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Internal Geophysics (Spatial Physics)

The mantle rotation pole position. A solar component

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ABSTRACT

The direction of the Earth's rotation axis with respect to the mantle has been studied for more than a century. The time variation of this direction is generally considered to be the sum of three components: the annual wobble, forced by the atmosphere, the Chandler wobble, a free oscillation with a period of 435 days, and the so-called drift of the mean pole. In the present paper, applying the singular spectrum analysis (SSA) technique, we uncover two more components, with smaller amplitude than the three first ones, but well identified, periodic with periods of 11 and 5.5 years, respectively, undoubtedly linked to solar activity. We interpret them tentatively as the result of an exchange of kinetic angular momentum between the atmosphere, in which a flow would be generated by solar activity, and the mantle. The order of magnitude of the required mean winds in the atmosphere computed in the frame of a schematic model is 1 ms^{-1} , compatible with the observed values of the meridional mean circulation.

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1. Introduction

The Earth's rotation axis, close to the polar inertia axis, presents, with respect to the mantle, variations at different time scales, from daily to secular. The trace of the rotation axis at the surface of the Earth is generally considered to be the sum of three components: the annual wobble, a forced oscillation due to an atmospheric excitation, the Chandler wobble, a free oscillation with a period of 435 days, whose maintaining mechanism is not yet fully understood, and the so-called drift of the mean pole (e.g., Chandler, 1891a; Chandler, 1891b; Gibert and Le Mouél, 2008; Hulot et al., 1996; Lambeck, 2005). In the present paper, we will add a fourth periodic component, that we attribute to solar activity.

In 1967, Karklin (1967) analyzed the International Latitude Service data from 1900 through 1959, and found

variations in the amplitude of both the Chandler and the annual oscillations with decadal periods (e.g., 10.3, 12.3, 17.7 years for Chandler) and amplitudes of a few 10^{-2} arcsec ($1 \text{ arcsec} \sim 4.85 \cdot 10^{-6} \text{ rad}$) that he attributed to Solar activity. Jady (1970) estimated the torque exerted by the solar wind on the Earth's magnetic dipole, M , using a simplified model of magnetosphere (a thin planar, perfectly conducting sheet). He found that the magnitude of the torque ($1.1 \cdot 10^{16} \text{ J}$) was smaller than that required to account for the change in the length of the day (l.o.d), but comparable to the lunar tidal torque ($2.7 \cdot 10^{16} \text{ joules}$). It seems that this mechanism has not been considered further since then.

2. The SSA analysis of the series of pole positions

We will use here the series of pole positions produced and issued by the International Earth Rotation Service (IERS; web site: <http://www.iers.org/IERS/EN/Science/EarthRotation/EarthRotation.html>) in the form of a yearly report (EOPC01); the 2015 EOPC01 report is considered. m_x

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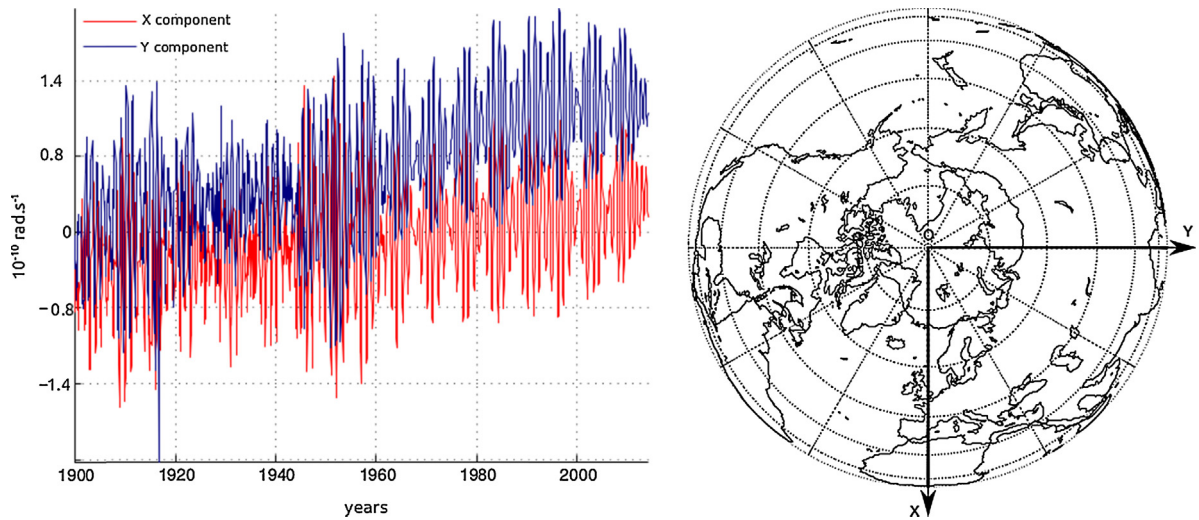


