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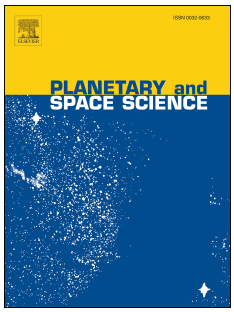
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Title: Tintigny meteorite: the first Belgian achondrite

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Abstract:

A late afternoon in February 1971, a meteorite impacted the rooftop of a house in Tintigny village in southern Belgium. Confirmed as a possible meteorite by the schoolteacher, the meteorite and its fall story did not leave the village. Finally, 46 years after the fall event, we got the opportunity to study and characterize this meteorite. In this work, we give a detailed report on its textural, mineralogical, whole-rock elemental and oxygen isotopic composition. Officially named as Tintigny, we classified it as an achondrite from howardite-eucrite-diogenite (HED) clan and more precisely a polymict eucrite. A brecciated basaltic rock believed to be originated from the surface of V-type asteroids namely the asteroid 4-Vesta. Tintigny has recorded the evidence of the impact metamorphism and metasomatism processes active on its parent body. Tintigny is one of the 39 eucrite falls known to date, and one of the 11 eucrites occurred in Europe. It is the fifth officially recognized meteorite and the first achondrite from Belgium. This report shows the importance of studying and accessing such a meteorite for further cosmochemical and planetary investigations and enriching our knowledge on the formation of HED meteorites and their parent bodi(es). In addition, it brings the attention to its importance as a scientific heritage that has to be properly understood and safeguarded for the generations of scientists, scholar, and amateurs to come

Keywords: Achondrite, Fall, Eucrite, Belgium, Tintigny

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INTRODUCTION

In February 1971 (precise date not recorded), Mr Eudore Schmitz was working in his barn in the village of Tintigny (southern Belgium, 49.683786°N, 5.532957°E) during the late afternoon when he heard a loud noise from the roof of the building. After going upstairs, he found a hole in a tile and a black stone on the barn floor. It was suggested that he burnt himself picking up the fragment, so he used some hay and then his hat to hold the stone. The schoolteacher of the village, Mr Albert Rossignon, confirmed that the stone was a meteorite and kept it, hoping that his identification would be confirmed during a subsequent investigation. The teacher later joined a religious seminary and became a priest. While he faithfully kept the meteorite and showed it from time to time to visitors and children, the stone and associated story never left the region. In 2017, after reading an article about recent Belgian meteorite recovery expeditions in Antarctica, he contacted Dr Vinciane Debaille, professor at the Université Libre de Bruxelles (ULB) and specialist in planetary sciences and meteoritics who recognized the stone as an achondritic meteorite. The meteorite was subsequently donated by Madam Germaine Mathus, widow of Mr Eudore Schmitz, and her children, Jean-Paul, Rita and Joseph Schmitz, to the Royal Belgian Institute for Natural Science (RBINS) and studied. While the meteorite is no longer complete due to handling of the stone by various people over the years, Father Rossignon affirms that the fusion crust was initially complete, with a piece of the tile originally stuck on the stone.

We have classified this meteorite as a polymict eucrite and Tintigny, its official name, has been approved by the Nomenclature Committee of the Meteoritical Society (Gattacceca et al., 2020). Tintigny is the fifth officially recognized meteorite and the first achondrite from the Belgian territory (Fig. 1). This meteorite is now on permanent open display at RBINS.

In this paper, details on the petrological, geochemical, and isotopic characteristics of Tintigny are reported and its formation processes studied.

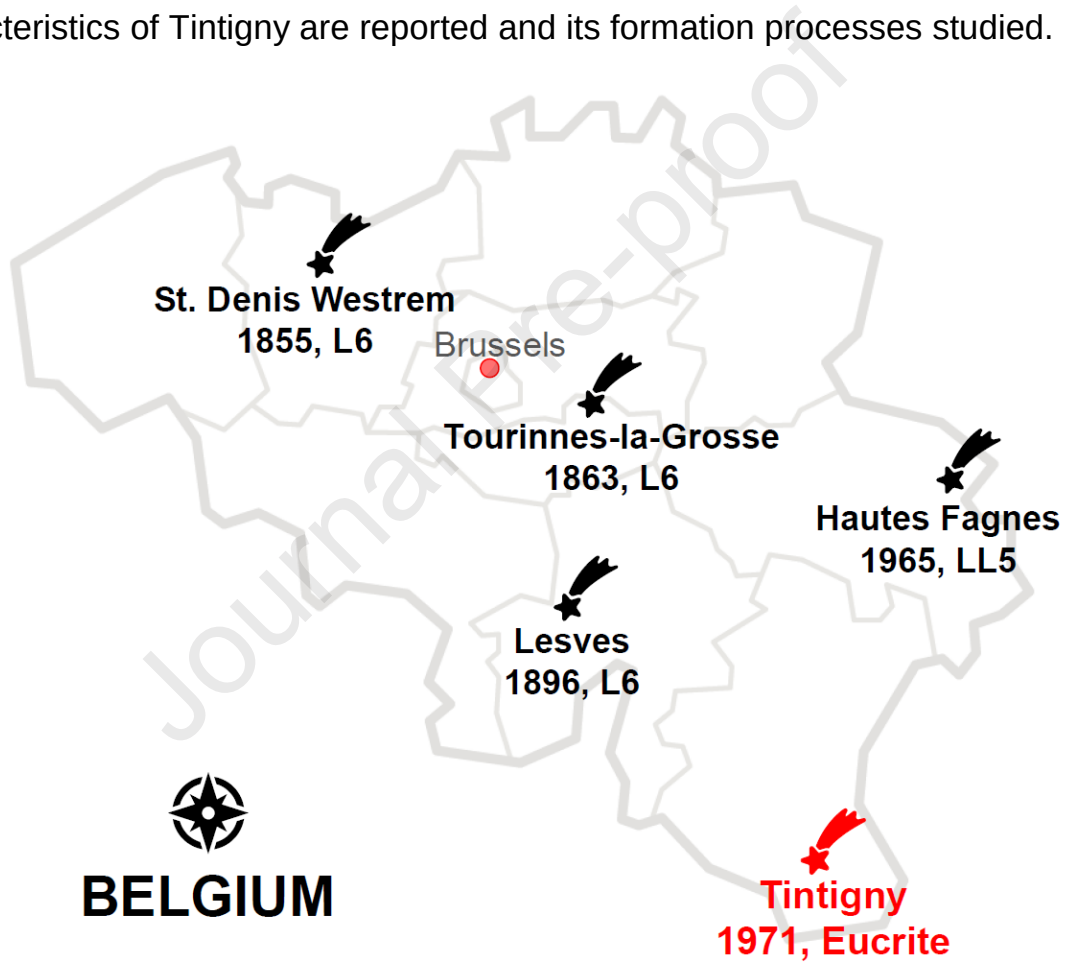


Figure 1: Meteorites of Belgium with fall/find years and types.

