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Lorrie A. Straka, Pasquier Noterdaeme, Raghunathan Srianand, Songkiat Nutalaya, Varsha P. Kulkarni, et al.. Galactic nebular lines in the fibre spectra of background QSOs: reaching a hundred QSO-galaxy pairs with spectroscopic and photometric measurements. *Monthly Notices of the Royal Astronomical Society*, 2015, 447, pp.3856-3872. 10.1093/mnras/stu2739 . insu-03644742

HAL Id: insu-03644742

<https://insu.hal.science/insu-03644742>

Submitted on 25 Apr 2022

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Galactic nebular lines in the fibre spectra of background QSOs: reaching a hundred QSO–galaxy pairs with spectroscopic and photometric measurements

Lorrie A. Straka,^{1★} Pasquier Noterdaeme,² Raghunathan Srianand,³
Songkiat Nutalaya,⁴ Varsha P. Kulkarni,⁵ Pushpa Khare,^{3†} David Bowen,⁶
Michael Bishof¹ and Donald G. York^{1,7}

¹Department of Astronomy and Astrophysics, University of Chicago, Chicago, IL 60637, USA

²Institut d’Astrophysique de Paris, UMR7095, CNRS-UMPC, 98bis bd Arago, F-75014 Paris, France

³Inter-University Centre for Astronomy and Astrophysics (IUCAA), Ganeshkhind, 411 007 Pune, India

⁴Department of Astronomy, University of Massachusetts, Amherst, MA 01003, USA

⁵Department of Physics and Astronomy, University of South Carolina, Columbia, SC 29208, USA

⁶Department of Physics and Astronomy, Princeton University, Princeton, NJ 08544, USA

⁷Enrico Fermi Institute, University of Chicago, Chicago, IL 60637, USA

Accepted 2014 December 22. Received 2014 December 20; in original form 2014 October 10

ABSTRACT

We present photometric and spectroscopic measurements of 53 QSO–galaxy pairs from the Sloan Digital Sky Survey (SDSS), where nebular emission lines from a $0 < z < 0.84$ foreground galaxy are detected in the fibre spectra of a background QSO, bringing the overall sample to 103 QSO–galaxy pairs detected in the SDSS. We here study the nature of these systems. Detected foreground galaxies appear at impact parameters between 0.37 and 12.68 kpc. The presence of oxygen and Balmer emission lines allows us to determine the emission line metallicities for our sample, which are on average supersolar in value. Star formation rates for our sample are in the range $0.01\text{--}12 M_{\odot} \text{ yr}^{-1}$. We utilize photometric redshift fitting techniques to estimate the M_{*} values of our galaxies ($\log M_{*} = 7.34\text{--}11.54$), and extrapolate this relationship to those galaxies with no imaging detections. Where available, we measure the absorption features present in the QSO spectrum due to the foreground galaxy and the relationships between their rest equivalent widths. We report an anticorrelation between impact parameter and $E(B - V)_{(g-i)}$, as well as a correlation between galaxy colour ($u - r$) and $E(B - V)_{(g-i)}$. We find that our sample is one of late-type, star-forming galaxies comparable to field galaxies in a similar redshift range, providing important clues to better understand absorption systems. These galaxies represent a sample of typical galaxies in the local Universe for which abundances, extinction, morphology, and absorption properties may be measured using background QSOs with great potential for follow-up observations.

Key words: galaxies: evolution – galaxies: photometry – quasars: absorption lines – cosmology: observations.

1 INTRODUCTION

Various techniques have been developed to uncover the links between gas and galaxies over cosmic history, approaching the problem from either side. For example, the gas causes a unique absorption signature to be imprinted in the spectra of background sources,

which, at low redshift, can be associated and correlated with the properties of foreground galaxies (e.g. Bergeron & Boisse 1991). At high redshift, absorption lines have been used to predict properties of the associated galaxies such as column density, absorption metallicity, abundance ratios, and gas kinematics (e.g. Ledoux et al. 1998; Wolfe, Prochaska & Gawiser 2003; Srianand et al. 2005) where direct detection of the associated galaxies remains challenging and limited to a small number of cases (e.g. Noterdaeme et al. 2012; Péroux et al. 2012; Fynbo et al. 2013). On the other hand, cross-correlating low-redshift galaxy and QSO position catalogues

* E-mail: straka.lorrie@gmail.com

† CSIR Emeritus Scientist.

can be used to subsequently study the properties of the gas (e.g. Gupta et al. 2010).

While absorption lines offer a powerful tool for studying the evolution of galaxies over a range of epochs, QSOs with known foreground galaxies do not necessarily feature the expected absorption signatures (e.g. Bowen, Blades & Pettini 1995; Tripp et al. 2005; Chen et al. 2010). Therefore, it is important to investigate the properties of those foreground galaxies known in emission in order to correlate them with the presence or absence of absorption, and with the properties of absorption features, if detected. Do these systems select the most metal-rich gas (e.g. Wild, Hewett & Pettini 2006; Nestor et al. 2008), and are these halo gas detections? The question yet remains how absorption systems seen in QSO spectra are related to the emission characteristics of the underlying galaxies.

Recently, narrow galactic emission lines present in the fibre spectra of background QSOs have been successfully utilized to probe galaxies at relatively small impact parameters. The techniques for galactic emission line searches are presented in Borthakur et al. (2010), Noterdaeme, Srianand & Mohan (2010, hereafter N10), York et al. (2012, hereafter Y12), and in Straka et al. (2013, hereafter S13). The selection of galaxies via emission lines is independent of its absorption characteristics and broad-band fluxes, bringing a valuable, independent view on the relationship between galaxies and absorption systems.

N10 presented a sample of [O III]-emitting galaxies on top of QSOs (GOTOQs). Therein, they focus on a subset of 17 systems at $z = 0.4\text{--}0.8$ for which the Mg II doublet is covered by the spectra. They measured lower limits on the star formation rate (SFR) in the range $0.2\text{--}20 M_{\odot} \text{ yr}^{-1}$ and detected strong Mg II, Mg I, and Fe II absorption lines in most cases. There is also a possible 2σ correlation between emission line metallicity and Mg II rest equivalent width (EW). By stacking spectra of a large sample of Mg II-selected systems, N10 also detected for the first time a correlation between Mg II EW and the fluxes in the [O II] and [O III] lines. However, the interpretation that it represents a correlation between Mg II EW and SFR (Ménard et al. 2011) may be too simplistic as differential fibre losses are at play (Lopez & Chen 2012). These are based on a stacked sample of higher redshift absorbers and the geometry of the absorption and emission is not clear (Ménard & Fukugita 2012).

Y12 dealt with the detection of GOTOQs in Sloan Digital Sky Survey (SDSS) DR5 through a search for strong H α and presented methods to deconvolve the QSO from the galaxy in the SDSS images. Building upon Y12, S13 utilized a multi-emission line search, simultaneously matching up to nine galactic emission lines in the optical. This paper showed that GOTOQs are more likely to be late-type, star-forming, and lightly reddened galaxies. Srianand et al. (2013, hereafter Sr13) presented results on one QSO–galaxy pair in 21-cm and optical data. They found that this line of sight has near-solar emission metallicity and dust extinction typical of translucent clouds. They also detect the $\lambda 6284$ diffuse interstellar band (DIB) and strong 21-cm absorption at the redshift of the galaxy.

Here, we present 28 additional detections of emission lines from foreground galaxies in QSO spectra, obtained by applying the procedures of N10 to H α . We perform spectroscopic and photometric measurements as in S13 for previously known GOTOQs (mostly from N10) for which such an analysis was not yet performed.

Results showing follow-up observations of such cases of GOTOQs may be found in Borthakur et al. (2010), Gupta et al. (2013), and Sr13. All three papers report the results of galactic interstellar absorption from H I 21 cm, Ca II H and K, and the Na I D lines.

Section 2 details our sample selection and addresses any bias in the sample. Section 3 details our data analysis. Section 4

reports our results. Finally, Section 5 summarizes our conclusions. Throughout, we have assumed the concordance cosmology ($H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.30$, $\Omega_{\Lambda} = 0.70$).

2 SAMPLE AND BIAS

2.1 Sample selection

We have systematically searched 102 418 QSO spectra from SDSS DR7 for intervening H α emission (excluding ~ 2000 for poor signal to noise) at $0 < z < 0.35$. The QSO continua for these spectra were iteratively fit by applying Savitzky–Golay filtering (Savitzky & Golay 1964), which performs a local polynomial regression while preserving relative maxima and minima (such as emission and absorption features). Then, the continuum-subtracted QSO spectra were cross-correlated with a template profile of H α +[N II] (referred to as ‘H α ’) generated from the SDSS galaxy template spectra. No signal-to-noise restrictions were placed on the H α search. About 350 H α candidates were returned from this search. Each candidate was then visually inspected for the presence of other galactic emission lines, such as H β and [O II] $\lambda 3727$. This provided us with 52 GOTOQs (including 17 $z > 0.4$ Mg II absorbers from N10), plus the 50 found with the independent searches from Y12 (H α) and S13 (multiple emission). In addition, one system has also been studied in Gupta et al. (2013), who present a study of the 21-cm absorption it produces in the spectrum of the background QSO. We include this system here due to its classification as a GOTOQ according to our criteria. This brings our total sample to 103 GOTOQs. We present spectroscopic and photometric measurements for the 53 remaining systems (the first 50 being published in Y12 or S13).

Throughout the paper, we primarily refer to the total sample of 103 GOTOQs. We will also refer to three sub-samples. From our total sample, all objects selected by H α will be represented in the figures as plus signs and referred to as the H α sample, all objects selected by [O III] will be represented as crosses in the figures and referred to as the [O III] sample, and those objects selected simultaneously by multiple lines will be represented as dots in the figures and referred to as the multisample.

Fig. 1(a) shows the redshift distribution of the total galaxy sample divided by detection criteria (H α , [O III], or multiple emission) and Fig. 1(b) shows the [O III] flux versus H α flux, again plotted by detection criteria. We find no discernible difference between the three SDSS GOTOQ samples in terms of parent population, indicating the selection methods are roughly equivalent. The three different search methods used here return slightly different sets of unique systems. However, we find distributions of various properties of these systems are similar. This implies that the three galaxy samples are drawn from the same population despite being selected through different search techniques, as can be seen in the figures.

Table 1 lists our sample of 103 GOTOQs: the SDSS plate, fibre, and MJD; the QSO RA and Dec.; references where the GOTOQ was first reported and where it was first photometrically treated in publication; and which of the three search functions returned them (allowing us to compare the three selection techniques). Those GOTOQs present here (53 of them) are labelled as this work. All subsequent tables and figures deal only with those targets treated in this work. Table 2 lists the index number for each of the 53 targets in this work, by which they will be referred to throughout, and the truncated four digit RA and Dec. ID for each target, the SDSS r -band point spread function (PSF) magnitude for the QSO, the Petrosian radius of the QSO, the five filter deconvolved magnitudes of the QSO, the reddening estimates $E(B - V)_{(g-i)}$ due to the

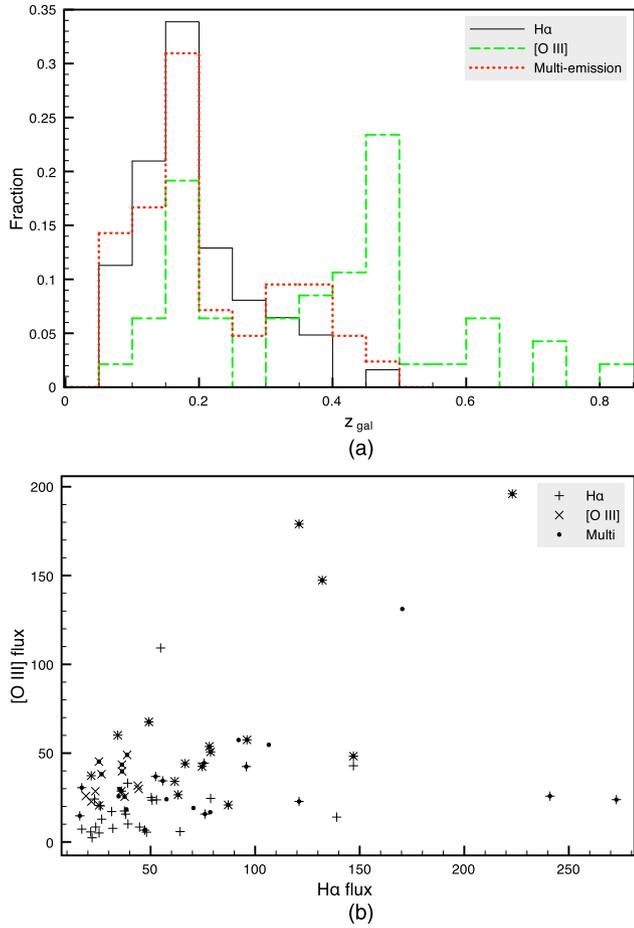


Figure 1. (a) The redshift distribution for our galaxy sample. Redshifts are determined by galactic emission in multiple lines. We see the same peaked distribution for the three search techniques at $z < 0.4$, with extended coverage for the [O III] search, which was able to reach up to $z = 0.8$. (b) [O III] flux versus H α flux for our sample. We see no difference in the detection techniques, which return roughly the same flux distribution in both H α and [O III].

foreground galaxy, and the number of other QSO Absorption Line System (QSOALS) along the line of sight to each QSO (excepting the galaxy studied in this work).

