

## RESEARCH LETTER

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## Key Points:

- We simultaneously image the La Soufrière de Guadeloupe lava dome with three muon telescopes aiming at the volcano from different locations
- We jointly invert the multitelescope muon data with gravity data to retrieve the three-dimensional density distribution in the dome
- We find a low-density, mechanically weak region where the hydrothermal system is most active, suggestive of an edifice collapse hazard

## Supporting Information:

- Supporting Information S1

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## Three-dimensional density structure of La Soufrière de Guadeloupe lava dome from simultaneous muon radiographies and gravity data

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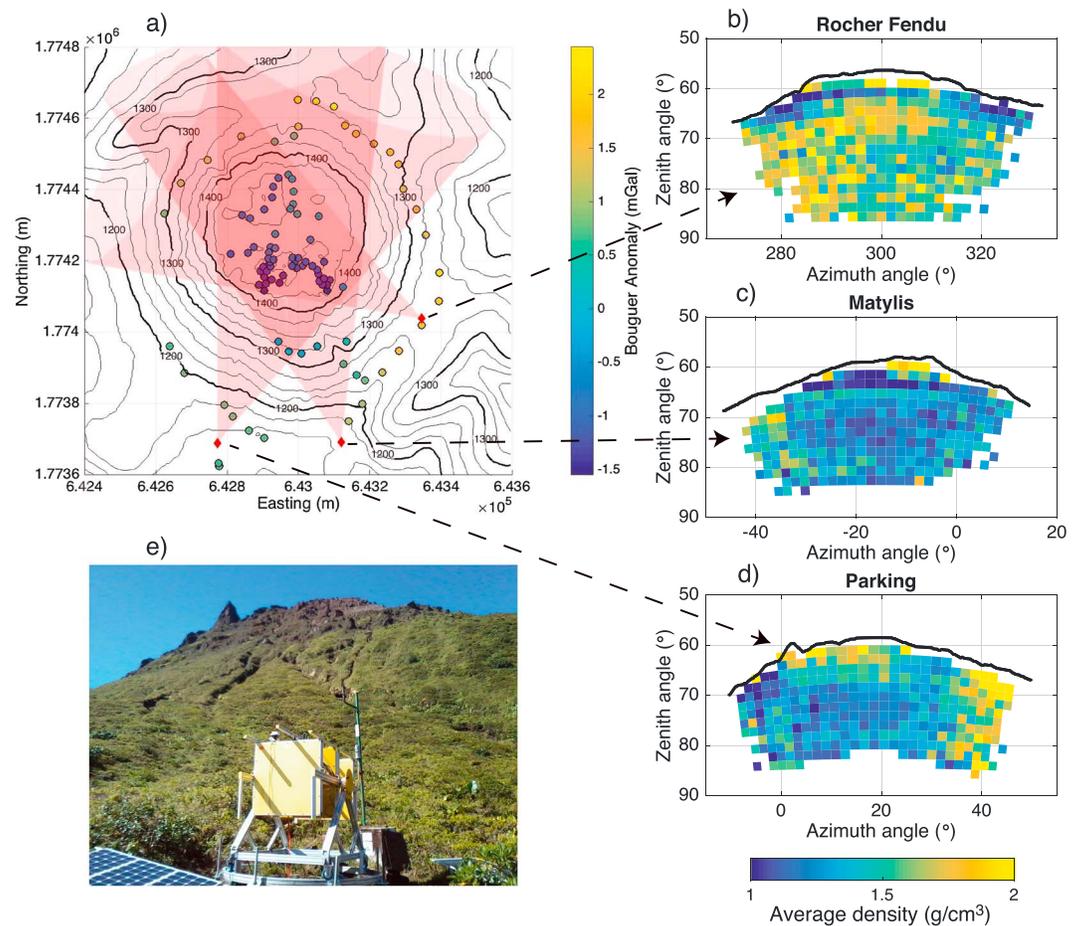
**Abstract** Muon imaging has recently emerged as a powerful method to complement standard geophysical tools in the understanding of the Earth's subsurface. Muon measurements yield a “radiography” of the average density along the muon path, allowing to image large volumes of a geological body from a single observation point. Here we jointly invert muon data from three simultaneous telescope acquisitions together with gravity data to estimate the three-dimensional density structure of the La Soufrière de Guadeloupe lava dome. Our unique data set allows us to achieve an unprecedented spatial resolution with this novel technique. The retrieved density model reveals an extensive, low-density anomaly where the most active part of the volcanic hydrothermal system is located, supporting previous studies that indicate this region as the most likely to be involved in a partial edifice collapse.