METHODOLOGY

Following examination of the whole rock using a stereomicroscope, representative fragments covering all petrographic textures were separated and embedded in resin at RBINS. Mineralogical and petrological studies were conducted using an optical microscope (ULB) and a PANalytical (Quanta 200) equipment Environmental Scanning Electron Microscope (ESEM) equipped with an EDAX (Apollo 10 silicon drift detector) Energy Dispersive Spectrometer (EDS) at the RBINS. The chemical composition of mineral phases were determined using a CAMECA SX50 electron microprobe (EMP) at the CAMPARIS facility (Paris), relying on a series of natural and synthetic standards and a later correction for SiO_2 calibration. The set instrument parameters include a focused electron beam (~ 1 µm in diameter), an accelerating voltage of 15 kV and a beam current of 10 nA. We aimed to analyze pyroxene and plagioclase with diverse compositions. Special care was considered to avoid grain boundaries, cracks etc. which can bias the quality of the analytical data. Whole-rock major and trace elemental concentrations were determined at the Laboratoire G-Time of the ULB, Belgium. Around 50 mg of sample representing the most common lithology (light grey) of the meteorite, was dissolved by alkaline fusion for major and trace element contents. Major elements were measured using a Thermo Fisher Scientific iCAP inductively coupled plasma-optical emission spectrometer (ICP-AES) at ULB with Y as an internal standard. Overall, the total reproducibility estimated based on United States Geological Survey reference material BHVO-2 is calculated to be better than 2% relative standard deviation (RSD). Trace elements were measured

using the Agilent 7700 Quadrupole-Inductively Coupled Plasma-Mass Spectrometer (Q-ICP-MS) operated with a He-filled collision cell at ULB. Indium was used as internal standard. The total reproducibility estimated based on USGS reference material BHVO-2 is calculated to be better than 10% relative standard deviation (RSD).

High-precision oxygen isotopic analysis of Tintigny was undertaken at the Open University (Milton Keynes, UK) using an infrared laser-assisted fluorination system (Miller et al., 1999; Greenwood et al., 2017). Measurements were made on a 200 mg aliquot of silicates from the bulk sample. Approximately 2 mg aliquots of samples and standards were loaded into a nickel sample block, which was then placed in a two-part chamber, made vacuum tight using a compression seal with a copper gasket and quick-release KFX clamp (Miller et al., 1999). A 3 mm thick BaF_2 window at the top of the chamber allows simultaneous viewing and laser heating of samples. Prior to analysis the sample chamber was heated overnight under vacuum to a temperature of about 70°C to remove any adsorbed moisture. Following overnight heating, the chamber was allowed to cool to room temperature and was then flushed with several aliquots of BrF_5. The system was then left to pump for at least a further 24 hours and oxygen isotopic analysis was only undertaken when the blank level reached <60 nanomoles of O_2. Sample heating in the presence of BrF_5 was carried out using a Photon Machines Inc. 50W infrared CO_2 laser (10.6 µm) mounted on an X-Y-Z gantry. Reaction progress was monitored by means of an integrated video system. After

fluorination, the released O_2 was purified by passing it through two cryogenic nitrogen traps and over a bed of heated KBr to remove any excess fluorine. The isotopic composition of the purified oxygen gas was analyzed using a Thermo Fisher MAT 253 dual inlet mass spectrometer with a mass resolving power of approximately 200.

Overall system precision, as defined by replicate analyses of our internal obsidian standard, is: ±0.053‰ for $\delta^{17}O$; ±0.095‰ for $\delta^{18}O$; ±0.018‰ for $\Delta^{17}O$ (2σ) (Starkey et al., 2016). Oxygen isotopic analyses are reported in standard δ notation, where $\delta^{18}O$ has been calculated as: $\delta^{18}O = [(^{18}O/^{16}O)_{sample}/(^{18}O/^{16}O)_{VSMOW} -1] 1000$ (‰) and similarly for $\delta^{17}O$ using the $^{17}O/^{16}O$ ratio. VSMOW is the international standard Vienna Standard Mean Ocean Water. $\Delta^{17}O$, which represents the deviation from the terrestrial fractionation line, has been calculated using the linearized format of Miller (2002):

$$\Delta^{17}O = 1000\ln(1+(\delta^{17}O/1000)) - \lambda 1000\ln(1+(\delta^{18}O/1000))$$ where $\lambda = 0.5247$.

RESULTS

Macroscopic description

As a result of the thermal effects experienced during atmospheric entry and impact through the roof of the building, the Tintigny meteorite finally fragmented into several pieces during its extensive manual handling alongside fractures, indicated by the occurrence of broken surfaces. The single recovered fragment weighs 210 g (Fig. 2).

The meteorite is partly covered by a shiny black fusion crust, which varies in thickness depending on the morphology of the underlying surface. Three main varieties of fusion crust are present: (i) the thickest (~ 1-2 mm) parts that occur in hollow areas of the surface and display a wavy texture comprising elevated ribbons of molten material (flow lines) (Fig. 2a); (ii) the main part that is considerably thinner (~ 0.5 mm) and indicates a full layer of molten material with a texture containing assemblages of crater-like pits (Fig. 2b); (iii) the intermediate part exhibiting a combination of these two variations, with a patchy texture occurring along the edges of the meteorite surface. Several cracks occur on the surface of the fusion crust. They correspond to deeper fractures visible where the surface is broken (Fig. 2c). At these sites, a light gray interior is revealed, composed of a fine-grained light-colored matrix hosting darker crystals and a cm-size dark grey clast (Fig 2d). Macroscopic observations indicate that Tintigny is a brecciated achondrite.

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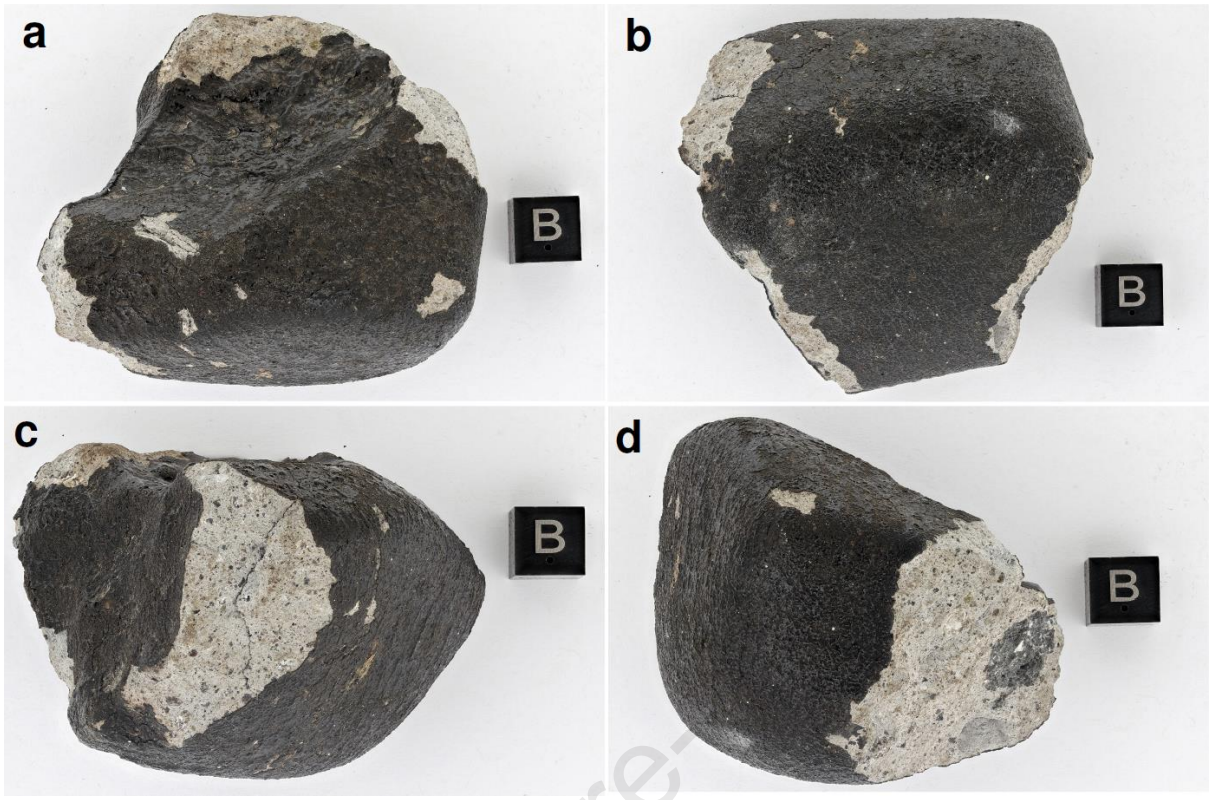