We have included online only a machine-readable table of all the measurements taken for the total sample. The first row includes the column headers. The columns for the table are as follows: QSO RA, QSO Dec., plate, MJD, fibre, QSO z , galaxy z , luminosity distance, three selection columns (H α , [O III], and multiline detection marked ‘1’ for each if they are found with these techniques), visual detection, Petrosian radius, PSF r -band magnitude of the QSO, QSO u magnitude, QSO g magnitude, QSO r magnitude, QSO i magnitude, QSO z magnitude, the number of additional intervening QSOALS, θ (arcsec), impact parameter (b ; kpc), galaxy u magnitude, $\log L/L_{\odot}^u$, galaxy g magnitude, galaxy r magnitude, $\log L/L_{\odot}^r$, L^*/L_{\odot}^r , galaxy i magnitude, $\log L/L_{\odot}^i$, galaxy z magnitude, $(u - r)$ galaxy colour, $E(B - V)_{g-r}$, H α flux, error, H β flux, error, [O III] flux (unresolved), error, [O III]a flux, error, [O III]b flux, error, [N II]a flux, error, [N II]b flux, error, [S II]a flux, error, [S II]b flux, error, R_{23l} , R_{23u} , O3N2, N2, $E(B - V)_{H\alpha/H\beta}$, $SFR_{H\alpha}$, $SFR_{[OIII]}$, $SFR_{H\alpha}^{\text{corrected}}$, $SFR_{[OIII]}^{\text{corrected}}$, $\log M_*$, $1\sigma M_*$ low bound, $1\sigma M_*$ upper bound, Ca II K EW, error, Ca II H EW, error, Na I D2 EW, error,

Na I D1 EW, error, Mg II $\lambda 2796$ EW, error, Mg II $\lambda 2804$ EW, error, Mg I $\lambda 2852$ EW, error, Fe II $\lambda 2587$ EW, error, and Fe II $\lambda 2600$ EW. Dashes in the cells indicate measurements that could not be made due to missing data (redshift range or missing data). Each flux and absorption measurement column has error columns following them. Fluxes are measured in 10^{-17} ergs s^{-1} cm^{-2} . Absorption lines are rest EW. Magnitude errors are less than 0.05 mag in each band. Limits are discussed in Section 3.1.

2.2 Bias

The discovery of this class of objects was serendipitous (Borthakur et al. 2010; N10; Y12). The follow-up to these discoveries (N10; Y12; S13) was aimed at finding interesting galaxies for which further studies (21 cm, QSOALS) would provide insights into the nature of galaxies. The primary purpose of this paper is to complete the catalogue of photometric measurements of the GOTOQ sample from SDSS.

There are several obvious biases. First, these galaxies show up because fibres were pointed at QSO candidates selected from the SDSS QSO candidate follow-up spectroscopy (Richards et al. 2003). Those that were real QSOs (Schneider et al. 2010, 105 783 QSOs as of DR7) show at least 100 galaxies at $z < 0.4$ which are included in the 3 arcsec diameter fibres of SDSS. The question arises whether these galaxies would be missed in a random galaxy survey without the background QSOs. Assuming a homogeneous distribution of galaxies in the sky, pointing the same number of fibers ($\sim 109\,000$) at random parts of the sky would presumably also give about the same number of galaxies to the same magnitude limit. We find galaxies down to a photometric magnitude limit of $r < 23$ and still detect strong spectroscopic emission in one hour of integration on SDSS. Therefore, we expect that in a random galaxy survey search the galaxies in our sample would not be passed over based on their emission line strengths, and there is no reason to expect our sample to differ from a random search except for cosmic small number statistics.

The second bias is that the continuum noise from the background QSOs actually limits the detection limit of the survey. The faintest galaxies would not be seen in the brightest QSOs (N10). However, since the number of QSOs detectable to the 19.1 mag photometric limit is roughly a factor of 6 per magnitude interval, this is a relatively small effect. This is derived from the fact that at low redshifts, any population uniformly distributed in space with a uniform luminosity distribution at all redshifts has six times more objects in an interval of one magnitude than in the previous interval of one magnitude. Simply, the very brightest QSOs are few and therefore the number of GOTOQs missed along the lines of sight to these QSOs is negligible to the total sample. See N10 for a complete discussion. From this bias, < 1 per cent of a true sample is missed, since at 17 mag there are already 1/36 as many QSOs as at 19 mag, and 1/216 as many at 16 mag.

The galaxies may have significant effect on the colours of the QSOs, causing the QSO to fall outside the QSO-selection colour-box. The same exposure time may detect the galaxy in a random fibre position on the sky. Additionally, if the galaxy is of sufficient brightness, it may make it difficult to determine a point source (such as the QSO).

3 DATA ANALYSIS

The sample herein has been treated in the same way as the samples from Y12 and S13, respectively. For more in depth discussion

Table 1. Catalogue of GOTOQs.^a

RA	Dec.	Plate	MJD	Fibre	z_{QSO}	z_{gal}	Detection ^b	Photo ^c	Search return ^d
00:13:42.44	−00:24:12.6	389	51795	274	1.64	0.1557	York+12	York+12	Y
00:16:37.07	+13:56:53.5	752	52251	43	2.57	0.0912	York+12	York+12	Y,S,N
00:45:15.45	−01:13:17.6	393	51794	122	1.59	0.1055	H α	This work	N
00:59:40.67	−00:09:46.2	1083	52520	105	2.39	0.3058	Straka+13a	Straka+13a	S
02:37:04.61	+27:52:39.7	2444	54082	194	0.97	0.3368	Straka+13a	Straka+13a	S
02:39:14.40	−07:05:57.1	455	51909	562	0.71	0.3423	Noterdaeme+10	York+12	Y,N
02:43:28.86	+00:38:31.2	807	52295	341	2.75	0.0279	Straka+13a	Straka+13a	S
02:59:42.42	+00:01:38.1	802	52289	244	2.20	0.0430	Straka+13a	Straka+13a	S
03:12:55.98	−00:14:00.0	1179	52637	99	1.03	0.1147	Straka+13a	Straka+13a	S
03:32:28.25	−00:14:34.7	1063	52591	236	1.74	0.0030	Straka+13a	Straka+13a	S

^aThe full table is available online.^bFirst published detection. H α indicates it is not yet published, but is first presented in this paper.York+12 detected in H α , Straka+13a detected in nine multiple galactic emission lines including H α , and Noterdaeme+10 detected in [O III].^cFirst published photometric measurement.^dNotes which automated searches returned these targets. York et al. (2012, Y), Straka et al. (2013, S), and Noterdaeme et al. (N).^eBorthakur et al. (2010).**Table 2.** Photometric data for QSOs.

Index	ID	SDSS quantities		Deconvolved apparent magnitudes					Reddening estimates		QSOALS ^c
		$m_{\text{PSF}(r)}$ ^a	Petrosian ^b	m_u	m_g	m_r	m_i	m_z	$\Delta(g-i)$	$E(B-V)_{(g-i)}$	
1	Q0045−0113	19.33	8.708	19.91	19.44	19.41	19.01	18.94	0.13	0.08	1
2	Q0808+0641	19.26	0.958	19.45	19.34	19.18	19.04	18.89	0.06	0.03	4
3	Q0811+2021	19.10	1.133	19.52	19.49	19.10	19.09	19.04	0.14	0.06	0
4	Q0902+1414	18.43	1.354	18.95	18.59	18.36	18.38	18.29	0.07	0.05	0
5	Q0914+3259	21.41	1.439	–	–	21.72	20.04	19.08	–	–	0
6	Q0929+3022	19.50	1.266	19.73	19.55	19.48	19.23	19.15	0.03	0.01	0
7	Q0950+5442	18.33	3.923	18.94	18.58	18.63	18.21	18.38	−0.14	−0.09	3
8	Q0952+0326	19.31	1.507	21.27	19.76	19.19	19.21	19.16	0.29	0.13	0
9	Q1012+1714	19.09	1.548	20.07	19.52	19.09	18.97	19.10	0.35	0.21	1
10	Q1013−0009	19.12	2.629	20.40	19.61	18.73	18.50	18.54	0.57	0.34	0

^aPSF fit.^bPetrosian radius measured by SDSS before deconvolution in arcseconds.^cThe number of intervening absorption systems found in the optical QSO spectra, excluding the system associated with the galaxy emission.

of these methods, please refer to these papers. In the process of completing the GOTOQ search, all observable properties of the galaxies and the absorption lines and extinctions were determined.

3.1 Broad-band flux measurements

To determine the photometric properties of the QSO and galaxy separate from one another, we have removed the QSO by PSF subtraction in order to avoid flux contamination to either. The PSF subtraction was performed with the Image Display Paradigm 3 (IDP3) software. Fig. 2 shows the SDSS tricolour composite thumbnail of each of our targets with the accompanying pre- and post-PSF subtracted SDSS data. The second panel shows contours of the QSO and galaxy, and the third panel shows the contours of just the galaxy after PSF subtraction of the QSO. The contours mark arbitrary flux levels in the QSO and galaxy in order to highlight their structure. Residuals at the position of the QSO are insignificant, being $< 1\sigma$ in the noise. The full figure can be found in the online version of the text. Fig. 3 shows SDSS tricolour composite thumbnails for those fields where we did not detect any galaxy after PSF subtraction of the QSO. Table 3 details the angular offsets of the detected galaxies from the QSOs along with the impact parameter b in kpc at the red-

shift derived from the galactic emission lines, assuming the adopted cosmology.

Reported apparent magnitudes are the asinh magnitudes as prescribed by Lupton, Gunn & Szalay (1999). Table 2 details the magnitudes calculated for the QSOs in all five filters ($ugriz$). Also presented in this table are r -band magnitude PSF model fits to the QSO from SDSS, which are fit and measured pre-deconvolution of the galaxy and QSO. Additionally, the Petrosian radius (\mathcal{R}_p) for each QSO is provided. The Petrosian radius is a measure of the size of an object, and values larger than about 1.4 arcsec indicate a profile that deviates from a point source. This could be an indicator of a galaxy's presence even if such a galaxy could not be discerned after QSO deconvolution. The galaxy might be too faint to discern in the photon noise of the wings of the QSO image (too small or too faint) or the object might have very low baryonic mass, but still have emission lines, with little stellar continuum (as in Rhode et al. 2013, for instance). Table 3 details the magnitudes and luminosities for the galaxies in all five filters as well as the galaxy ($u-r$) colour. All magnitudes for QSOs and galaxies have errors of less than 0.05 mag. Where no galaxy is detected, we report magnitude limits of the following: $m_u > 23.60$, $m_g > 24.45$, $m_r > 24.00$, $m_i > 23.50$, and $m_z > 23.50$.

