^{-im\cdot\Omega(t-\tau)} [\psi_m^e(\mathbf{J}, \tau) + \psi_m^s(\mathbf{J}, \tau)]. \end{aligned} \quad (23)$$

Note that equation (22) can be re-arranged as

$$\frac{\partial F_0}{\partial t} = \frac{\partial}{\partial \mathbf{J}} \cdot \left[\mathbf{D}(\mathbf{J}) \cdot \frac{\partial F_0}{\partial \mathbf{J}} \right], \text{ with } \mathbf{D}(\mathbf{J}) = \sum_m D_m(\mathbf{J}) \mathbf{m} \otimes \mathbf{m},$$

an anisotropic tensor diffusion matrix. Using the basis decomposition introduced in equation (15), the diffusion coefficients from

equation (23) take the form

$$D_m(\mathbf{J}, t) = \sum_{p,q} \psi_m^{(p)} \psi_m^{(q)*} \left[a_q^*(t) + b_q^*(t) \right] \times \int_{-\infty}^t d\tau e^{-im \cdot \Omega(t-\tau)} \left[a_p(\tau) + b_p(\tau) \right]. \quad (24)$$

Expressing the temporal coefficients $a_p(t)$ and $b_p(t)$ via their Fourier transforms, we obtain

$$D_m(\mathbf{J}, t) = \frac{1}{(2\pi)^2} \sum_{p,q} \psi_m^{(p)} \psi_m^{(q)*} \int d\omega \left[a_q^* + b_q^* \right](\omega) e^{i\omega t} \times \int_{-\infty}^t d\tau e^{-im \cdot \Omega(t-\tau)} \int d\omega' \left[a_p + b_p \right](\omega') e^{-i\omega' \tau}. \quad (25)$$

The amplification relation (20) allows us to rewrite equation (25) as

$$D_m(\mathbf{J}, t) = \frac{1}{(2\pi)^2} \sum_{p,q} \sum_{p_1, q_1} \psi_m^{(p)} \psi_m^{(q)*} \int d\omega e^{i\omega t} \left[\mathbf{I} - \widehat{\mathbf{M}}(\omega) \right]_{qq_1}^{-1} \widehat{b}_{q_1}^*(\omega) \times \int_{-\infty}^t d\tau e^{-im \cdot \Omega(t-\tau)} \int d\omega' e^{-i\omega' \tau} \left[\mathbf{I} - \widehat{\mathbf{M}}(\omega') \right]_{pp_1}^{-1} \widehat{b}_{p_1}(\omega'). \quad (26)$$

2.6 Statistical expectation

The final stage of the derivation is to introduce the statistics of the external perturbations. Indeed, our previous calculation corresponds to the response of the system to a *given* particular perturbation history: $t \mapsto \mathbf{b}(t)$. Let us now denote the ensemble average operation on such different realizations as $\langle \cdot \rangle$. As the global underlying background gravitational potential is assumed to be stationary, the mapping $(\mathbf{x}, \mathbf{v}) \mapsto (\boldsymbol{\theta}, \mathbf{J})$ remains the same for the different realizations, so that the operations of derivation or integration with respect to $\boldsymbol{\theta}$ and \mathbf{J} commute with the ensemble average. The diffusion equation (22), when ensemble averaged, takes the form

$$\frac{\partial \langle F_0 \rangle}{\partial t} = \sum_m \mathbf{m} \cdot \frac{\partial}{\partial \mathbf{J}} \left[\left\langle D_m(\mathbf{J}) \mathbf{m} \cdot \frac{\partial F_0}{\partial \mathbf{J}} \right\rangle \right]. \quad (27)$$

A priori, the gradient $\partial F_0 / \partial \mathbf{J}$, cannot be taken out of the ensemble average operation. However, we intend to describe the effect of an averaged fluctuation on a given F_0 representing a mean disc. Then, one may assume that the quasi-stationary distribution function F_0 , its gradients and therefore the response matrix $\widehat{\mathbf{M}}$ do not change significantly from one realization to another, so that they can be taken out of the ensemble average operation. This means that we assume there exists a mean response for the secular distribution, $F_0 = \langle F_0 \rangle$, de-correlated from the perturbations, so that we have $\langle D_m(\mathbf{J}) \mathbf{m} \cdot \partial F_0 / \partial \mathbf{J} \rangle = \langle D_m(\mathbf{J}) \rangle \mathbf{m} \cdot \partial F_0 / \partial \mathbf{J}$. We also suppose that the time evolution of the exterior perturbing potential is stationary and therefore introduce the corresponding temporal autocorrelation function defined as

$$\mathbf{C}_{kl}(t_1 - t_2) = \langle b_k(t_1) b_l^*(t_2) \rangle, \quad (28)$$

where the exterior perturbation is also assumed to be of zero mean. The autocorrelation \mathbf{C} connects the temporal coefficients \mathbf{b} , whereas the diffusion coefficients from equation (26) involve the Fourier transformed ones $\widehat{\mathbf{b}}$, so that one needs to compute $\langle \widehat{b}_k(\omega) \widehat{b}_l^*(\omega') \rangle$.

One can straightforwardly show that

$$\langle \widehat{b}_k(\omega) \widehat{b}_l^*(\omega') \rangle = 2\pi \delta_D(\omega - \omega') \widehat{\mathbf{C}}_{kl}(\omega), \quad (29)$$

where $\widehat{\mathbf{C}}$ is the temporal Fourier transform of the autocorrelation of the external potential. Using this result in the ensemble averaged expression (26) yields

$$\langle D_m(\mathbf{J}, t) \rangle = \frac{1}{2\pi} \sum_{p,q} \psi_m^{(p)} \psi_m^{(q)*} \int d\omega \int_{-\infty}^0 d\tau' e^{-i(\omega - m \cdot \Omega)\tau'} \times \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1} \cdot \widehat{\mathbf{C}} \cdot \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1} \Big|_{pq}(\omega), \quad (30)$$

where we performed the change of variables $\tau' = \tau - t$. One should note that when ensemble averaged, the diffusion coefficients are (explicitly) independent of t .² In order to shorten temporarily the notations, we introduce $\widehat{\mathbf{L}} = \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1} \cdot \widehat{\mathbf{C}} \cdot \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1}$. In equation (30), one has to evaluate an expression of the form

$$\frac{1}{2\pi} \int_{-\infty}^{+\infty} d\omega \widehat{\mathbf{L}}(\omega) \int_{-\infty}^0 d\tau' e^{-i(\omega - m \cdot \Omega)\tau'} = \frac{i}{2\pi} \int_{-\infty}^{+\infty} d\omega \frac{\widehat{\mathbf{L}}(\omega)}{\omega - m \cdot \Omega}, \quad (31)$$

where in the integration over τ' we only kept the term for $\tau' = 0$, by adding an imaginary part to the frequency ω so that $\omega = \omega + i0^+$, ensuring the convergence for $\tau' \rightarrow -\infty$. The remaining integral over ω can be evaluated using Plemelj formula

$$\frac{1}{x \pm i0^+} = \mathcal{P} \left(\frac{1}{x} \right) \mp i\pi \delta_D(x), \quad (32)$$

where \mathcal{P} denotes Cauchy principal value. Therefore, equation (31) becomes

$$(31) \propto \frac{i}{2\pi} \mathcal{P} \int_{-\infty}^{+\infty} d\omega \frac{\widehat{\mathbf{L}}(\omega)}{\omega - m \cdot \Omega} + \frac{1}{2} \widehat{\mathbf{L}}(m \cdot \Omega). \quad (33)$$

The final step of the derivation is to show that the principal value term present in equation (33) has no impact on the secular diffusion. Indeed, using the expression (23) of the diffusion coefficients, one can show that they satisfy $D_{-m}(\mathbf{J}) = D_m^*(\mathbf{J})$. As a consequence, as we are summing on all the modes \mathbf{m} , the diffusion equation (22) may be rewritten under the form

$$\frac{\partial F_0}{\partial t} = \sum_m \mathbf{m} \cdot \frac{\partial}{\partial \mathbf{J}} \left[\text{Re} [D_m(\mathbf{J})] \mathbf{m} \cdot \frac{\partial F_0}{\partial \mathbf{J}} \right]. \quad (34)$$

From equations (19) and (28), we know that the response matrix and the autocorrelation matrix are hermitian so that $\widehat{\mathbf{M}}^* = \widehat{\mathbf{M}}^t$ and $\widehat{\mathbf{C}}^* = \widehat{\mathbf{C}}^t$. As a consequence, the matrix $\widehat{\mathbf{L}}$ is also hermitian. Since $\text{Re}(D_m) = (D_m + D_m^*)/2$, starting from equation (33), we immediately obtain

$$\langle \text{Re} [D_m(\mathbf{J})] \rangle = \frac{1}{2} \sum_{p,q} \psi_m^{(p)} \psi_m^{(q)*} \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1} \cdot \widehat{\mathbf{C}} \cdot \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1} \Big|_{pq}(\mathbf{m} \cdot \Omega), \quad (35)$$

so that the full secular diffusion equation takes the form

$$\frac{\partial F_0}{\partial t} = \sum_m \mathbf{m} \cdot \frac{\partial}{\partial \mathbf{J}} \left[\mathbf{m} \cdot \frac{\partial F_0}{\partial \mathbf{J}} \sum_{p,q} \frac{1}{2} \psi_m^{(p)}(\mathbf{J}) \psi_m^{(q)*}(\mathbf{J}) \times \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1} \cdot \widehat{\mathbf{C}} \cdot \left[\mathbf{I} - \widehat{\mathbf{M}} \right]^{-1} \Big|_{pq}(\mathbf{m} \cdot \Omega) \right]. \quad (36)$$

² Though they depend on the secular time-scale via the variation of F in $\widehat{\mathbf{M}}$.

Equation (36) is the main result of this section. In Appendix A, we present an alternative derivation of these diffusion coefficients based on Hamilton's equations from which we recover the exact same diffusion equation. The 1/2 factor recovered via these two complementary approaches was skipped in the calculation presented in Pichon & Aubert (2006), because of an error in the bounds of half-temporal integrations, similar to the one present in equation (31). This derivation is valid in any dimensions, provided the underlying system is integrable. One may also note that in the homogeneous limit, equation (36) reduces to the secular diffusion equation obtained in Chavanis (2012a, 2013). In the next section, we will restrict ourselves to 2D configurations and make further assumptions in order to simplify equation (35) into a one-dimensional quadrature.

3 THIN TEPID DISCS AND WKB LIMIT

One difficulty for the implementation of the secular diffusion equation (36) is to simultaneously have an explicit mapping $(\mathbf{x}, \mathbf{v}) \mapsto (\boldsymbol{\theta}, \mathbf{J})$ to the angle-actions coordinates, and be able to evaluate the diffusion coefficients given by equation (35), which require to invert the response matrix $[\mathbf{I} - \hat{\mathbf{M}}]$. In order to deal with the *non-locality* of Poisson's equation, we also have to explicitly introduce potential basis elements, $\psi^{(p)}$, as in equation (14), to compute the response matrix from equation (19). To ease these calculations in a 2D axisymmetric disc, one may rely on the WKB assumption (Liouville 1837; Toomre 1964; Kalnajs 1965; Lin & Shu 1966; Palmer, Papaloizou & Allen 1989), which assumes that the perturbations and self-responses will take the form of tightly wound spirals, which in turn allows us to write Poisson's equation *locally*. Considering only such perturbations sums up to introducing basis elements with specific properties as detailed later on.

3.1 Epicyclic approximation and isothermal DF

In order to explicitly build up a mapping $(\mathbf{x}, \mathbf{v}) \mapsto (\boldsymbol{\theta}, \mathbf{J})$ for an axisymmetric disc, we assume that the disc is sufficiently cold and therefore rely on the so-called epicyclic approximation.