### 1. Introduction

The La Soufrière de Guadeloupe volcano (Lesser Antilles) is an explosive subduction volcano holding one of the most hazardous volcanic hydrothermal systems in the world [Loughlin *et al.*, 2015]. It currently experiences an increasing unrest, raising concern on possible pore-fluid overpressurization and flank instability. Knowledge of the density distribution in its upper structure (lava dome) is crucial for hazard assessment since it can help to determine mechanically weak regions that may result in flank collapse. Volcano density structures are traditionally studied with gravity data, and 3-D models solely based on these data are highly nonunique and low resolution, especially since field conditions often make it difficult to achieve a good data coverage.

Since a decade, muon imaging has been developed by several teams to probe the density structure of various geological bodies [e.g., Nagamine *et al.*, 1995; Tanaka *et al.*, 2001; Lesparre *et al.*, 2010]. The principles of the method are common to those of X-ray medical imaging, where particles are absorbed differentially according to the matter's density along their travel path [Nagamine *et al.*, 1995]. The method exploits the properties of muon interaction with ordinary matter, in particular, a small cross section and a quasi-linear propagation path [Nagamine *et al.*, 1995]. Experimental muon trajectories are inferred from a series of pixelated particle detectors, such that the lava dome is sampled in hundredths of directions delimited by a cone (Figure 1a), and scanned typically in a few weeks [Lesparre *et al.*, 2010]. The number of muons detected is used to compute the rock average density for each trajectory, resulting in 2-D average-density radiographies (Figures 1b–1d).

Among all projects using the muon imaging technique, DIAPHANE has been a pioneer since almost a decade in designing, producing, and deploying portable, autonomous, and robust detectors (so-called muon “telescopes”) on the slopes of active volcanoes: the La Soufrière de Guadeloupe in the Lesser Antilles [Lesparre *et al.*, 2012b; Jourde *et al.*, 2013, 2016], the Mayon in the Philippines, and the Etna in Italy [Carbone *et al.*, 2014]. Despite difficult field conditions, data acquisition is continuous, deadtimeless with a high duty cycle (more than 95%), and only limited by electrical power blackouts during the worst weather periods. Details on the methods and detectors developed by DIAPHANE may be found in Lesparre *et al.* [2012c], Jourde *et al.* [2013], and Marteau *et al.* [2014, and references therein].



**Figure 1.** (a) Location of muon telescopes (diamonds) and gravity stations (circles) used to reconstruct the La Soufrière de Guadeloupe lava dome density distribution. Gravity stations are colored according to their Bouguer anomaly. Red cones indicate the region of the dome sampled by each telescope. (b–d) Muon radiographies obtained with the three telescopes. Each pixel corresponds to a particular line of sight of the telescope. The average density is computed from the opacity, which is inferred from the number of muons detected, and the total length of rock crossed by the muons. (e) The muon telescope at the “Parking” site.

Information about the 3-D density distribution of a geological body may be obtained by combining more than one telescope acquisitions [e.g., Tanaka *et al.*, 2010] or jointly inverting the muon data with gravity data [e.g., Nishiyama *et al.*, 2014, 2017]. Nishiyama *et al.* [2017] jointly inverted data from a single muon telescope with gravity data to study a volcanic lava dome. Their model has a limited spatial resolution due to the few data available (30 muon data points and 30 gravity data points). A study on the resolving kernels of muon and gravity data joint inversion by Jourde *et al.* [2015] suggests that in the case of an ideal coverage of the geological body by several muon telescopes, gravity data do not significantly improve the inversion results. However, this might not be the case when the muon radiographies are obtained from a limited number of locations.

Here we jointly invert muon data from three acquisition sites together with gravity data points (Figure 1a) to study the 3-D density distribution in the La Soufrière lava dome. We first present the data sets, the forward modeling, and the inversion method developed to process the muon and gravity data. We emphasize the sensitivity of the inversion to the regularization scheme and its relation with the distance between the measurement point and each voxel. In a second part we perform a synthetic test to cross-check the methodology and constrain the regularization parameters. The last part presents the results obtained in the 3-D reconstruction of the dome and a discussion on the model’s coherence with known geological structures and also identified limitations of the present data processing.