Figure 2: Tintigny meteorite from different angles. Note the different textures of fusion crust, the occurrence of fractures, and a dark grey clast in d. Image credit: RBINS.

Microscopic description

Using optical and electron microscopy, Tintigny exhibits a brecciated sub-ophitic basaltic texture mainly composed of plagioclase/maskelynite and clinopyroxene (Fig. 3). These minerals occur both as large crystals (>50µm) and as smaller (<50 µm) ones, the latter mainly composing the clastic matrix. Accessory minerals include troilite, ilmenite, chromite, (Fe,Ni) metal, and silica. Most of the minerals exhibit well-delineated edges, enhancing the clastic texture of the rock. At least two generations of shock fractures are visible:

153 those specific to clasts and large crystals, and those cross-cutting both the
154 larger grains and the matrix materials.

155 In addition to the main sub-ophitic texture, at least three distinct textures are
156 present in specific clasts. These clasts are different from their host based on
157 their texture and degree of equilibrium recorded in the composing minerals.
158 Figure 3a,4e shows a ~ 1 mm long clast with a sub-ophitic texture. Elongated
159 pyroxenes are unequilibrated and display Fe enriched rims. The elongation
160 directions of pyroxene crystals control the alignment of plagioclase. Abundant
161 cracks perpendicular to pyroxene elongation are visible. They are absent from
162 the adjacent plagioclase crystals, suggesting elevated levels of mechanical
163 stress prior plagioclase crystallization. Some of these plagioclase crystals are
164 cut by randomly oriented cracks, which are filled by Fe-rich pyroxenes. This is
165 related to fusion crust formation. Another clast with a similar texture is visible in
166 Fig. 3f.

167 Under the electron microscope, the dark clast visible on the broken surface
168 (Fig. 2d) displays a melt rock texture (Fig. 3e,4f). The groundmass is a mixture
169 of quenched pyroxene and plagioclase and only a few larger grains including
170 (Fe,Ni) metal, ilmenite, chromite, and an intergrowth of plagioclase-pyroxene
171 are present.

172 A single ~ 0.5 mm clast and several smaller grains exhibit a symplectitic
173 mixture of pyroxene, fayalite, and silica (Fig. 3a,c,d and Fig. 4b,d).

174 A small number of pyroxenes show evidence of Fe enrichment along veins
175 and crystal rims (Fig. 3f and Fig. 4b,e). These fractures are mostly limited to

individual grains and do not occur in the whole rock, indicating their formation was prior to the assembly of Tintigny on its parent body.