The natural coordinates for an axisymmetric galactic disc are the polar coordinates (R, ϕ) , with their associated momenta (p_R, p_ϕ) . Within such coordinates, the stationary Hamiltonian of the system reads

$$H_0(R, \phi, p_R, p_\phi) = \frac{1}{2} \left[p_R^2 + \frac{p_\phi^2}{R^2} \right] + \psi_0(R), \quad (37)$$

where ψ_0 is the axisymmetric stationary background potential within the disc. The first action of the system is the angular momentum J_ϕ defined as

$$J_\phi = L_z \equiv \frac{1}{2\pi} \oint d\phi p_\phi = p_\phi = R^2 \dot{\phi}. \quad (38)$$

As soon as the value of J_ϕ is imposed, one obtains a new equation of motion for the R variable given by

$$\ddot{R} = -\frac{\partial \psi_{\text{eff}}}{\partial R}, \quad (39)$$

where the effective potential is defined as

$$\psi_{\text{eff}}(R) = \psi_0(R) + \frac{J_\phi^2}{2R^2}. \quad (40)$$

The main idea behind the epicyclic approximation is to approximate the radial motion as an harmonic oscillation. For a given value of J_ϕ , we define implicitly the guiding radius R_g as

$$0 = \left. \frac{\partial \psi_{\text{eff}}}{\partial R} \right|_{R_g} = \left. \frac{\partial \psi_0}{\partial R} \right|_{R_g} - \frac{J_\phi^2}{R_g^3}, \quad (41)$$

so that $R_g(J_\phi)$ corresponds to the radius for which stars with an angular momentum equal to J_ϕ evolve on circular orbits. For a stationary potential, the mapping between R_g and J_ϕ is bijective and unambiguous (up to the sign of J_ϕ). We define the azimuthal frequency $\Omega(R_g)$ and the epicyclic frequency $\kappa(R_g)$ as

$$\begin{cases} \Omega^2(R_g) = \left. \frac{1}{R_g} \frac{\partial \psi_0}{\partial R} \right|_{R_g} = \frac{J_\phi^2}{R_g^4}, \\ \kappa^2(R_g) = \left. \frac{\partial^2 \psi_{\text{eff}}}{\partial R^2} \right|_{R_g} = \left. \frac{\partial^2 \psi_0}{\partial R^2} \right|_{R_g} + 3 \frac{J_\phi^2}{R_g^4}. \end{cases} \quad (42)$$

A Taylor expansion at first order near R_g of equation (39) shows that R satisfies the differential equation $\ddot{R} = -\kappa^2(R - R_g)$, which is the evolution equation of an harmonic oscillator centred on R_g . We introduce the amplitude A of the radial oscillations and define the radial action J_r as

$$J_r \equiv \frac{1}{2\pi} \oint dR p_R = \frac{1}{2} \kappa A^2. \quad (43)$$

The case $J_r = 0$ corresponds to circular orbits. The larger J_r , the wider the radial oscillations of the star. One should note that within the epicyclic approximation, the two intrinsic frequencies Ω and κ are only function of the angular momentum J_ϕ and do not depend on the radial action J_r . Finally, one can show (Lynden-Bell & Kalnajs 1972; Palmer 1994; Binney & Tremaine 2008) that the mapping between (R, ϕ, p_R, p_ϕ) and $(\theta_R, \theta_\phi, J_r, J_\phi)$ takes at first order the form

$$\begin{cases} R = R_g + A \cos(\theta_R), \\ \phi = \theta_\phi - \frac{2\Omega}{\kappa} \frac{A}{R_g} \sin(\theta_R). \end{cases} \quad (44)$$

Within this approximation, one can easily parametrize plausible stationary distribution functions for a galactic disc, defined as functions of the actions (J_ϕ, J_r) . Indeed, we suppose that the stationary distribution F_0 of the disc is a Schwarzschild distribution function (or locally *isothermal*) given by

$$F_0(R_g, J_r) = \frac{\Sigma(R_g) \Sigma(R_g)}{\pi \kappa(R_g) \sigma_r^2(R_g)} \exp \left[-\frac{\kappa(R_g) J_r}{\sigma_r^2(R_g)} \right], \quad (45)$$

where $\Sigma(R_g)$ is the surface density of the disc and $\sigma_r^2(R_g)$, which depends on the position in the disc, encodes the typical radial velocity dispersion of the stars at a given radius. The larger σ_r^2 , the hotter the disc and the more stable it is.

3.2 The WKB basis

In order to use a WKB approach in the secular diffusion equation (22), one needs to introduce explicitly a basis of density potentials from which the WKB hypothesis will follow.

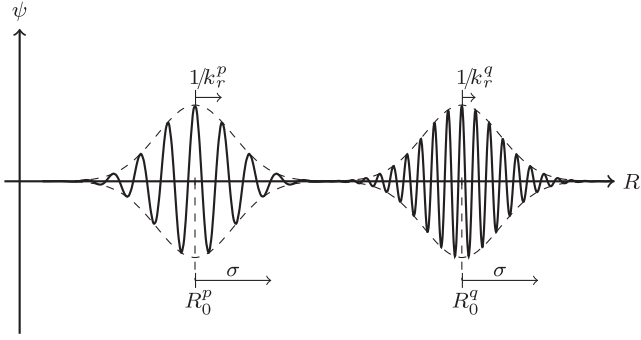


Figure 1. Two WKB basis elements. Each Gaussian is centred around a radius R_0 . The typical extension of the Gaussian is given by the decoupling scale σ , and they are modulated at the radial frequency k_r .

3.2.1 Definition of the basis elements

We define on the plane the following basis of potential functions, well suited to represent tightly wound spirals

$$\psi^{[k_\phi, k_r, R_0]}(R, \phi) = \mathcal{A} e^{i(k_\phi \phi + k_r R)} \mathcal{B}_{R_0}(R), \quad (46)$$

where the functions $\mathcal{B}_{R_0}(R)$ are of the form

$$\mathcal{B}_{R_0}(R) = \frac{1}{(\pi\sigma^2)^{1/4}} \exp\left[-\frac{(R - R_0)^2}{2\sigma^2}\right]. \quad (47)$$

To our knowledge, this is the first time that such tightly wound basis elements have been introduced in the context of discs secular dynamics. The central radius R_0 is the radius around which the Gaussian \mathcal{B}_{R_0} is centred, k_ϕ is an azimuthal number representing the angular component of the basis elements, and k_r corresponds to the radial frequency of the potential element. Here, σ is a scale-separation parameter ensuring the biorthogonality of the basis elements (see below). The radial dependence of the basis elements is illustrated in Fig. 1. The amplitude \mathcal{A} is not yet defined and will be chosen in order to guarantee the correct normalization of the basis. The unusual normalization of \mathcal{B}_{R_0} will ensure that \mathcal{A} is independent of σ and will also allow us to naturally introduce Dirac deltas $\delta_D(R - R_0)$ in some of the next calculations.

3.2.2 Associated surface density elements

Equation (14) requires to have a biorthogonal potential basis. We now determine the surface density basis elements associated with the potentials from equation (46). For 2D razor thin discs, we are in the presence of a discontinuity leading to a surface density $\Sigma(R, \phi)$. In order to satisfy Poisson's equation, we extend our potential to the z -axis via the following Ansatz:

$$\psi^{[k_\phi, k_r, R_0]}(R, \phi, z) = \mathcal{A} e^{i(k_\phi \phi + k_r R)} \mathcal{B}_{R_0}(R) Z(z). \quad (48)$$

Injecting such an expression into Poisson's equation in vacuum $\Delta \psi^{[k_\phi, k_r, R_0]} = 0$, we obtain after some algebra

$$-\frac{Z''}{Z} = -k_r^2 \left[1 - \frac{i}{k_r R} + 2i \frac{R - R_0}{\sigma^2} \frac{1}{k_r} + \frac{R - R_0}{R} \frac{1}{(\sigma k_r)^2} + \frac{1}{(\sigma k_r)^2} + \frac{k_\phi^2}{(k_r R)^2} - \left[\frac{R - R_0}{\sigma^2} \frac{1}{k_r} \right]^2 \right]. \quad (49)$$

In order to obtain a simple expression for the surface density basis elements, we introduce additional WKB-like assumptions, so that the terms appearing in equation (49) are all negligible in front of 1.

First of all, we assume that the spirals are tightly wound so that we have

$$k_r R \gg 1. \quad (50)$$

Moreover, introducing the typical size R_{sys} of the system, we add the supplementary constraint

$$k_r \sigma \gg \frac{R_{\text{sys}}}{\sigma}. \quad (51)$$

In this limit, assuming that k_ϕ remains of the order of the unity, equation (49) becomes

$$\frac{Z''}{Z} = k_r^2. \quad (52)$$

Therefore, we conclude that within the WKB approximation, the extended 3D potential can be written as

$$\psi^{[k_\phi, k_r, R_0]}(R, \phi, z) = \psi^{[k_\phi, k_r, R_0]}(R, \phi) e^{-k_r |z|}, \quad (53)$$

where we added absolute value on z in order to respect the boundaries conditions of the potential at $z = \pm\infty$, where the potential has to tend to 0. Such a potential introduces a discontinuity of $\partial\psi/\partial z$ at the plane $z = 0$, consistent with the given surface density. Gauss theorem for the discontinuities at a plane may be written as

$$\Sigma(R, \phi) = \frac{1}{4\pi G} \left[\lim_{z \rightarrow 0^+} \frac{\partial\psi}{\partial z} - \lim_{z \rightarrow 0^-} \frac{\partial\psi}{\partial z} \right]. \quad (54)$$

We immediately conclude that the surface density associated with a given potential element $\psi^{[k_\phi, k_r, R_0]}$ is given by

$$\Sigma^{[k_\phi, k_r, R_0]}(R, \phi) = -\frac{|k_r|}{2\pi G} \psi^{[k_\phi, k_r, R_0]}(R, \phi). \quad (55)$$

3.2.3 Biorthogonality condition

The next step of the definition of the WKB basis is to ensure that this basis is biorthogonal as in equation (14). Indeed, it has to satisfy the property

$$\delta_{k_\phi^p}^{k_\phi^q} \delta_{k_r^p}^{k_r^q} \delta_{R_0^p}^{R_0^q} = - \int dR R d\phi \left[\psi^{[k_\phi^p, k_r^p, R_0^p]}(R, \phi) \right]^* \Sigma^{[k_\phi^q, k_r^q, R_0^q]}(R, \phi).$$

The r.h.s. of this expression becomes

$$\frac{|k_r^q|}{2\pi G} \mathcal{A}_p \mathcal{A}_q \frac{1}{\sqrt{\pi\sigma^2}} \int d\phi e^{i(k_\phi^q - k_\phi^p)\phi} \times \int dR R e^{i(k_r^q - k_r^p)R} \exp\left[-\frac{(R - R_0^p)^2}{2\sigma^2}\right] \exp\left[-\frac{(R - R_0^q)^2}{2\sigma^2}\right]. \quad (56)$$