## 2. Data Acquisition and Processing

Muon data were simultaneously acquired by three muon telescopes at the “Parking,” “Matylis,” and “Rocher Fendu” sites (Figure 1a) between June and August 2015 and for a duration of 63, 36, and 25 days, respectively. The measurement period was enough to obtain representative statistics of the outgoing muon flux in each direction. We applied the methodology proposed by *Jourde et al.* [2013] to filter out the upward flux of muons, which otherwise introduces a significant bias in the density estimation. The resulting average density as a function of the muon direction relative to the telescope orientation is presented in Figures 1b–1d. The telescopes used in this analysis (Figure 1e) count  $31 \times 31$  lines of sight, but in practice the number of data points per muon telescope is lower since some lines of sight point above the volcano or to regions where the rock thickness is too high. In our case, a total of 931 muon data points are obtained combining the three telescopes. We estimate the error for each muon data point as a combination of the measurement error and a model error related to the approximations done to solve the forward problem [*Jourde*, 2015].

The muon radiographies display density values that vary between  $0.7$  and  $3 \text{ g cm}^{-3}$ , with average densities of  $1.58$  (Rocher Fendu),  $1.32$  (Matylis), and  $1.43$  (Parking)  $\text{g cm}^{-3}$ . These values are low compared to the typical density of  $2.7 \text{ g cm}^{-3}$  found for andesitic rocks at La Soufrière [*Komorowski et al.*, 2008]. Rock alteration and dissolution caused by hydrothermal fluids and the presence of large fractures and cavities in the dome may strongly reduce the bulk density of a lava dome as reported by *Ball et al.* [2015] and *Voight et al.* [2002]. Indeed, using a new analysis of data from the last (1976) phreatic eruption [cf. *Rosas-Carbalaj et al.*, 2016], we reestimated the total ejected volume of this eruption to be  $5\text{--}6 \times 10^6 \text{ m}^3$ . This large amount of nonmagmatic material implies that an extensive volume of the dome is highly altered and that large cavities prevail. *Gailler et al.* [2013] estimate the average density of the La Soufrière complex to be  $1.8 \text{ g cm}^{-3}$  based on regional gravity surveys. A low average density is thus not surprising; however, we also point out the existence of a diffusive flux of muons, which can bias the estimated density toward lower values [*Nishiyama et al.*, 2014, 2016]. This phenomenon is currently being studied with dedicated simulations, which show that in the case of La Soufrière it can effectively shift the inferred densities to lower values. We therefore treat the densities displayed in the radiographies as relative rather than absolute. Low- and high-density regions can be clearly identified in the three radiographies, even though no spatial smoothing has been applied to the data processed.

A total of 109 gravity measurements were done between March 2014 and February 2015 with a Scintrex CG-5 gravimeter. Measurements have a precision of the order of  $10 \text{ }\mu\text{Gal}$ . Details of the data processing prior to inversion, including drift, tides, and ellipsoid and topography corrections are given by *Jourde* [2015]. After processing, we retained 103 good-quality data points to perform the inversion. The correlation of the Bouguer anomaly with respect to topography evaluated for different trial densities [*Jourde*, 2015; *Nettleton*, 1939] is minimum for an optimal Earth density of  $1.75 \text{ g cm}^{-3}$  [*Parasnis and Cook*, 1952]. This value is smaller than the estimate given by *Gunawan* [2005] in their regional study comprising data points from the whole island of Guadeloupe, but we use our local estimation to focus on the density distribution restricted to the La Soufrière lava dome. The resulting Bouguer gravity anomalies are presented in Figure 1a. Lowest gravity anomalies are found in the southern region of the summit, whereas high gravity anomalies are localized at the north and west sides of the dome’s base.

## 3. Forward Modeling and Inversion

The muon forward problem involves calculating a convolution of the incident muon open-air flux, their energy loss along the trajectory inside the target, and the telescope acceptance [e.g., *Lesparre et al.*, 2012a]. The relevant parameters are, on the target side, the opacity,  $\rho = \int_L \rho(\xi) d\xi$  (where  $\rho$  is the density and  $L$  the particle path length), and on the probe side, the incident muon flux. The differential muon flux associated with the line  $\mathbf{t} = (\mathbf{r}, \varphi, \theta)$  ( $\mathbf{r}$  is the measurement point and  $\varphi$  and  $\theta$  are the azimuth and zenith angles of the trajectory, respectively) may be expressed as a function  $\delta\phi_{\mathbf{t}} = \frac{\partial^3 \phi}{\partial \Omega \partial S}(\rho, \varphi, \theta) (\text{s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1})$  ( $\Omega$  is the solid angle and  $S$  the surface) that accounts for the muon flux reaching the instrument. This muon flux can be computed analytically at the surface of the Earth [e.g., *Tang et al.*, 2006; *Shukla*, 2016], but here we use a more precise model obtained with the Corsika simulation code [*Heck et al.*, 1998] which accounts for the altitude of the lava dome. To simulate the energy loss of the muons interacting with matter for a given opacity, we use the Monte Carlo transport code Geant4 [*Agostinelli et al.*, 2003]. Finally, to estimate the outgoing flux from the number

of muon hits measured, we use an experimental acceptance function, calibrated on the field for each telescope [Lesparre et al., 2012a].