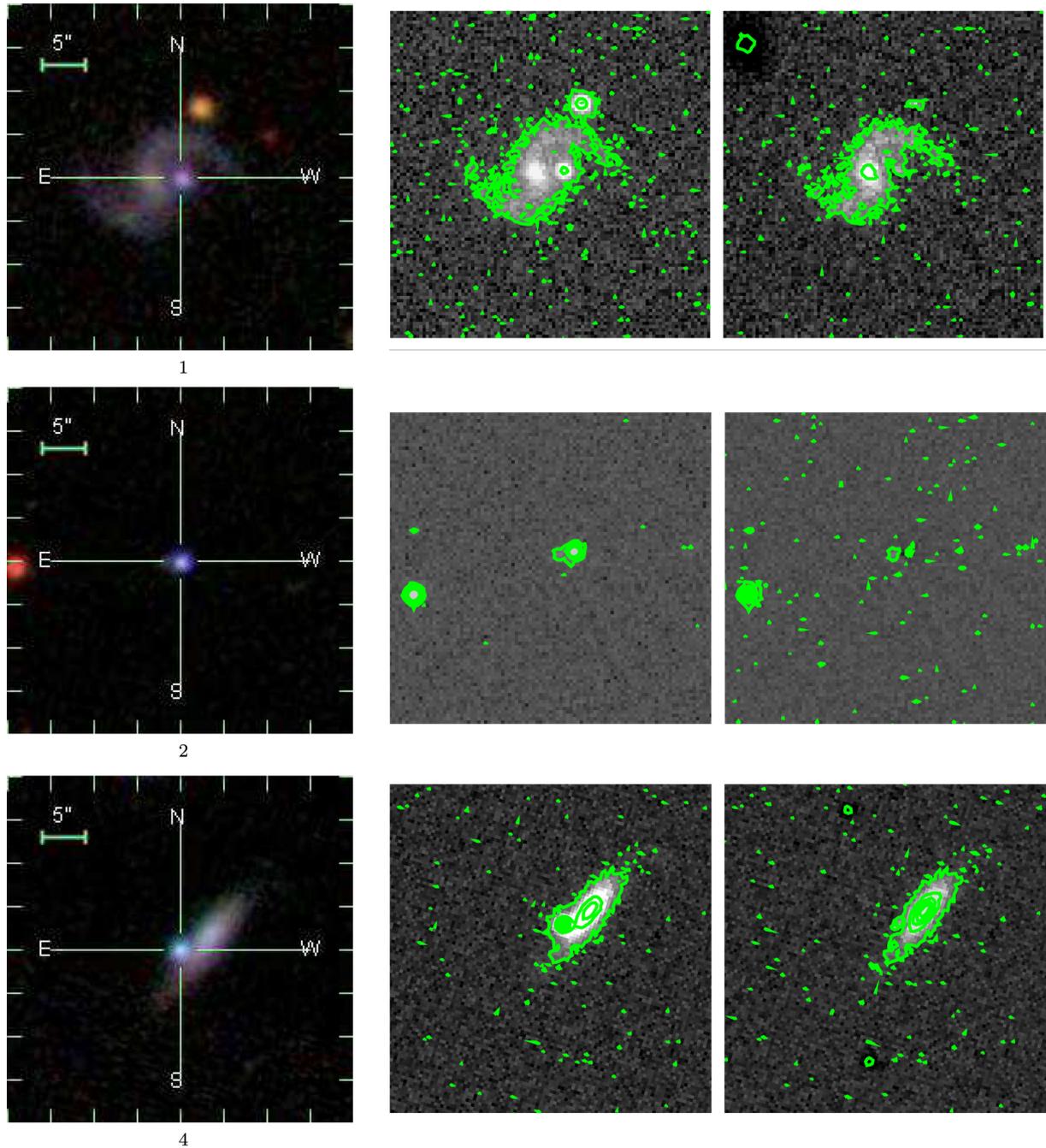


Figure 2. SDSS multicolour images of the 36 fields with a QSO intercepting a low- z galaxy detected in imaging, ordered by the increasing RA and increasing numerical index. The scale of the images is indicated in the upper-left-hand corner. The orientation is north up and east left. On the right, r -band images of the QSO and galaxy before (left) and after (right) PSF subtraction of the QSO. Contours mark arbitrary steps in flux in the image to mark the boundaries of the galaxy and QSO. The full figure may be seen on the web.

We have also determined the dust and reddening estimates for the QSOs and galaxies. The observer-frame colour excess ($\Delta(g-i)$) is calculated by taking the colour difference ($g-i$) for the QSO and comparing it to the median ($g-i$) for QSOs at the same redshift, taken from Schneider et al. (2007). This value allows us to estimate the reddening of the QSO due to dust in the intervening emission galaxy. Additionally, from this $\Delta(g-i)$ value, we have calculated the $E(B-V)_{(g-i)}$ value (the absorber rest-frame colour excess) according to the prescription of York et al. (2006). Several of the QSOs in our sample have multiple absorption line systems.

Therefore, reddening values cannot be attributed solely to the galaxy in question in these cases.

3.2 Emission line measurements

We have performed the same measurements as in Papers I and II of this series in order to compare our samples consistently. Therefore, parameters for each of the emission lines found in the spectra were measured with the `SPLIT` package within `IRAF`. Fig. 4 shows the SDSS spectra for each of our targets. Determined flux values can be

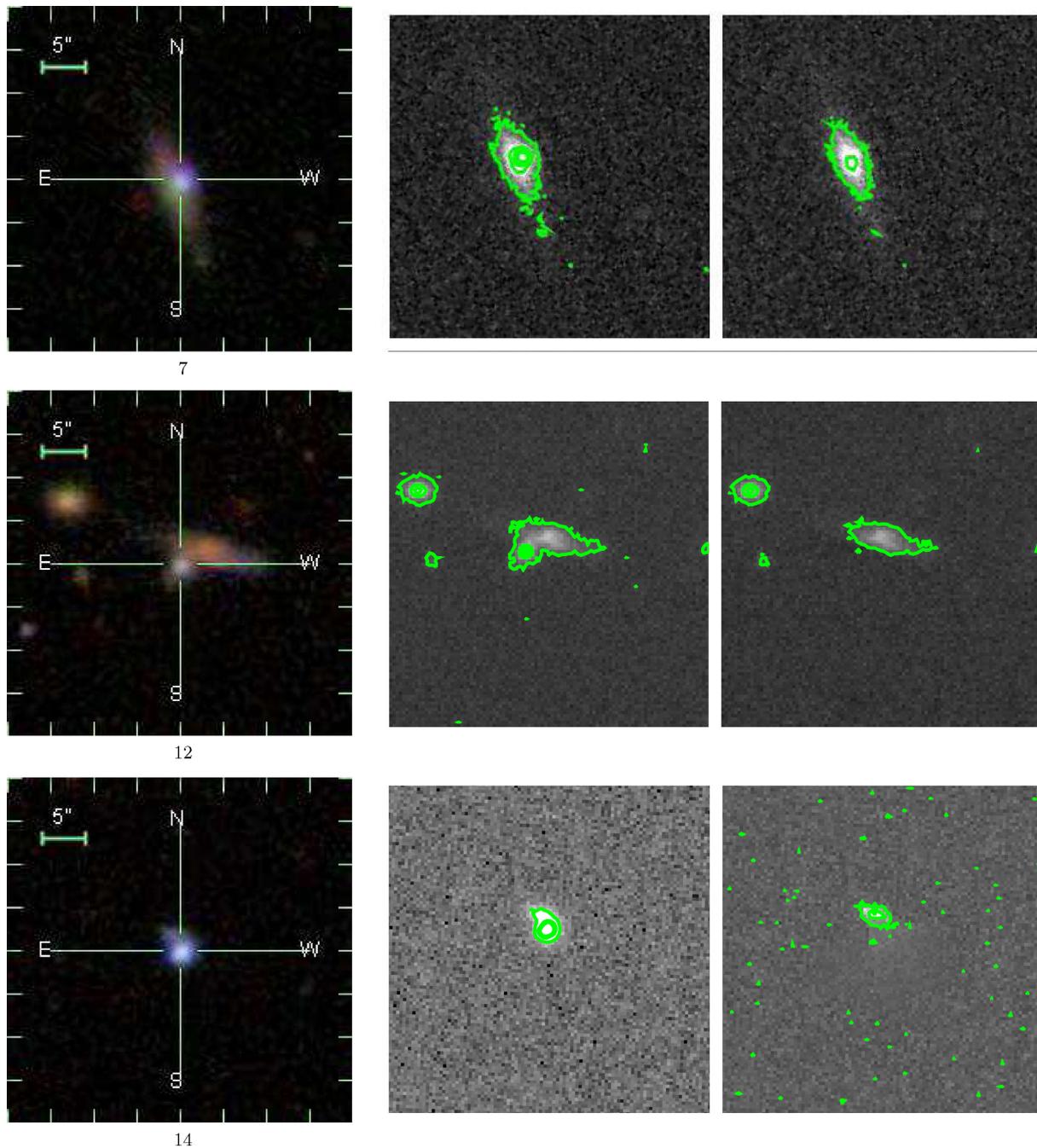


Figure 2 – continued

found in Table 5 measured in 10^{-17} ergs cm^{-2} s^{-1} . Wavelengths are measured in Angstroms. Error values are determined by measuring the standard deviation in a featureless region of the spectra near the emission line and multiplying this value by the square root of the number of pixels occupied by the emission line. These are available for the interested reader.

Integrated SFRs from $\text{H}\alpha$ and $[\text{O II}] \lambda 3727$ emission were calculated as prescribed in Kennicutt (1998). The $\text{H}\alpha/\text{H}\beta$ values allow us to calculate the Balmer decrement extinction correction to SFR, as well as the reddening values for the Small Magellanic Cloud extinction curve (see York et al. 2006), given by $E(B - V)_{\text{H}\alpha/\text{H}\beta}$ (see S13). These are given in Table 5 where we have reported the uncorrected SFRs in addition to the SFR values corrected for extinction in the

foreground galaxy via the Balmer decrement. We also correct for Milky Way extinction.

From the measured emission line fluxes, we can determine the emission line metallicity via several techniques, namely the indices R_{23} (Pagel et al. 1979), N2, and O3N2 (Pettini & Pagel 2004).

$$\log R_{23} = \log \left(\frac{[\text{O II}] \lambda 3727 + [\text{O III}] \lambda \lambda 4959, 5007}{\text{H}\beta} \right) \quad (1)$$

$$\text{N2} = \log \left(\frac{[\text{N II}] \lambda 6584}{\text{H}\alpha} \right) \quad (2)$$

$$\text{O3N2} = \log \left[\left(\frac{[\text{O III}] \lambda 5007}{\text{H}\beta} \right) \right] / \left[\left(\frac{[\text{N II}] \lambda 6584}{\text{H}\alpha} \right) \right]. \quad (3)$$