Fig. 1. Polar motion data series (EOPC01) provided by the International Earth Rotation Service. The X and Y components, m_x , m_y , correspond to the red and blue lines respectively. m_x and m_y are given in $\text{rad}\cdot\text{s}^{-1}$.

and m_y are the components of the pole position in a tangent plane Oxy at the north pole (or in the equatorial plane). O is a conventional origin, the axis Ox is in the Greenwich meridian, the axis Oy is in the 90°E meridian; the sampling interval is 0.6 month. Data m_x and m_y cover a period from 1900 to 2015 (see Fig. 1); they will be counted in $\text{rad}\cdot\text{s}^{-1}$, like the Earth's spin itself. We analyze those series through singular spectrum analysis SSA (Vautard and Ghil, 1989) (see Appendix). The components of the pole motion extracted by SSA are ranked according to their “importance” (that is their amplitude, in a broad sense): first the annual oscillation (SSA component 1, prograde and retrograde), then the Chandler term (SSA component 2, wobble), then with the same “importance” the trend $y(t)$ that can be seen as the component of the mean pole drift (SSA component 3). Two other components, with an “importance” two times smaller, are identified mainly along the axis m_y (they are smaller, by about one order of magnitude, along the axis m_x):

- an oscillating component with segments containing quasi-pure sinusoidal variations with a 11-year period and a characteristic amplitude of $\sim 0.5 \cdot 10^{-12} \text{ rad}\cdot\text{s}^{-1}$, which we note m_y^{11} , and a central segment presenting a sine variation with a 22-year period; the amplitude of the latter component is $\sim 1.4 \cdot 10^{-12} \text{ rad}\cdot\text{s}^{-1}$; we note it m_y^{22} (SSA component 4, see Fig. 2). During the reconstruction process of each component obtained with SSA, we regroup them according to the magnitude of their associated eigenvalues (see Appendix). That is why, in this specific case, an oscillating variation composed of two distinct periods is reconstructed;
- a $m_y^{5.5}$ component with a smaller characteristic amplitude ($\sim 4 \cdot 10^{-13} \text{ rad}\cdot\text{s}^{-1}$), displaying 5.5-year modulated periodic sine oscillations (SSA component 5, see Fig. 3).

3. A tentative interpretation

We consider here that the periodic oscillations (with periods 22 years, 11 years and 5.5 years) detected in the

series of pole positions are manifestations of solar activity for lack of any reasonable alternative. In this short note – devoted essentially to unveiling these periodicities – we do not attempt to propose a full explanation of their presence in the rotation pole series, but only to evoke some elements of a possible mechanism. First, a direct electromagnetic torque exerted by the solar wind on the Earth's main magnetic field, *i.e.* on the conducting core (the dipole according to Jady, 1970), transmitted to the mantle through some kind of core–mantle coupling (*e.g.*, Jackson, 1997; Jault and Le Mouél, 1991) could be envisioned. We have mentioned in the introduction the first attempt made by Jady, concerning I.o.d., but this process appears to lead to insufficient torques. A second mechanism, which we now elaborate somewhat further, involves an action of solar activity on the atmosphere, resulting in flows, or winds, and a transfer of angular momentum from the atmosphere to the mantle.