The integration on ϕ is straightforward and gives a term equal to $2\pi \delta_{k_\phi^p}^{k_\phi^q}$. In order to be able to integrate on R , we now need to impose new WKB-like assumptions to justify the biorthogonality of the basis. We introduce the spatial Fourier transform \mathcal{F} with respect to R , and the difference of the two radial frequencies $\Delta k_r = k_r^q - k_r^p$. Equation (56) requires to integrate an expression of the form

$$\int dR e^{i\Delta k_r R} \mathcal{B}_{R_0^p}(R) \mathcal{B}_{R_0^q}(R) = \int dk' \mathcal{F}[\mathcal{B}_{R_0^p}](k') \mathcal{F}[\mathcal{B}_{R_0^q}](\Delta k_r - k') \propto \int dk' \exp\left[-\frac{k'^2}{2/\sigma^2}\right] \exp\left[-\frac{(\Delta k_r - k')^2}{2/\sigma^2}\right] e^{-i(R_0^p k' + R_0^q [\Delta k_r - k'])}, \quad (57)$$

where we used the property that the Fourier transform of a product is given by the convolution of the Fourier transforms and that the Fourier transform of a Gaussian of spread σ is a Gaussian of spread $1/\sigma$. In expression (57), we note that if $|k'| \gg 1/\sigma$ or $|\Delta k_r - k'| \gg 1/\sigma$, then the product of the two terms can be considered to be negligible. We will therefore suppose that one of the conditions

$$\Delta k_r \gg \frac{1}{\sigma} \quad \text{or} \quad \Delta k_r = 0, \quad (58)$$

holds. Under this assumption, the term is non-zero only for $k_r^p = k_r^q$. Finally, it remains to prove that non-zero terms are only obtained when $R_0^p = R_0^q$. The peaks of the two Gaussians in equation (56) can be considered as sharp and separated if $\Delta R_0 = R_0^p - R_0^q$ satisfies the condition

$$\Delta R_0 \gg \sigma \quad \text{or} \quad \Delta R_0 = 0. \quad (59)$$

To sum up our assumptions so far, we should consider peak-radius R_0 , spread σ and radial frequencies k_r such that

$$\Delta R_0 \gg \sigma \gg \frac{1}{\Delta k_r}. \quad (60)$$

To these conditions, one must also add the constraints obtained in equations (50) and (51) via Poisson's equation. With these assumptions, we can ensure that we must necessarily have $k_\phi^p = k_\phi^q$, $k_r^p = k_r^q$ and $R_0^p = R_0^q$, in order to have a non-zero term. The last step is to explicitly calculate the amplitude \mathcal{A} of the basis elements. Indeed, starting from equation (56), we have the condition

$$-\mathcal{A}^2 \frac{|k_r|}{G} \frac{1}{\sqrt{\pi}\sigma} \int dR R \exp\left[-\frac{(R-R_0)^2}{\sigma^2}\right] = -1, \quad (61)$$

which may be rewritten as

$$1 = \mathcal{A}^2 \frac{|k_r|}{G} \frac{R_0}{2} \left[1 + \operatorname{erf}\left[\frac{R_0}{\sigma}\right] + \frac{1}{\sqrt{\pi}} \frac{\sigma}{R_0} e^{-R_0^2/\sigma^2}\right]. \quad (62)$$

Using the assumptions made in equation (60), we immediately conclude that $R_0/\sigma \gg 1$, so that $\operatorname{erf}[R_0/\sigma] \sim 1$ and $\sigma/(\sqrt{\pi}R_0) \exp[-R_0^2/\sigma^2] \ll 1$. We therefore finally obtain the expression of the amplitude of the basis potentials as

$$\mathcal{A} = \sqrt{\frac{G}{|k_r| R_0}}. \quad (63)$$

3.2.4 Fourier development in angles

The diffusion equation involves terms of the form $\psi_m^{(p)}(\mathbf{J})$, that we will now evaluate for the WKB basis, equation (46). Using the epicyclic mapping from equation (44), we need to estimate

$$\begin{aligned} \psi_m^{[k_r, k_r, R_0]}(\mathbf{J}) &= \frac{1}{(2\pi)^2} \int d\theta_\phi \int d\theta_R e^{-im_\phi\theta_\phi} e^{-im_r\theta_R} e^{ik_\phi\theta_\phi} e^{ik_r R_g} \\ &\times e^{i[k_r A \cos(\theta_R) - k_\phi \frac{2\Omega}{\kappa} \frac{A}{R_g} \sin(\theta_R)]} \mathcal{A} \mathcal{B}_{R_0}(R_g + A \cos(\theta_R)). \end{aligned} \quad (64)$$

The integration on θ_ϕ is straightforward and is equal to $2\pi\delta_{k_\phi}^{m_\phi}$. Looking at the dependence in θ_R within the complex exponential, we can write

$$k_r A \cos(\theta_R) - k_\phi \frac{2\Omega}{\kappa} \frac{A}{R_g} \sin(\theta_R) = H_{k_\phi}(k_r) \sin(\theta_R + \theta_R^0),$$

where we have defined

$$H_{k_\phi}(k_r) = A \sqrt{k_r^2 + k_\phi^2 \left[\frac{2\Omega}{\kappa R_g}\right]^2}; \quad \theta_R^0 = \tan^{-1} \left[-\frac{\kappa}{2\Omega} \frac{k_r R_g}{k_\phi} \right]. \quad (65)$$

Thanks to our WKB assumptions, an approximation of the amplitude term $H_{k_\phi}(k_r)$ and the phase-shift θ_R^0 is possible. Indeed, both of these terms involve an expression of the form $2k_\phi \times \Omega/\kappa \times 1/(k_r R_g)$. Yet, we made the assumption that $k_r R_g \gg 1$. Moreover, we know that for typical galaxies $1/2 \leq \Omega/\kappa \leq 1$ (Binney & Tremaine 2008). Assuming that k_ϕ is of the order of unity, we obtain the approximations

$$H_{k_\phi}(k_r) \simeq A |k_r| \simeq \sqrt{\frac{2J_r}{\kappa}} |k_r|; \quad \theta_R^0 \simeq -\frac{\pi}{2}. \quad (66)$$

We also supposed that the radial oscillations are small, so that the epicyclic amplitude satisfies $A \ll R_g$. Thanks to this assumption, we may replace $\mathcal{B}_{R_0}(R_g + A \cos(\theta_R))$ by $\mathcal{B}_{R_0}(R_g)$, keeping the dependence on A only in the complex exponential. This is a crucial step to be able to integrate explicitly on θ_R . We also introduce the Bessel functions of the first kind \mathcal{J}_ℓ which satisfy the property

$$e^{iH_{k_\phi}(k_r) \sin(\theta_R + \theta_R^0)} = \sum_{\ell \in \mathbb{Z}} \mathcal{J}_\ell[H_{k_\phi}(k_r)] e^{i\ell(\theta_R + \theta_R^0)}. \quad (67)$$

It is then possible to perform explicitly the integration on θ_R in equation (64), which is equal to $2\pi \delta_\ell^{m_r}$, so that only one Bessel function remains. We finally obtain the expression of the Fourier transform in angles of the basis elements

$$\psi_m^{[k_\phi, k_r, R_0]}(\mathbf{J}) = \delta_{m_\phi}^{k_\phi} e^{ik_r R_g} e^{im_r \theta_R^0} \mathcal{A} \mathcal{J}_{m_r}[H_{k_\phi}(k_r)] \mathcal{B}_{R_0}(R_g). \quad (68)$$

3.3 Estimation of the response matrix

Using the explicit WKB potential basis introduced in equation (46), one can now estimate the matrix response from expression (19). The approximation obtained in equation (66) allows us to simplify the phase-shift terms, so that equation (19) becomes

$$\begin{aligned} \hat{\mathbf{M}}_{[k_\phi^p, k_r^p, R_0^p], [k_\phi^q, k_r^q, R_0^q]}(\omega) &= (2\pi)^2 \sum_m \int d^2 \mathbf{J} \frac{\mathbf{m} \cdot \partial F_0 / \partial \mathbf{J}}{\omega - \mathbf{m} \cdot \boldsymbol{\Omega}} \delta_{m_\phi}^{k_\phi^p} \delta_{m_\phi}^{k_\phi^q} e^{iR_g[k_r^q - k_r^p]} \mathcal{A}_p \mathcal{A}_q \\ &\times \mathcal{J}_{m_r}(H_{k_\phi}(k_r^p)) \mathcal{J}_{m_r}(H_{k_\phi}(k_r^q)) \mathcal{B}_{R_0^p}(R_g) \mathcal{B}_{R_0^q}(R_g). \end{aligned} \quad (69)$$

One should note that this expression is similar to equation (56). Indeed, using the assumptions from equation (60), we are able to ensure that only the diagonal coefficients of the response matrix are different from 0. First of all, the azimuthal Kronecker symbols impose that $k_\phi^p = k_\phi^q$. There is however a slight complication in the calculation because of the presence of additional terms depending on R_g . In order to sketch the proof of this statement, we introduce the function

$$h(R_g) = \left| \frac{dJ_\phi}{dR_g} \right| \frac{\mathbf{m} \cdot \partial F_0 / \partial \mathbf{J}}{\omega - \mathbf{m} \cdot \boldsymbol{\Omega}} \mathcal{A}_p \mathcal{A}_q \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^p \right] \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^q \right]. \quad (70)$$

This function captures all the additional R_g dependence appearing in equation (69). Using the change of variables $J_\phi \mapsto R_g$, the integral on J_ϕ which has to be evaluated in equation (69) takes the form

$$\int dR_g h(R_g) e^{iR_g[k_r^q - k_r^p]} \exp \left[-\frac{(R_g - R_0^p)^2}{2\sigma^2} - \frac{(R_g - R_0^q)^2}{2\sigma^2} \right]. \quad (71)$$

Using the assumption from equation (59) relative to the possible values of ΔR_0 in the WKB basis, one can note that the product

of the two Gaussians in R_g imposes $R_0^p = R_0^q$ in order to have a non-zero contribution. The previous expression then becomes

$$(71) \propto \int dR_g h(R_g) e^{iR_g[k_r^q - k_r^p]} \exp\left[-\frac{(R_g - R_0)^2}{\sigma^2}\right]. \quad (72)$$

Using the same argument as in equation (57), we can rewrite the Fourier transform in R as the convolution of two radial Fourier transforms so that it becomes

$$(71) \propto \int dk' \mathcal{F}[h](k') \exp\left[-\frac{(\Delta k_r - k')^2}{4/\sigma^2}\right]. \quad (73)$$

We now use equation (58) relative to the possible values of Δk_r in the WKB basis. If we suppose that $\Delta k_r \neq 0$, the width of the Gaussian from equation (73) imposes that the integration will only probe the contribution of $\mathcal{F}[h]$ in the neighbourhood of $k' \sim \Delta k_r \gg 1/\sigma$. We assume that the radial Fourier transform of the function h is such that it is mainly *focused* in the frequency region $|k| \lesssim 1/\sigma$, meaning that for a typical galactic disc, the main frequencies of radial variations of h are inferior to $1/\sigma$. Under this assumption of slow radial variation within the disc, one can see that non-zero contributions can only be obtained for $\Delta k_r = k_r^p - k_r^q = 0$. As a consequence, we have shown that within the WKB approximation the response matrix is diagonal. For these diagonal coefficients, it only remains to evaluate explicitly the integrals over J_ϕ and J_r in order to obtain the expression of the response matrix eigenvalues. This calculation is presented in Appendix B. Within the assumption that the galactic disc is tepid, the eigenvalues of the response matrix take the form

$$\widehat{\mathbf{M}}_{\left[k_\phi^p, k_r^p, R_0\right], \left[k_\phi^q, k_r^q, R_0\right]}(\omega) = \delta_{k_r^p, k_r^q}^{\delta_{k_\phi^p, k_\phi^q}} \frac{2\pi G \Sigma |k_r^p|}{\kappa^2 (1 - s^2)} \mathcal{F}(s, \chi), \quad (74)$$

where χ and s are, respectively, defined in equations (B6) and (B10) and $F(s, \chi)$ is the reduction factor introduced in equation (B11). This amplification eigenvalue is in full agreement with the seminal works from Kalnajs (1965) and Lin & Shu (1966), which independently derived the WKB dispersion relation for stellar discs.³

3.4 Estimation of the diffusion coefficients

The expression (35) of the diffusion coefficients shows that the diffusion coefficients require the evaluation of $[\mathbf{I} - \widehat{\mathbf{M}}]^{-1}$. In order to simplify the notations, we will denote our potential basis with only one index so that

$$\psi^{(p)} = \psi^{[k_\phi^p, k_r^p, R_0^p]}. \quad (75)$$

Equation (74) shows that the response matrix is diagonal in the WKB approximation. We therefore introduce the eigenvalues of $\widehat{\mathbf{M}}$ as

$$\lambda_p \equiv \widehat{\mathbf{M}}_{pp}. \quad (76)$$

The matrix $[\mathbf{I} - \widehat{\mathbf{M}}]^{-1}$ is then diagonal and reads

$$[\mathbf{I} - \widehat{\mathbf{M}}]_{pq}^{-1} = \delta_p^q \frac{1}{1 - \lambda_p}. \quad (77)$$

Thanks to these diagonal coefficients, the expression of the diffusion coefficients from equation (35) becomes

$$D_m(\mathbf{J}) = \frac{1}{2} \sum_{p,q} \psi_m^{(p)} \psi_m^{(q)*} \frac{1}{1 - \lambda_p} \frac{1}{1 - \lambda_q} \widehat{\mathbf{C}}_{pq}(\mathbf{m} \cdot \boldsymbol{\Omega}). \quad (78)$$

³ For nice introductions to the WKB dispersion relation in stellar discs, see section 6.2.2 of Binney & Tremaine (2008) and section 1.4.2 of Binney (2013).