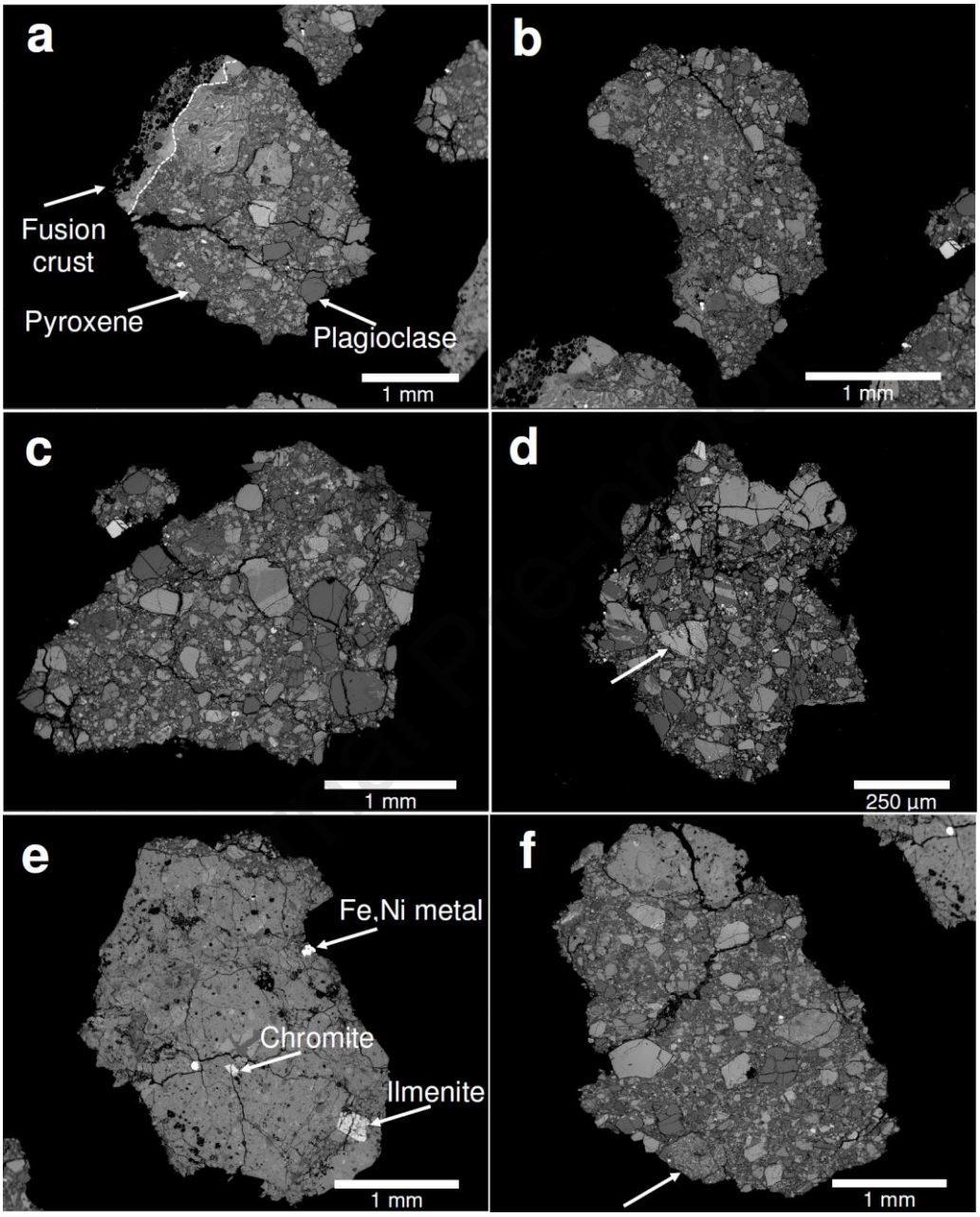
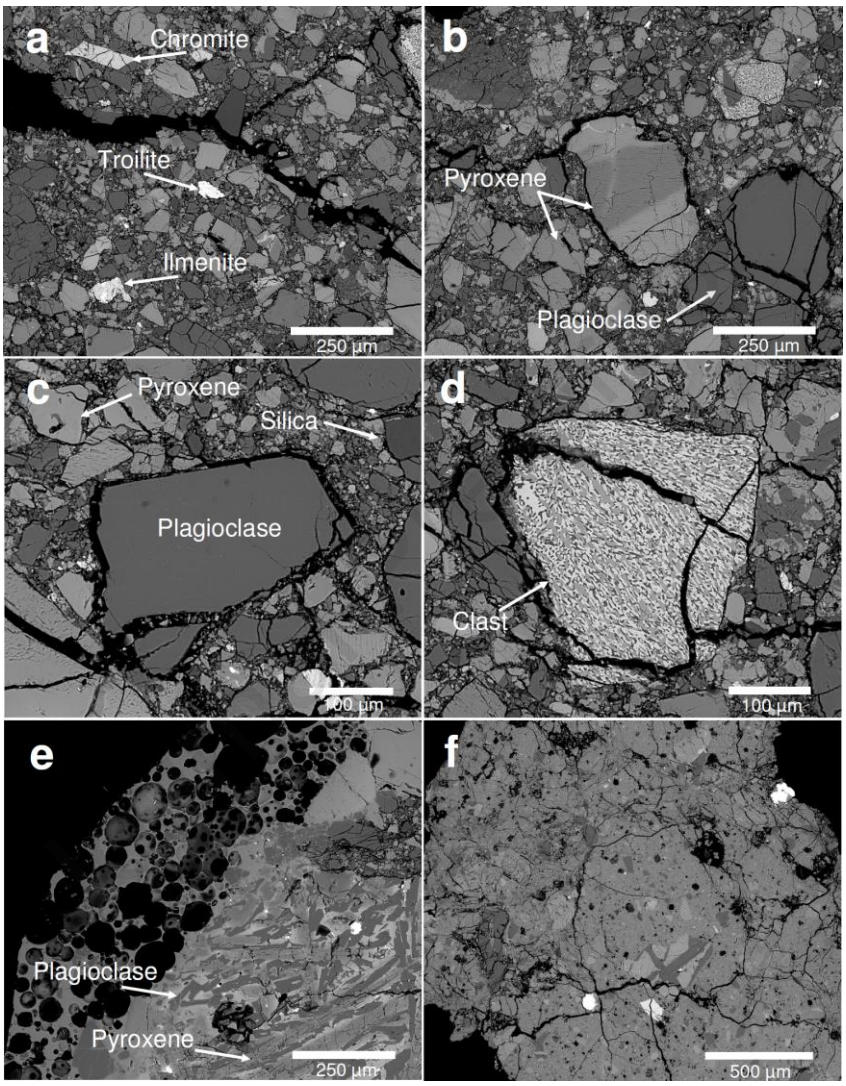


Figure 3: Backscattered electron microscope images of Tintigny. The presence of different textures suggests Tintigny to be a brecciated eucrite. a) A sub-ophitic clast adjacent to the fusion crust shows a texture distinct from that of the host. In the center, a clast with a symplectitic texture is visible. b) Typical texture of Tintigny. A pyroxene with exsolution lamellae is visible. c) Variety of pyroxene textures including exsolution lamella and symplectitic assemblages. d) Partial transformation of pyroxene to a symplectitic assemblage

185 is visible in one clast (shown by arrow). e) Dark grey clast (also visible in Fig. 2d), showing
186 textural evidence of melt rock. f) The occurrence of different clasts. Note Fe-rich veins in a ~
187 500 µm pyroxene in upper left. A clast with relatively larger crystals (shown with arrow) with
188 similar sub-ophitic texture to the clast shown in 3a and 4e.



189

190 **Figure 4:** a) Accessory minerals in Tintigny. b) Pyroxene showing a variety of textures
191 (exsolution, twinning, symplectitic, etc.), suggesting a complex thermal history experienced
192 by the Tintigny. c) Clastic texture of Tintigny formed as a result of impact metamorphism on
193 the parent body surface. d) A symplectitic assemblage of hedenbergite-fayalite-silica. e)
194 Alignment of plagioclase along the surface of pyroxene. Note the higher concentration of Fe
195 in pyroxene rims. f) A closer view of the dark grey clast (Fig. 2d) shows lower number of
196 crystals and a shock-melted texture.

197 **Mineral chemistry**

198 Table 1 summarizes the chemical composition of pyroxene and plagioclase
199 crystals analyzed randomly in different fragments of Tintigny. Mineral
200 chemistry calculations of pyroxene end members show ranges from 8.5 to 60.7
201 mol% for enstatite, 30.1 to 70.0 mol% for ferrosilite, and 2.6 to 38.4 mol% for
202 wollastonite. Based on these values, most pyroxenes in Tintigny are pigeonite
203 and augite (Morimoto 1988; Marshall 1996) (Fig. 5a). The Fe/Mn ratios of
204 pyroxenes range from 27.1 to 39.3, with the highest ratio observed in pyroxene
205 from the symplectitic clast (Pyx #5, Table 1). Fe/Mn and Fe/Mg ratios in low-
206 Ca pyroxene (Wo <10, Pyx #1,2,3,6,7,10) are 30.2±4.4 and 0.8±0.3,
207 respectively. These ratios in high-Ca pyroxene (n=8) are 34.3±3.7 for Fe/Mn
208 and 2.6±2.4 for Fe/Mg. The average pyroxene Fe/Mn ratio for all pyroxene is
209 32.5±4.4 (SD, n=14). Fe/Mg ranges from 0.6 to 8.2, with an average value of
210 1.8±2.0 (SD, n=14). Considering pyroxene Fe/Mn ranges of 40±11, 62±18,
211 32±6, and 30±2 for basaltic rocks from the Earth, Moon, Mars, and 4Vesta
212 (eucrites), respectively, and based on our data, particularly those of low-Ca
213 pyroxene, Tintigny falls in the range of basaltic eucrites (Papike et al., 2003).
214 We believe the higher standard deviation of our data results from a higher
215 diversity and relatively lower number of the analyzed minerals.

216 The anorthite content of four analyzed plagioclase ranges from 75.8 to 90.3
217 mol%. Plagioclase in the symplectitic clast, with 75.8 mol% anorthite, is less
218 calcic (and more sodic) than the host (Fig. 5b). Excluding this clast, the
219 anorthite percentage averages to 86.8±3.4 mol% (SD, n=3). This value is in

range of both vestan (87±2 mole%) and lunar rocks (89±3 mole%) (Papike et al., 2003), however the Fe/Mn ratio of pyroxene indicates that Tintigny is a member of the HED suit.