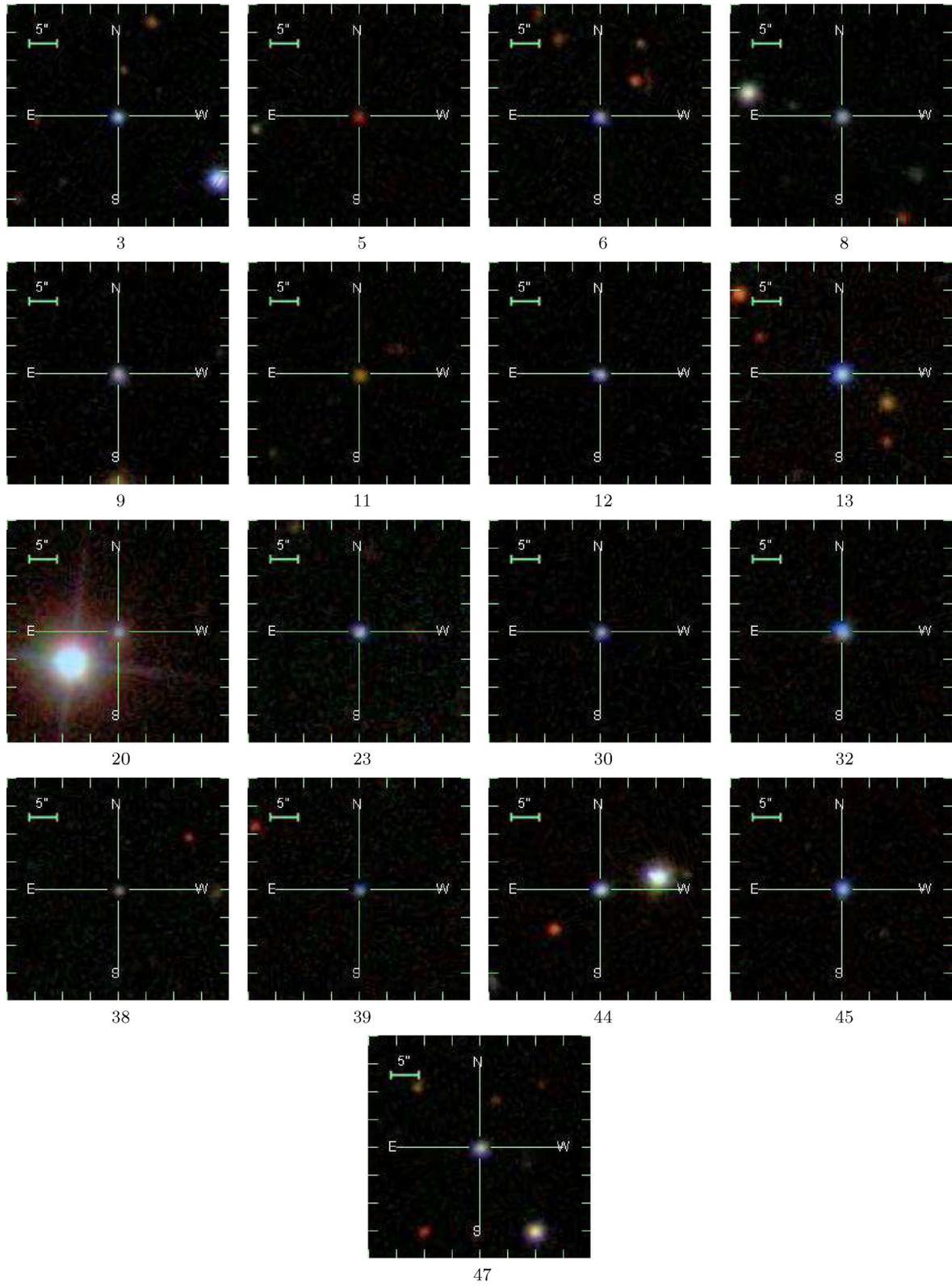


Figure 3. GOTOQs without visually confirmed galaxies. These are candidates for dwarf and dark galaxies.

Table 3. Photometric data for deconvolved galaxies.

Index	θ^a (arcsec)	b (kpc)	m_u	$\log L/L_{\odot u}$	m_g	m_r	$\log L/L_{\odot r}$	$L^*/L_{\odot r}$	m_i	$\log L/L_{\odot i}$	m_z	$(u-r)$	$\log M_*$
1	3.43	6.61	18.90	10.36	17.94	17.54	10.24	1.07	17.29	10.28	16.96	1.36	$9.62^{+0.78}_{-0.02}$
2	0.09	0.53	23.33	8.59	21.94	22.64	8.20	0.01	21.45	8.62	21.02	0.69	$9.94^{+0.01}_{-0.06}$
3	–	–	–	–	–	–	–	–	–	–	–	–	–
4	3.60	3.59	18.28	10.61	17.86	17.41	10.29	1.21	17.03	10.38	16.93	0.87	$9.65^{+0.03}_{-0.82}$
5	–	–	–	–	–	–	–	–	–	–	–	–	8.72^b
6	–	–	–	–	–	–	–	–	–	–	–	–	8.77^b
7	1.08	0.98	20.07	9.14	19.01	18.14	9.24	0.11	18.14	9.24	17.64	1.93	$8.80^{+0.15}_{0.06}$
8	–	–	–	–	–	–	–	–	–	–	–	–	9.0^b
9	–	–	–	–	–	–	–	–	–	–	–	–	7.48^b
10	3.30	6.03	20.08	9.89	18.70	18.17	9.99	0.60	17.62	10.15	17.17	1.91	$10.48^{+0.02}_{-0.13}$

^aDash means that no offset could be determined. Errors on magnitudes are within 0.05 mag.

^bWe use the relationship between $\log M_*$ and $\log \text{SFR}$ for our detected systems to extrapolate the $\log M_*$ values for those systems with no imaging detection. Errors for these measurements are within 1 dex.

From these values, we determine the analytic $12 + \log(\text{O}/\text{H})$ values prescribed by McGaugh (1991) and Pettini & Pagel (2004). Table 5 lists these values. A full discussion of these measurements can be found in the sections to follow. Limits for our emission measurements are available in the online table.

4 RESULTS

4.1 Photometry

Table 2 lists the PSF r -band magnitude for each QSO, as well as the Petrosian radius \mathcal{R}_p , which can be used as an indicator of the presence of a galaxy. Petrosian radii greater than ~ 1.4 arcsec are indicative of the presence of an extended object. Seeing can greatly affect the Petrosian radius, but with few exceptions the majority of our detections have $\mathcal{R}_p > 1.4$ arcsec. Interestingly, all but two of our fields with undetected galaxies also have $\mathcal{R}_p > 1.2$ arcsec. This indicates that the galaxies we seek are small and may fall within the PSF of the QSO. High-resolution instrumentation in combination with adaptive optics will be required to separate these galaxies from their QSO counterparts. This will be addressed below.

Fig. 5 shows a histogram of each of the samples, with a total sample histogram shown in solid black line. The total sample from all three searches (the total sample refers to all 103 GOTOs together) yields a mean $(u-r)$ value of 1.14 and a median of 1.26. Strateva et al. (2001) present the colour separation of galaxy types in SDSS data. In comparison to our colour histogram, fig. 7 of Strateva et al. (2001) displays the analogous histogram for $(u-r)$ values for their spectroscopic sample. The dividing colour value between morphological types according to Strateva et al. (2001) is $(u-r)=2.22$. Comparison with our own total sample shows that over 80 per cent of galaxies in our combined sample fall in the range of late-type galaxies. For the total sample, over 90 per cent fall in the range of late-type galaxies. Two of our targets from this work have $(u-r)>2.22$ (target numbers 28 and 38), indicating they may be good candidates for a green valley study of red spirals or blue ellipticals. One other target from S13 (Q0259+0001) has $(u-r)=2.64$, bringing the total to three in the green valley region.

The L_r^* values are listed in Table 3. Values $<0.2L_r^*$ are indicative of dwarf galaxies. By this definition, over half of our sample herein could potentially be dwarf galaxies. The connection with dwarf haloes and dark galaxies is discussed below.

4.1.1 Stellar mass estimates

Following the measurement of the five-band magnitudes for our systems, we have estimated the stellar mass of each detected galaxy using the photometric redshift code HYPERZ (Bolzonella, Miralles & Pello 2000). The code allows for fixed spectroscopic redshifts to be input (known from our selection criteria), constraining the spectral energy distribution (SED) model fit to the known redshift. The code then performs a reduced χ^2 analysis to obtain the best-fitting SED using stellar population synthesis models from Bruzual & Charlot (2003). Recent work by Taylor et al. (2015) shows that photometric redshift software calculations are not significantly improved by the inclusion of near-infrared data points beyond the initial five-band optical data points. In fact, they may decrease the robustness of the calculations. We therefore believe our stellar mass calculations from our five-band photometry are as robust as possible. Our estimates for M_* are in Table 3.

Fig. 6 shows the relationship we find between our HYPERZ stellar masses and the SFR measured within the SDSS spectral fibre (taken as a lower limit as the total galaxy area may not fall entirely within the fibre). All SFR are the $\text{H}\alpha$ Balmer decrement-corrected SFR where available, except in the $z > 0.4$ cases where we use the $[\text{O II}]$ SFR. We see good agreement with previous trends in $\log M_*$ – $\log \text{SFR}$ from Elbaz et al. (2007) and Tasca et al. (2014). Overplotted are the relationships at $z \sim 0$ and $z \sim 1$ from Elbaz et al. (2007). We note that our SFR are higher than those cited therein, which is attributed to our selection criteria. Since we have required strong $\text{H}\alpha$ or $[\text{O III}]$ in our sample, we naturally select stronger star forming regions than in Elbaz et al. (2007). We observe a median trend of

$$\log \text{SFR} = 2.12 + 0.95 \log \frac{M_*}{10^{11}}. \quad (4)$$

We use this equation to extrapolate the estimated stellar masses of those galaxies with no imaging detections in our sample. These are depicted in Fig. 6 as circles, also colour coded by redshift. These points show a clear increase in $\log M_*$ – $\log \text{SFR}$ with redshift. In order to illustrate the reason these galaxies may not have been detected, we also plot horizontal and vertical lines at the median boundary for each redshift bin. This makes it obvious that for all but the highest redshift non-detections, these galaxies lie below the median $\log M_*$ – $\log \text{SFR}$ for detected galaxies in the same sample. This is particularly evident in the $0.1 < z < 0.19$ bin where the orange circles are all significantly below the median. In this case, we are confident these galaxies were not detected due to their low

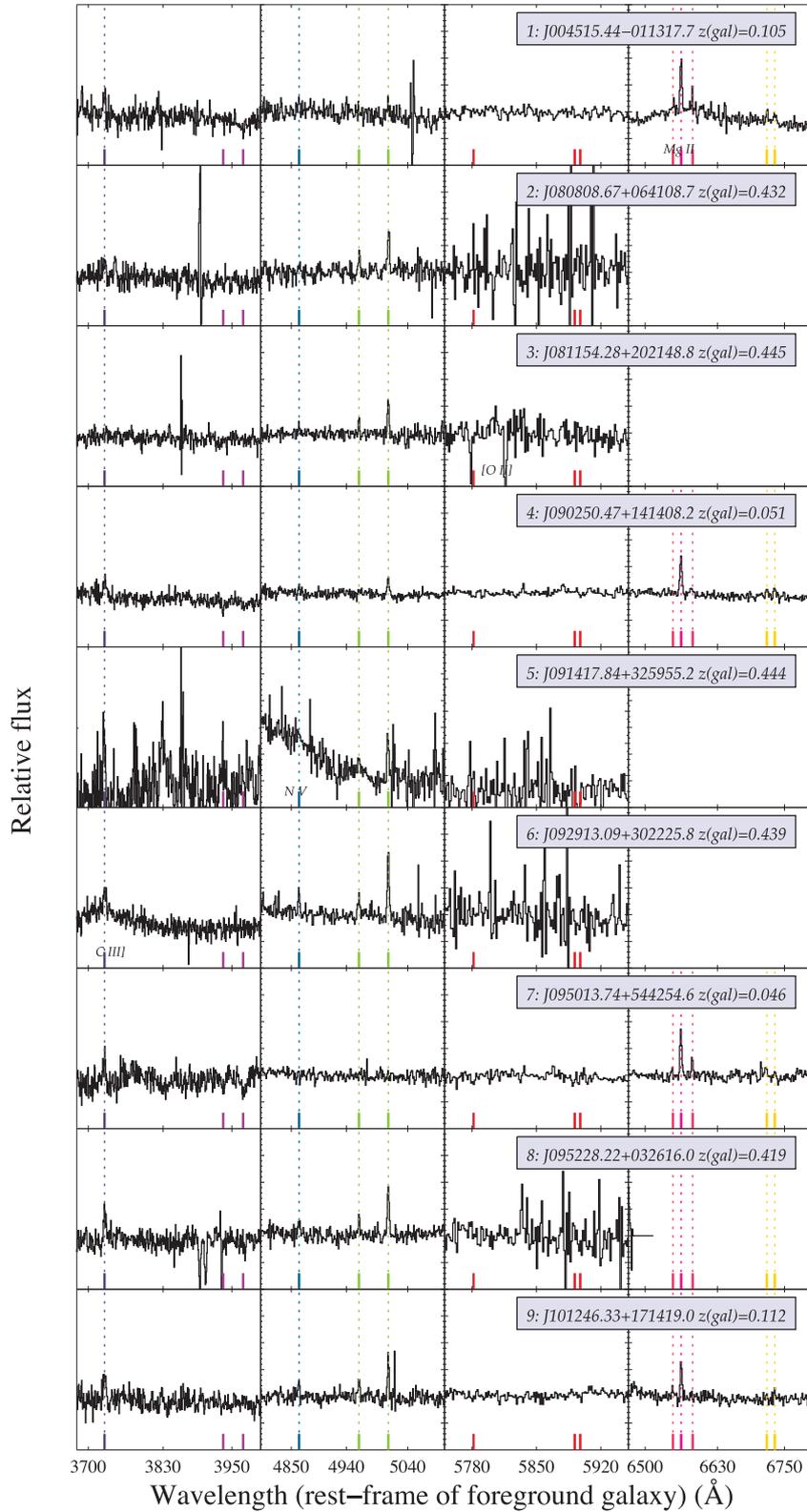


Figure 4. SDSS spectra of our targets. The coloured tick marks at the bottom of each spectra mark the expected positions of each emission and absorption feature. In order, these are [O II]; Ca II; H β ; [O III]a,b; the 5780 Å diffuse interstellar band (DIB) feature; Na I; [N II]a; H α ; [N II]b; and [S II]a,b. Broad emission lines labelled at the base of the spectra (such as Mg II in the first spectra) indicate QSO emission. Other absorption features not marked are from systems at other redshifts. The remaining spectra figures may be seen in the web version of this article.