Numerical tests using Geant4 show that the relation between observed opacities and outgoing flux for each telescope line of sight can be treated as linear in the  $\pm 2$  standard deviations range. Our linear inversion fits the data within the  $\pm 1$  standard deviation range. We thus use the relation between observed outgoing flux and opacities for  $\pm 2$  standard deviations to estimate the Jacobian around this point. We then need to compute the opacity resulting from a 3-D density distribution. For this we discretize the lava dome in cubes of  $8 \times 8 \times 8 \text{ m}^3$  using a digital elevation model with a 5 m precision and compute the relative contribution of each cube to the total opacity for each muon trajectory.

For the gravity case, we decouple the problem between the region close to the measurements and the region outside, used to compute the topography correction [cf. Jourde, 2015]. In the vicinity of the volcano we use a polyhedral discretization based on the 5 m resolution digital elevation model. We compute each polyhedra contribution to each gravity station with the program by D'Urso [2013, 2014].

The muon and gravity data are sensitive to the same physical parameter; thus, a joint inversion can be achieved by combining their forward kernels. We use the same model parametrization for the gravity and muon problems such that

$$G \begin{bmatrix} \rho_\mu \\ \Delta\rho \end{bmatrix} = \begin{bmatrix} G_g \\ G_\mu \end{bmatrix} \begin{bmatrix} \rho_\mu \\ \Delta\rho \end{bmatrix} = \begin{bmatrix} \mathbf{d}_g \\ \mathbf{d}_\mu \end{bmatrix} = \mathbf{d}, \quad (1)$$

where  $G$  is the forward kernel,  $\mathbf{d}$  the data vector, subscripts  $g$  and  $\mu$  denote the gravity and muon case, respectively, and  $\rho_\mu$  is the density distribution discretized in cubes of  $16 \times 16 \times 16 \text{ m}^3$  (forward kernels are built such that the same discretization is used). The scalar parameter  $\Delta\rho$  accounts for a possible density offset between the gravity- and muon-inferred models due to the negative density bias caused by the diffusion of low-energy muons on the volcano surface. Accounting for this offset, the density models are related through  $\rho_g = \rho_\mu + \Delta\rho$ . Our linear inversion minimizes an objective function consisting of a weighted data misfit term and a model regularization term:

$$\phi(\mathbf{m}) = (\mathbf{d} - \mathbf{Gm})^T C_d^{-1} (\mathbf{d} - \mathbf{Gm}) + \epsilon^2 (\mathbf{m} - \mathbf{m}_{\text{prior}})^T C_\rho^{-1} (\mathbf{m} - \mathbf{m}_{\text{prior}}), \quad (2)$$