At this stage, we use the property from equation (29) to rewrite $\widehat{\mathbf{C}}_{pq}$ as a function of the basis coefficients \widehat{b}_p and \widehat{b}_q^* . Remembering that the basis elements $\psi_m^{(p)}$ and the matrix eigenvalues λ_p do not change from one realization to another, one can rewrite equation (78) under the form

$$D_m(\mathbf{J}) = \left\langle \frac{1}{2\pi} \int d\omega' \frac{1}{2} \sum_{p,q} \psi_m^{(p)}(\mathbf{J}) \psi_m^{(q)*}(\mathbf{J}) \times \frac{1}{1 - \lambda_p} \frac{1}{1 - \lambda_q} \widehat{b}_p(\mathbf{m} \cdot \boldsymbol{\Omega}) \widehat{b}_q^*(\omega') \right\rangle. \quad (79)$$

It is important here to note that the eigenvalues λ_p, λ_q and the basis coefficient \widehat{b}_p are both evaluated at the intrinsic frequency $\mathbf{m} \cdot \boldsymbol{\Omega}$, whereas \widehat{b}_q^* has to be evaluated at the *integrated* frequency ω' . In the upcoming calculations, in order to shorten the notations, when obvious, the frequencies of evaluation will not be written. Using the expressions of the basis elements in the WKB approximation from equation (68), we can write

$$D_m(\mathbf{J}) = \left\langle \frac{1}{2\pi} \int d\omega' \sum_{k_r^p, k_r^q, R_0^p, R_0^q} \frac{1}{2} \frac{G}{\sqrt{R_0^p R_0^q}} \frac{1}{\sqrt{|k_r^p k_r^q|}} \times \mathcal{J}_{m_r} \left[\frac{\sqrt{2J_r}}{\kappa} k_r^p \right] \mathcal{J}_{m_r} \left[\frac{\sqrt{2J_r}}{\kappa} k_r^q \right] e^{iR_g(k_r^p - k_r^q)} \frac{1}{1 - \lambda_p} \frac{1}{1 - \lambda_q} \times \frac{1}{\sqrt{\pi\sigma^2}} \exp\left[-\frac{(R_g - R_0)^2}{2\sigma^2}\right] \exp\left[-\frac{(R_g - R_0^q)^2}{2\sigma^2}\right] \widehat{b}_p \widehat{b}_q^* \right\rangle, \quad (80)$$

where we already got rid of the sum over k_ϕ^p and k_ϕ^q , since the Fourier transform of the WKB basis elements from equation (68) imposes to have

$$m_\phi = k_\phi^p = k_\phi^q. \quad (81)$$

In the expression (80), we also neglected the phase terms in $m_r \theta_0^R$ and simplified the value at which the Bessel functions have to be evaluated using the approximation introduced in equation (66).

In order to obtain an expression independent from the choice of the basis (i.e. the precise value of σ), we will now replace the coefficients \widehat{b}_p by their expressions in terms of the *true* exterior potential function ψ^e , which is completely independent of the choice of the basis. As the potential basis in the WKB approximation is bi-orthogonal, the temporal Fourier transform of the basis coefficients is given by

$$\widehat{b}_p(\omega) = - \int d^2\mathbf{x} [\Sigma^{(p)}(\mathbf{x})]^* \widehat{\psi}^e(\mathbf{x}, \omega), \quad (82)$$

where the hat $\widehat{}$ corresponds to the temporal Fourier transform defined in equation (11). Using the expression of the surface density basis from equation (55), we obtain

$$\widehat{b}_p = \int dR R \int d\phi \frac{|k_r^p|}{2\pi G} \mathcal{A} e^{-i|k_r^p R + k_\phi^p \phi} \mathcal{B}_{R_0^p}(R) \widehat{\psi}^e(R, \phi).$$

The integration on ϕ is straightforward and leads to a term equal to $2\pi \widehat{\psi}^e_{k_\phi^p}(R)$, where the presence of the index k_ϕ^p corresponds to the Fourier transform with respect to the physical angle ϕ , using the same conventions as in equation (8). We may now write

$$\widehat{b}_p = \sqrt{\frac{|k_r^p|}{G R_0^p}} \frac{1}{(\pi\sigma^2)^{1/4}} \int dR R \exp\left[-\frac{(R - R_0^p)^2}{2\sigma^2}\right] e^{-iR k_r^p} \widehat{\psi}^e_{k_\phi^p}(R). \quad (83)$$

This integration should be interpreted as the radial Fourier transform at the frequency k_r^p of the exterior potential in the region close to R_0^p . Since the integrand contains a Gaussian in R of spread σ , we may take the term in R out of the integral and consider it to be equal to R_0^p . We now define the local Fourier transform of the exterior potential on a restricted region of radius (Gabor 1946) as

$$\widehat{\psi}_{k_\phi, k_r}^e[R_0] = \frac{1}{2\pi} \int dR \widehat{\psi}_{k_\phi}^e(R) \exp\left[-\frac{(R-R_0)^2}{2\sigma^2}\right] e^{-i(R-R_0)k_r}. \quad (84)$$

This definition is motivated by the fact that if we consider the case of an uniform perturbation $\psi^e = 1$, then its local Fourier transform is independent of R_0 . One may note that this definition is not independent of the decoupling scale σ , but as we will see later on, it is the relevant quantity in order to obtain diffusion coefficients independent of this *ad hoc* parameter. Thanks to this definition, the basis coefficients from equation (83) become

$$\widehat{b}_p = \sqrt{\frac{|k_r^p| R_0^p}{G}} \frac{2\pi}{(\pi\sigma^2)^{1/4}} e^{-iR_0^p k_r^p} \widehat{\psi}_{k_\phi, k_r^p}^e[R_0^p]. \quad (85)$$

We recall the notation used for the exterior potential in the previous expression. The index k_ϕ^p corresponds to the azimuthal Fourier transform with respect to the physical angle ϕ , and the index k_r^p corresponds to the *local* radial Fourier transform with respect to the physical radius R in the neighbourhood of R_0^p , as defined in equation (84). The diffusion coefficients from equation (80) are then given by

$$\begin{aligned} D_m(\mathbf{J}) = & \left\langle \frac{1}{2\pi} \int d\omega' \sum_{k_r^p, k_r^q, R_0^p, R_0^q} \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^p \right] \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^q \right] \right. \\ & \times e^{i(R_g - R_0^p)k_r^p} e^{-i(R_g - R_0^q)k_r^q} \frac{1}{1 - \lambda_p} \frac{1}{1 - \lambda_q} \\ & \times \frac{2\pi}{\sigma^2} \exp\left[-\frac{(R_g - R_0^p)^2}{2\sigma^2}\right] \exp\left[-\frac{(R_g - R_0^q)^2}{2\sigma^2}\right] \\ & \left. \times \widehat{\psi}_{m_\phi, k_r^p}^e[R_0^p] \widehat{\psi}_{m_\phi, k_r^q}^{e*}[R_0^q] \right\rangle. \quad (86) \end{aligned}$$

One can note that the gravitational constant G has disappeared, since the dependence on the strength of the gravity is now *hidden* in the units of ψ^e . One should also recall that the previous expression has to be evaluated at the resonance frequency, so that $\omega = \mathbf{m} \cdot \boldsymbol{\Omega}$, except for $\widehat{\psi}_{m_\phi, k_r^q}^{e*}[R_0^q]$ which has to be evaluated at the frequency ω' . The main step of the simplification is now to replace the discrete sums on the basis index k_r^p , k_r^q , R_0^p and R_0^q by continuous integrals. One should indeed now recall that our potential basis elements are made of three different index. Here, k_ϕ is a discrete index which must necessarily be equal to m_ϕ , so that it is absent from the sums, k_r is a continuous index, whose value has to belong to $]1/\sigma; \dots[$, because of the approximations made in equation (60), and finally R_0 whose values belong to $]\sigma; \dots[$. We must also comply with the two assumptions (58) and (59) about the distance Δk_r and ΔR_0 between two consecutive elements of the basis. In order to get rid of the sum over the discrete index, we will use Riemann formula $\sum f(x)\Delta x \simeq \int dx f(x)$, with Δx controlling the distance between two consecutive elements. The dependences with the two radial frequencies k_r^p , k_r^q and the two radii R_0^p , R_0^q are such that the sums on the index p can be completely disentangled from the sums on the index q . In order

to emphasize the gist of the calculation, the diffusion coefficients from equation (86) may be written under the form

$$D_m(\mathbf{J}) = \left\langle \frac{1}{2\pi} \int d\omega' g(\mathbf{m} \cdot \boldsymbol{\Omega}) g^*(\omega') \right\rangle, \quad (87)$$

where $g(\omega)$ is defined as

$$g(\omega) = 2\pi \sum_{k_r^p, R_0^p} g_s(k_r^p, R_0^p, \omega) e^{i(R_g - R_0^p)k_r^p} \mathcal{G}(R_g - R_0^p). \quad (88)$$

In equation (88), $g_s(k_r^p, R_0^p, \omega)$ encompasses all the *slow* dependences of the diffusion coefficients with respect to the position R_0 and the radial frequency k_r so that

$$g_s(k_r^p, R_0^p, \omega) = \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^p \right] \frac{1}{1 - \lambda_{k_r^p}} \widehat{\psi}_{m_\phi, k_r^p}^e[R_0^p, \omega], \quad (89)$$

and $\mathcal{G}(R_g - R_0^p)$ is a normalized Gaussian given by

$$\mathcal{G}(R_g - R_0^p) = \frac{1}{\sqrt{2\pi\sigma^2}} \exp\left[-\frac{(R_g - R_0^p)^2}{2\sigma^2}\right]. \quad (90)$$

In the discrete sum from equation (88), the basis elements are separated by constant step distances ΔR_0 and Δk_r . We suppose that generally k_r^p and R_0^p are given by

$$\begin{cases} k_r^p = n_k \Delta k_r, \\ R_0^p = R_g + n_r \Delta R_0, \end{cases} \quad (91)$$

where n_k is a strictly positive integer and n_r is an integer that can be both positive and negative. One can note in equation (88) the presence of a rapidly evolving complex exponential which may cancel out the diffusion coefficients if the basis step distances are not chosen carefully. Injecting the dependences from equation (91) in the complex exponential from equation (88), one can see that we have to sum terms of the form

$$\exp\left(i(R_g - (R_g + n_r \Delta R_0))n_k \Delta k_r\right) = \exp\left(-in_r n_k \Delta R_0 \Delta k_r\right).$$

As a consequence, since $n_r n_k$ is an integer, in order to have no contributions from the complex exponential term, one has to choose step distances so that

$$\Delta R_0 \Delta k_r = 2\pi. \quad (92)$$

This choice of step distances, imposed by the complex exponential term, corresponds to a critical sampling (Daubechies 1990), which allows us when performing the change to continuous expression in equation (88) to leave out the complex exponential. This transformation is a subtle stage of the calculation, since we require our step distances ΔR_0 and Δk_r to be simultaneously *large* to comply with the WKB constraints from equation (60) and *small* to allow the use of Riemann sum formula. As the radial Gaussian $\mathcal{G}(R_g - R_0^p)$ is sufficiently peaked and correctly normalized, one can replace it by $\delta_D(R_g - R_0^p)$. The integration on R_0^p can then be immediately performed to obtain

$$g(\omega) = \int dk_r^p g_s(k_r^p, R_g, \omega). \quad (93)$$

Using this result in equation (87), we finally obtain the expression of the diffusion coefficients as

$$\begin{aligned} D_m(\mathbf{J}) = & \left\langle \frac{1}{2\pi} \int d\omega' \int dk_r^p \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^p \right] \frac{1}{1 - \lambda_{k_r^p}} \widehat{\psi}_{m_\phi, k_r^p}^e[R_g] \right. \\ & \left. \times \int dk_r^q \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^q \right] \frac{1}{1 - \lambda_{k_r^q}} \widehat{\psi}_{m_\phi, k_r^q}^{e*}[R_g] \right\rangle. \quad (94) \end{aligned}$$