3.1. Schematic atmospheric flows

Since the vertical component of large scale winds in the atmosphere is small compared with the horizontal one, and the thickness of the different layers is small compared with the Earth's radius r , we assume velocities $\vec{u}$ to be purely horizontal and adopt a thin layer approximation (Chedin et al., 1982). A layer of thickness h , between radii a and b , is replaced by a surface layer of density $H = \int_a^b \rho(r) dr$ ($\text{kg}\cdot\text{m}^{-2}$). The horizontal divergence-free velocity field $\vec{u}_H(\theta, \varphi, t)$ can be represented by an expansion in elementary toroidal vectors,

$$\vec{u}_H(\theta, \varphi, t) = \sum_n t_n^{m,c}(t) \vec{\theta}_n^{m,c} + t_n^{m,s}(t) \vec{\theta}_n^{m,s} \quad (1)$$

with:

$$\vec{\theta}_n^{m(c,s)} = \vec{OP} \wedge \vec{\nabla}_H P_n^m(\cos\theta, \sin\theta) \quad (2)$$

$\vec{OP} = \vec{r}$ (O the Earth's center), $P_n^m(\cos\theta)$ is the associated Legendre function of the first kind. The t_n^m are coefficients

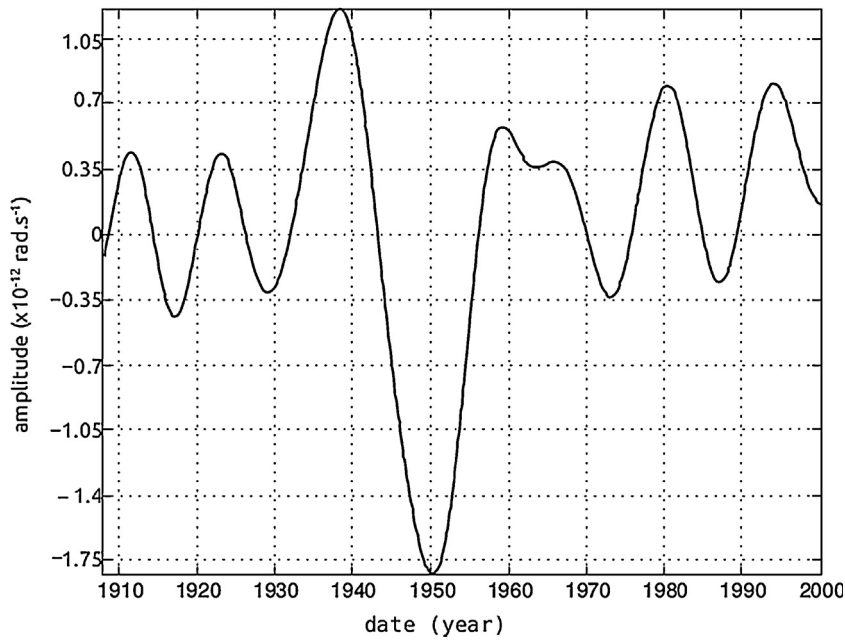


Fig. 2. Oscillating signal extracted, using SSA, from m_y : “component 4” displays quasi-pure oscillations with a 11-yr period, and, in the central part, one oscillation with a 22-yr period.

in $\text{m}\cdot\text{s}^{-1}$. $\vec{\nabla}_H$ is the horizontal gradient. Now, we are interested in the angular momenta of $\rho\vec{u}_H$ around the two equatorial axes Ox and Oy (see Fig. 1). In the frame of our approximation,

$$\sigma_y(t) = H \int_0^{2\pi} \int_0^\pi \vec{r} \wedge \vec{u}_H(\theta, \varphi, t) r^2 \sin\theta \, d\theta \, d\varphi \quad (3)$$

$\times (\text{in } \text{kg}\cdot\text{m}^2\cdot\text{s}^{-1})$

$\hat{y}$ is the unit vector of Oy axis; a similar expression holds for $\sigma_x(t)$. It is easily seen that only the elementary velocity vector,

$$t_1^{1s}(t)\vec{\theta}_1^{1s} = t_1^{1s}(t)\vec{r} \wedge \vec{\nabla}_H \sin\theta \sin\varphi \quad (4)$$

contributes to the angular momentum of $\vec{u}_H$ around Oy. We will write this term in the equivalent form,