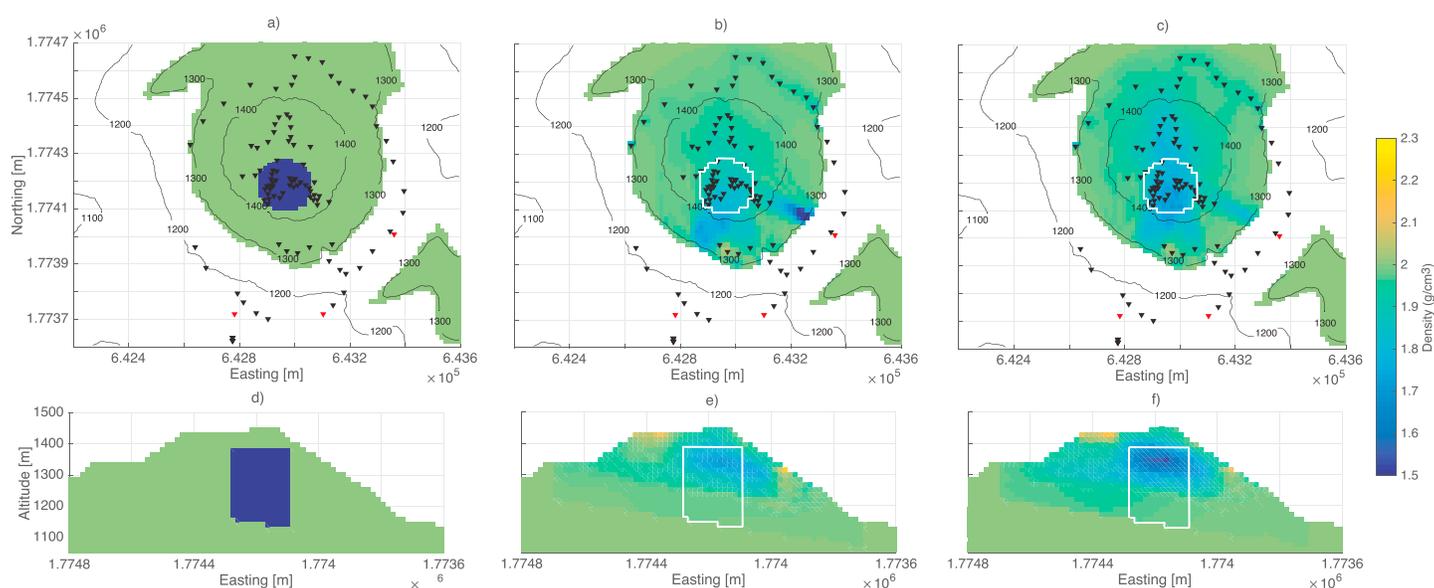
where  $\mathbf{m} = [\rho_\mu, \Delta\rho]$ ,  $\mathbf{m}_{\text{prior}}$  is the a priori density model and density shift,  $C_d^{-1}$  is the data covariance matrix,  $C_\rho^{-1}$  is the model regularization matrix, and  $\epsilon^2$  the trade-off parameter, which establishes the relative weight of the regularization in the cost function [Menke, 2012]. In practice, we adjust this parameter to obtain the largest possible contribution of the model regularization while fitting the data to the error level. To determine the prior model, we searched for the homogeneous density value that minimizes the density offset. We found that a value equal to the Earth density used to compute the Bouguer anomalies results in a minimum offset. Accordingly, we used a prior value of 0 for the density offset. The model regularization consists of a combination of smoothness constraints, achieved by penalizing the difference in density between adjacent cubes, and damping, which favors density models that are close to  $\mathbf{m}_{\text{prior}}$ .

Being sensitive to the opacity, muon data have no inherent information on the density distribution along the muon path. Similarly, gravity data have no inherent depth resolution. It is known that in the regularized inversion of potential field data the lack of depth information causes the retrieved structures to concentrate near the surface. This problem can be overcome by using a scaling of the regularization matrix to counteract the natural decay of the sensitivity with depth. Since the original application on the inversion of magnetic data by Li and Oldenburg [1996], this scaling has been applied to other types of potential field data. Here we combine a depth scaling as introduced by Li and Oldenburg [1996] to compensate for the decay of sensitivity in the gravity data, with a distance-to-telescope scaling to compensate for the same effect in the muon data. Details on these scalings are given in the supporting information [Boulanger and Chouteau, 2001; Ghalehnoee et al., 2017]. We emphasize that this is the first time that such scalings are implemented in an inversion of muon data. Previous studies deal with this problem by reducing the amount of damping in the zone where anomalies are expected, thus severely biasing the result of the inversion.

## 4. Results and Interpretation

### 4.1. Synthetic Tests

Before inverting the real data, synthetic tests are performed to assess the influence of the regularization scaling and the resolution that can be expected. Figures 2a and 2d show slices of one of the models used to



**Figure 2.** Results of the synthetic data inversions used to study the scaling applied to the regularization matrix. Columns correspond to different slices of the density models. (a, d) True model used to generate the synthetic data. (b, e) Density model obtained from an inversion without applying any scaling to the regularization matrix. (c, f) Density model obtained from an inversion using the regularization matrix scaling. Both models fit the data to the same level.