In this expression, all the radial functions have to be evaluated at the position R_g and at the temporal frequency $\mathbf{m} \cdot \boldsymbol{\Omega}$, except for $\widehat{\psi}_{m_\phi, k_r^q}^*$ which is evaluated at the frequency ω' . The eigenvalues λ_{k_r} are given by equation (74) and read

$$\lambda_{k_r}[R_g, \mathbf{m} \cdot \boldsymbol{\Omega}] = \frac{2\pi G \Sigma |k_r|}{\kappa^2 (1 - s^2)} \mathcal{F}(s, \chi). \quad (95)$$

Note that, as requested, in equation (94), all the dependences in σ have disappeared, so that the value of these diffusion coefficients is independent of the precise choice of the WKB basis. One can finally introduce the autocorrelation of the external perturbation \widehat{C}_ψ as

$$\begin{aligned} \widehat{C}_\psi[m_\phi, \omega, k_r^p, k_r^q, R_g] \\ = \frac{1}{2\pi} \int d\omega' \left\langle \widehat{\psi}_{m_\phi, k_r^p}^*[R_g, \omega] \widehat{\psi}_{m_\phi, k_r^q}^*[R_g, \omega'] \right\rangle, \end{aligned} \quad (96)$$

so that the expression (94) of the diffusion coefficients takes the form

$$\begin{aligned} D_m(\mathbf{J}) = \int dk_r^p \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^p \right] \frac{1}{1 - \lambda_{k_r^p}} \int dk_r^q \mathcal{J}_{m_r} \left[\sqrt{\frac{2J_r}{\kappa}} k_r^q \right] \\ \times \frac{1}{1 - \lambda_{k_r^q}} \widehat{C}_\psi[m_\phi, \mathbf{m} \cdot \boldsymbol{\Omega}, k_r^p, k_r^q, R_g]. \end{aligned} \quad (97)$$

We may finally assume that the external perturbations are spatially quasi-stationary, so that we have

$$\begin{aligned} \left\langle \widehat{\psi}_{m_\phi}^e[R_1, t_1] \widehat{\psi}_{m_\phi}^{e*}[R_2, t_2] \right\rangle \\ = \mathcal{C}[m_\phi, t_1 - t_2, R_1 - R_2, (R_1 + R_2)/2], \end{aligned} \quad (98)$$

where the dependence of \mathcal{C} with respect to $(R_1 + R_2)/2$ is supposed to be weak. As demonstrated in Appendix C, one can then show that

$$\begin{aligned} \left\langle \widehat{\psi}_{m_\phi, k_r^1}^*[R_g, \omega_1] \widehat{\psi}_{m_\phi, k_r^2}^{e*}[R_g, \omega_2] \right\rangle \\ = 2\pi \delta_D(\omega_1 - \omega_2) \delta_D(k_r^1 - k_r^2) \widehat{\mathcal{C}}[m_\phi, \omega_1, k_r^1, R_g]. \end{aligned} \quad (99)$$

Using this autocorrelation function *diagonalized* both in ω and k_r , the expression of the diffusion coefficients from equation (94) finally takes the form

$$D_m(\mathbf{J}) = \int dk_r \frac{\mathcal{J}_{m_r}^2 \left[\sqrt{\frac{2J_r}{\kappa}} k_r \right]}{[1 - \lambda_{k_r}]^2} \widehat{\mathcal{C}}[m_\phi, \mathbf{m} \cdot \boldsymbol{\Omega}, k_r, R_g]. \quad (100)$$

Equation (100) is the main result of this section. The corresponding anisotropic tensor diffusion coefficient reads

$$\mathbf{D} = \sum_m \mathbf{m} \otimes \mathbf{m} \int dk_r \frac{\mathcal{J}_{m_r}^2 \left[\sqrt{\frac{2J_r}{\kappa}} k_r \right]}{[1 - \lambda_{k_r}]^2} \widehat{\mathcal{C}}[m_\phi, \mathbf{m} \cdot \boldsymbol{\Omega}, k_r, R_g].$$

One may sometimes simplify further equation (100) when the function $k_r \mapsto \lambda_{k_r}$ is a sharp function reaching a maximum value $\lambda_{\max}(R_g, \omega = \mathbf{m} \cdot \boldsymbol{\Omega})$, for $k_r = k_{\max}(R_g, \omega)$, with a characteristic spread given by Δk_λ . Under this assumption of so-called *small denominators*, the previous expression of the diffusion coefficients can be approximated as

$$D_m(\mathbf{J}) = \Delta k_\lambda \frac{\mathcal{J}_{m_r}^2 \left[\sqrt{\frac{2J_r}{\kappa}} k_{\max} \right]}{[1 - \lambda_{k_{\max}}]^2} \widehat{\mathcal{C}}[m_\phi, \mathbf{m} \cdot \boldsymbol{\Omega}, k_{\max}, R_g]. \quad (101)$$

One should note here that the autocorrelation of the external perturbation \mathcal{C} , which *sources* the diffusion coefficients $D_m(\mathbf{J})$ depends on four different parameters: the azimuthal wavenumber m_ϕ , the

location in the disc via R_g , the radial frequency k_{\max} of the most amplified tightly wound spiral at this position, and finally the local intrinsic frequency $\mathbf{m} \cdot \boldsymbol{\Omega}$.

4 DISCUSSION AND CONCLUSION

Starting from Boltzmann's collisionless equation expressed in angle-actions coordinates and relying on a time-scale decoupling, we derived in equation (36) a diffusion equation describing the long-term evolution of a perturbed self-gravitating collisionless system.⁴ This general formalism is appropriate to capture the *nature* of a collisionless system (via its natural frequencies and susceptibility) as well as its *nurture* via the structure of the power-spectrum of the external perturbations. Hence, it yields the ideal framework in which to study the long-term evolution of such system.

When applying this Fokker–Planck diffusion equation to an infinitely thin galactic disc, we used two main approximations. We first assumed the disc to be tepid. Having orbits with small radial oscillations justified the use of the epicyclic approximation, allowing us to explicitly build up in equation (44) a mapping between the physical coordinates (\mathbf{x}, \mathbf{v}) and the angle-actions coordinates $(\boldsymbol{\theta}, \mathbf{J})$. Another important consequence of the epicyclic development is to allow for a direct determination of the local frequencies of the system Ω and κ , as in equation (42). Being able to *localize* the resonances is crucial in this formalism, since the diffusion coefficients from equation (36) show that both the susceptibility of the system via $[\mathbf{I} - \widehat{\mathbf{M}}]$ and the external perturbing power spectrum via $\widehat{\mathbf{C}}$ have to be evaluated at the intrinsic frequency $\mathbf{m} \cdot \boldsymbol{\Omega}$. The second approximation involves an explicit WKB basis introduced in equation (46). It allowed us to obtain in equation (74) a diagonal response matrix, as if gravity was only *local*. Thanks to the assumption of radial decoupling, the WKB approximation led to equation (100), a simple quadrature for the diffusion coefficients, with which it is straightforward to identify the physically relevant modes. Such simplification provides useful insight into the physical processes at work, e.g. the relevant resonances, their loci and their relative strengths.

The formalism of secular resonant dressed orbital diffusion and its WKB limit is implemented in the companion paper (Fouvry & Pichon 2015, hereafter [Paper II](#)) to recover the formation of resonant ridges in action-space when an isolated stellar Mestel disc (Mestel 1963) is left evolving for hundreds of dynamical times. The development of such ridges has been shown to originate from a resonant mono-dimensional diffusion, specifically enhanced in restricted locations in the disc. It captures the respective roles and importance of various parameters of the system. Indeed, [Paper II](#) illustrates on an example that the self-gravity of the disc (via the amplification eigenvalues λ), its susceptibility (via the anisotropic diffusion coefficients $D_m(\mathbf{J})$), its inhomogeneity (via the gradients $\partial F_0 / \partial \mathbf{J}$), its *temperature* (via σ_r^2), its physical structure (via the introduction of tapering functions representing, respectively, the bulge and the outer edge of the disc), and the detail of the source of perturbations (via the power spectrum of ψ^e), all contributes non-negligibly to the appearance of resonant ridges. Such features have been observed both in numerical experiments (Sellwood 2012) and in the Solar neighbourhood (Wielen 1977; Dehnen 1998; Nordström et al. 2004; Famaey et al. 2005; Aumer & Binney 2009; McMillan 2011).

The WKB assumption can also be used to study the collisional evolution of a self-gravitating disc containing a finite number of

⁴ Appendix A also shows how this diffusion equation is obtained via a different route involving Hamilton's equations.

substructures. Indeed, in Fouvry, Pichon & Chavanis (2015), the same *local* WKB approach will be applied to the Lenard–Balescu non-linear equation (Balescu 1960; Lenard 1960; Weinberg 1998; Heyvaerts 2010; Chavanis 2012b), which accounts for self-driven orbital secular diffusion of a self-gravitating system induced by an intrinsic shot noise due to the discreteness of the system. Possible cases of applications of this approach are the secular diffusion of giant molecular clouds in galactic disc, the secular migration of planetesimals in protoplanetary discs, or even the long-term evolution of population of stars within the Galactic centre.

For self-gravitating systems which do not take the form of an infinitely thin disc, for which the epicyclic approximation and the WKB assumption may be relevant, the formalism of secular forcing can still be used. The diffusion equation (36) could for instance describe the secular diffusion of dark matter cusps in galactic centres induced by perturbations from stochastic feedback processes originating from the baryonic disc (Fouvry et al. 2015). Given a detailed characterization of the perturbations induced by e.g. the cosmic environment, one could also study their long-term effects on a typical self-gravitating collisionless galactic disc. Indeed, in the context of the upcoming *GAIA* mission, this externally induced secular evolution is thought to be a compelling approach to describe the radial migration of stars and its impact on the observed metallicity gradients (Sellwood & Binney 2002; Roškar et al. 2008; Schönrich & Binney 2009; Solway, Sellwood & Schönrich 2012; Minchev, Chiappini & Martig 2013). It may also be applied to describe the secular diffusion of accretion streams within the Galactic halo. Finally, an extension of the formalism of Section 3 to discs with a finite thickness might allow us to understand the process of disc thickening.

ACKNOWLEDGEMENTS

JBF thanks the Gaia Research for European Astronomy Training Programme for travel funding and the department of theoretical physics in Oxford for hospitality. CP and JBF thank the Institute of Astronomy, Cambridge, for hospitality while this investigation was initiated. We thank J. Binney and P. H. Chavanis for detailed comments. This work is partially supported by the Spin(e) grants ANR-13-BS05-0005 of the French *Agence Nationale de la Recherche* and by the LABEX Institut Lagrange de Paris (under reference ANR-10-LABX-63) which is funded by ANR-11-IDEX-0004-02.