$$t_1^{1s}(t)\vec{\theta}_1^{1s} = \omega(t)y\hat{r} \wedge \vec{r} \quad (5)$$

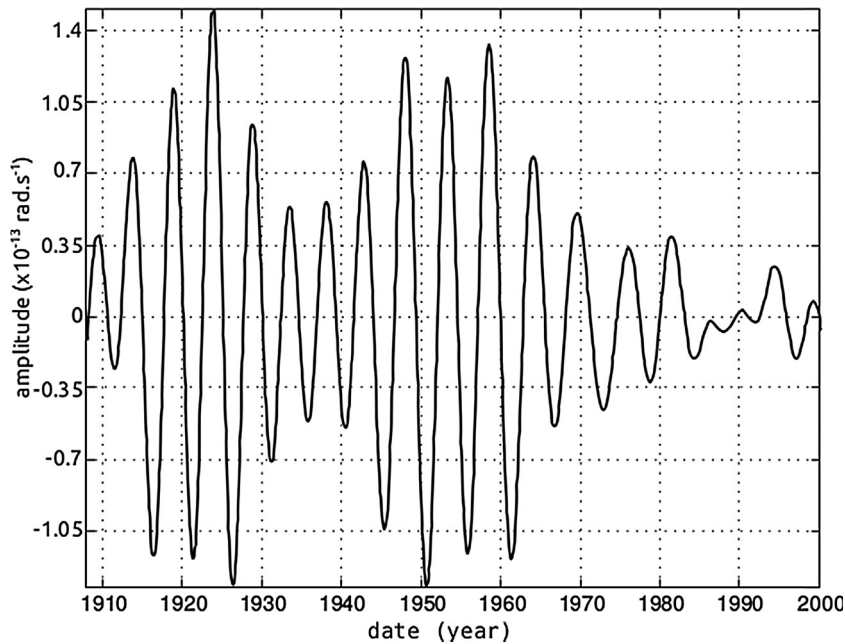


Fig. 3. Oscillating signal extracted, using SSA, from m_y : “component 5” displays quasi-pure oscillations with a 5.5-yr period.

$\omega(t)$ in rad.s^{-1} and $t_1^1 = r\omega_y$, which makes it clear that this term is simply a solid rotation around Oy. The same holds for ω_x and $\omega_x \wedge \vec{r}$.

3.2. Estimates of atmospheric winds

Applying SSA analysis (see [Appendix](#)) to the evolution of m_y , the equatorial components of the mantle rotation axis in the 90° east meridian, we have found three periodic components with periods 11 years, 22 years, and 5.5 years, with amplitudes directly measurable in [Figs. 2 and 3](#) (in $10^{-12} \text{ rad s}^{-1}$). These components are about one order of magnitude smaller for m_x . Multiplying these amplitudes by the equatorial inertia momentum of the mantle ($A = 8 \cdot 10^{37} \text{ kg m}^2$), we obtain the amplitudes of the corresponding periodic variations of the equatorial angular momenta of the mantle,

$$\begin{aligned} Am_y^{11} &\sim 3.04 \cdot 10^{25} e^{i\Omega t}, \quad \Omega = \frac{2\pi}{11 \cdot \pi \cdot 10^7} \\ Am_y^{5.5} &\sim 0.56 \cdot 10^{25} e^{i\Omega t}, \quad \Omega = \frac{2\pi}{5.5 \cdot \pi \cdot 10^7} \quad (\text{in } \text{kg} \cdot \text{m}^2 \cdot \text{s}^{-1}) \\ Am_y^{22} &\sim 8.80 \cdot 10^{25} e^{i\Omega t}, \quad \Omega = \frac{2\pi}{22 \cdot \pi \cdot 10^7} \end{aligned} \quad (6a)$$

The same computations give for the x component:

$$\begin{aligned} Am_x^{11} &\sim 7.02 \cdot 10^{24} e^{i\Omega t}, \quad \Omega = \frac{2\pi}{11 \cdot \pi \cdot 10^7} \\ Am_x^{5.5} &\sim 0.12 \cdot 10^{24} e^{i\Omega t}, \quad \Omega = \frac{2\pi}{5.5 \cdot \pi \cdot 10^7} \quad (\text{in } \text{kg} \cdot \text{m}^2 \cdot \text{s}^{-1}) \\ Am_x^{22} &\sim 5.63 \cdot 10^{24} e^{i\Omega t}, \quad \Omega = \frac{2\pi}{22 \cdot \pi \cdot 10^7} \end{aligned} \quad (6b)$$