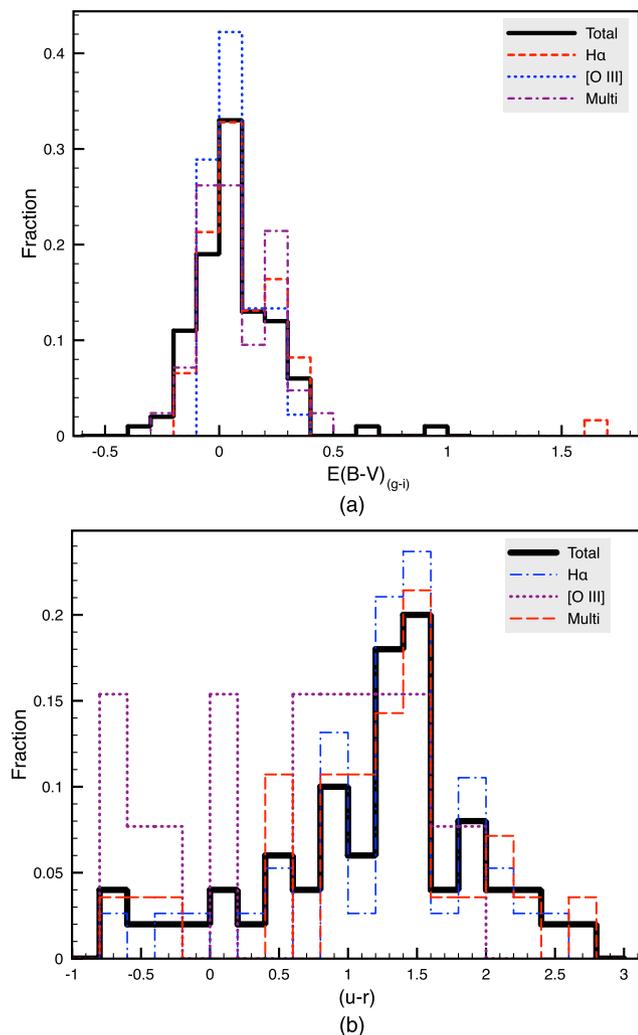


Figure 5. (a) Histogram of $E(B - V)_{(g-i)}$. (b) Histograms of galaxy colour for each of our individual samples and the combined sample. The red dash-dotted line is for the H α -selected sample. The purple dashed line is for the multi-emission selected sample. The blue dotted line indicates the histogram for the sample new to this paper. The black solid line is for the combination of all three samples.

mass and SFR. In the case of $0.3 < z < 0.39$, the non-detections are scattered about the median, making it more difficult to estimate the cause. As we look to increasing redshift, in the case of the $z > 0.5$ bin we see all of our non-detections are significantly above the median for detections in this sample. However, the small sample size of this bin limits the statistical analysis we can make.

4.2 Dust measurements along QSO sight lines

The total sample of GOTOQs has a mean $\Delta(g - i)$ of 0.11 and a median value of 0.07. From the $\Delta(g - i)$ values, we determine $E(B - V)_{(g-i)}$. We have assumed $R_V = 3.1$. The total sample has a mean $E(B - V)_{(g-i)}$ of 0.06 and a median of 0.05. We have included in Table 2 the number of additional intervening absorption systems detected in the SDSS spectra, excluding any absorption systems at the same redshift as the emission. Interestingly, the most reddened QSOs in our sample do not correspond to those QSOs with the most absorption systems, indicating the dust for these systems is very probably associated with the galaxy detected in emission.

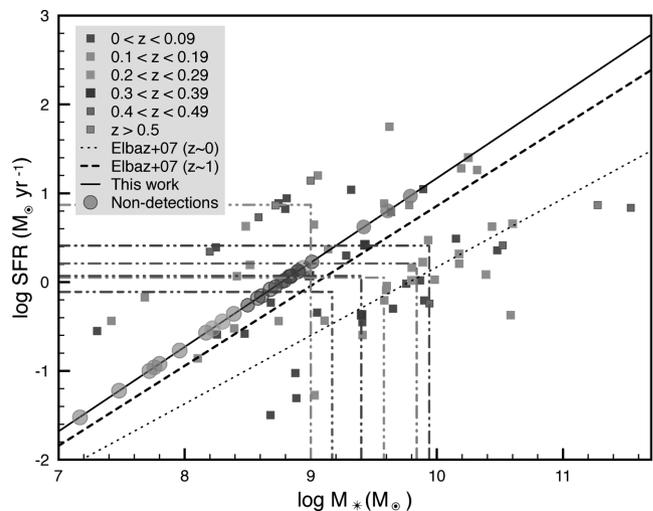


Figure 6. Stellar mass of our image-detected GOTOQs as a function of their H α SFR are shown as squares. We bin them in 0.1 redshift increments, shown by the different colours. The circles show our extrapolated stellar masses for those galaxies with no imaging detections. We see no real trend in this small sample with redshift for the galaxies with detections, but it is clear that those galaxies with no detections have increasing estimated masses with increasing redshift. We have plotted the observed trend of stellar mass with SFR from Elbaz et al. (2007). The $z \sim 1$ trend line has a better fit to our median trend line.

In fact, three of our most reddened QSOs from this work (IDs $(E(B - V))$: 10 (0.34), 35 (0.36), and 42 (0.37)) have no additional intervening absorption systems detected, while one other highly reddened system (49 (0.95)) has only one additional intervening system. Fig. 5 shows the histogram of extinction values of the combined sample. It can be seen that the centre of the histogram is offset from zero, indicating a higher reddening value for those QSOs with foreground galaxies on average. York et al. (2006) report an average $E(B - V)_{(g-i)} = 0.013$ for their Mg II selected sample of 809 QSOs at higher redshift. As indicated by York et al. (2006), values of $E(B - V)$ this high coming from higher redshift galaxies would produce easily detectable Mg II absorption lines, statistically. As in S13, we see a spread in the histogram of ~ 0.2 , likely caused by error introduced in the deconvolution of the QSO from the galaxy and the intrinsic QSO colour variations.

A Kolmogorov–Smirnov (KS)-test between the total sample (103 data points) and York et al. (2006) sample (809 data points) yields a discrepancy of 0.374 and a P -value of 0, indicating these two samples are not from the same parent population. This difference may be due to differences in galaxy type; our galaxies are thought to be primarily dusty late-type chosen by their emission signatures, while the Mg II absorbers from York et al. were chosen by absorption and their emission properties are unknown. A study of the SFR in the York et al. Mg II sample or a study in Mg II absorption (requiring UV observations) of our GOTOQs would help determine the differences in these two samples. We find our total sample to be much more reddened than the sample of York et al. (2006). If we limit the comparison to only those systems for which we detect Mg II absorption (as in York et al.), we find a discrepancy of 0.338 with a P -value of 8.601 per cent. However, our sample of Mg II absorbers contained herein is small (only 17 compared to the York et al. sample of 806) and may not give a very accurate result. It is also possible that some of our GOTOQs with galaxies at $z < 0.4$ have Mg II absorption, but it is outside the spectrograph range of SDSS.

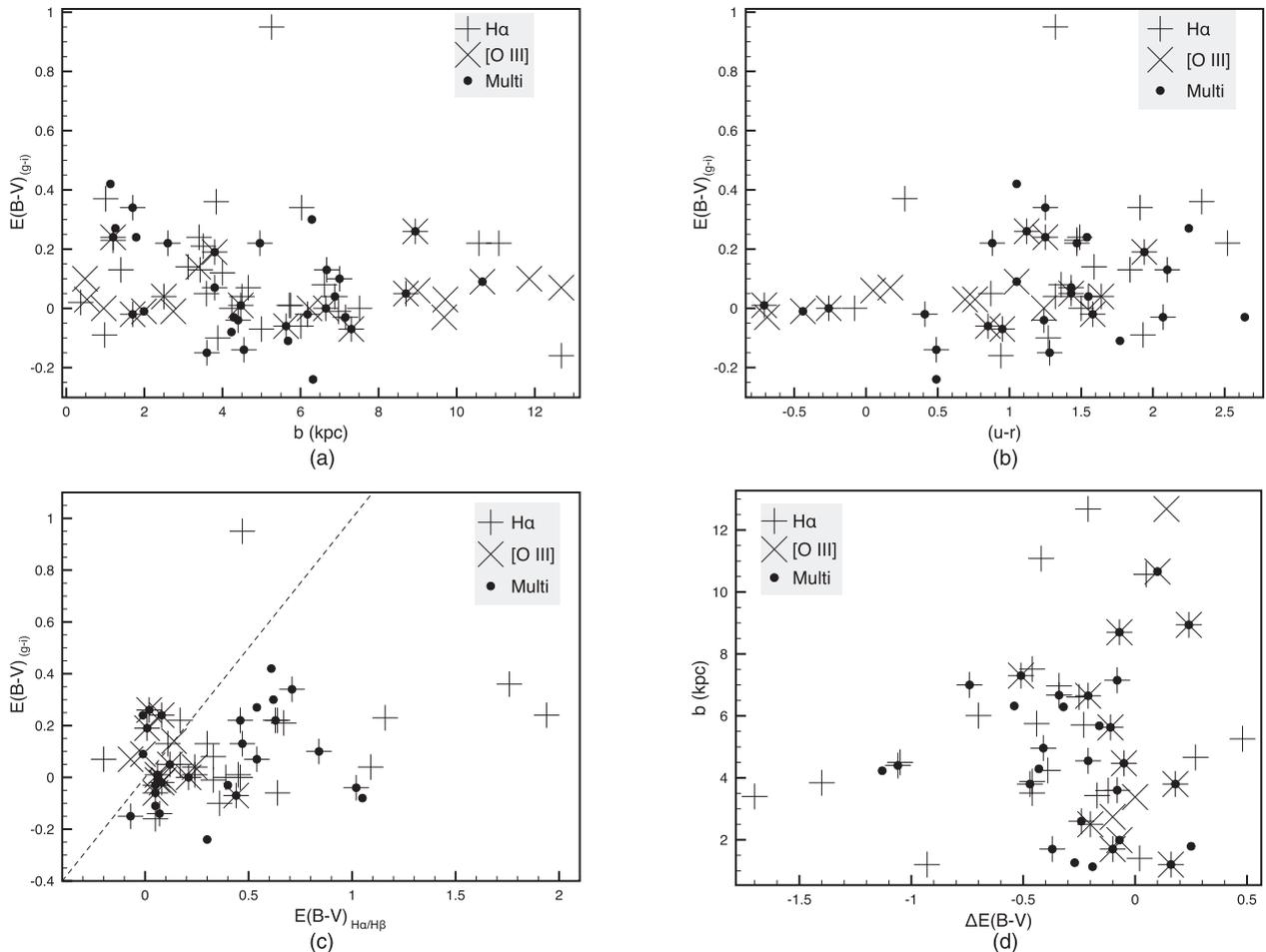


Figure 7. (a) $E(B - V)_{(g-i)}$ against impact parameter. (b) Reddening against galaxy colour. (c) Both reddening values plotted against one another. (d) The difference in our reddening values against impact parameter.