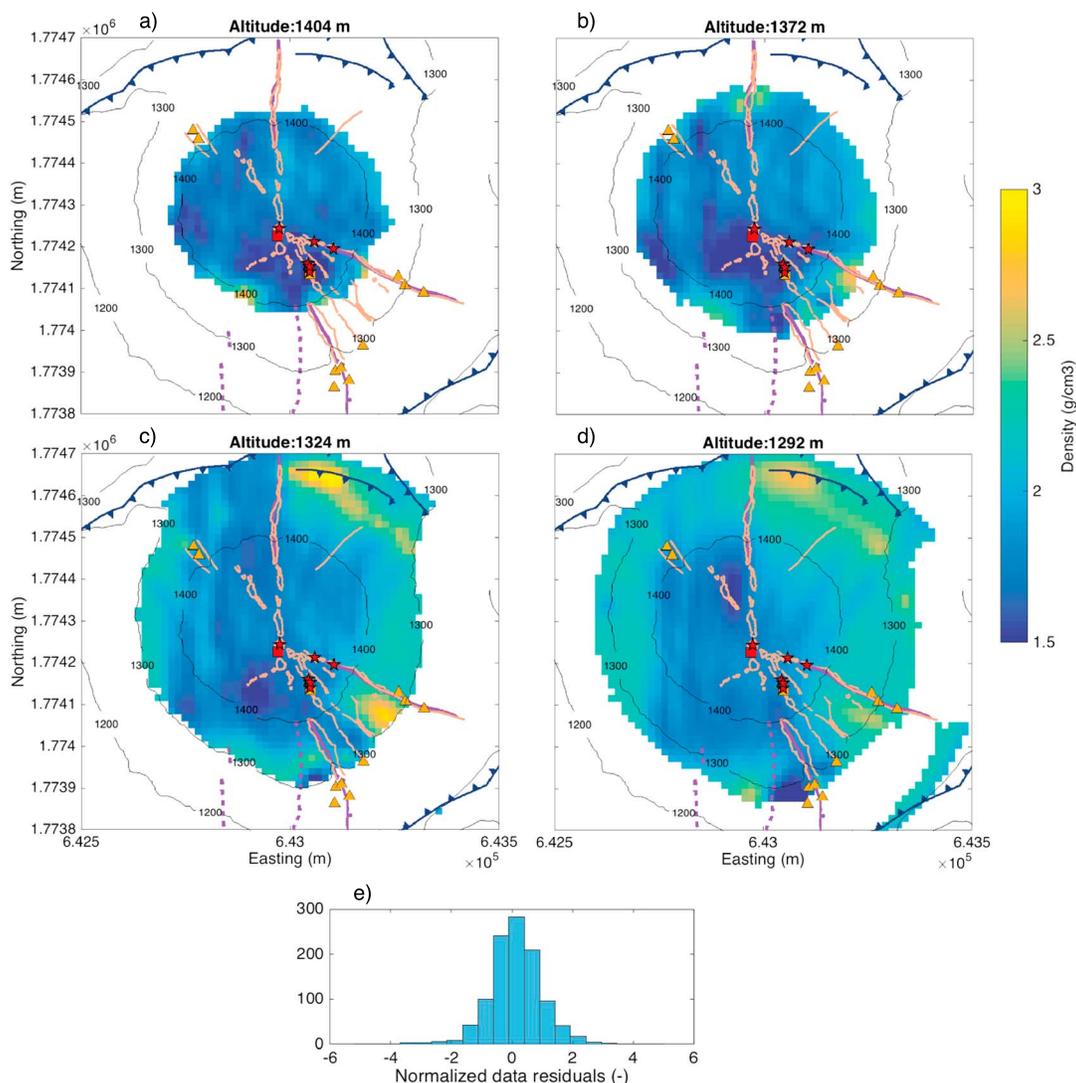
generate the synthetic data contaminated by errors with the same standard deviation as estimated for the real data. Without scaling, the inversion places the largest density contrasts close to the muon telescopes (Figures 2b and 2e). In turn, the matrix scaling places the density contrasts where the true anomaly is located (Figures 2c and 2f). Other tests with smaller and larger radii of the cylinder and with its center close to one of the muon telescopes did not change these results. Since the three telescopes point northward, there is a more pronounced underdetermination in the north-south direction which results in a spreading of the density anomaly in this direction. Inversions without the gravity data (tests not shown) confirm the importance of these data in the delimitation of the density anomalies in cases where the spatial coverage of the telescopes is not complete.

The retrieved anomaly has a shorter depth extension than the true one (Figures 2d and 2f). Indeed, the position and inclination of the telescopes determine a plane above which the density is explored. In our case, the telescope in the Parking site is the one with the lowest altitude. Its lowest plane can be inferred from Figure 2f where the anomaly begins. We therefore interpret the density models only above this plane. Furthermore, given the large nonuniqueness of the problem and the regularization used, we only interpret our model in terms of low- and high-density regions without asserting their precise delimitations.