REFERENCES

- Aumer M., Binney J. J., 2009, *MNRAS*, 397, 1286
 Balescu R., 1960, *Phys. Fluids*, 3, 52
 Chavanis P.-H., 2013, in Falcón-Barroso J., Knapen J. H., eds, *Dynamics of Secular Evolution*. Cambridge Univ. Press, Cambridge, p. 259
 Binney J., Lacey C., 1988, *MNRAS*, 230, 597
 Binney J., Tremaine S., 2008, *Galactic Dynamics*, 2nd edn. Princeton Univ. Press, Princeton, NJ
 Born M., 1960, *The Mechanics of the Atom*. F. Ungar Pub. Co., New York
 Chavanis P. H., 2012a, *Eur. Phys. J. Plus*, 127, 19
 Chavanis P.-H., 2012b, *Physica A: Stat. Mech. Appl.*, 391, 3680
 Chavanis P.-H., 2013, *Eur. Phys. J. Plus*, 128, 126
 Daubechies I., 1990, *IEEE Trans. Inf. Theory*, 36, 961
 Dehnen W., 1998, *AJ*, 115, 2384
 Dubois Y. et al., 2014, *MNRAS*, 444, 1453
 Earn D. J. D., Sellwood J. A., 1995, *ApJ*, 451, 333
 Evans N. W., Read J. C. A., 1998, *MNRAS*, 300, 106
 Famaey B., Jorissen A., Luri X., Mayor M., Udry S., Dejonghe H., Turon C., 2005, *A&A*, 430, 165

- Fouvry J. B., Pichon C., 2015, *MNRAS*, 449, 1982 (Paper II)
 Fouvry J. B., Pichon C., Chavanis P. H., 2015, submitted
 Gabor D., 1946, *J. Inst. Electr. Eng.: Radio Commun. Eng.*, 93, 429
 Goldstein H., 1950, *Classical Mechanics*. Addison-Wesley, Reading, MA
 Goodman J., 1988, *ApJ*, 329, 612
 Gradshteyn I. S., Ryzhik I. M., 2007, *Table of Integrals, Series, and Products*. Elsevier Academic Press, Amsterdam
 Heyvaerts J., 2010, *MNRAS*, 407, 355
 Jalali M. A., Hunter C., 2005, *ApJ*, 630, 804
 Jeans J. H., 1915, *MNRAS*, 76, 70
 Kalnajs A. J., 1965, PhD thesis. Univ. Harvard
 Kalnajs A. J., 1976, *ApJ*, 205, 745
 Kalnajs A. J., 1977, *ApJ*, 212, 637
 Lenard A., 1960, *Ann. Phys.*, 10, 390
 Lin C. C., Shu F. H., 1966, *Proc. Natl. Acad. Sci.*, 55, 229
 Liouville J., 1837, *J. Math. Pures. Appl.*, 1, 16
 Lynden-Bell D., Kalnajs A. J., 1972, *MNRAS*, 157, 1
 Ma C.-P., Bertschinger E., 2004, *ApJ*, 612, 28
 McMillan P. J., 2011, *MNRAS*, 418, 1565
 Mestel L., 1963, *MNRAS*, 126, 553
 Minchev I., Chiappini C., Martig M., 2013, *A&A*, 558, A9
 Nardini C., Gupta S., Ruffo S., Dauxois T., Bouchet F., 2012, *J. Stat. Mech.: Theory Exp.*, 12, 10
 Nordström B. et al., 2004, *A&A*, 418, 989
 Palmer P., 1994, *Astrophysics and Space Science Library, Stability of Collisionless Stellar Systems: Mechanisms for the Dynamical Structure of Galaxies*. Springer, Dordrecht
 Palmer P. L., Papaloizou J., Allen A. J., 1989, *MNRAS*, 238, 1281
 Pichon C., Aubert D., 2006, *MNRAS*, 368, 1657
 Pichon C., Cannon R. C., 1997, *MNRAS*, 291, 616
 Roškar R., Debattista V. P., Stinson G. S., Quinn T. R., Kaufmann T., Wadsley J., 2008, *ApJ*, 675, L65
 Schönrich R., Binney J., 2009, *MNRAS*, 399, 1145
 Sellwood J. A., 2012, *ApJ*, 751, 44
 Sellwood J. A., Athanassoula E., 1986, *MNRAS*, 221, 195
 Sellwood J. A., Binney J. J., 2002, *MNRAS*, 336, 785
 Solway M., Sellwood J. A., Schönrich R., 2012, *MNRAS*, 422, 1363
 Toomre A., 1964, *ApJ*, 139, 1217
 Vauterin P., Dejonghe H., 1996, *A&A*, 313, 465
 Weinberg M. D., 1991, *ApJ*, 368, 66
 Weinberg M. D., 1993, *ApJ*, 410, 543
 Weinberg M. D., 1998, *MNRAS*, 297, 101
 Weinberg M. D., 2001a, *MNRAS*, 328, 311
 Weinberg M. D., 2001b, *MNRAS*, 328, 321
 Wielen R., 1977, *A&A*, 60, 263
 Zang T. A., 1976, PhD thesis, Massachusetts Institute of Technology

APPENDIX A: STATISTICAL APPROACH VIA HAMILTON'S EQUATION

We now derive the statistical expression (36) of the secular diffusion coefficients using a different method based on Hamilton's equations and inspired from Binney & Lacey (1988). We will indeed quantify the temporal rate of change of the actions, represented by \mathbf{J} . The main difference between the following calculation and that made in Binney & Lacey (1988) is that we take explicitly into account the self-gravity of the system which leads to the appearance of a self-perturbing potential ψ^s triggered by ψ^c . Starting from the Hamiltonian introduced in equation (1), and using the Fourier development in angles as in equation (8), Hamilton's equations, $\dot{\theta} = \partial H / \partial \mathbf{J}$ and $\dot{\mathbf{J}} = -\partial H / \partial \theta$, take the form

$$\begin{cases} \dot{\theta} = \Omega + \sum_m e^{im\theta} \frac{\partial}{\partial \mathbf{J}} [\psi_m^s + \psi_m^c], \\ \dot{\mathbf{J}} = -i \sum_m m e^{im\theta} [\psi_m^s + \psi_m^c]. \end{cases} \quad (\text{A1})$$

As we aim to describe the *wandering* in action-space of the particles, we introduce a limited development of the change in actions and angles of the form

$$\begin{cases} \boldsymbol{\theta}(t) = \boldsymbol{\theta}_0 + \boldsymbol{\Omega}t + \Delta\boldsymbol{\theta}(t), \\ \mathbf{J}(t) = \mathbf{J}_0 + \Delta\mathbf{J}(t). \end{cases} \quad (\text{A2})$$

In order to solve this system of coupled differential equations, we will proceed step-by-step, by including gradually the perturbative terms in $\psi^e + \psi^s$. First of all, one must note that the unperturbed orbits follow the straight-line trajectories $(\boldsymbol{\theta}, \mathbf{J}) = (\boldsymbol{\theta}_0 + \boldsymbol{\Omega}t, \mathbf{J}_0)$. Then, the first-order term in action $\Delta\mathbf{J}$ is given by

$$\Delta\mathbf{J}(T) = \int_0^T dt \dot{\mathbf{J}}(t), \quad (\text{A3})$$

where $\dot{\mathbf{J}}$ is given by Hamilton's equations (A1), where all the occurrences of $\boldsymbol{\theta}(t)$ and $\mathbf{J}(t)$ are replaced by the expressions obtained for the unperturbed orbits. After a time T , the shift in action at first order is therefore given by

$$\Delta\mathbf{J}(T) = -i \sum_m \mathbf{m} \int_0^T dt [\psi_m^s(\mathbf{J}_0, t) + \psi_m^e(\mathbf{J}_0, t)] e^{im \cdot (\boldsymbol{\theta}_0 + \boldsymbol{\Omega}t)}.$$

We introduce the operation of angle-average on the initial phase $\boldsymbol{\theta}_0$ as

$$\{F\}_{\boldsymbol{\theta}_0} = \frac{1}{(2\pi)^d} \int d\boldsymbol{\theta}_0 F(\boldsymbol{\theta}_0). \quad (\text{A4})$$

In order to characterize the *wandering* in action-space, one has to study the behaviour of the *square* of the perturbation $\Delta\mathbf{J}$. From Binney & Lacey (1988), we know the relation between the *wandering* $\Delta\mathbf{J}$ in action-space and the diffusion coefficient appearing in the Fokker–Planck equation, which is given by

$$\mathbf{D}_{ij}(\mathbf{J}_0) = \frac{1}{2T} \{\Delta J_i \Delta J_j\}_{\boldsymbol{\theta}_0}(T), \quad (\text{A5})$$

where the diffusion equation has been written under the compact form

$$\frac{\partial F}{\partial t} = \sum_{i,j} \frac{\partial}{\partial J_i} \left[\mathbf{D}_{ij}(\mathbf{J}) \frac{\partial F}{\partial J_j} \right]. \quad (\text{A6})$$

Taking an average over the initial phases $\boldsymbol{\theta}_0$, using the fact that $\psi_{-m}^{e/s} = [\psi_m^{e/s}]^*$ and projecting the result on the biorthogonal potential basis $\psi^{(p)}$, one can write

$$\begin{aligned} \{\Delta J_i \Delta J_j\}_{\boldsymbol{\theta}_0} &= \sum_m \sum_{p,q} \mathbf{m}_i \mathbf{m}_j \psi_m^{(p)}(\mathbf{J}_0) \psi_m^{(q)*}(\mathbf{J}_0) \\ &\times \int_0^T dt_1 \int_0^T dt_2 [a_p(t_1) + b_p(t_1)] [a_q^*(t_2) + b_q^*(t_2)] e^{im \cdot \boldsymbol{\Omega}(t_1 - t_2)}. \end{aligned} \quad (\text{A7})$$

In order to have an expression which only depends on the exterior potential ψ^e , we use the convolution relation (20), written as an amplification relation, to obtain

$$\begin{aligned} \{\Delta J_i \Delta J_j\}_{\boldsymbol{\theta}_0} &= \sum_m \sum_{p,q} \sum_{k,l} \mathbf{m}_i \mathbf{m}_j \psi_m^{(p)}(\mathbf{J}_0) \psi_m^{(q)*}(\mathbf{J}_0) \\ &\times \int_0^T dt_1 \int_0^T dt_2 \int_0^{t_1} d\tau_1 [\mathbf{I} - \mathbf{M}]_{pk}^{-1}(t_1 - \tau_1) b_k(\tau_1) \\ &\times \int_0^{t_2} d\tau_2 [\mathbf{I} - \mathbf{M}]_{ql}^{-1*}(t_2 - \tau_2) b_l^*(\tau_2) e^{im \cdot \boldsymbol{\Omega}(t_1 - t_2)}. \end{aligned} \quad (\text{A8})$$

This four-dimensional integral is transformed using the change of variables

$$u_1 = t_1 - \tau_1; u_2 = t_2 - \tau_2; v_1 = \tau_1 + \tau_2; v_2 = \tau_1 - \tau_2. \quad (\text{A9})$$

It is straightforward to check that the Jacobian of this transformation is 2, so that equation (A8) becomes

$$\begin{aligned} \{\Delta J_i \Delta J_j\}_{\boldsymbol{\theta}_0} &= \sum_m \sum_{p,q} \sum_{k,l} \mathbf{m}_i \mathbf{m}_j \psi_m^{(p)}(\mathbf{J}_0) \psi_m^{(q)*}(\mathbf{J}_0) \\ &\times \frac{1}{2} \int_0^T du_1 [\mathbf{I} - \mathbf{M}]_{pk}^{-1}(u_1) e^{im \cdot \boldsymbol{\Omega}u_1} \\ &\times \int_0^T du_2 [\mathbf{I} - \mathbf{M}]_{ql}^{-1*}(u_2) e^{-im \cdot \boldsymbol{\Omega}u_2} \int_{-(T-u_2)}^{T-u_1} dv_2 \\ &\times \int_{|v_2|}^{2T-u_1-u_2-|v_2+u_1-u_2|} dv_1 b_k \left[\frac{v_1 + v_2}{2} \right] \\ &\times b_l^* \left[\frac{v_1 - v_2}{2} \right] e^{im \cdot \boldsymbol{\Omega}v_2}. \end{aligned} \quad (\text{A10})$$