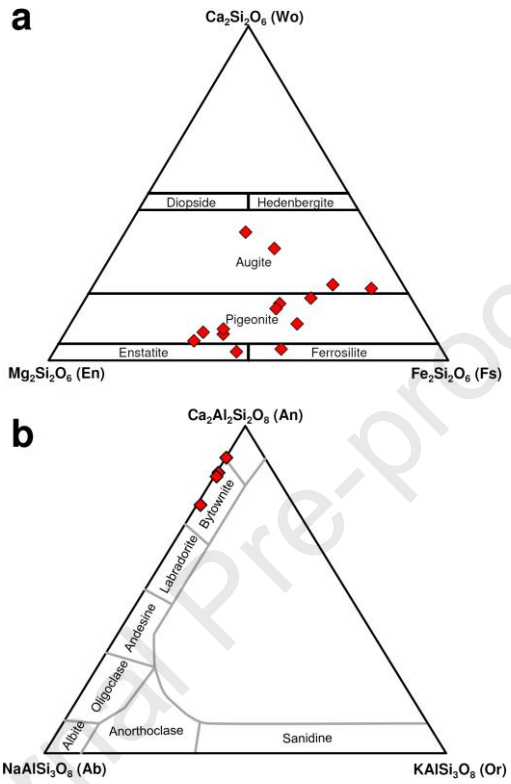
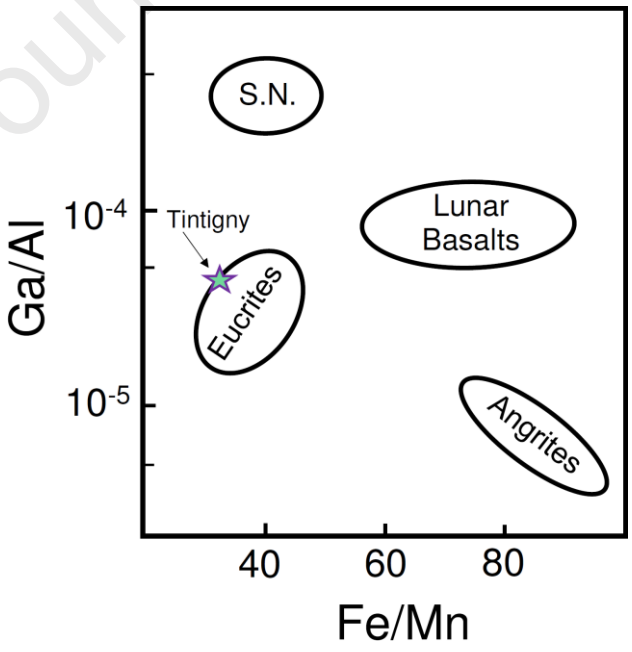


Figure 5: Pyroxene and plagioclase compositions in Tintigny.

Whole-rock chemical composition

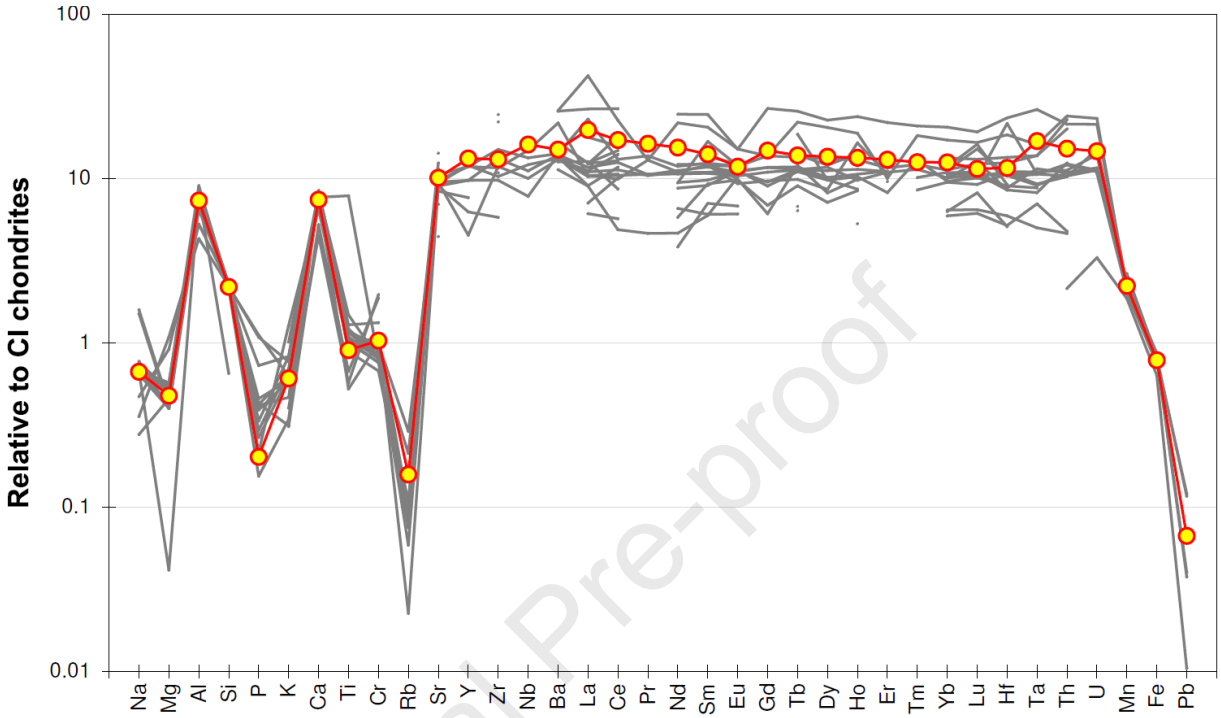
Table 2 shows the major and trace element concentrations of Tintigny. Based on bulk rock Fe/Mn vs. Fe/Mg ratios, three distinct zones for chondrites, lunar rocks, and howardite-eucrite-diogenite (HED)/Martian meteorites can be defined (Goodrich & Delaney, 2000). For Tintigny, the bulk rock Fe/Mn and Fe/Mg ratios are 33.9 and 3.1, respectively. These values overlap with those measured for HED and Martian meteorites. To discriminate between different

233 types of basaltic achondrites, Fe/Mn has previously been combined with other
234 useful ratios such as Ga/Al (Barrat et al., 2003). The Ga/Al ratio of Tintigny is
235 4.17×10^{-5}, fully in range of those of eucrites (Fig. 6). The CI-normalized
236 elemental concentrations for Tintigny are compared to those of 18
237 noncumulate eucrites in Fig 7. The latter plot indicates the strong similarities
238 between the chemical composition of Tintigny and that of noncumulate
239 eucrites. This similarity to eucrites is also evident based on various
240 combinations of major and trace elements for HED. On a binary plot of Ca
241 versus Mg, Tintigny overlaps with eucrites, but is distinct from howardites or
242 diogenites (Fig. 8). Similar behavior occurs in Sm versus Mg and Yb versus La
243 plots. Based on the abundance of TiO_2 (0.63%) and FeO/MgO ratio (2.66),
244 Tintigny is a member of non-cumulate eucrites (Fig. 9) (Barrat et al., 2003).
245