#### 4.2. Inversion of Real Data

The density model resulting from the joint inversion of muon and gravity real data is shown in Figures 3a–3d. The root mean square error equals 1, and the fit is of the same quality for both the muon and gravity data (Figure 3e). The offset parameter equals  $0.47 \text{ g cm}^{-3}$ , and the average density is of  $2.09 \text{ g cm}^{-3}$ , a more realistic value than  $1.6 \text{ g cm}^{-3}$  retrieved from the muon data alone and consistent with the average densities of Figures 1b–1d. Inverting the gravity and muon data sets without the density offset parameter produces a model with density contrasts similar to the model shown in Figure 3 down to an altitude of  $\sim 1300 \text{ m}$ . However, below this altitude and in order to improve the fit of the gravity data, the inversion places an artificial high-density region that mimics the role of the offset parameter.

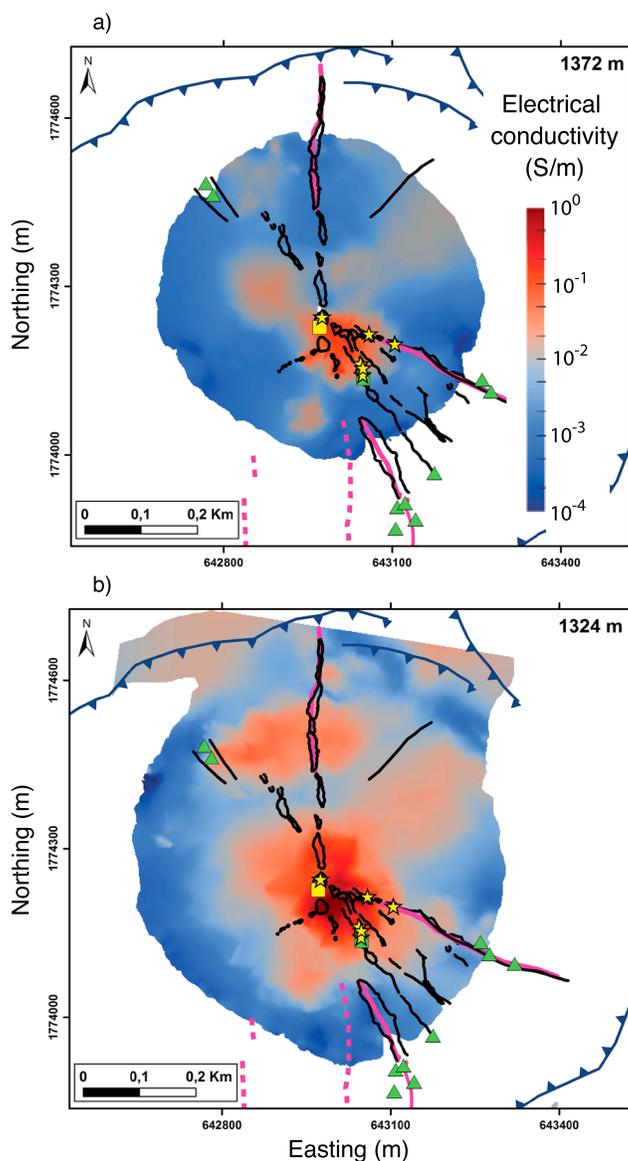
In the presence of strong topography variations, relative gravity data can be used to retrieve an absolute density model [e.g., Linde *et al.*, 2014]. Simple synthetic tests with anomalies like the one shown in Figure 2 suggest that the altitude differences of  $\sim 150 \text{ m}$  between the dome's summit and base are sufficient to retrieve absolute density values. Extending this conclusion to complex density distribution is not straightforward, and we prefer to restrain our interpretations on relative density variations, although the densities retrieved by the inversion fall in the expected range (see Figure 3).



**Figure 3.** (a-d) Horizontal slices of the 3-D density model obtained from the joint inversion of field muon and gravity data. A structural map is superposed to facilitate interpretation (orange lines: fractures; violet lines: faults; blue lines: collapse scars; orange symbols: past activity; red symbols: present activity; triangles: hydrothermal fluid exurgence; stars: active fumaroles; squares: boiling acid ponds). (e) Histogram of the data residuals (gravity and muons) normalized by their error standard deviation. Residuals add to a RMS error of 1.

The most extensive low-density region is found below the southern part of the summit (Figures 3a–3c), where the most presently active fumaroles are located. This region counts several deep (up to 122 m) pits, craters, and fractures that were active during the last six phreatic eruptions, as well as evidence of rock alteration [Rosas-Carbajal *et al.*, 2016]. The low-density anomaly is well resolved since the three telescopes are oriented to cover this part of the dome where many gravity measurements were done (cf. Figure 1a). The low-density anomaly clearly extends down to  $z = 1324$  m but becomes unrecognizable at  $z = 1292$  m level because of a limited resolution at this altitude level that is only scanned by the telescope located at the Parking site. As a consequence, the clear low-density anomaly observed for  $z > 1324$  m blurs into a faint low-density anomaly spread through the imaged region (Figure 1d).