Let us come back to our simplified flow in the atmosphere. We consider that the mantle-atmosphere system is isolated (neglecting the changes in the small angular momentum of the core); the changes in the equatorial angular momentum of the mantle are attributed to exchanges of momentum between the mantle and an atmospheric layer to which solar activity communicates motions with solar periods and the geometry given by formula (5). And, aiming only at orders of magnitude, we simply write that changes in angular momenta of the (schematic) atmospheric flows, σ_y^{11} , $\sigma_y^{5.5}$ and σ_y^{22} , balance the corresponding mantle angular momentum periodic variations Am_y^{11} , $Am_y^{5.5}$ and Am_y^{22} (the same for the smaller x components). Now, the expressions of the σ are (taking e.g., σ_y and the 11-yr period),

$$\begin{aligned} \sigma_y^{11} &= H \\ &= H \omega_y^{11} r^4 \int_0^{2\pi} \int_0^\pi (1 - \sin^2 \theta \sin^2 \varphi) d\theta d\varphi e^{i2\pi t / 11 \cdot \pi \cdot 10^7} \quad (7) \\ &\sim 14.8 H \omega_y^{11} r^4 e^{i2\pi t / 11 \cdot \pi \cdot 10^7} \quad (\text{in } \text{kg} \cdot \text{m}^2 \cdot \text{s}^{-1}) \end{aligned}$$

Equating σ_y^{11} and Am_y^{11} (and dropping the exponential factor), we obtain:

$$H \omega_y^{11} \sim Am_y^{11} / 14.8 r^4 \quad (8)$$

The numerical value of Am_y is given by (6). The same expressions hold for the other lines $\omega_y^{5.5}$ and ω_y^{22} , and the x components. Of course, the angular velocity ω cannot be found independently of H .

4. Discussion

Wind velocities in the atmosphere have long been measured (e.g., [Drob et al., 2015](#); [Pichugina et al., 2012](#); [Talaat and Lieberman, 2010](#)). Zonal wind velocities can reach 60 ms^{-1} ([Fleming et al., 1990](#); [Gross et al., 2005](#); [Salby and Callaghan, 2000](#)). They do not contribute to the t_1^1 components. Meridional flows do contribute, but they are much smaller, of the order of a few ms^{-1} ([Gudiksen et al., 1968](#); [Nastrom et al., 1982](#); [Oort and Rasmusson, 1970](#)). Let us consider the periodic flows computed in Section 3, in the case when the layer concerned by the flow is the troposphere. Then $H \sim 10^4 \text{ kg m}^{-2}$, and from formula (8), we get:

$$\omega_y^{11} \sim 10^{-4} \times 3.04 \times \frac{10^{25}}{14.8} (6.5 \cdot 10^6)^4 \sim 10^{-7} \text{ s}^{-1} \quad (9)$$

In the same way, $\omega_y^{5.5} \sim 0.18 \cdot 10^{-7} \cdot \text{s}^{-1}$, $\omega_y^{22} \sim 2.7 \cdot 10^{-7} \cdot \text{s}^{-1}$.

In terms of linear velocities u_H , the orders of magnitude, obtained simply by multiplying ω by $r = 6.5 \cdot 10^6 \text{ m}$, are

$$u^{11} \sim 0.6 \text{ m} \cdot \text{s}^{-1}, \quad u^{5.5} \sim 0.11 \text{ m} \cdot \text{s}^{-1}, \quad \text{and } 1.8 \text{ m} \cdot \text{s}^{-1} \quad (10)$$

The periodic flows computed in Section 3 in the frame of our schematic model take the form of rigid equatorial rotations, a kind of image of the equatorial rotations of the solid mantle. They have magnitudes compatible with measured meridional flows in the troposphere. Such flows, if present, might be difficult to evidence directly, taking into account, in particular, the long (solar) periods involved. In fact, as noticed in [Le Mouél et al. \(2010\)](#), the solid mantle would play the role of an integrator, revealing these periodic large-scale motion variations (a velocity field is transformed into a rotation vector). We can call also upon stratospheric winds. Indeed, up to now, we have skipped a major step of the proposed process, the means by which solar activity can affect circulation in the atmosphere. We recall that the view according to which the solar cycle can affect the climate is not favorably received, for the main reason that during a solar cycle total solar irradiance varies by only 10^{-3} . But, first, the different wavelengths are not modulated by the same factor during a cycle. Second, electromagnetic irradiance is not the only acting power to take into consideration; one may call upon solar wind, galactic cosmic rays, and induced effects on ionospheric electric currents, acting in particular on clouds ([Gray et al., 2010](#)). We elaborate only briefly on the mechanism we invoke above, which requires interactions, couplings, between different layers of the atmosphere. Variations in solar UV radiation influence directly the ionosphere; links between processes having their seats in the ionosphere and in the stratosphere have been “demonstrated” ([Rishbeth, 2006](#)). “Transfers of energy and momentum happen” ([Forbes, 2000](#); [Gray et al., 2010](#)). In fact, much interest is presently devoted to the