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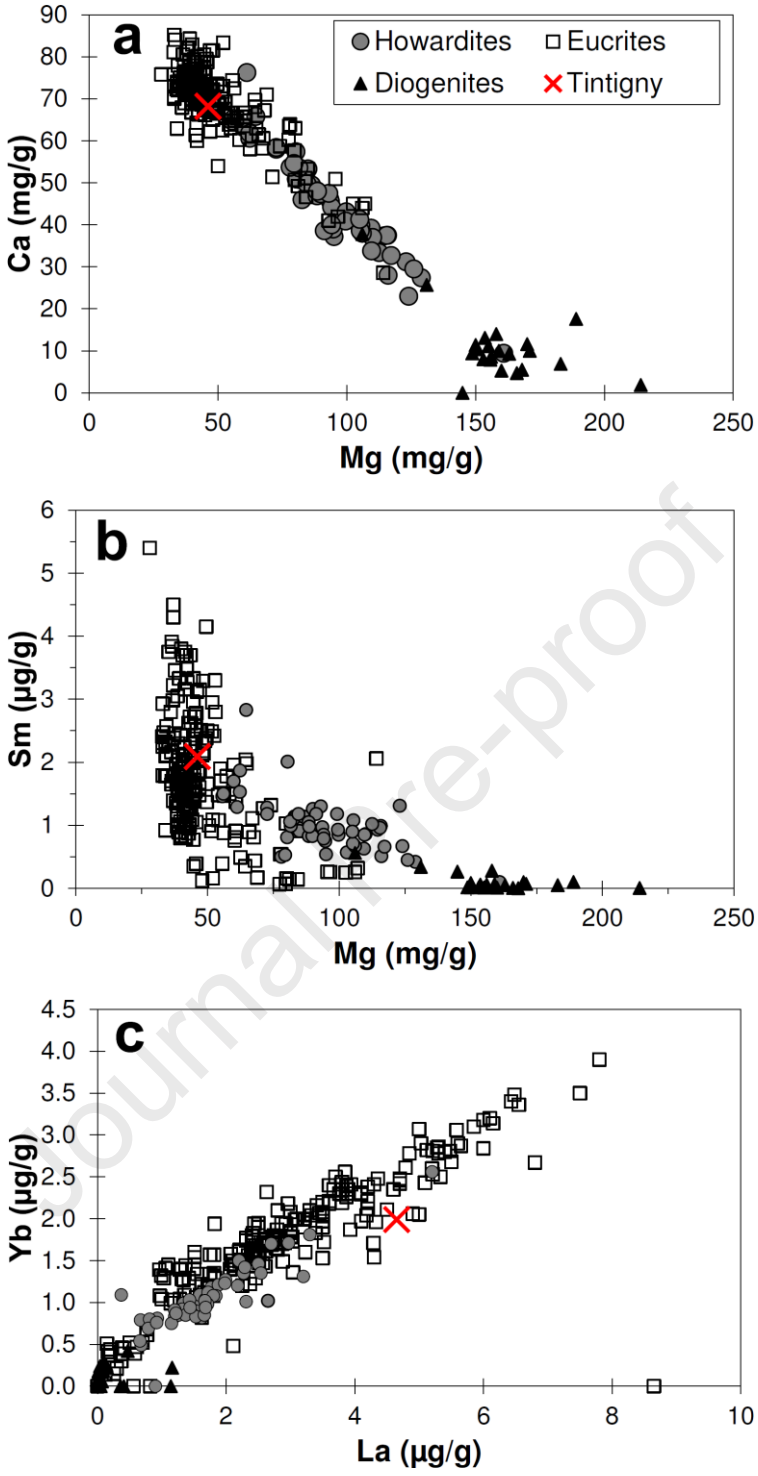
247 **Figure 6:** Ga/Al vs. Fe/Mn of Tintigny in comparison to other basaltic meteorites. Plot
248 modified from Barrat et al. (2003). S.N. stands for shergottites and nakhlites clan of Martian
249 meteorites.



250

251 **Figure 7:** CI-normalized whole-rock chemical composition of Tintigny and 18 noncumulate
252 eucrites (Béréba, Bouvante, Cachari, Chervony Kut, Elephant Moraine A79004,A79005,
253 Haraiya, Ibitira, Junzac, Juvinas, Millbillillie, Moore County, Pasamonte, Pomozdino, Yamato
254 794002,791573,82049,82202). Eucrite data from Kitts & Lodders (1998), Mittlefehldt (2015).
255 Average CI chondrite data are from Wasson & Kallemeyn (1988). Lithophile, siderophile, and
256 chalcophile elements are shown with increasing atomic number, respectively.

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257

258 **Figure 8:** The abundance of Mg vs. Ca, Mg vs. Sm., and La vs. Yb for Tintigny in relation to
259 howardite-eucrite-diogenite (HED) meteorites. Compilation of HED data by Mittlefehldt
260 (2015).

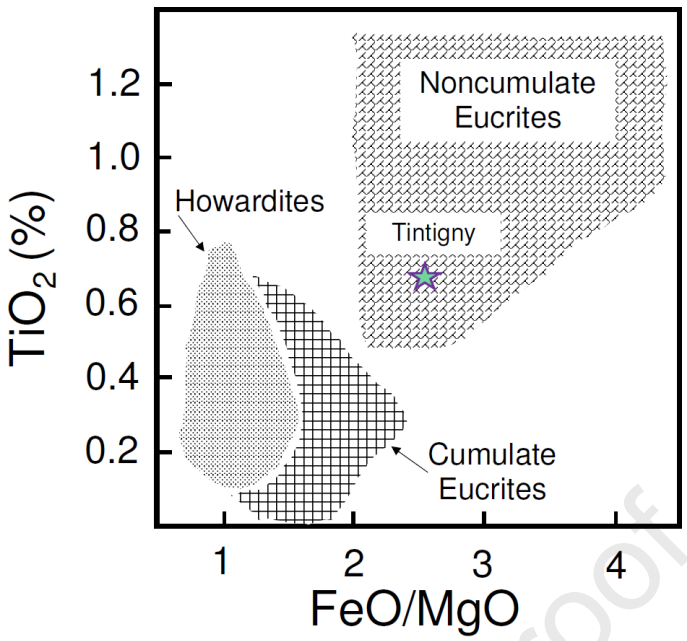


Figure 9: TiO_2 vs. FeO/MgO of Tintigny in comparison to howardite and eucrites. Plot modified from Barrat et al. (2003).

Oxygen isotopic composition

Two replicate analyses Tintigny gave the following oxygen isotopic: $\delta^{17}O$ = 1.723 ± 0.018 (1σ); $\delta^{18}O$ = 3.756 ± 0.041 (1σ) and $\Delta^{17}O$ = -0.246 ± 0.003 (1σ). The oxygen isotope data for Tintigny are plotted in relation to the HED data of Greenwood et al. (2017) in Fig. 10 and this shows that the meteorite lies close to the Eucrite Fractionation Line (EFL) defined by eucrite and diogenite falls ($\Delta^{17}O$ = -0.240). In addition, the $\delta^{18}O$ value of Tintigny plots centrally in the field of eucrite analyses. The oxygen isotopic data is therefore consistent with the classification of Tintigny as a eucrite.