The upper, well-resolved part of this anomaly coincides with the electrically conductive region found by Rosas-Carbajal *et al.* [2016] with the inversion of electrical resistivity tomography (ERT) data. Slices of the resulting electrical conductivity model at the same depths as in Figures 3b and 3d) are presented in Figure 4. While the high conductivity values in this region may have been explained by the presence of highly conductive fluids (high temperature and salinity), the low density confirms that this is also a highly altered, porous,



**Figure 4.** Horizontal slices of the 3-D electrical conductivity model by *Rosas-Carbajal et al.* [2016] at the same altitude as slices in Figures 3b and 3c. High conductivity values correspond to hydrothermally altered, hot-fluid-saturated rocks. Low conductivity values correspond to either cold, unaltered rocks or regions where large cavities are present. See Figure 3 for map references.

and thus mechanically weak zone, reinforcing the likelihood of a partial edifice collapse hazard [*Rosas-Carbajal et al.*, 2016]. Interestingly, the Tarissan boiling acid pond (red square in figures) where the largest amount of magmatic fluids arrive to the surface is located in the northern limit of this low-density, high-conductivity region.

More limited low-density regions are found in the west southwest close to the summit (Figures 3a and 3b) and in the extreme south and north center region at  $\sim 1300$  m (Figure 3d). The southernmost region is constrained to the north by gravity data (cf. Figure 1a), and it probably corresponds to open fractures in the prolongation of La Ty fault. This is supported by low conductivity values in the ERT model and historical evidence for exurgence of hot acid hydrothermal fluid from this location (see orange triangle in Figure 3d) during the 1976 phreatic eruption. The low-density anomaly subparallel to the large north-south fracture is located below a smaller, north-west fracture. Even if located relatively far for the muon telescopes, the structure is well resolved thanks to the gravity points acquired in its vicinity. *Kuster and Silve* [1997] report on a large cavity located exactly in this region.

High-density anomalies are found at the northeastern base of the dome (Figures 3c and 3d), toward the southeast flank (Figures 3b–3d), and in the south flank, west of the low-density region (Figures 3a–3d). In the north, the anomaly correlates to the position of the gravity measurements and follow the footpath along the base of the lava dome. This anomaly may be structurally controlled as it follows the detachment plane of an ancient collapse [Rosas-Carbajal *et al.*, 2016]. Given the lack of data outside this region we do not make further interpretations. In the south southwest (Figures 3c and 3d) and in the northern extreme at  $z = 1372$  m (Figure 3b) the high density coincides with low electrical conductivity anomalies (Figures 4a and 4b), revealing an extensive volume of unaltered dome rock. The southern part of this anomaly belongs to a “bulge,” which rests on top of a highly altered zone that reaches the surface in the south and thus constitutes a potential detachment plane [cf. Rosas-Carbajal *et al.*, 2016]. Likewise, the high-density region in the southeast flank corresponds to nonaltered lava dome spines.

## 5. Conclusions

We obtained the first 3-D density model of a volcano lava dome from the joint inversion of simultaneous muon radiographies and gravity data. Several density anomalies are detected in La Soufrière, especially an extensive low-density zone located below the southern part of the summit where most of the increasingly active fumaroles are observed. These results reinforce the inference from previous studies that the southern flank of the La Soufrière lava dome is highly porous and mechanically weak, as well as saturated with hot acid fluid. The associated rock alteration and dissolution promotes instability and raises the likelihood of a partial collapse. Our density model constitutes the basis for undertaking monitoring of density changes using simultaneous muon measurements from different points around the lava dome. Future work will focus on improving the spatial coverage to refine the model resolution. New recently installed muon telescopes on the west and northern parts of the volcano are fitted with a smaller pixel support ( $63 \times 63$  lines of sight). Theoretical work on the behavior of the diffusive muon flux will be crucial to estimate absolute density and its uncertainty. Joint inversion with other types of geophysical data such as ERT data will be also explored. This work demonstrates the potential of muon data to move from 2-D “radiographies” to dynamic 3-D imaging of geological structures. This is of paramount importance given the limitation of conventional volcano monitoring in tracking the internal state of the system with full spatial coverage.

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