For the combined sample, we find an anticorrelation between impact parameter (b) and $E(B - V)_{(g-i)}$, with a Spearman value of $r_s = -0.27$ at the $\alpha = 0.03$ significance level, indicating higher dust content in regions closer to the centre of the galaxies. Fig. 7(a) shows the plot for these variables. For the total sample, we also find a significant correlation between $(u - r)$ and $E(B - V)_{(g-i)}$, with a Spearman value of $r_s = 0.32$ at the $\alpha = 0.02$ significance level. This is perhaps expected, given that systems with higher dust content will appear more red and those with lower extinction values will be unaffected in their colour, remaining blue. Fig. 7(b) plots our data for this relationship.

4.3 Galactic emission line measurements

For the redshift range $z < 0.4$, it is possible to detect $H\alpha$, $H\beta$, $[O II]$ $\lambda 3727$, $[O III]$ $\lambda \lambda 4960, 5008$, $[N II]$ $\lambda \lambda 6550, 6585$, and $[S II]$ $\lambda \lambda 6718, 6733$. Table 5 reports our emission line detections for each of our systems. Note that we do not apply any correction to these values for reddening or for the fact that the entire galaxy may not fall within the spectral fibre. Therefore, these flux values are likely lower limits. In this redshift range, all but one of our systems have detected $H\alpha$ (Q1329+6304, ID 32). This object is near the upper limit of the redshift range and has a noisy spectrum in the region of $H\alpha$. For five galaxies at $z < 0.4$ (target IDs 26, 27, 40, 41, and 42), we do not detect $[O III]$ $\lambda 5008$. However, each of these five have detected $H\alpha$ as a primary indicator. All targets with $z > 0.4$ (17 out of 53 targets)

have detected $[O II]$. Additionally, only targets 3, 10, 28, 32, 43, and 47 lack $[O II]$ detections.

Table 4 lists the star formation derived using $H\alpha$ and $[O II]$ luminosities. In the case of $z < 0.4$, we also detected $H\beta$ in 25 cases. In these cases, we derive the dust extinction by comparing the Balmer ratio with the expected value of 2.85. The dust-corrected SFRs are also given in Table 4.

From $H\alpha/H\beta$, we have calculated the $E(B - V)_{H\alpha/H\beta}$. We find a mean value for the total sample of 0.36 with a median of 0.24. In comparing $E(B - V)_{H\alpha/H\beta}$ to $E(B - V)_{(g-i)}$, we find a correlation of $r_s = 0.25$ with a significance level of $\alpha = 0.04$. As can be seen from Fig. 7(c), for most of our sample $E(B - V)_{H\alpha/H\beta}$ is higher, suggesting that on average the reddening is higher around $H II$ regions spread throughout the galaxy (the source of the integrated emission lines) than along the QSO line of sight. We also plot the difference between these two variables, $\Delta E(B - V)$, versus impact parameter, which can be found in Fig. 7(d). We find no significant relationship between them.

With few exceptions, the majority of our SFRs are below $5 M_{\odot} \text{ yr}^{-1}$. Kobulnicky et al. (2003) present the SFR values for their sample of field galaxies. At redshifts $0.28 < z < 0.4$, the average SFR is $1.06 M_{\odot} \text{ yr}^{-1}$, while the average SFR in the range $0.4 < z < 0.81$ is $6.01 M_{\odot} \text{ yr}^{-1}$. The average SFR for their entire sample of 64 field galaxies in the range $0.28 < z < 0.81$ is $4.74 M_{\odot} \text{ yr}^{-1}$. This value is significantly higher than our average of $2.44 M_{\odot} \text{ yr}^{-1}$ over the range of $0 < z < 0.4$, but includes a much

Table 4. Star formation rates.

Index	$E(B - V)_{H\alpha/H\beta}$	$SFR_{H\alpha}^a$ ($M_{\odot} \text{ yr}^{-1}$)	$SFR_{[O III]}^a$ ($M_{\odot} \text{ yr}^{-1}$)	$SFR_{H\alpha}^b$ ($M_{\odot} \text{ yr}^{-1}$)	$SFR_{[O III]}^b$ ($M_{\odot} \text{ yr}^{-1}$)
1	0.33	0.08	0.17	0.18	0.67
2	–	–	0.58	–	–
3	–	–	–	–	–
4	0.17	0.02	0.04	0.03	0.07
5	–	–	0.89	–	–
6	–	–	1.01	–	–
7	–	–	–	–	–
8	–	–	1.69	–	–
9	–0.10	0.06	0.24	0.04	0.16
10	–	0.07	–	–	–

^aSFR uncorrected for extinction.^bSFR corrected for extinction using $H\alpha/H\beta$.**Table 5.** Emission line strengths for galaxies in front of QSOs and emission line metallicities.

ID	$f_{H\alpha}^a$	$f_{H\beta}$	$f_{[O III]}$	$f_{[O III]a}$	$f_{[O III]b}$	$f_{[N II]a}$	$f_{[N II]b}$	$f_{[S II]a}$	$f_{[S II]b}$	R_{23l}	R_{23u}	$12 + \log(O/H)$ O3N2	N2
1	37.73	9.74	43.49	– ^b	4.02	15.68	6.51	8.56	14.79	–	–	8.86	8.91
2	–	13.97	6.04	14.82	20.54	–	–	–	–	7.27	8.91	–	–
3	–	5.28	–	8.67	20.65	–	–	–	–	–	–	–	–
4	50.32	15.01	43.02	–	25.03	–	18.83	–	–	–	–	8.67	8.91
5	–	–	8.75	3.33	15.99	–	–	–	–	–	–	–	–
6	–	10.69	10.16	11.39	29.20	–	–	–	–	7.63	8.77	–	–
7	–	–	–	–	–	–	–	–	–	–	–	–	–
8	–	8.63	19.01	6.53	22.23	–	–	–	–	7.89	8.77	–	–
9	21.68	8.24	53.96	20.28	22.74	4.94	–	–	6.16	8.59	8.14	–	–
10	38.10	–	–	5.91	15.68	7.96	11.44	20.39	24.52	–	–	–	8.92

^aFlux measured in $10^{-17} \text{ ergs s}^{-1} \text{ cm}^{-2}$.^bIndicates no emission line detected or emission line is redshifted out of range of the spectrograph.

higher redshift range where the SFR is clearly much higher. Their average in the range $0.28 < z < 0.4$ is comparable to our average, and therefore we suggest that our sample of galaxies is comparable to other samples of field galaxies. Our sample is comparable in SFR to other studies of field galaxies as well (Lilly, Carollo & Stockton 2003; Mouhcine et al. 2006). We require near-infrared spectroscopy on those galaxies at $z > 0.4$ to obtain $H\alpha$ SFR values for a more direct comparison between these two samples.

Table 5 lists our emission line metallicity values. Emission line metallicities were determined using three indices, including R_{23} (Pagel et al. 1979; Kobulnicky, Kennicutt & Pizagno 1999; Gharanfoli et al. 2007), N2, and O3N2 (Pettini & Pagel 2004; Gharanfoli et al. 2007). Where there exists data to determine the ratio of $[O III] \lambda 5008/[N II] \lambda 6585$, we can eliminate the degeneracy in the R_{23} index. Where data does not exist for this ratio, we list the lower and upper branch of $12 + \log(O/H)$ for R_{23} . Following the prescriptions of McGaugh (1991) and Kobulnicky et al. (1999), we have calculated the values for $12 + \log(O/H)$ for N2 and O3N2 using analytic expressions. We adopted the solar value for $12 + \log(O/H)_{\odot}$ to be 8.69 (Asplund et al. 2009). Our values show that our galaxies are generally below the solar metallicity as measured by R_{23} , but generally above solar as measured by O3N2 and N2. We note that our emission line measurements are lower limits, given the fact that the entire galaxy likely does not fall within the SDSS spectral fibre and therefore the fibre does not collect all of the galaxy emission. In comparison with the sample from N10, we see that our values for the lower and upper branches of R_{23} are comparable, as expected.

In the sample from Kobulnicky et al. (2003), discussed above ($0.28 < z < 0.81$), the average emission line metallicity is found to

be $12 + \log(O/H)$ of 8.64. Lilly et al. (2003) present a sample of 66 galaxies in the redshift range $0.47 < z < 0.92$. In general, we find our R_{23} upper branch values to be significantly below those found in Lilly et al. (2003), but our lower branch values are above theirs. This suggests our sample has a roughly comparable metallicity as theirs, and therefore our galaxy sample is comparable to galaxies at higher redshifts.

Fig. 8 shows the data points from our sample of 103 GOTOQs (which have detected emission in $H\alpha$, $H\beta$, $[O III]$, and $[N II]$) on a BPT (Baldwin–Phillips–Terlevich; Baldwin, Phillips & Terlevich 1981) diagram. The BPT diagram indicates the ionization mechanism of nebular gas in the region observed. The dashed curve represents the division between starburst galaxies and active galactic nuclei (AGN) from Kewley et al. (2001). The dotted line is the same division, but updated by Kauffmann et al. (2003). Additionally, it is accepted that the general definition of Seyfert galaxies is $\log([O III]/H\beta) > 0.48$ and $\log([N II]/H\alpha) > -0.22$, and the definition for LINERs (low-ionization nuclear emission regions) is $\log([O III]/H\beta) < 0.48$ and $\log([N II]/H\alpha) > -0.22$ (Osterbrock 1989; Kewley et al. 2001; Kauffmann et al. 2003; Kewley et al. 2013). The horizontal and vertical lines represent these boundaries.

According to these parameters, all of the galaxies apart from two in our sample with these four emission lines fall within the normal $H II$, indicating gas is ionized here through normal star formation processes. The two points falling slightly within the AGN region are within the error bars to be in the $H II$ region as we would expect. Further inspection of their emission lines shows no indication of broadening typical of AGN emission. Follow-up spectroscopy will further constrain the positions of these galaxies on the plot and

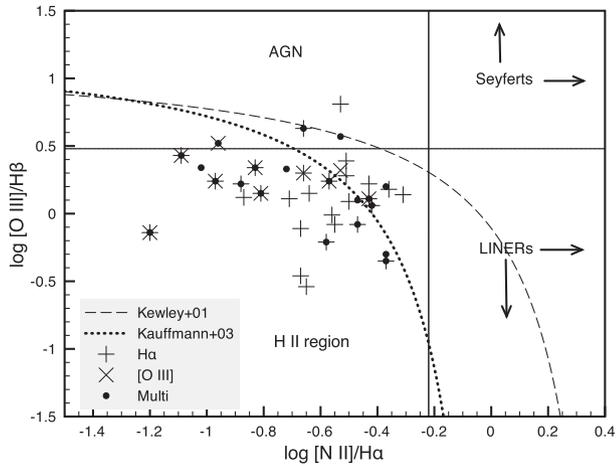


Figure 8. BPT diagram (Baldwin et al. 1981) of all three samples in this series. The figure shows that our systems are nearly all normal H II regions, with a few that may be candidates for normal AGN. The dotted line denotes the model of Kauffmann et al. (2003), while the dashed line is the model of Kewley et al. (2001) for the division between starburst galaxies and AGN. The cross-hairs mark the division between Seyferts and LINERs.

perhaps add new points to the plot through detections of emission lines (particularly [N II]) which are currently undetected.

4.4 Absorption line measurements

Measurements of the galaxy absorption features for Ca II K, Ca II H, Na I D2, Na I D1, the Mg II doublet, Mg I, and Fe II found in the QSO spectra are reported in Table 6 for this work. This table also lists other absorption features detected in the optical in the footnotes, including Mn II, Cr I, Fe II, Ti II, Fe II, and CH, where available. EW errors are in the range 0.3–0.6 Å. The extension of our search to redshifts up to $z \sim 0.8$ allows us to detect these absorption features in some systems in this work, whereas in Y12 and S13 they are not within the range of the spectrograph due to redshift constraints. Additionally, the relative velocity shift between the emission and absorption lines is listed. As in N10, we find the Δv values to be quite small (all but three $< 100 \text{ km s}^{-1}$), which is consistent with the

circular velocities of typical galaxies and suggests the absorbing gas is bound to the emission line galaxy.

Fig. 9 shows our data points for the rest EW of Ca II K against impact parameter. We find no significant correlation between these two variables. Similarly, we find no significant correlation between the rest EW of Ca II K and $E(B - V)_{(g-i)}$, displayed in Fig. 9.