We introduce as in equation (27), the operation of ensemble average over many realizations denoted with $\langle \cdot \rangle$. As in equation (28), we assume that the exterior perturbing potential is a stationary random process. One can then perform an ensemble average of the expression (A10), while assuming, as in equation (30), that the response matrix coefficients can be taken out of the ensemble average operation. We obtain

$$\begin{aligned} \langle \{\Delta J_i \Delta J_j\}_{\boldsymbol{\theta}_0} \rangle &= \sum_m \sum_{p,q} \sum_{k,l} \mathbf{m}_i \mathbf{m}_j \psi_m^{(p)}(\mathbf{J}_0) \psi_m^{(q)*}(\mathbf{J}_0) \\ &\times \frac{1}{2} \int_0^T du_1 [\mathbf{I} - \mathbf{M}]_{pk}^{-1}(u_1) e^{im \cdot \boldsymbol{\Omega}u_1} \\ &\times \int_0^T du_2 [\mathbf{I} - \mathbf{M}]_{ql}^{-1*}(u_2) e^{-im \cdot \boldsymbol{\Omega}u_2} \\ &\times \int_{-(T-u_2)}^{T-u_1} dv_2 \mathbf{C}_{kl}(v_2) e^{im \cdot \boldsymbol{\Omega}v_2} \\ &\times (2T - u_1 - u_2 - |v_2 + u_1 - u_2| - |v_2|). \end{aligned} \quad (\text{A11})$$

The next important step of the calculation is to compare T with the various autocorrelation times of the system. The first one is T_{corr}^{ψ} describing the typical autocorrelation time of the realizations of the external perturbations. Two values of the potential perturbations separated by a time larger than T_{corr}^{ψ} can be considered as independent. The second autocorrelation time-scale is $T_{\text{corr}}^{\mathbf{M}}$, which describes the typical autocorrelation time of the response matrix \mathbf{M} and could be called the *look-back* time. From the expression used in equation (A8), one can note that the values of the self-response coefficients are obtained via a non-Markovian mechanism, where the past values are amplified thanks to the response matrix. However, the self-gravitating system cannot have an *infinite memory*, so that only the sufficiently recent past values should play a role in this amplification. As a consequence, during the amplification process, only the past behaviour for a time interval of the order of $T_{\text{corr}}^{\mathbf{M}}$ is relevant and amplified, so that $T_{\text{corr}}^{\mathbf{M}}$ represents the depth with which the self-response mechanism can probe past values. We finally suppose that the time T for which the wandering in phase-space is studied satisfies the comparison relations

$$T \gg T_{\text{corr}}^{\psi}; T \gg T_{\text{corr}}^{\mathbf{M}}. \quad (\text{A12})$$

As a consequence, the integration boundaries appearing in (A11) become

$$\begin{aligned} \langle \{\Delta J_i \Delta J_j\}_{\theta_0} \rangle &= \sum_m \sum_{p,q} \sum_{k,l} \mathbf{m}_i \mathbf{m}_j \psi_m^{(p)}(\mathbf{J}_0) \psi_m^{(q)*}(\mathbf{J}_0) \\ &\times \frac{1}{2} \int_0^{T_{\text{corr}}^M} du_1 [\mathbf{I} - \mathbf{M}]_{pk}^{-1}(u_1) e^{i\mathbf{m} \cdot \Omega u_1} \\ &\times \int_0^{T_{\text{corr}}^M} du_2 [(\mathbf{I} - \mathbf{M})_{ql}^{-1}]^*(u_2) e^{-i\mathbf{m} \cdot \Omega u_2} \\ &\times \int_{-T_{\text{corr}}^{\psi}}^{T_{\text{corr}}^{\psi}} dv_2 \mathbf{C}_{kl}(v_2) e^{i\mathbf{m} \cdot \Omega v_2} \\ &\times (2T - u_1 - u_2 - |v_2 + u_1 - u_2| - |v_2|). \end{aligned} \quad (\text{A13})$$

Thanks to the assumptions (A12), one can see that the last term of equation (A11) can be approximated by $2T$. The remaining integrations can then be seen as truncated temporal Fourier transforms, so that equation (A11) becomes

$$\begin{aligned} \langle \{\Delta J_i \Delta J_j\}_{\theta_0} \rangle &= T \sum_m \sum_{p,q} \sum_{k,l} \mathbf{m}_i \mathbf{m}_j \psi_m^{(p)}(\mathbf{J}_0) \psi_m^{(q)*}(\mathbf{J}_0) \\ &\times [\mathbf{I} - \widehat{\mathbf{M}}]_{pk}^{-1}(\mathbf{m} \cdot \Omega) [(\mathbf{I} - \widehat{\mathbf{M}})_{ql}^{-1}(\mathbf{m} \cdot \Omega)]^* \widehat{\mathbf{C}}_{kl}(\mathbf{m} \cdot \Omega). \end{aligned} \quad (\text{A14})$$

The last step of the simplification is to recall that equation (19) guarantees that $\widehat{\mathbf{M}}^* = \widehat{\mathbf{M}}^t$, so that using equation (A5), we finally obtain the expression of the diffusion coefficients from equation (A6) which read

$$\begin{aligned} \mathbf{D}_{ij}(\mathbf{J}_0) &= \frac{1}{2} \sum_m \sum_{p,q} \mathbf{m}_i \mathbf{m}_j \psi_m^{(p)}(\mathbf{J}_0) \psi_m^{(q)*}(\mathbf{J}_0) \\ &\times [(\mathbf{I} - \widehat{\mathbf{M}})^{-1} \cdot \widehat{\mathbf{C}} \cdot (\mathbf{I} - \widehat{\mathbf{M}})^{-1}]_{pq}(\mathbf{m} \cdot \Omega). \end{aligned} \quad (\text{A15})$$

With this approach, we recover the same diffusion coefficients as the ones obtained in equation (36) via the quasi-linear approach presented in the main text.

APPENDIX B: WKB RESPONSE MATRIX

We estimate the value of the diagonal response matrix coefficients introduced in equation (69) within the WKB approximation. Using the definition of $h(R_g)$ from equation (70) and the fact that J_ϕ is an increasing function of R_g , the integration on J_ϕ in equation (69) takes the form

$$\int dR_g h(R_g) \frac{1}{\sqrt{2\pi(\sigma/\sqrt{2})^2}} \exp\left[-\frac{(R_g - R_0)^2}{2(\sigma/\sqrt{2})^2}\right]. \quad (\text{B1})$$

One should note that in this expression, we have a Gaussian of spread $\sigma/\sqrt{2}$ in R_g , correctly normalized in order to have an integral over R_g equal to 1. Assuming that this Gaussian is sufficiently peaked, we may replace it by $\delta_D(R_g - R_0)$, so that the integral on J_ϕ can be dropped. For a Schwarzschild distribution function as in equation (45), we then obtain

$$\begin{aligned} \widehat{\mathbf{M}}_{[k_\phi, k_r, R_0], [k_\phi, k_r, R_0]}(\omega) &= (2\pi)^2 \mathcal{A}^2 \left| \frac{dJ_\phi}{dR_g} \right|_{R_0} \frac{\Omega \Sigma}{\pi \kappa \sigma_r^2} \sum_{m_r} \frac{1}{\omega - m_r \kappa - k_\phi \Omega} \end{aligned}$$

$$\begin{aligned} &\times \left\{ \left[-m_r \frac{\kappa}{\sigma_r^2} + k_\phi \frac{\partial}{\partial J_\phi} \left[\ln \left(\frac{\Omega \Sigma}{\kappa \sigma_r^2} \right) \right] \right] \int dJ_r e^{-\frac{\kappa J_r}{\sigma_r^2}} \mathcal{J}_{m_r}^2(H_{k_\phi}(k_r)) \right. \\ &\left. - k_\phi \frac{\partial}{\partial J_\phi} \left[\frac{\kappa}{\sigma_r^2} \right] \int dJ_r J_r e^{-\frac{\kappa J_r}{\sigma_r^2}} \mathcal{J}_{m_r}^2(H_{k_\phi}(k_r)) \right\}. \end{aligned} \quad (\text{B2})$$

In order to perform the integration on J_r , the first step is to notice that the only dependence on J_r in the Bessel terms is in $H_{k_\phi}(k_r)$, through $A = \sqrt{2J_r/\kappa}$. Therefore, we will use the two integration formula (see formula 6.615.1 from Gradshteyn & Ryzhik 2007)

$$\begin{cases} \int_0^{+\infty} dJ_r e^{-\alpha J_r} \mathcal{J}_{m_r}^2(\beta \sqrt{J_r}) = \frac{e^{-\beta^2/2\alpha}}{\alpha} \mathcal{I}_{m_r} \left[\frac{\beta^2}{2\alpha} \right], \\ \int_0^{+\infty} dJ_r J_r e^{-\alpha J_r} \mathcal{J}_{m_r}^2(\beta \sqrt{J_r}) = \frac{e^{-\beta^2/2\alpha}}{\alpha^2} \left\{ -\frac{\beta^2}{2\alpha} + 1 + |m_r| \right\} \mathcal{I}_{m_r} \left[\frac{\beta^2}{2\alpha} \right] \\ \quad + \frac{\beta^2}{2\alpha} \mathcal{I}_{|m_r|+1} \left[\frac{\beta^2}{2\alpha} \right]. \end{cases} \quad (\text{B3})$$

where $\alpha > 0$, $\beta > 0$, and $m_r \in \mathbb{Z}$. First, let us write out explicitly the dependence of $\mathcal{H}_{k_\phi}(k_r)$ with J_r . From equations (43) and (65), we find that

$$\mathcal{H}_{k_\phi}(k_r) = \sqrt{J_r} \beta, \quad (\text{B4})$$

where β is defined as

$$\beta = \sqrt{\frac{2}{\kappa} \left[k_r^2 + k_\phi^2 \left[\frac{2\Omega}{\kappa R_g} \right]^2 \right]} \simeq \sqrt{\frac{2}{\kappa}} k_r. \quad (\text{B5})$$

This approximate expression has been obtained using the same approximation as in equation (66). We also introduce the notation

$$\chi = \frac{\sigma_r^2}{\kappa^2} \left[k_r^2 + k_\phi^2 \left[\frac{2\Omega}{\kappa R_g} \right]^2 \right] \simeq \frac{\sigma_r^2 k_r^2}{\kappa^2}. \quad (\text{B6})$$

We are now able to compute the integrals on J_r from equation (B2), to obtain

$$\begin{aligned} \widehat{\mathbf{M}}_{[k_\phi, k_r, R_0], [k_\phi, k_r, R_0]}(\omega) &= (2\pi)^2 \mathcal{A}^2 \left| \frac{dJ_\phi}{dR_g} \right|_{R_0} \frac{\Omega \Sigma}{\pi \kappa \sigma_r^2} \sum_{m_r} \frac{1}{\omega - m_r \kappa - k_\phi \Omega} \\ &\times \left\{ e^{-\chi} \frac{\sigma_r^2}{\kappa} \mathcal{I}_{m_r}[\chi] \left[-m_r \frac{\kappa}{\sigma_r^2} + k_\phi \frac{\partial}{\partial J_\phi} \left[\ln \left(\frac{\Omega \Sigma}{\kappa \sigma_r^2} \right) \right] \right] \right. \\ &\left. - k_\phi \frac{\partial}{\partial J_\phi} \left[\frac{\kappa}{\sigma_r^2} \right] e^{-\chi} \frac{\sigma_r^4}{\kappa^2} \left[(1 + |m_r| - \chi) \mathcal{I}_{m_r}[\chi] + \chi \mathcal{I}_{|m_r|+1}[\chi] \right] \right\}. \end{aligned} \quad (\text{B7})$$