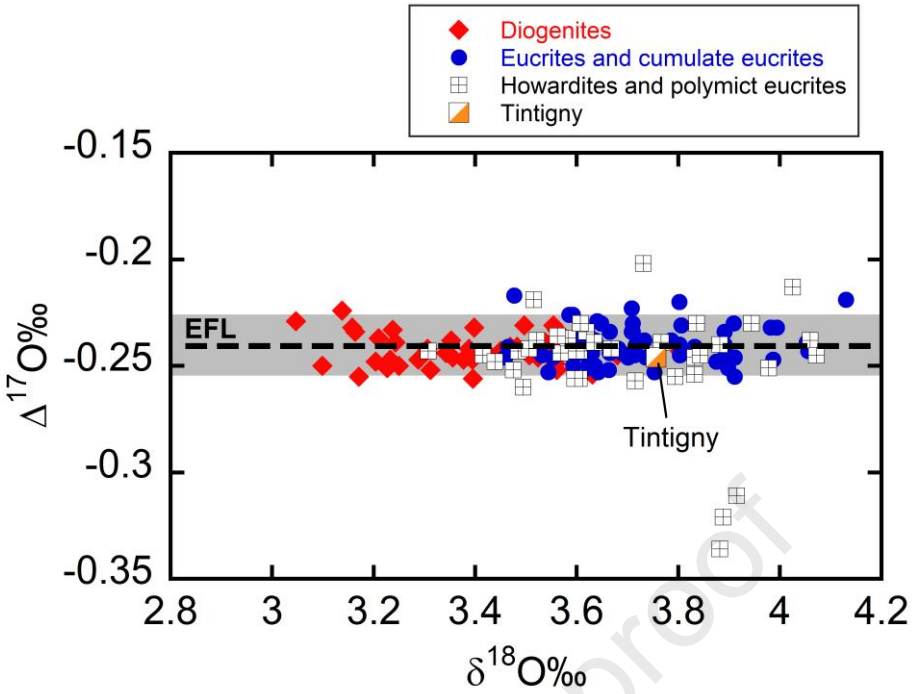


Figure 10: Oxygen isotopic composition of Tintigny in comparison with HED meteorites. HED data from Greenwood et al. (2017).

DISCUSSION AND CONCUSIONS

In this work, we used petrography, mineralogy, and whole-rock major and trace-element chemistry to classify the Tintigny meteorite. In addition, we have undertaken oxygen isotope analysis which has confirmed that Tintigny is a member of the HED clan. Based on our studies we conclude that Tintigny is a non-cumulate eucrite, specifically a polymict basaltic eucrite.

Eucrites, together with howardites and diogenites, form the HED clan that constitutes 74.6% by number of all achondrites, and 3.6% by number of all meteorites in meteorite collections worldwide (Meteoritical Bulletin Database (https://www.lpi.usra.edu/meteor/metbull.php), accessed February 2021).

Based on the spectral data obtained in the laboratory and their comparison with data measured by ground-based observatories and the results from the NASA Dawn mission, HEDs are thought to originate from differentiated asteroids with V-type spectra, and in particular 4-Vesta (McSween et al., 2010; Moskovitz et al. 2010). Postcrystallization events on the parent body such as thermal metamorphism, metasomatism, shock metamorphism, and space weathering have led to the formation of rocks with complex geological histories (Yamaguchi et al., 1994; Takeda et al., 1985; Warren et al., 2014). As described earlier, Tintigny has also been affected by these processes, in particular thermal and shock metamorphism and possibly metasomatism as recorded in some grains (as Fe enrichment along veins and crystal rims). Many HED meteorites are brecciated, with a general distinction between monomict and polymict breccias (Delaney & Prinz, 1984; Mittlefehldt et al., 2013, Zucolotto et al., 2018). These breccias result from large-scale impact events of the Solar System bodies and witness to the importance of impact processing of the surface and subsurface of planetary bodies. The lack of chemical zoning in most minerals (except in some clasts), the presence of Fe,Ni metal exsolutions in pyroxene, and sub-solidus exsolution of augite lamellae within pigeonite hosts (Fig. 4) are indicative of parent body thermal metamorphism (Righter & Drake, 1997; Vollmer et al., 2020). The formation of clasts with symplectitic texture has been suggested to result from the breakdown of metastable pyroxene in gabbroic eucrites (Patzer & McSween, 2012; Seddiki et al., 2013). Barrat et al. (2011) and Warren et al. (2014) linked

311 the formation of Fe enrichment along the veins and crystal rims to fluid-driven
312 alteration on the parent body surface.

313 Diogenites are mostly orthopyroxene cumulates that formed in plutons
314 which crystallized at varying levels within the crust. Eucrites are mostly basaltic
315 rocks that formed at faster cooling rates, most likely as a result of
316 emplacement either on the surface or at shallow depths. According to their
317 textural and compositional characteristics eucrites can be further divided into
318 basaltic rocks (mostly monomict breccias), cumulate gabbros, and polymict
319 eucrites made of different eucritic textures (without diogenitic clasts) like
320 Tintigny (Vollmer et al., 2020). Barrat et al. (2003) also divides them to
321 cumulate and noncumulate eucrites. Howardites formed as the products of
322 impact events on the parent body surface, occurring as breccias made from
323 different percentages of eucritic and diogenitic materials.

324 Meteoritical Bulletin Database (https://www.lpi.usra.edu/meteor/metbull;
325 August 2021) lists 2447 HED meteorites encompassing 1516 eucrites, 399
326 howardite, and 532 diogenites. Out of this number, only 69 of the HED are
327 *falls,* i.e., meteorites that have been observed falling and have been collected
328 soon after their impact, avoiding the detrimental effects of terrestrial
329 weathering, which are common in meteorite *finds* (meteorites without any fall
330 record) and even in some cases in *falls* if not collected immediately (Walker et
331 al*.,* 2018*;* Pourkhorsandi, 2018; Pourkhorsandi et al*.,* 2019*)*. Including Tintigny,
332 only 39 eucrite falls are known to date, 11 of them occurred in Europe and
333 Tintigny being the only one from Belgium. This highlights the importance of

classification and accessibility of such a meteorite for further cosmochemical and planetary studies and enriching our knowledge on the formation of HED meteorites and their parent bodi(es). In addition to its scientific importance, we emphasize the importance of the discovery of a historical meteorite fall in bringing attention to national scientific heritage that has to be properly understood and safeguarded for the generations of scientists, scholar, and amateurs to come (Franza & Pratesi, 2021).