The reader should refer to N10 for a full treatment of the Mg II properties of this sample. Fe II $\lambda\lambda 2586, 2600$ can be seen in the SDSS spectrum at $z > 0.47$. Of our 53 targets, seven are in this redshift range. Of these, six have detected Fe II absorption (17, 23, 25, 45, 47, and 51). Similarly, Mg I $\lambda 2852$ can only be seen at $z > 0.33$; 20 of our 53 targets are in this range; however, only 12 have detected Mg I absorption (2, 3, 6, 8, 22, 23, 24, 29, 45, 47, 50, and 51).

Those systems with all three detected Mg II, Mg I, and Fe II with $\text{EW} > 1 \text{ \AA}$ fit the criteria set by Rao, Turnshek & Nestor (2006) for selecting Damped Lyman α (DLA) systems. From our sample, targets number 23, 25, 45, 47, and 51 all fit this description and will require follow-up UV observations to determine if they are indeed DLA systems. Target number 17 may also be a good candidate, with only Mg I below the threshold. N10 note that it is possible based on the absorption properties, all 17 of the Mg II detected systems could be candidates for DLA systems and 21-cm absorption. However, based on measurements of strong Mg II and Fe II in Meiring et al. (2007, 2008, 2009) with Ly α measurement well below the DLA limit, it is unlikely this trend suggested in Rao et al. (2006) will be supported. We note, however, that virtually none of the QSOs presented here have sufficient radio fluxes for 21-cm searches with current telescope technology. Additional absorption features found in these systems are noted in the footnotes of Table 6.

In this work, we find 35 of our 53 targets have some form of absorption. 10 of these 35 have no detectable galaxy in imaging after PSF subtraction (IDs 3, 6, 8, 13, 20, 23, 32, 34, 39, and 47). These targets may be good candidates for dark galaxies (Matsuoka et al. 2012; Rhode et al. 2013). This leaves 18 targets with no detectable absorption, only seven of which have no visible galaxy in imaging (IDs 5, 9, 11, 12, 30, 38, and 44). Comparing the images with detections with the absorption detections, it seems absorption occurs more frequently in the discs of galaxies, while when the impact of the QSO is higher (more considerably in the halo region), absorption is likely to be missed.

Table 6. Rest-frame EW (W) for absorption interstellar lines in foreground galaxies.

Index	z_{abs}	Δv (km s^{-1})	Ca II K ($W_{\lambda 3933}$)	Ca II H ($W_{\lambda 3969}$)	Na I D2 ($W_{\lambda 5890}$)	Na I D1 ($W_{\lambda 5896}$)	Mg II ($W_{\lambda 2796}$)	Mg II ($W_{\lambda 2804}$)	Mg I ($W_{\lambda 2852}$)	Fe II ($W_{\lambda 2587}$)	Fe II ($W_{\lambda 2600}$)
1	0.1059	115.2	<0.39	<0.63	0.36	0.33	–	–	–	–	–
2	0.4327	72.9	<0.63	<0.65	<1.53	<1.53	2.42	1.94	0.80	–	–
3	0.4448	32.2	<0.42	<0.34	<1.05	<1.05	0.98	1.05	0.30	–	–
4	0.0504	–18.4	0.80	0.39	<0.35	<0.35	–	–	–	–	–
5	–	–	<4.23	<3.58	<3.38	<3.38	–	–	–	–	–
6 ^a	0.4391	–58.1	<0.76	<0.67	1.94 ^f	1.94 ^f	1.99	1.49	1.47	–	–
7	0.0461	0.00	1.6	1.8	–	–	–	–	–	–	–
8	–	–	<0.56	<0.45	<1.54	<1.54	2.00	0.60	2.50	–	–
9	–	–	<0.50	<0.60	<0.40	<0.40	–	–	–	–	–
10	–	–	<0.78	<0.78	<0.55	<0.55	–	–	–	–	–

^aAlso: CH, $W_{\lambda 4234} = 1.19$

^bAlso: Mn II, $W_{\lambda 2594.50} = 0.77$; Cr I, $W_{\lambda 3579.70} = 0.68$; Cr I, $W_{\lambda 3594.50} = 0.63$

^cAlso: Ti II, $W_{\lambda 3242.90} = 1.03$

^dAlso: Ti II, $W_{\lambda 3242.90} = 0.47$

^eAlso: Mn II, $W_{\lambda 2606.50} = 0.32$; Ti II, $W_{\lambda 3384.70} = 0.58$; Fe II, $W_{\lambda 2344.20} = 1.81$; Fe II, $W_{\lambda 2374.50} = 1.55$; Fe II, $W_{\lambda 2382.80} = 2.09$

^fThese absorption features are blended and measured together.

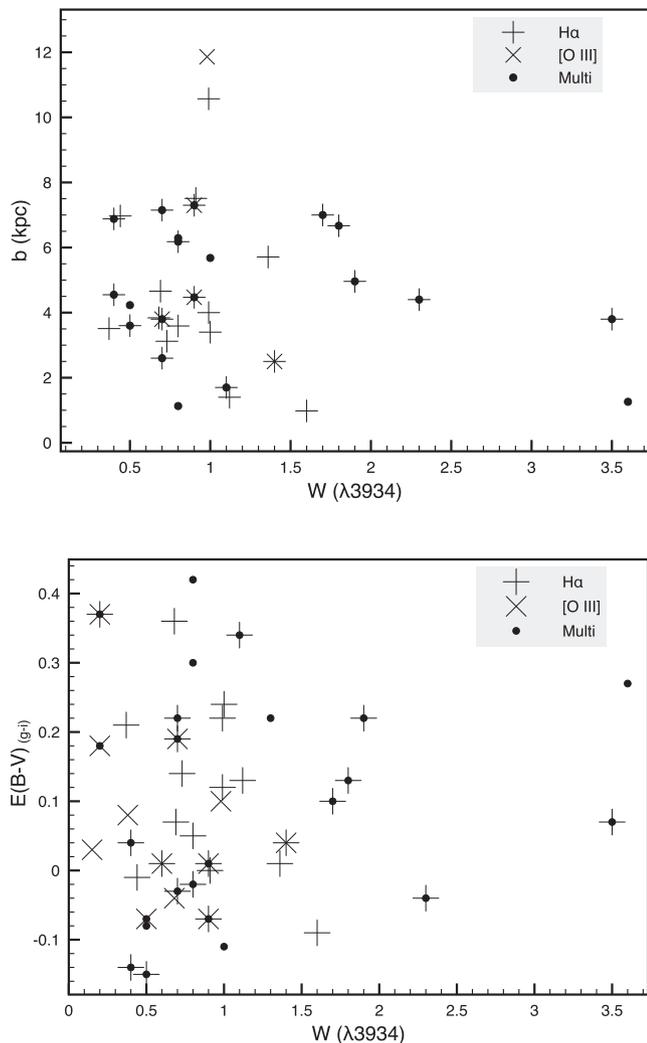


Figure 9. Ca II K rest EW versus impact parameter and Ca II K rest EW versus $E(B - V)_{(g-i)}$.

4.4.1 Number density of absorption lines

We estimate the number density of Ca II, Na I, Mg II, and Mg I by calculating the relative redshift coverage for each of these absorbers in our sample.

For Ca II, we find a ΔZ of 8.58, leading to a $dN/dz \sim 0.20$. Though we only have 43 Ca II absorbers in our sample, this is higher than dN/dz from Richter et al. (2011), who from 18 strong ($\log N_{\text{Ca II}} > 11.65$) Ca II absorbers find a dN/dz of ~ 0.117 (after bias correction). We have not selected our systems based on any prior knowledge of their absorption properties and these numbers indicate we have not selected an overabundance of Ca II absorbers in our sample.

Through similar calculation, we find dN/dz of Mg II for our sample of 18 absorbers to be ~ 0.50 . Again, though limited, this is in agreement with recent studies of Mg II dN/dz such as Evans et al. (2013, 0.449 ± 0.003 , $0.40 < z < 1.14$) and Nestor, Turnshek & Rao (2005, 0.60, $z \sim 0.3$) for strong Mg II absorbers ($W_0^{2796} > 1.0$). We cover a redshift range of $0.36 < z < 0.79$ with our Mg II absorbers. Mg II absorbers are known to probe a wide array of galactic environments, including tracing outflows, out to impact parameters of 250 kpc (e.g. Churchill et al. 1999, 2012; Lundgren et al. 2012). It is therefore no surprise that galaxies known to intervene with background QSOs show strong Mg II absorption at

100 per cent completion. We refer the reader to N10 for a more complete assessment of the Mg II properties in this sample.

4.4.2 Correlations between absorption line EW

For our total sample of 103 GOTOQs, we look at the relationship between the strengths of the various absorption lines. In all cases, more data points are needed to effectively constrain the relationships. However, we present here the preliminary analysis. We see only two possible correlations: $W_0^{\lambda 3933}$ (Ca II K) with $W_0^{\lambda 2852}$ (Mg I) and $W_0^{\lambda 2796}$ (Mg II) with $W_0^{\lambda 2600}$ (Fe II).

Five systems show both Ca II K and Mg I 2852 (36 per cent of Mg I have Ca II). This relationship has a Spearman correlation value of 0.70 with an alpha significance of 0.19. However, this is at best a weak correlation and requires many more points for validation. This weak correlation has a relationship of $W_0^{\lambda 2852} = 2.31 W_0^{\lambda 3933} - 0.26$.

Six of our GOTOQs present with both Mg II and Fe II (100 per cent of Fe II have Mg II). Even with only six points, we see a preliminary tight correlation between Mg II 2796 and Fe II 2600 EW. This relationship has a Spearman correlation value of 0.77, corresponding to an alpha significance of 0.07. We note that both of these relationships are hindered by our limited low-redshift coverage, since Mg II, Mg I, and Fe II are only obtainable in the SDSS spectra at $z > 0.39$, $z > 0.37$, and $z > 0.5$, respectively. This correlation has a relationship of $W_0^{\lambda 2600} = 0.6 W_0^{\lambda 2796} + 0.58$.

We calculate the relationships between each of the other absorption lines, but find no significant correlation between them. The most significant are the following.

(i) 14 GOTOQ systems have both Ca II K and Na I D2 (33 per cent of Ca II absorbers have Na I). This results in a Spearman value of 0.25, with alpha significance of 0.46 between the two rest-frame EW. There is considerable scatter between the two.

(ii) Five systems with both Mg I 2852 and Fe II 2600 (36 per cent of Mg I absorbers have Fe II).

(iii) 14 systems with both Mg II 2796 and Mg I 2852 (78 per cent of Mg II absorbers have Mg I).

(iv) Six systems with both Ca II K and Mg II 2796, which is 33 per cent of Mg II absorbers have Ca II K. This is much higher than the incidence rate found by Sardane, Turnshek & Rao (2014, 3 per cent of Mg II absorbers have Ca II) and we find no significant correlation between the two in our sample. Sardane et al. (2014) detect a positive correlation between the strengths of these two absorbers, albeit with a large spread. Our sample, however, is too small for any statistical comparison. We note that our six systems fall well within the trend of fig. 14 of Sardane et al. (2014). The higher incidence rate in our sample could be due to our bias towards very short offsets from the galaxy centre, where, in these inner regions, we might expect higher Ca II abundance relative to Mg II and less Ca II ionization.

Our sample has too few systems with both Na I D2 and Mg II, Ca II K and Fe II, Na I D2 and Mg I, and Na I D2 and Fe II to make any sensible statistical analysis regarding these absorption relationships.