coupling between the stratosphere and the ionosphere up to altitudes of several hundred kilometers. Coupling between stratosphere and ionosphere has been evidenced from a number of different sources (e.g., Baldwin and Dunkerton, 2001, Goncharenko et al., 2010); Kuroda and Koder (2005), in a long paper devoted to solar influence on climate, write “*Understanding the physical mechanisms connecting the lower and upper atmosphere will be an essential part of future models...*” and “*Considering the upper atmosphere as a part of a complex and evolving Earth system is a necessary step for further advances...*”. Our findings imply some form of coupling between the solid Earth and the atmosphere. Such coupling is established through the observations of a good correlation between some l.o.d. variations and the angular momentum of the atmosphere for periods up to a few years. Indeed, some l.o.d. variations are attributed to exchanges of angular momentum between the mantle and the atmosphere (e.g., in Lambeck, 2005).

5. Conclusion

We have found evidence of periodic solar components in the mantle pole position time series (with periods of 5.5 years, 11 years and 22 years): oscillations occur mainly along the equatorial 90° east meridian, with an amplitude of $0.5 \cdot 10^{-12} \text{ rad s}^{-1}$ for the 11 year-oscillations, of $1.5 \cdot 10^{-13} \text{ rad s}^{-1}$ for the 5.5-year oscillation, and of $2 \cdot 10^{-12} \text{ rad s}^{-1}$ for the 22-year oscillations. We argue that this oscillation (wobble) of the solid mantle can result from an exchange of angular momentum with an equatorial flow in the terrestrial atmosphere, an image of mantle libration. With this hypothesis, we obtain an order of magnitude of 1 m s^{-1} for the flows (wind velocities), compatible with meridional velocities measured in the atmosphere. These solar-generated atmospheric winds would not be small.

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Appendix A. Singular spectrum analysis

A thorough account of singular spectrum analysis (SSA) is given for instance in Vautard and Ghil (1989) and Vautard et al. (1992). SSA is a powerful, non-parametric, time series analysis tool that relies on the Karhunen–Love spectral decomposition (Kittler and Young, 1973). We summarize it here for the reader who would not be familiar with the method.

Consider the real-valued time series $\mathbf{x} = \{x_n; n = 1, \dots, N\}$; SSA decomposes vector $\mathbf{x}$ into a sum of physical components that may be subsequently identified as trends and quasi-periodic oscillations. The starting point is to embed $\mathbf{x}$ into an L -dimensional vector space, by using lagged copies of $\mathbf{x}$ L that

are $K = N - L + 1$ long. In other words, form the following “trajectory” (Hankel) matrix $\mathbf{X}$:

$$\mathbf{X} = \begin{pmatrix} x_1 & \cdots & x_k \\ \vdots & \ddots & \vdots \\ x_L & \cdots & x_n \end{pmatrix}$$

The next step consists in performing the singular value decomposition (Golub and Kahan, 1965) of the trajectory matrix $\mathbf{X}\mathbf{X}^T$; we obtain a collection of eigen-triplets $\{\lambda_i; i = 1; \dots; L\}$ of $\mathbf{X}\mathbf{X}^T$, which are ranked in the decreasing order of magnitude ($\lambda_1 \geq \lambda_2 \geq \dots \geq \lambda_L \geq 0$). The eigen-triplets are linked by $V_i = \mathbf{X}^T U_i / \sqrt{\lambda_i}$. Once this decomposition is made, the last stage consists in the construction of groups of components. Each group is a subset of matrices X_i ; let $I = \{i_1, \dots, i_p\}$, be a group of indices $\{i_1, \dots, i_p\}$, the corresponding matrix $X_I = X_{i_1} + \dots + X_{i_p}$, thus the matrix becomes $\mathbf{X} = X_{I_1} + \dots + X_{I_d}$. The choice of the indices in each group is done empirically; for that, we may find it interesting to group the eigen-triplets that have common eigenvalues.

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