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443 **Tables:**

444 Table 1: The analyzed pyroxene and plagioclase compositions (in wt%) from Tintigny.

Mineral	Pyx #1	Pyx #2	Pyx #3	Pyx #4	Pyx #5	Pyx #6	Pyx #7	Pyx #8	Pyx #9	Pyx #10	Pyx #11	Pyx #12	Pyx #13	Pyx #14	Plg #1	Plg #2	Plg #3	Plg #4
SiO_2	52.57	52.53	52.64	49.91	51.26	49.09	51.03	49.27	49.81	53.01	46.09	47.68	49.81	45.54	47.20	48.20	45.55	50.14
Al_2O_3	1.70	0.61	0.83	0.59	0.82	0.68	0.94	1.24	0.81	1.15	1.98	0.97	1.04	1.17	32.25	33.23	34.16	31.01
TiO_2	0.26	0.15	0.11	0.39	0.51	0.42	0.20	0.47	0.71	0.15	1.11	0.74	0.67	0.92	0.00	0.00	0.01	0.03
FeO	20.87	27.24	23.95	32.26	17.91	32.54	23.61	27.97	28.50	20.46	32.58	31.56	23.21	37.84	0.63	0.57	0.51	0.63
Cr_2O_3	0.84	0.22	0.61	0.23	0.31	0.30	0.93	0.41	0.35	0.74	0.24	0.21	0.33	0.18	b.d.l.	b.d.l.	0.01	b.d.l.
MnO	0.76	0.74	0.86	1.06	0.45	0.91	0.85	0.88	0.88	0.74	0.87	0.82	0.75	1.03	b.d.l.	0.02	0.01	0.05
MgO	19.40	17.28	17.69	10.33	10.51	12.96	17.50	10.72	11.49	20.84	5.37	7.88	8.77	2.58	0.06	0.02	0.05	0.04
CaO	3.98	1.21	3.69	4.85	17.80	1.55	4.44	7.52	7.00	2.80	9.63	8.15	15.25	9.08	17.00	16.83	17.99	14.65
Na_2O	0.02	0.03	0.01	0.01	0.04	b.d.l.	0.01	0.07	0.07	b.d.l.	0.17	0.10	0.05	0.01	1.52	1.64	1.06	2.48

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K_2O	b.d.l.	0.03	0.02	0.01	0.01	0.01	b.d.l.	0.02	0.02	0.01	0.01	0.05	0.01	b.d.l.	0.08	0.10	0.02	0.15
NiO	0.05	0.06	0.03	0.03	0.07	0.15	b.d.l.	0.08	0.03	0.01	0.04	b.d.l.	0.07	0.07	0.01	0.01	b.d.l.	0.05
Total	100.40	100.10	100.44	99.67	99.69	98.61	99.51	98.65	99.67	99.91	98.09	98.16	99.96	98.42	98.75	100.62	99.37	99.23
En (mol%)	57.1	51.7	52.4	32.4	31.5	40.1	51.6	33.7	35.3	60.7	17.6	25.1	26.8	8.5	-	-	-	-
Fs (mol%)	34.5	45.7	39.8	56.7	30.1	56.5	39.0	49.3	49.2	33.4	59.8	56.3	39.8	70.0	-	-	-	-
Wo (mol%)	8.4	2.6	7.9	10.9	38.4	3.4	9.4	17.0	15.5	5.9	22.6	18.6	33.5	21.5	-	-	-	-
Fe/Mn	27.1	36.3	27.5	30.1	39.3	35.3	27.4	31.4	32.0	27.3	37.0	38.0	30.6	36.3	-	-	-	-
Fe/Mg	0.6	0.9	0.8	1.7	1.0	1.4	0.8	1.5	1.4	0.6	3.4	2.2	1.5	8.2	-	-	-	-
An (mol%)	-	-	-	-	-	-	-	-	-	-	-	-	-	-	85.7	84.5	90.3	75.8
Ab (mol%)	-	-	-	-	-	-	-	-	-	-	-	-	-	-	13.8	14.9	9.6	23.2
Or (mol%)	-	-	-	-	-	-	-	-	-	-	-	-	-	-	0.5	0.6	0.1	0.9

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450 Table 2: Major and trace element composition of Tintigny eucrite.

451 Concentrations are reported in mg/g and µg/g for major and trace elements,

452 respectively.

	Tintigny	RSD (%)	BHVO-2[1]	RSD (%)	BVHO-2[2]
Na	3.26	2	16.0	0.7	*16.4*
Mg	46.1	1	43.5	2	*43.6*
Al	63.0	0.8	71.7	0.5	*71.6*
Si	229	1	235	0.4	*233*
P	0.206	1	1.17	2	*1.2*
K	0.339	0.6	4.29	0.8	*4.30*
Ca	68.3	0.7	81.3	1	*81.7*
Ti	3.78	1	16.4	1	*16.3*
Cr	2.73	0.4	0.315	1	*0.280*
Mn	4.21	1	1.33	1	*1.3*
Fe	142.3	1	89.1	0.4	*86.3*
Ga	2.63	12	22.57	1	*21.7*
Rb	0.351	49	8.35	5	*9.80*
Sr	79.7	1	394	1	*389*
Y	19.0	1	25.2	0.4	*28.0*
Zr	49.6	1	162	1	*172*
Nb	4.34	2	17.2	1	*18*
Ba	34.4	2	134	1	*130*
La	4.65	3	16.0	1	*15.2*
Ce	10.5	0	37.6	1	*38.0*
Pr	1.51	4	5.25	2	*5.34*
Nd	7.04	3	24.1	3	*25.0*

Sm	2.09	5	6.18	4	*6.20*
Eu	0.660	2	2.10	2	*2.07*
Gd	2.91	2	6.40	3	*6.30*
Tb	0.490	2	0.910	1	*0.940*
Dy	3.32	2	5.34	2	*5.31*
Ho	0.727	1	0.994	3	*1.04*
Er	2.08	4	2.52	4	*2.54*
Tm	0.310	3	0.337	8	*0.335*
Yb	1.99	3	2.00	2	*2.00*
Lu	0.280	11	0.274	7	*0.280*
Hf	1.39	2	4.32	4	*4.47*
Ta	0.270	10	1.22	2	*1.40*
Th	0.441	6	1.17	2	*1.25*
U	0.121	3	0.407	5	*0.412*
Pb	0.162	17	0.769	1	*1.65*

453 [1]: BHVO-2 basalt reference material analyzed during same session. [2]:

454 [2]Compilled literature value from GeoReM database (http://georem.mpch-

455 mainz.gwdg.de/)

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Highlights

- Tintigny meteorite fell in February 1971 in a village in southern Belgium.
- Based on our multimethod study, we classified it as a polymict eucrite.
- This type of meteorites are achondrites from howardite-eucrite-diogenite (HED) clan.
- HED meteorites are believed to originate from the surface of the asteroid 4-Vesta.
- Tintigny is one of the 39 eucrite falls known to date.

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AUTHORSHIP STATEMENT

Manuscript title: Tintigny meteorite: the first Belgian achondrite

All persons who meet authorship criteria are listed as authors, and all authors certify that they have participated sufficiently in the work to take public responsibility for the content, including participation in the concept, design, analysis, writing, or revision of the manuscript. Furthermore, each author certifies that this material or similar material has not been and will not be submitted to or published in any other publication before its appearance in the Planetary and Space Science journal.

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Competing interests

The authors declare no competing interests.

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