4.5 Correlations between emission and absorption strength

We test to see if there are any correlations between emission and absorption strength in our sample. We compare our strongest emission lines (H α , [O II] $\lambda 3727$, and [O III] $\lambda 5008$) with the strongest absorbers within redshift range (Ca II K and Mg II $\lambda 2796$). We are not able to determine the relationship between H α and Mg II due to

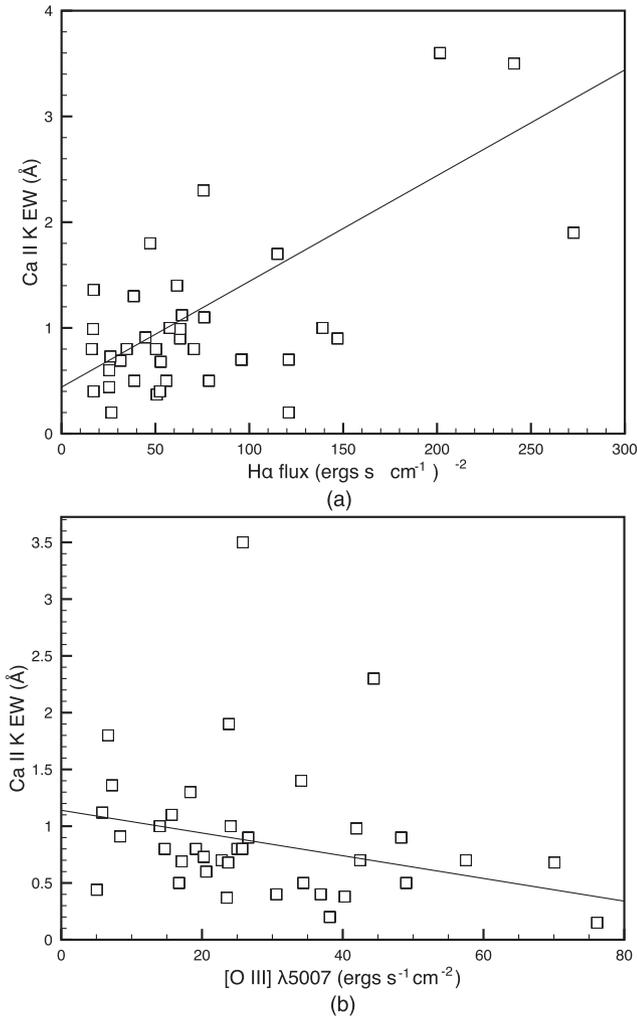


Figure 10. (a) Ca II K rest EW versus $H\alpha$ flux. We observe a positive correlation between these two parameters. (b). Ca II K rest EW versus [O III] $\lambda 5008$. We also observe an anticorrelation between these two parameters. Both the correlations return a significance of 0.03. The plotted lines show our best fit to each of the sets of data. Please see the text for more details.

redshift restrictions. We find no correlation between [O II] and Mg II or [O III] and Mg II for our limited sample. We find a positive correlation at the 0.03 significance level for $H\alpha$ and Ca II K (Spearman rho of 0.35), and also an anticorrelation between [O III] and Ca II K at the 0.03 significance level (Spearman rho of -0.30). The best fit to these relationships is as follows:

$$W_0^{\lambda 3933} = 0.01 F_{H\alpha} + 0.44 \quad (5)$$

and

$$W_0^{\lambda 3933} = -0.01 F_{[OIII]} + 1.14. \quad (6)$$

Fig. 10 plots our data for these relationships, with the best-fitting trend lines overplotted. Strong $H\alpha$ is an indicator of active star formation, as shown by the SFRs measured for this sample. Likewise, Ca II is a tracer of dusty environments, such as star-forming regions, indicating active star formation in dusty regions will have both strong nebular emission and Ca II absorption in related ratios. The opposite trend for [O III] flux may be explained by [O III] not being a direct indicator of SFR, but instead its dependence on n_e and T_e .

4.6 Dwarf galaxies and dark haloes?

As previously suggested in S13, there is evidence that a small number of our galaxies may be candidates for dark galaxies and a significant fraction of the sample are low-mass dwarf galaxies. Recent studies by Matsuoka et al. (2012) show that images as deep as $m_R = 30$ do not detect the source which is initially detected in H I 21-cm data. They additionally do not detect any nebular emission lines from this dark region, though the optically detected component of this galaxy to the northeast (NE) does have strong emission detected. Additional data on the 21-cm emission of these continuum sources from the Very Large Array and Arecibo Legacy Fast ALFA (ALFALFA) survey show neutral hydrogen masses as low as $10^5 M_\odot$. However, in our sample we detect strong nebular emission for all of our targets based on the selection criteria. The findings of Matsuoka et al. (2012) indicate then that our galaxies with no imaging detection are not candidates for dark galaxies, but instead typical of low-mass dwarf galaxies.

The SFR for Leo P (Rhode et al. 2013) is $\log(\text{SFR}) = -4.27$, which is much lower than anything we found in our sample. However, Leo P lies above the L_B -SFR relation for dark galaxies and therefore, while it is not a candidate for a dark galaxy, may be indicative of the type of objects we looking at in our sample of non-detected galaxies.

The non-detection of these objects is most probably in large part due to their redshift. The mean and median redshift of our detected galaxies is 0.18 and 0.13, respectively. However, the mean and median redshift of our undetected galaxies is much higher: 0.31 and 0.34, respectively. The higher- z nature of these low-mass objects naturally makes them more difficult to detect in imaging. These objects in particular would not have otherwise been found in traditional galaxy surveys, making them unique to our search technique and sample. If we compare galaxies in our sample with similar SFR, we can see there is a strong dependence on redshift whether these galaxies are detected or not, implying if we moved a detected galaxy to the higher redshift of a non-detected galaxy with similar SFR, we would not detect this galaxy either. We have also compared the $\log M_*$ values for these non-detected galaxies with our detections at similar redshift. We find that, in addition to having higher than average redshift, they also have lower than average $\log M_*$.

Currently, the number of observed dwarf galaxies in the universe falls far short of what is predicted by models (e.g. Klypin et al. 1999; Moore, Ghigna & Governato 1999). Additionally, the suppression of star formation in dwarf galaxies is not well understood. These data then show that it is possible to have a gas-rich dwarf galaxy, even at masses as low as $10^5 M_\odot$, and still detect strong emission and star formation. These qualitative characteristics are comparable to what we have found for our sample of galaxies, strongly suggesting that those fields with undetected galaxies in imaging within our sample fall into the category of low-mass dwarf galaxies.

It has been suggested in the past that H_2 density correlates much more closely with star formation than total gas density of a system (Bigiel et al. 2008; Leroy et al. 2008). Kuhlen et al. (2012) and (Kuhlen, Madau & Krumholz 2013) provide model evidence that the suppression of star formation in dwarf and dark galaxies is better controlled by H_2 gas and results in ~ 80 per cent dark haloes with $M_h < 10^9 M_\odot$ and $\text{SFR} < 1 M_\odot \text{ yr}^{-1}$. Thus, it would be interesting to obtain 21-cm and CO data for any of our low-mass dwarf galaxy candidates to see how they compare with these recent findings.

5 CONCLUSION

Here, we report 103 galaxies intervening with background QSOs detected in imaging, spectral emission, and absorption. The advantage of this sample is that it is entirely random with respect to the foreground galaxies. The standard procedure of detection of the some $>200\,000$ QSO absorption line systems known (Lundgren et al., in preparation; York et al., in preparation) is to note the absorbers in QSO spectra, then look for foreground galaxies. The sample here represents over 40 per cent of confirmed absorber host galaxies, i.e. those galaxies known to have QSOs probing hypothetical haloes of <200 kpc in size (even though we find the impact parameters in our survey to all be within <12 kpc). There is no particular reason to believe that QSOs with such close impact parameters would not probe the more extended regions of any haloes associated with these galaxies along the line of sight. The discrepancy in numbers is striking, but possibly just caused by the fact that galaxies as faint as those in our sample (103) have not been reached in the limited searches to date (Lundgren et al. 2012).

Follow-up of galaxies in this sample should provide a good comparison with high- z samples, to test the current paradigm that haloes of galaxies are the cause of QSO absorption lines. Differences of the detected properties of the QSOALS in the two samples may, at any rate, shed light on the detailed structure of galaxy haloes.

We have here presented a unique sample of 53 galaxies found intervening between background QSOs in the SDSS DR7 at $z < 0.8$ detected primarily via their H α and [O III] galactic emission, bringing our total sample size to 103. We confirmed these galaxies by measuring eight other galactic emission lines where present. The emission and absorption properties of each of these systems has been reported, showing we are studying a sample of late-type galaxies with low SFR. Our targets are shown to be comparable with field galaxies at low redshifts in terms of reddening and SFR. Many of the galaxies in our sample appear to be dwarfs in nature, given their low r -band luminosity.

We find that the SFR for our targets are, with few exceptions, extremely low. They are, however, in agreement with surveys of field galaxies in the literature. We have also determined emission line metallicity estimates for each of our targets. These values range from sub-solar to supersolar, with the majority being sub-solar in value. These values are also in agreement with studies of field galaxy emission line metallicities. In addition, we have estimated the $\log M_*$ values for our galaxies and extrapolated the relationship to those galaxies with no imaging detections, finding good agreement with the literature. These non-detections are likely low-mass dwarf galaxies.

We find a significant difference between our total sample and the Mg II selected sample presented in York et al. (2006). Comparing the reddening values of our Mg II absorbers in the total sample to the sample of York et al. (2006), we find an increased likelihood that they could be from the same population, however, the probability is still very small. We do detect an anticorrelation between impact parameter and $E(B - V)_{(g-i)}$ at the $\alpha = 0.03$ significance level, as well as a correlation between $E(B - V)_{(g-i)}$ and $(u - r)$ at the $\alpha = 0.2$ significance level. Additionally, we find a correlation between $E(B - V)_{(g-i)}$ and $E(B - V)_{\text{H}\alpha/\text{H}\beta}$ at the $\alpha = 0.04$ significance level. The absorption features show that five (potentially six) of our systems are candidates for DLA systems along the QSO line of sight, with Mg II, Mg I, and Fe II EW $> 1 \text{ \AA}$.

We estimate the dN/dz values for our Ca II and Mg II absorbers and find compared to other studies of these features we do not have an oversampling, indicating GOTOQs are not preferentially traced

by these systems more than expected. Both of these absorption lines are found to trace dusty, star-forming regions and a wide range of galactic environments, respectively. We have also compared absorption line strengths and found positive correlation between Ca II and Mg I and also between Mg II and Fe II. In comparing emission line strength with absorbers, we find a positive correlation between Ca II and H α and an anticorrelation between Ca II and [O III].

We are currently reducing optical integral field spectroscopy of many of our targets in order to determine the total SFR and kinematics of the systems. Additionally, we plan to obtain deeper, higher resolution imaging of those systems with no imaging detections. Also planned are UV spectra of our targets to obtain data on Lyman α absorption, and Mg II absorption where $z < 0.4$. These data will allow a direct comparison with higher z Mg II absorbers and other absorption systems.

ACKNOWLEDGEMENTS

LAS and VPK would like to acknowledge partial funding from NSF grant AST/0908890 (PI Kulkarni). VPK also acknowledges partial support from NSF grant AST/1108830. LAS and DGY acknowledge support from the John Templeton Foundation through a grant to DGY. We thank the referee for the constructive comments which improved this paper. Funding for the SDSS and SDSS-II has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, the US Department of Energy, the National Aeronautics and Space Administration, the Japanese Monbukagakusho, the Max Planck Society, and the Higher Education Funding Council for England. The SDSS website is <http://www.sdss.org/>. The SDSS is managed by the Astrophysical Research Consortium (ARC) for the Participating Institutions. The Participating Institutions are The University of Chicago, Fermilab, the Institute for Advanced Study, the Japan Participation Group, The Johns Hopkins University, the Korean Scientist Group, Los Alamos National Laboratory, the Max-Planck-Institute for Astronomy (MPIA), the Max-Planck-Institute for Astrophysics (MPA), New Mexico State University, University of Pittsburgh, University of Portsmouth, Princeton University, the United States Naval Observatory, and the University of Washington.

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SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

Table 1. Catalogue of GOTOQs.¹

Table 2. Photometric data for QSOs.

Table 3. Photometric data for deconvolved galaxies.

Table 4. Star formation rates.

Table 5. Emission line strengths for galaxies in front of QSOs and emission line metallicities.

Table 6. Rest-frame EW (W) for absorption interstellar lines in foreground galaxies.

Figure 2. SDSS multicolour images of the 36 fields with a QSO intercepting a low- z galaxy detected in imaging, ordered by the increasing RA and increasing numerical index.

Figure 4. SDSS spectra of our targets (<http://mnras.oxfordjournals.org/lookup/suppl/doi:10.1093/mnras/stu2739/-/DC1>).

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