In order to simplify this expression, we recall that we have the property $\mathcal{I}_{-m_r}(\chi) = \mathcal{I}_{m_r}(\chi)$. Therefore, in equation (B7), we have to study four types of sums on m_r , which may be simplified as

$$\begin{cases} \sum_{m_r \in \mathbb{Z}} \frac{m_r \mathcal{I}_{m_r}[\chi]}{\omega' - m_r \kappa} = -\frac{2}{\kappa} \sum_{m_r=1}^{+\infty} \frac{\mathcal{I}_{m_r}[\chi]}{1 - [\omega'/m_r \kappa]^2}, \\ \sum_{m_r \in \mathbb{Z}} \frac{\mathcal{I}_{m_r}[\chi]}{\omega' - m_r \kappa} = \frac{\mathcal{I}_0[\chi]}{\omega'} + \frac{2}{\omega'} \sum_{m_r=1}^{+\infty} \frac{\mathcal{I}_{m_r}[\chi]}{1 - [m_r \kappa/\omega']^2}, \end{cases} \quad (\text{B8})$$

$$\left\{ \begin{aligned} \sum_{m_r \in \mathbb{Z}} \frac{|m_r| \mathcal{I}_{m_r}[\chi]}{\omega' - m_r \kappa} &= \frac{2}{\omega'} \sum_{m_r=1}^{+\infty} \frac{m_r \mathcal{I}_{m_r}[\chi]}{1 - [m_r \kappa / \omega']^2}, \\ \sum_{m_r \in \mathbb{Z}} \frac{\mathcal{I}_{|m_r|+1}[\chi]}{\omega' - m_r \kappa} &= \frac{\mathcal{I}_1[\chi]}{\omega'} + \frac{2}{\omega'} \sum_{m_r=1}^{+\infty} \frac{\mathcal{I}_{m_r+1}[\chi]}{1 - [m_r \kappa / \omega']^2}, \end{aligned} \right. \quad (\text{B9})$$

where we use $\omega' = \omega - k_\phi \Omega_\phi$. We define the dimensionless parameter s as

$$s = \frac{\omega - k_\phi \Omega}{\kappa}. \quad (\text{B10})$$

We also introduce the reduction factor $\mathcal{F}(s, \chi)$ (Kalnajs 1965; Lin & Shu 1966) and similar functions $\mathcal{G}(s, \chi)$, $\mathcal{H}(s, \chi)$ and $\mathcal{I}(s, \chi)$ defined as

$$\left\{ \begin{aligned} \mathcal{F}(s, \chi) &= 2(1 - s^2) \frac{e^{-\chi}}{\chi} \sum_{m_r=1}^{+\infty} \frac{\mathcal{I}_{m_r}[\chi]}{1 - [s/m_r]^2}, \\ \mathcal{G}(s, \chi) &= 2(1 - s^2) \frac{e^{-\chi}}{\chi} \left[\frac{1}{2} \frac{\mathcal{I}_0[\chi]}{s} + \frac{1}{s} \sum_{m_r=1}^{+\infty} \frac{\mathcal{I}_{m_r}[\chi]}{1 - [m_r/s]^2} \right], \\ \mathcal{H}(s, \chi) &= 2(1 - s^2) \frac{e^{-\chi}}{\chi} \frac{1}{s} \sum_{m_r=1}^{+\infty} \frac{m_r \mathcal{I}_{m_r}[\chi]}{1 - [m_r/s]^2}, \\ \mathcal{I}(s, \chi) &= 2(1 - s^2) \frac{e^{-\chi}}{\chi} \left[\frac{1}{2} \frac{\mathcal{I}_1[\chi]}{s} + \frac{1}{s} \sum_{m_r=1}^{+\infty} \frac{\mathcal{I}_{m_r+1}[\chi]}{1 - [m_r/s]^2} \right]. \end{aligned} \right. \quad (\text{B11})$$

Moreover, we notice that we can use the simplification $\partial/\partial J_\phi [\kappa/\sigma_r^2] \sigma_r^4/\kappa^2 = -\partial/\partial J_\phi [\sigma_r^2/\kappa]$, and that thanks to equation (42), one can also explicitly compute

$$\left| \frac{dJ_\phi}{dR_g} \right|_{R_0} = \frac{R_g^3 \kappa^2}{2J_\phi} \Big|_{R_0} = \frac{R_0 \kappa^2}{2\Omega}. \quad (\text{B12})$$

Finally, using the expression of the amplitude of the basis potentials from equation (63), we obtain a detailed expression of the matrix coefficients as

$$\begin{aligned} \widehat{\mathbf{M}}_{[k_\phi^p, k_r^p, R_0], [k_\phi^q, k_r^q, R_0]}(\omega) &= \delta_{k_\phi^p, k_\phi^q} \delta_{k_r^p, k_r^q} \frac{2\pi G \Sigma |k_r|}{\kappa^2 (1 - s^2)} \\ &\times \left\{ \mathcal{F}(s, \chi) + k_\phi^p \frac{\sigma_r^2}{\kappa} \frac{\partial}{\partial J_\phi} \left[\ln \left(\frac{\Omega \Sigma}{\kappa \sigma_r^2} \right) \right] \mathcal{G}(s, \chi) \right. \\ &\left. + k_\phi^p \frac{\partial}{\partial J_\phi} \left[\frac{\sigma_r^2}{\kappa} \right] [(1 - \chi)\mathcal{G}(s, \chi) + \mathcal{H}(s, \chi) + \chi \mathcal{I}(s, \chi)] \right\}. \end{aligned} \quad (\text{B13})$$

where one must remember that within the WKB approximation, the response matrix is diagonal. For a tepid disc, we may neglect some of the terms appearing in equation (B13). A tepid disc corresponds to a disc where the orbits possess a small radial energy, so that all the orbits are close to circular orbits. It also implies that $|\partial F_0/\partial J_r| \gg |\partial F_0/\partial J_\phi|$. For a Schwarzschild distribution function, the typical spread in J_r is of the order of σ_r^2/κ , so that we may consider equation (B13) as a limited development in σ_r^2/κ and $\partial/\partial J_\phi [\sigma_r^2/\kappa]$. Therefore, for a tepid disc the diagonal coefficients of the response matrix finally take the form given in equation (74).

APPENDIX C: AUTOCORRELATION DIAGONALIZATION

Let us now show how the hypothesis of spatially quasi-stationarity of the external perturbations introduced in equation (98) leads to a *diagonalization* of the autocorrelation with respect to the radial frequencies k_r , as shown in equation (99). In order to shorten the notations, we do not write anymore the dependence with respect to the azimuthal number m_ϕ , and the exterior perturbation will be noted as $\psi = \psi^e$. As a consequence, the assumption of temporal and quasi-spatial stationarity from equation (98) takes the form

$$\langle \psi[R_1, t_1] \psi^*[R_2, t_2] \rangle = \mathcal{C}[t_1 - t_2, R_1 - R_2, (R_1 + R_2)/2]. \quad (\text{C1})$$

Equation (94) for the diffusion coefficients requires us to study the term $\langle \widehat{\psi}_{k_r^1}[R_g, \omega_1] \widehat{\psi}_{k_r^2}^*[R_g, \omega_2] \rangle$. Using the definition of the temporal Fourier transform from equation (11) and the local radial Fourier transform from equation (84), we may rewrite it as

$$\begin{aligned} &\langle \widehat{\psi}_{k_r^1}[R_g, \omega_1] \widehat{\psi}_{k_r^2}^*[R_g, \omega_2] \rangle \\ &= \frac{1}{4\pi^2} \int dt_1 dt_2 dR_1 dR_2 e^{i\omega_1 t_1} e^{-i\omega_2 t_2} \langle \psi[R_1, t_1] \psi^*[R_2, t_2] \rangle \\ &\quad \times g[R_g - R_1] g[R_g - R_2] e^{-i(R_1 - R_g)k_r^1} e^{i(R_2 - R_g)k_r^2}, \end{aligned} \quad (\text{C2})$$

where $g[R]$ is defined as

$$g[R] = \exp[-R^2/(2\sigma^2)]. \quad (\text{C3})$$

We now use the assumption from equation (C1) relative to the radial dependences of the perturbation autocorrelation, and the change of variables

$$\begin{cases} u_r = t_1 + t_2 & ; & v_r = t_1 - t_2, \\ u_r = \frac{1}{2}(R_1 + R_2) & ; & v_r = R_1 - R_2. \end{cases} \quad (\text{C4})$$

This transformation is of determinant 2, so that equation (C2) becomes

$$\begin{aligned} &\langle \widehat{\psi}_{k_r^1}[R_g, \omega_1] \widehat{\psi}_{k_r^2}^*[R_g, \omega_2] \rangle \\ &= \frac{1}{8\pi^2} \int du_r dv_r du_r dv_r e^{i\frac{\omega_1 - \omega_2}{2} u_r} e^{i\frac{\omega_1 + \omega_2}{2} v_r} \\ &\quad \times e^{-i(k_r^1 - k_r^2)u_r} e^{-i\frac{k_r^1 + k_r^2}{2} v_r} e^{iR_g(k_r^1 - k_r^2)} \\ &\quad \times g[R_g - u_r - v_r/2] g[R_g - u_r + v_r/2] \mathcal{C}[v_r, v_r, u_r]. \end{aligned} \quad (\text{C5})$$

The integration on u_r is straightforward and is equal to $2\pi\delta_D((\omega_1 - \omega_2)/2)$. The integration on v_r is then direct and gives $\widehat{\mathcal{C}}[\omega_1, v_r, u_r]$. Finally, we note that the product of the two Gaussians in equation (C5) can be rewritten in order to disentangle the dependences on u_r and v_r to read

$$g[R_g - u_r - v_r/2] g[R_g - u_r + v_r/2] = g[\sqrt{2}(R_g - u_r)] g[v_r/\sqrt{2}], \quad (\text{C6})$$

where the presence of $\sqrt{2}$ comes from the definition of the g function introduced in equation (C3). One can then rewrite equation (C5) as

$$\begin{aligned} &\langle \widehat{\psi}_{k_r^1}[R_g, \omega_1] \widehat{\psi}_{k_r^2}^*[R_g, \omega_2] \rangle \\ &= \frac{1}{2\pi} e^{iR_g(k_r^1 - k_r^2)} \delta_D(\omega_1 - \omega_2) \int dv_r e^{-i\frac{k_r^1 + k_r^2}{2} v_r} g[v_r/\sqrt{2}] \\ &\quad \times \int du_r \widehat{\mathcal{C}}[\omega_1, v_r, u_r] g[\sqrt{2}(R_g - u_r)] e^{-i(k_r^1 - k_r^2)u_r}. \end{aligned} \quad (\text{C7})$$

As we have assumed that the function $u_r \mapsto \widehat{\mathcal{C}}[\omega_1, v_r, u_r]$ is a slowly varying function, we may take it out of the integration on u_r and

evaluate it as $\widehat{\mathcal{C}}[\omega_1, v_r, R_g]$. The remaining integration on u_r can then be computed and reads

$$\begin{aligned} & \int du_r g[\sqrt{2}(R_g - u_r)] e^{-i(k_r^1 - k_r^2)u_r} \\ &= \sqrt{\pi}\sigma e^{-iR_g(k_r^1 - k_r^2)} \exp\left[-\frac{(k_r^1 - k_r^2)^2}{4\sigma^2}\right] \\ &= 2\pi \delta_D(k_r^1 - k_r^2) e^{-iR_g(k_r^1 - k_r^2)}, \end{aligned} \quad (\text{C8})$$

where we replaced the Gaussian in $k_r^1 - k_r^2$ by a Dirac delta, while paying a careful attention to the correct normalization. As a consequence, equation (C7) becomes

$$\begin{aligned} & \langle \widehat{\psi}_{k_r^1}[R_g, \omega_1] \widehat{\psi}_{k_r^2}^*[R_g, \omega_2] \rangle = \delta_D(\omega_1 - \omega_2) \delta_D(k_r^1 - k_r^2) \\ & \times \int dv_r e^{-ik_r^1 v_r} \widehat{\mathcal{C}}[\omega_1, v_r, R_g] g[v_r/\sqrt{2}]. \end{aligned} \quad (\text{C9})$$

Because of the definition from equation (84), the presence of the factor $1/\sqrt{2}$ corresponds to the change $\sigma \rightarrow \sqrt{2}\sigma$ so that the remaining integral on v_r may be interpreted as a local radial Fourier transform centred around the position $v_r = 0$. Therefore, we straightforwardly obtain the *diagonalized* expression introduced in equation (99).

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.